

# Source constraints from $v$ & $p$ with GCOS

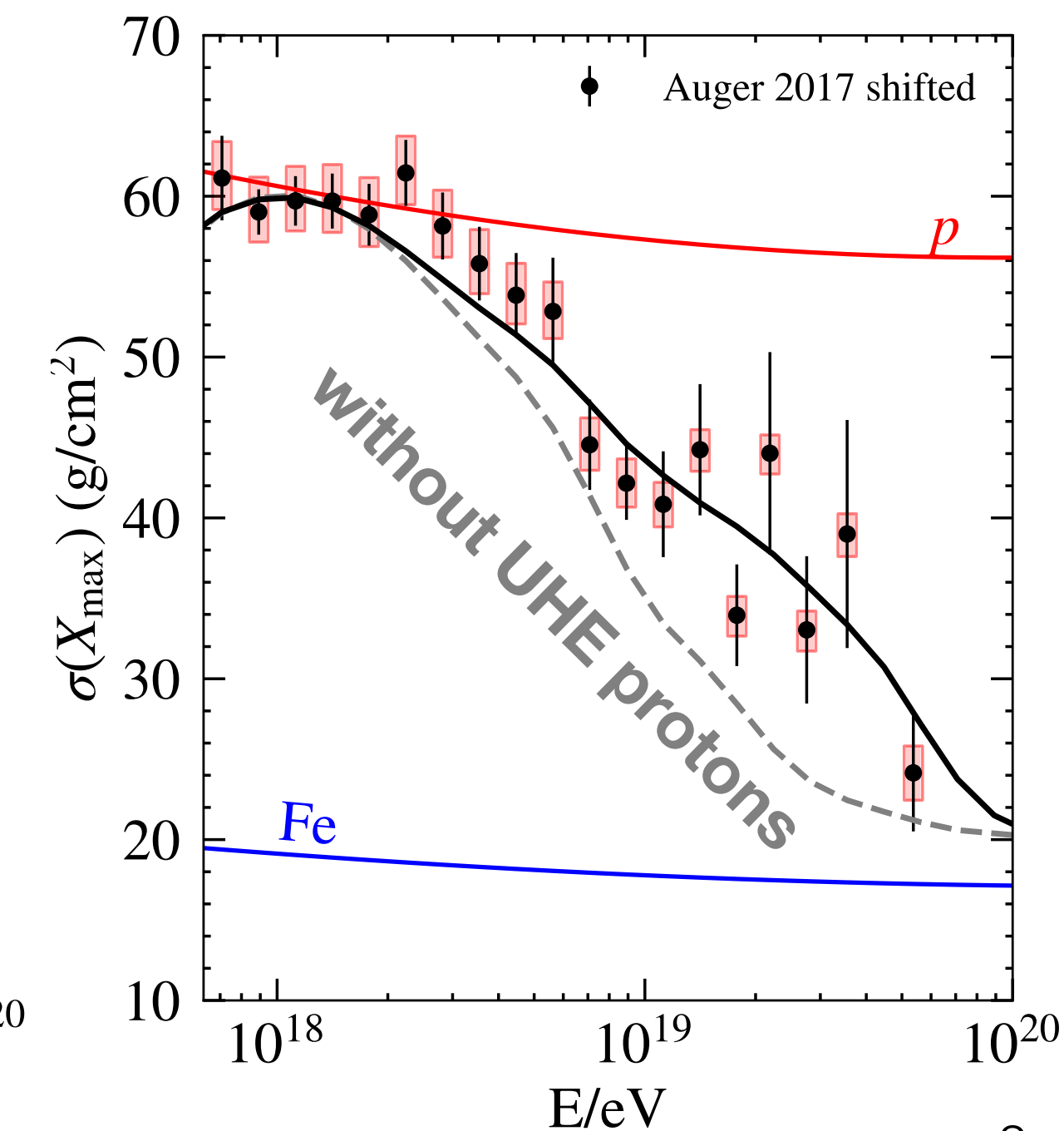
