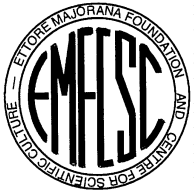


Historical introduction

Basic properties of Cosmic Rays



«ETTORE MAJORANA» FOUNDATION AND CENTRE FOR SCIENTIFIC CULTURE

INTERNATIONAL SCHOOL OF COSMIC-RAY ASTROPHYSICS
«MAURICE M. SHAPIRO»

23rd Course: “Multi-Messenger Astroparticle Physics” 20 – 28 July 2024

PRESIDENT AND DIRECTOR OF THE CENTRE: PROFESSOR A. ZICHICHI

DIRECTORS OF THE COURSE: PROFESSORS J.R. HÖRANDEL, T. STANEV, R. SPARVOLI - J.P. WEFEL (director emeritus)



Nikhef



Jörg R. Hörandel

RU Nijmegen, Nikhef, VU Brussel

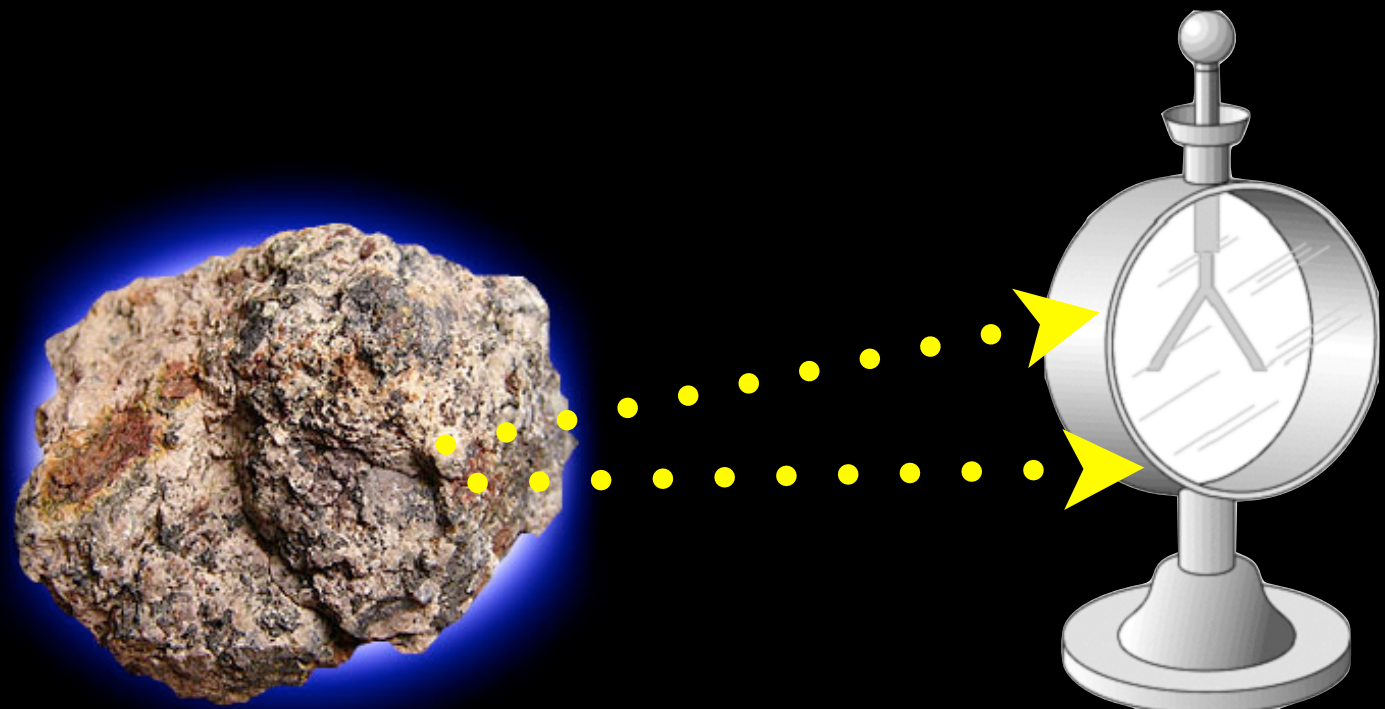
<http://particle.astro.ru.nl>

Discovery of Radioactivity

Discovery of Radioactivity