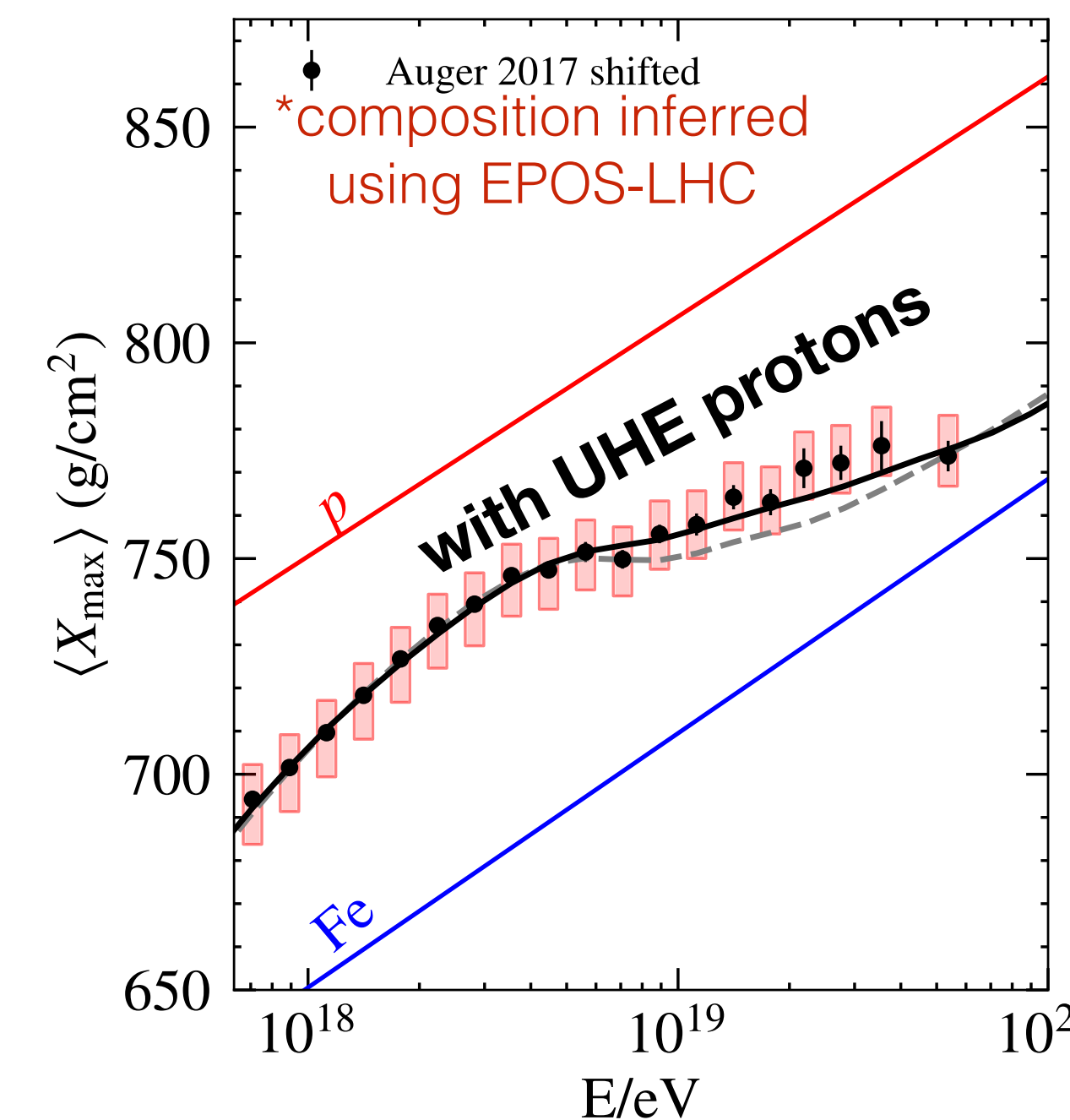
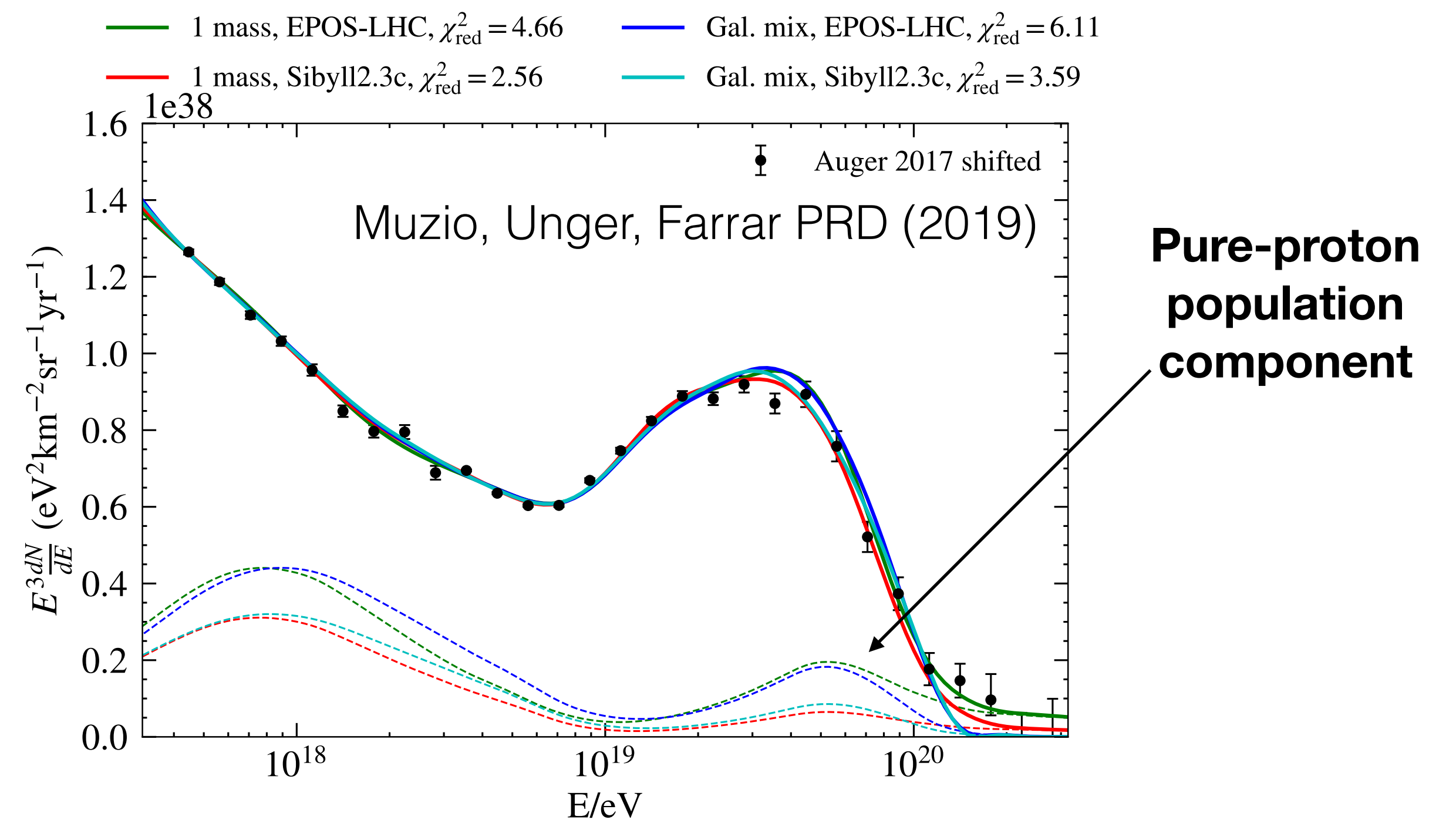
Marco Muzio