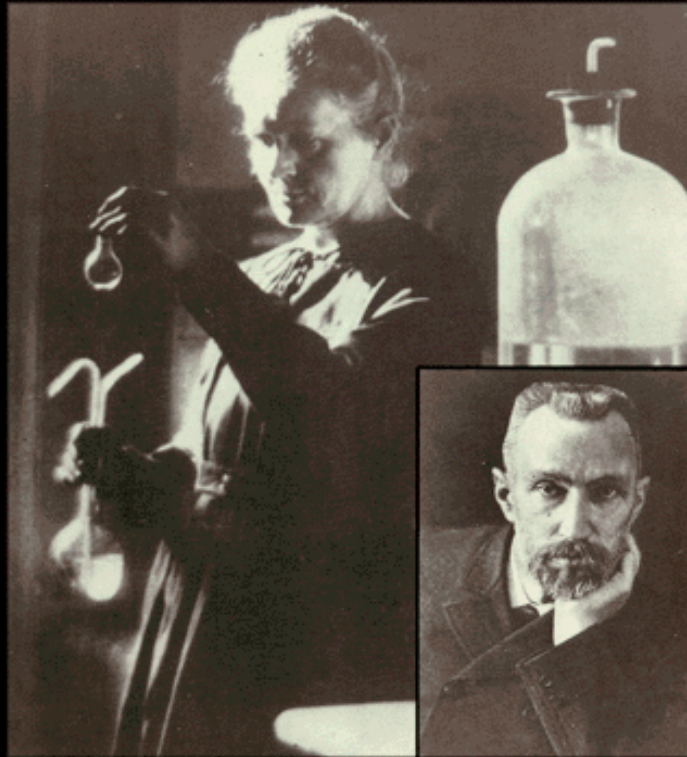
Discovery of Radioactivity



Discovery of Radioactivity

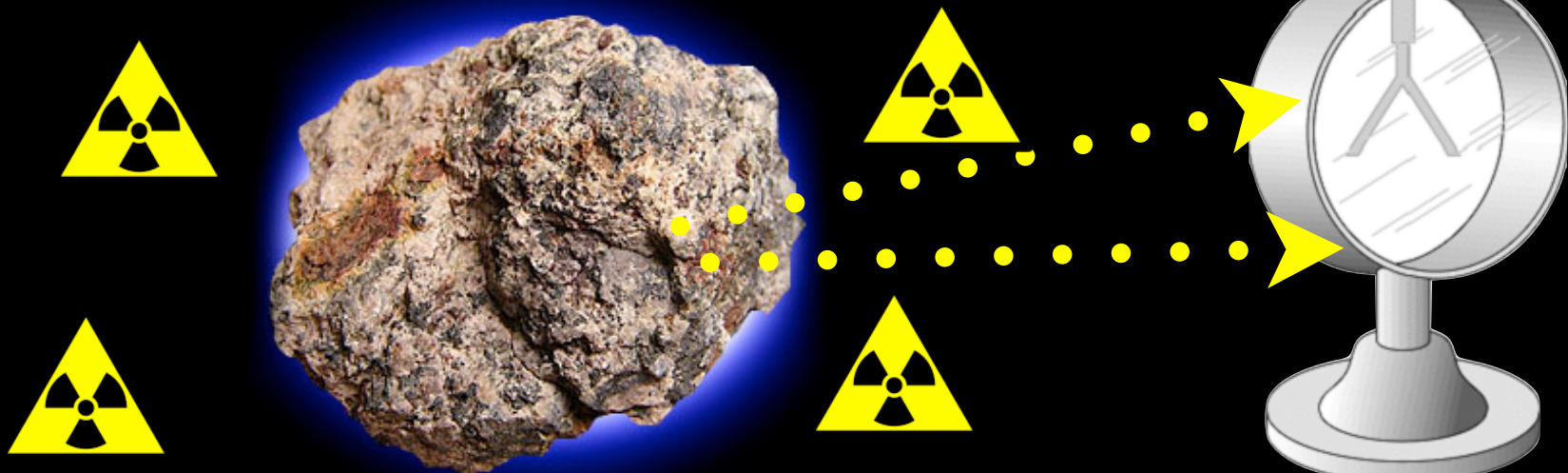


Henri Becquerel



Marie & Pierre Curie

Nobel Prize
1903



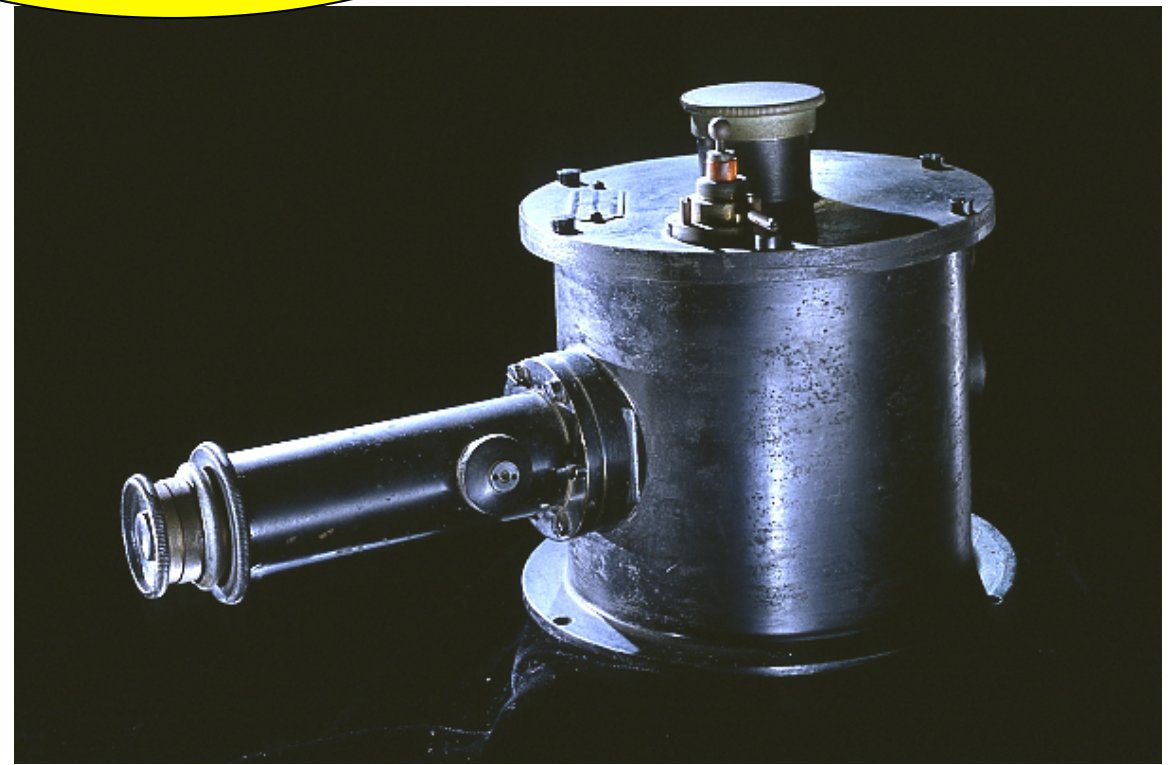
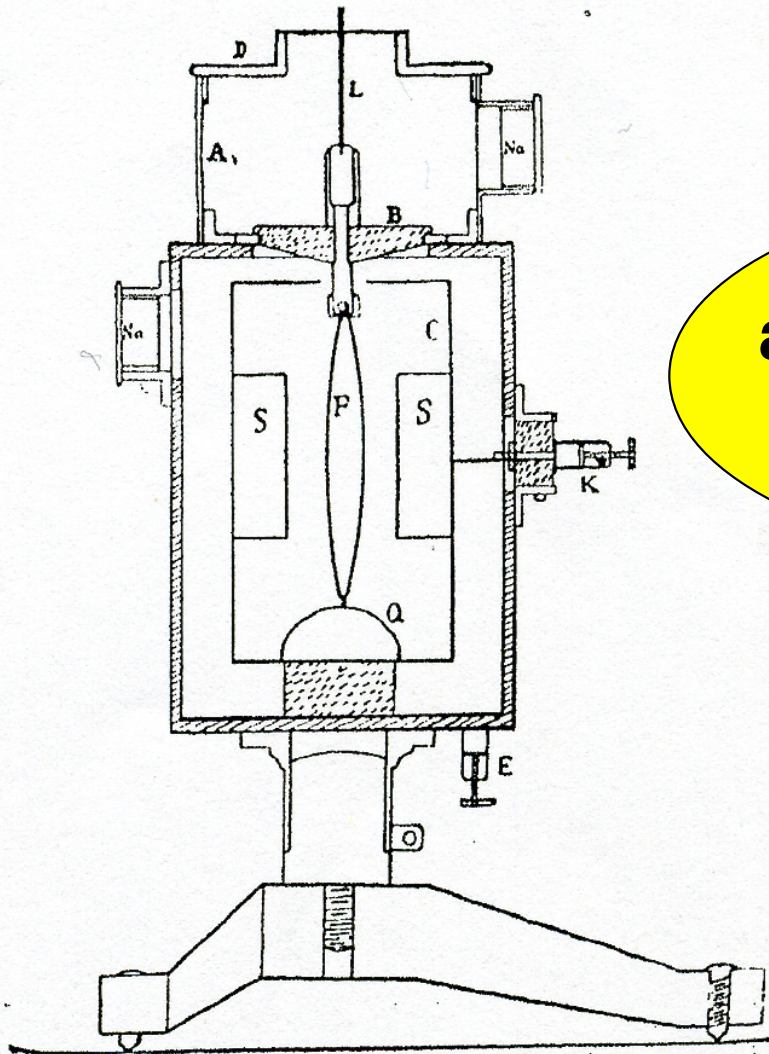
Ein neues Elektrometer für statische Ladungen.