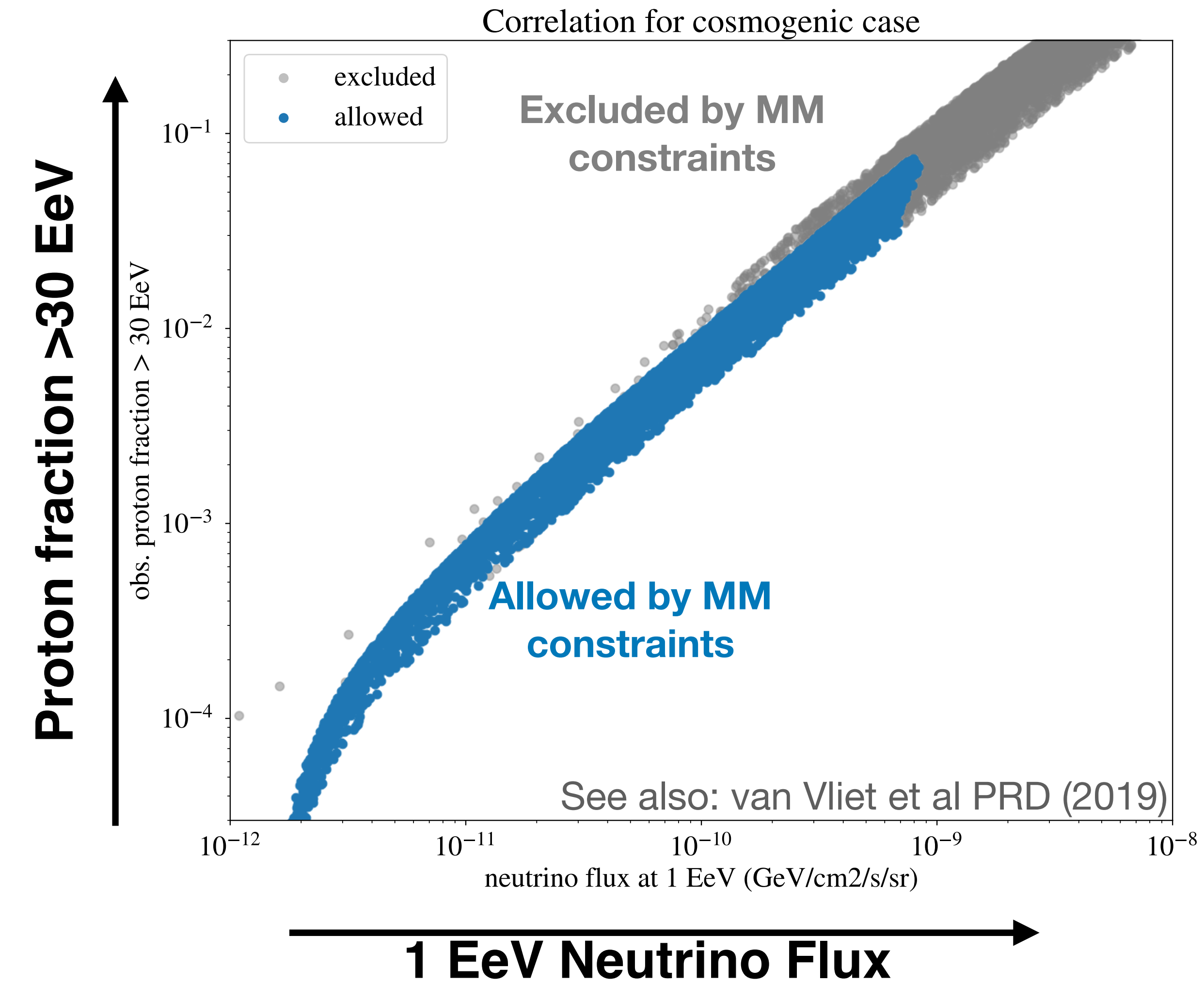
**PennState**

# This study

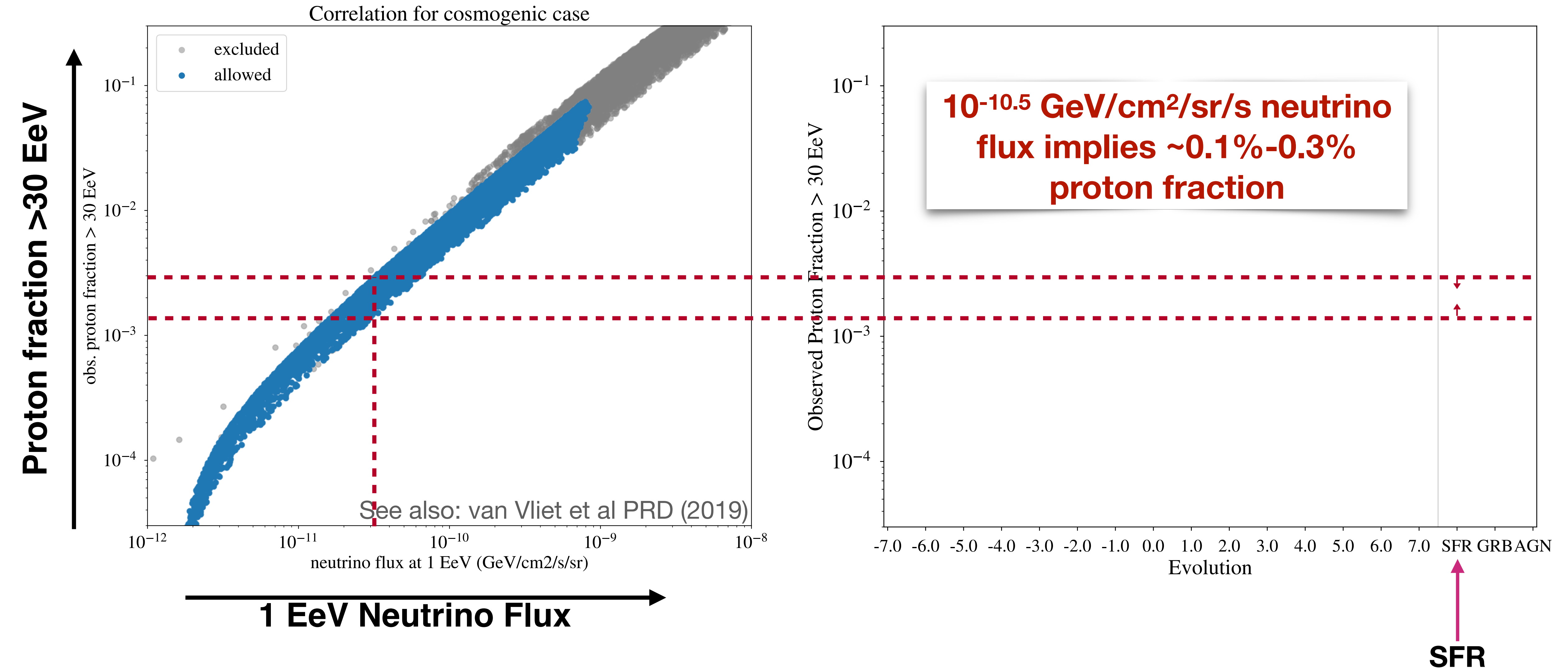
- CR source model: Unger-Farrar-Anchordoqui (UFA)
- Two source populations:
  - Baseline population, fitting most of UHECR data
    - Fixed to best-fit parameters for given evolution
  - Pure-proton population, producing protons  $>10$  EeV
    - Model parameters uniformly sampled, explore full range of model predictions
    - Exclude models that violate UHECR, neutrino, & gamma-ray data
- **To what extent can UHE protons & secondary neutrinos constrain source evolution if source interactions are significant?**



# Example: SFR, cosmogenic neutrinos only

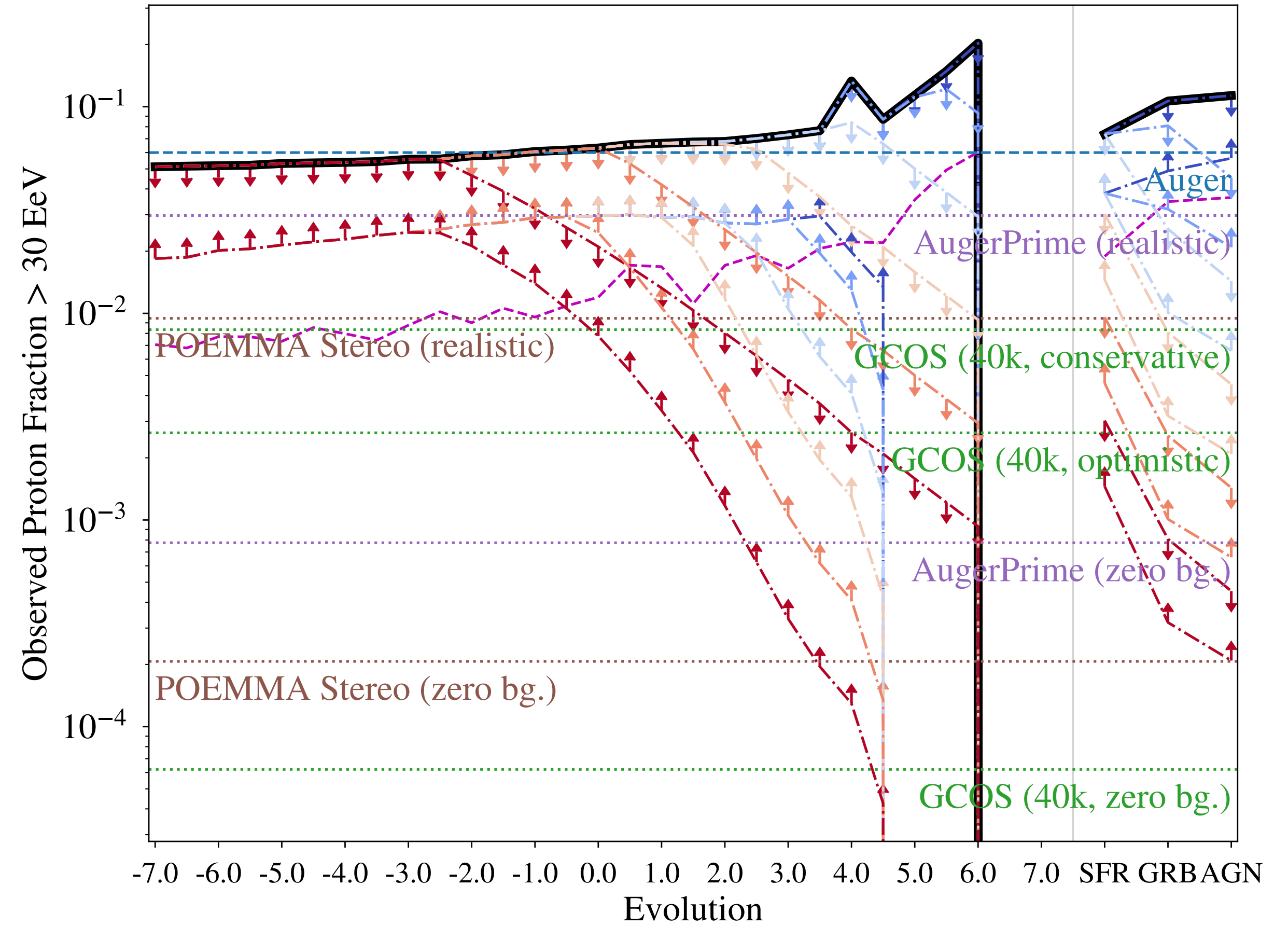


# Example: SFR, cosmogenic neutrinos only



# Cosmogenic neutrinos only

$$\xi(z) = \begin{cases} (1+z)^m & z \leq z_0 \\ (1+z_0)^m e^{-(z-z_0)} & z > z_0 \end{cases}$$