Dritte Mitteilung¹⁾.

Von Th. Wulf.

Mitteilung enthält einige
weiter beschriebenen Appa-
raturerhöhung seiner Transport-

**a new electrometer
for static charges**





**Sir J.J.Thomson
Nobel Prize 1906**



Conduction of electricity through gases (1928):

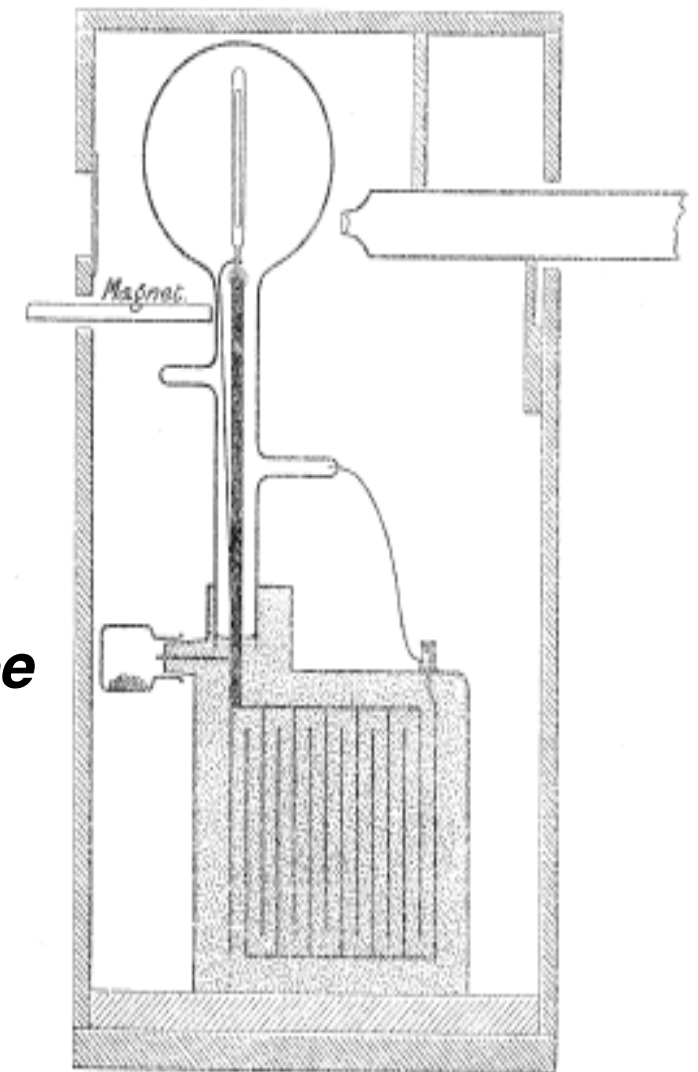
It would be one of the romances of science if these obscure and prosaic minute leakages of electricity from well-insulated bodies should be the means by which the most fundamental problems in the evolution of the cosmos came to be investigated.