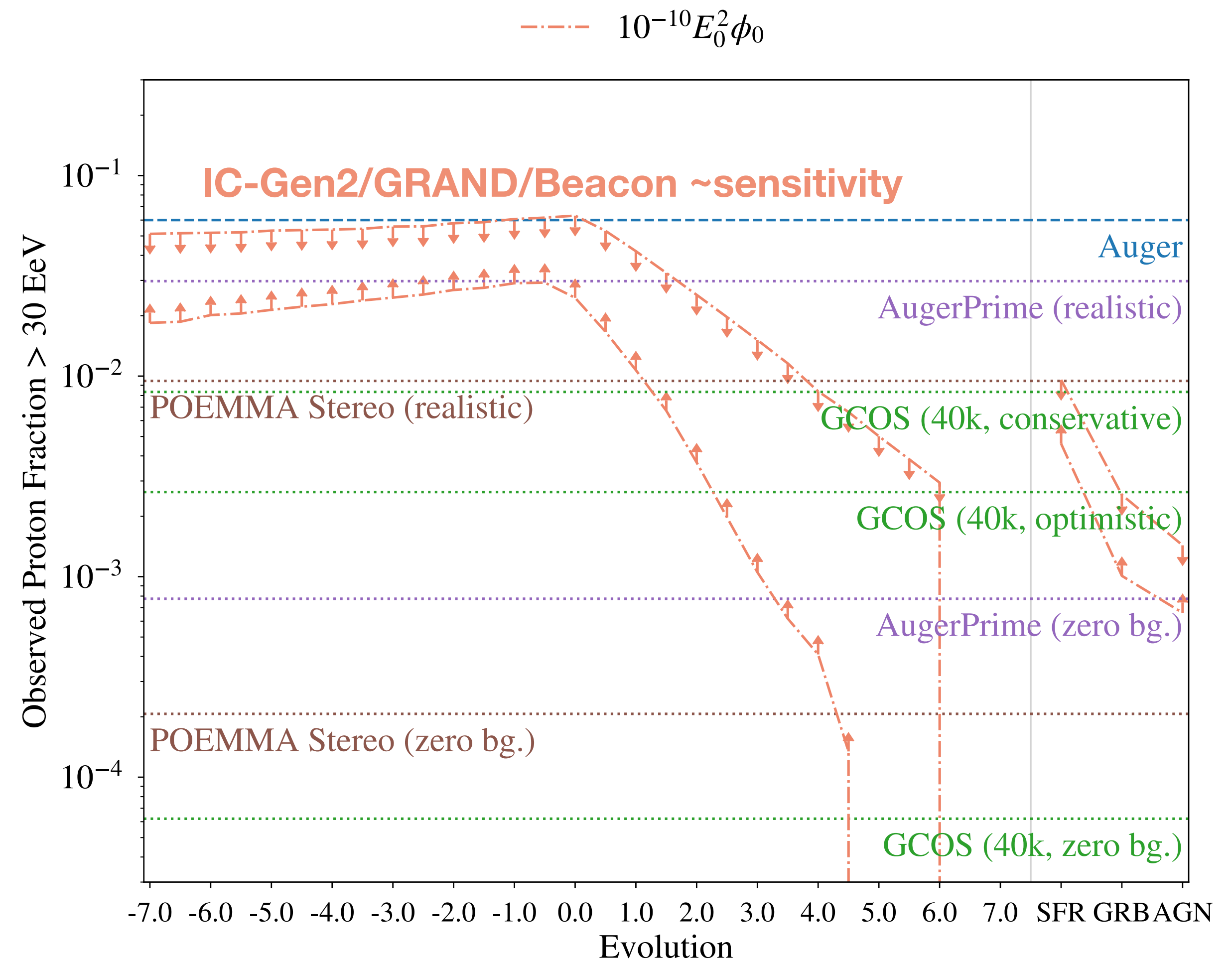
**Source Evolution**  
 $E_0^2 \phi_0 = \text{GeV}/\text{cm}^2/\text{s}/\text{sr}$



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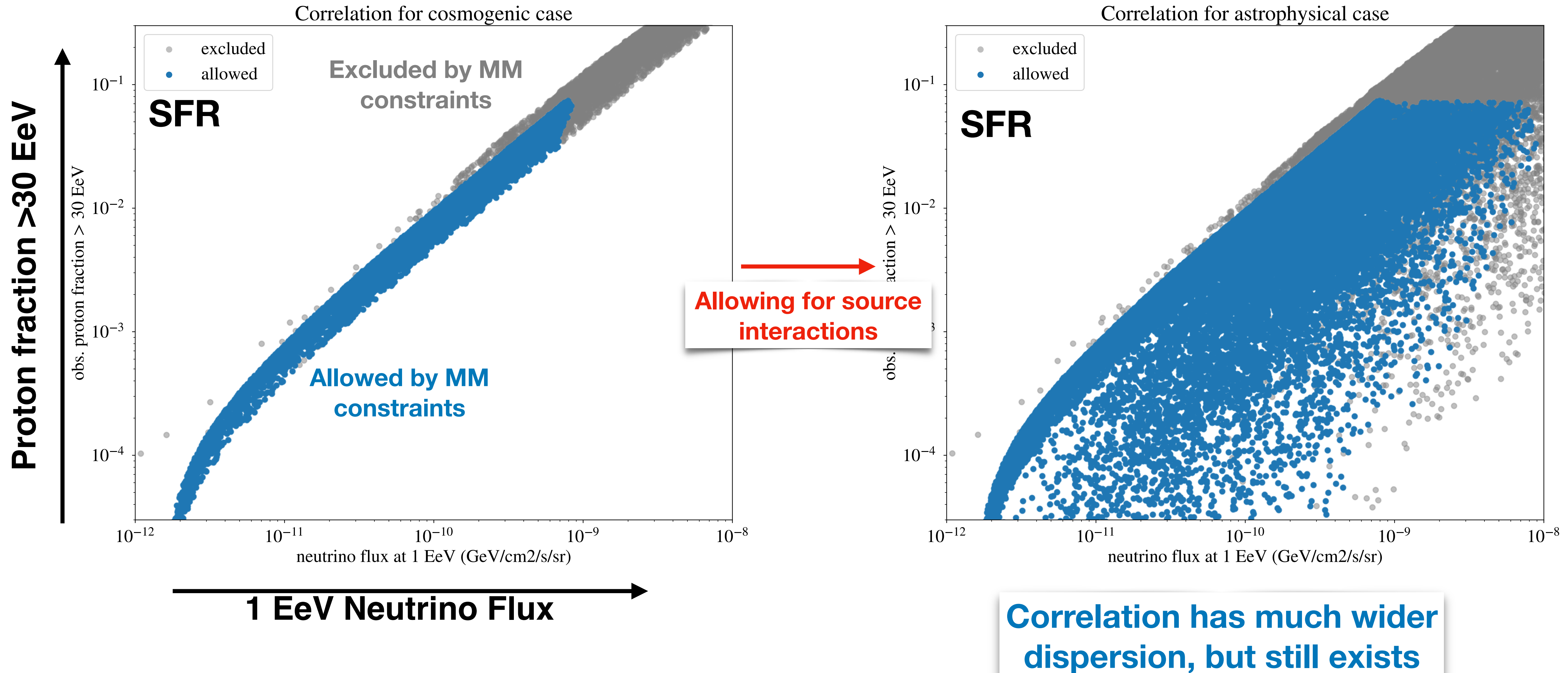
- Combined UHE neutrino & proton fraction measurement can constrain source evolution
- **IC-Gen2 Radio/GRAND200k/Beacon** + **GCOS** measurement will give source evolution for:
  - $m < 4$ , conservative case (p-He MF~0.3)
  - $m < 6/\text{SFR}$ , optimistic case (p-He MF~0.7)
  - GRB/AGN, background-free case