Detector used by Wilson to investigate ionization of air

“the continuous production of ions in dust-free air could be explained as being due to radiation from sources outside our atmosphere, possibly radiation like Röntgen rays or cathode rays, but of enormously greater penetrating power”

C T R Wilson, Proc Roy Soc A 68 (1901) 151



**on the origin of
gamma radiation in the
atmosphere**

Über den Ursprung der in der Atmosphäre
vorhandenen γ -Strahlung.

Von Th. Wulf.

Tabelle I.

Strahlung der Wände von Gebäuden.

Ort	Material	Alter	Strahlung Ionen pro ccm u. Sekunde
Abtei Maria Laach bei Andernach a. Rh. . .	} Vulkanisch Tuff	} 50 Jahre	13,7
Valkenburg, Colleg, Holland-L.			
Löwen, Colleg, Belgien	Ziegelsteine	15 "	5,7
Namur, Colleg N.D. de la paix, Belgien . . .	Ziegelsteine	—	8,0
Wynandsrade Kasteel, Holland	Ziegelsteine	ca. 100	3,7
	Ziegelsteine	200 Jahre	0,0

Nur in dem alten holländischen Kasteel Wynandsrade, vor fast 200 Jahren aus Ziegelsteinen erbaut, zeigte sich kein Unterschied in der Strahlung im Zimmer und im Freien. — Am stärksten war die Strahlung in Maria Laach in einem

Man kann den Inhalt dieser Arbeit kurz so zusammenfassen. Es wird über Versuche berichtet, welche beweisen, daß an dem Beobachtungsort die durchdringende Strahlung von primär radioaktiven Substanzen verursacht wird, welche in den obersten Erdschichten liegen, bis etwa 1 m unter der Oberfläche.

Wenn ein Teil der Strahlung aus der Atmosphäre stammt, so ist er doch so klein, daß er sich mit den gebrauchten Mitteln nicht nachweisen ließ.