**Source Evolution**

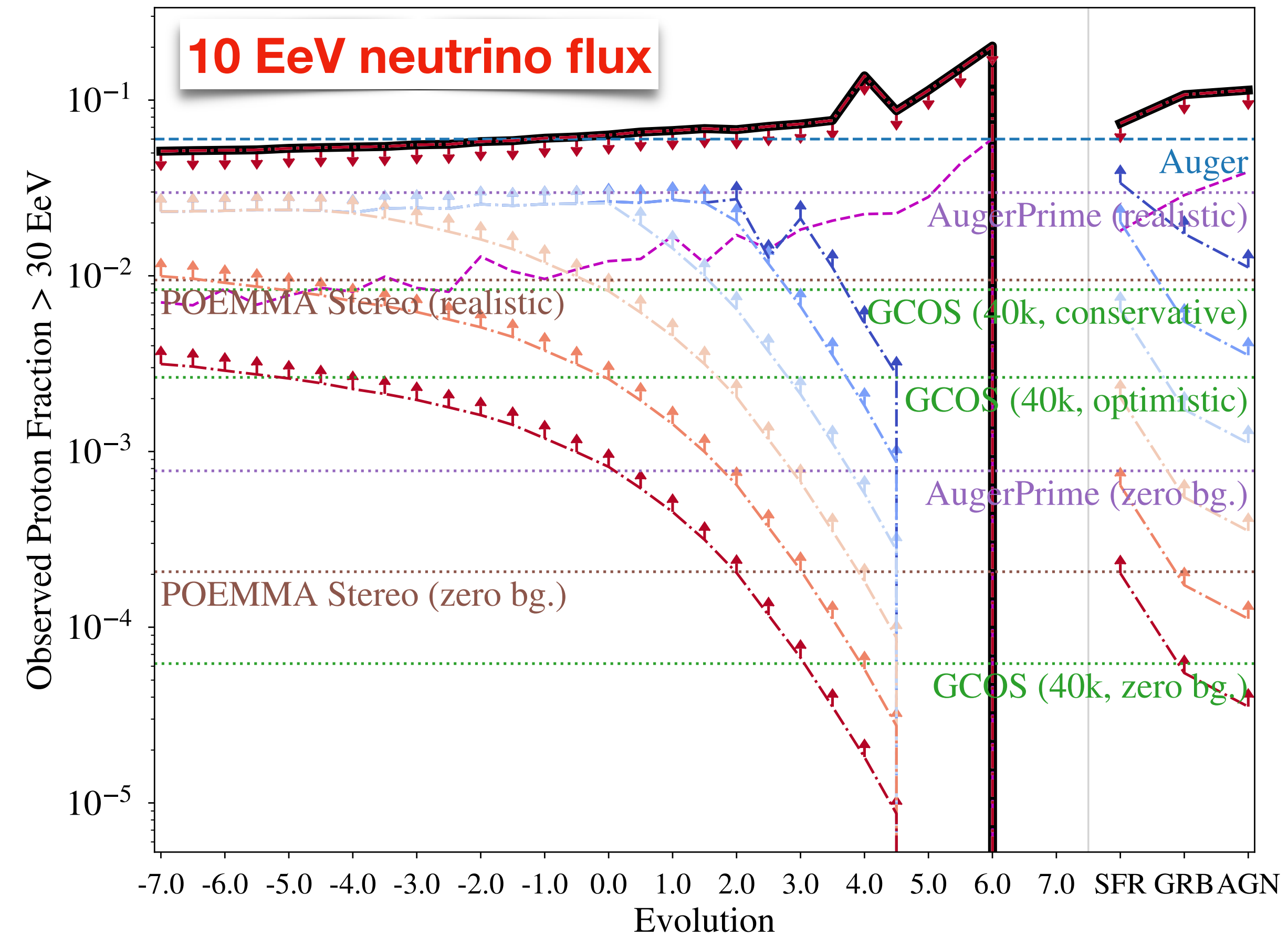
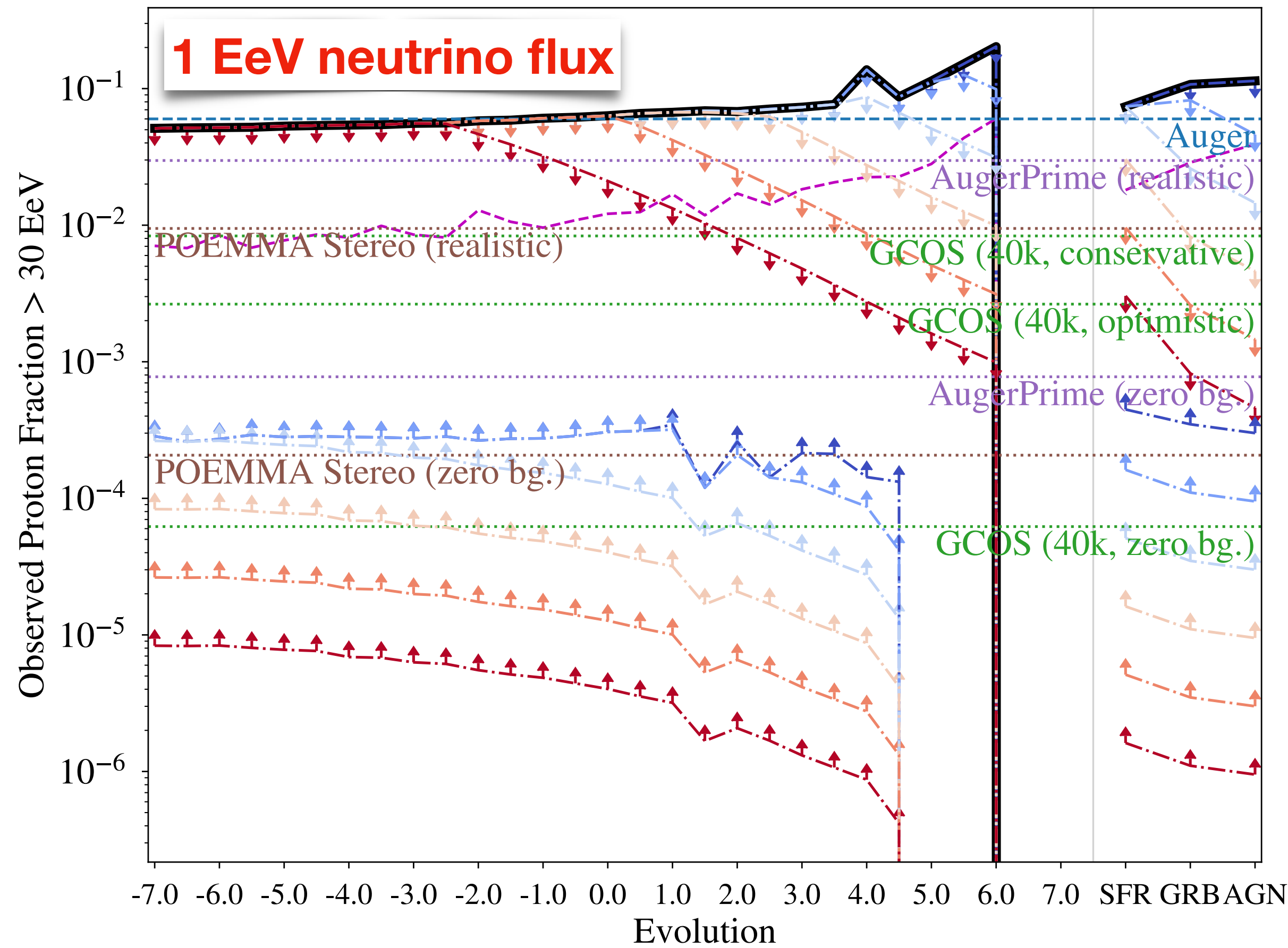
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# Effect of Source Interactions



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**1 EeV neutrino flux → upper-bound on  $m$**   
**10 EeV neutrino flux → lower-bound on  $m$**



# Summary

- GCOS will constrain observed proton fraction:
  - He contamination case:  $O(0.1\%)$ 
    - Conservative merit factor  $\sim 0.3$  gives constraints comparable to POEMMA
  - Background free case:  $O(0.001\%)$
- **GCOS & planned neutrino experiments will strongly constrain UHECR source evolution**
  - Cosmogenic only case:  $m$  is measured
  - General case: constrain  $m$  with 1 EeV & 10 EeV neutrino measurements