Die zeitlichen Schwankungen in der γ -Strahlung liegen

**the radiation originates
from the soil
maybe a small contribution
from the atmosphere**



~1910



Theodor Wulf

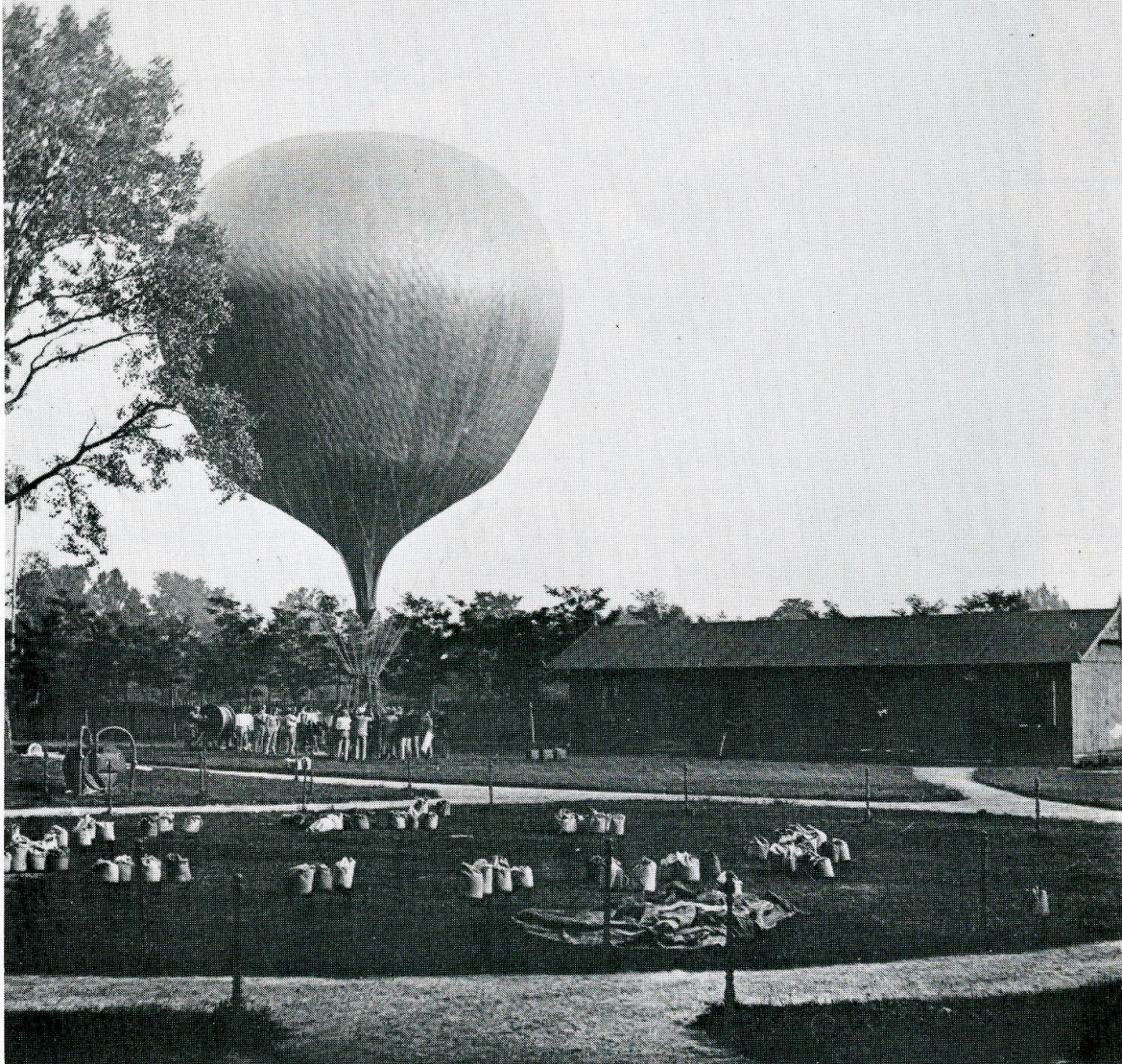
**1909: Soddy & Russel:
attenuation of gamma rays
follows an exponential law**

$$I = I_0 e^{-\mu L}$$

Discovery of Cosmic Rays

Viktor Franz Hess

7. August 1912



Aeronautisches Gelände im Wiener Prater, von dem aus V. F. Hess in den Jahren 1911/12 seine ersten Freiballon-Forschungsfahrten unternommen hatte. (Courtesy of Heeresgeschichtliche Museum, Vienna)

Discovery of Cosmic Rays

Viktor Franz Hess 7. August 1912



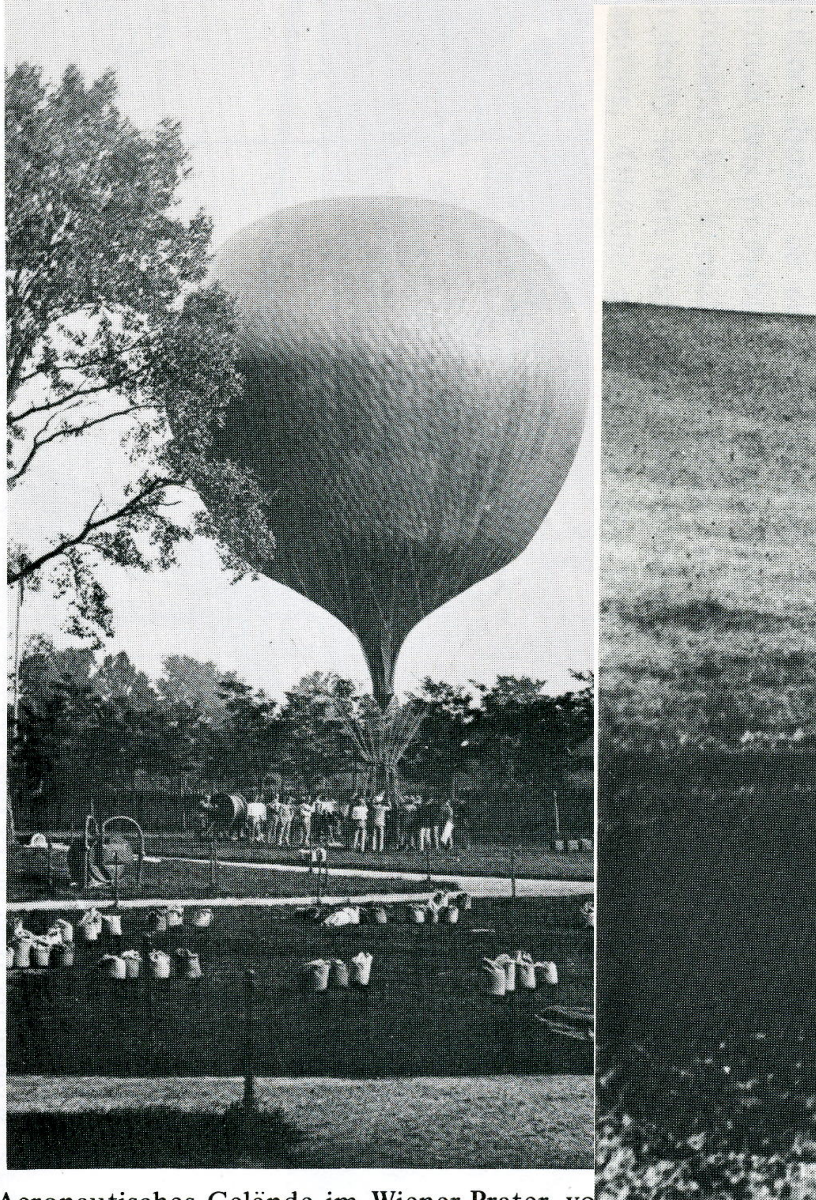
Aeronautisches Gelände im Wiener Prater, vor seinen ersten Freiballon-Forschungsfahrten und schichtliche Mu

Hess on gondola in 1912 probably in test flight. The date and place is not clear at present.

Discovery of Cosmic Rays

Viktor Franz Hess

7. August 1912



Aeronautisches Gelände im Wiener Prater, vor seine ersten Freiballon-Forschungsfahrten u schichtliche Mu

Hess on



Route des Entdeckungsfluges der kosmischen Strahlung.

Discovery of Cosmic Rays

Viktor Franz Hess

7. August 1912

Early cosmic-ray work published in German

Jörg R. Hörandel

Citation: [AIP Conf. Proc. 1516](#), 52 (2013); doi: 10.1063/1.4792540

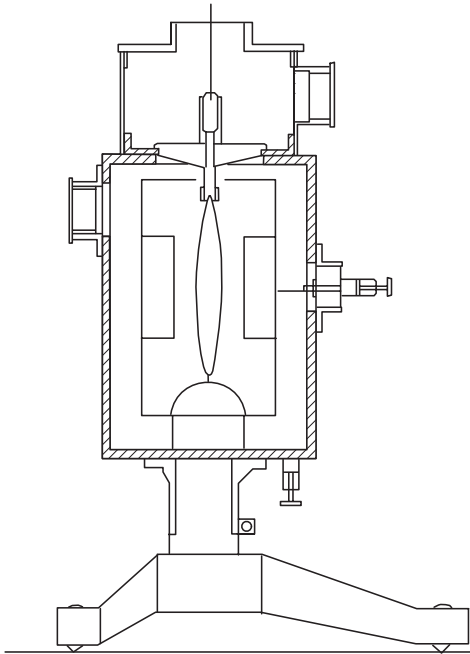


FIGURE 1. *Left:* Electrometer after Th. Wulf [5]. *Right:* Two grandsons of V.F. Hess revealing a plaque to commemorate the discovery of cosmic rays on August 7th, 2012, close to the presumed landing site of V.F. Hess in Pieskow close to Berlin. It reads: "To commemorate the discovery of cosmic rays. On 7 August 1912 landed the Austrian physicist Victor F. Hess with a hydrogen balloon close to Pieskow. On the journey from Lower-Bohemia he reached an altitude of 5300 m and he proved the existence of a penetrating, ionizing radiation from outer space. For the discovery of cosmic rays V.F. Hess has been awarded the Nobel Prize in Physics in 1936. The participants of the symposium '100 years cosmic rays', Bad Saarow-Pieskow, 7 August 2012".



Aeronautisches Gelände im Wiener Prater, von
seine ersten Freiballon-Forschungsfahrten u
schichtliche Mu

Hess on

Route des Entdeckungsfluges der kosmischen Strahlung.

Aus der Abteilung für Geophysik, Meteorologie und Erdmagnetismus:

Viktor F. Hess (Wien), Über Beobachtungen der durchdringenden Strahlung bei sieben Freiballonfahrten.

Im Vorjahre habe ich bereits Gelegenheit gehabt, zwei Ballonfahrten zur Erforschung der durchdringenden Strahlung zu unternehmen: über die erste Fahrt wurde...

on the observation of the penetrating radiation during 7 balloon campaigns

7. Fahrt (7. August 1912).

Ballon: „Böhmen“ (1680 cbm Wasserstoff).
 Meteorolog. Beobachter: E. W. Pf.

Führer: Hauptmann W. Hoffory.
 Lufterlekt. Beobachter: V. F. Hess.

	Mittlere Höhe	Beobachtete Strahlung						Temp.	Relat. Feucht. Proz.
		absolut	relativ m	Apparat 1	Apparat 2	Apparat 3			
				q_1	q_2	q_3	reduz. q_3		
1	15h 15—16h 15	156	0	17,3	12,9	—	—	} 1 1/2 Tag vor dem Aufstiege (in Wien)	
2	16h 15—17h 15	156	0	15,9	11,0	18,4	18,4		
3	17h 15—18h 15	156	0	15,8	11,2	17,5	17,5		
4		1700	1400	15,8	14,4	—	—		
5		2750	2500	17,3	17,3	—	—	+6,4 ⁰	
6		3850	3600	19,8	19,8	—	—	+1,4 ⁰	
7		4800	4700	40,7	31,0	—	—	-6,8 ⁰	
		(4400—5350)	—	—	—	—	—	-9,8 ⁰	
8	10h 45—11h 15	4400	4200	28,1	22,7	—	—	—	
9	11h 15—11h 45	1300	1200	(9,7)	11,5	—	—	—	
10	11h 45—12h 10	250	150	11,9	10,7	—	—	+16,0 ⁰	
11	12h 25—13h 12	140	0	15,0	11,6	—	—	(nach der Landung in Pieskow, Brandenburg)	

hydrogen!

altitude

intensity

Aus der Abteilung für Geophysik, Meteorologie
und Erdmagnetismus:

Viktor F. Hess (Wien), Über Beobachtungen
der durchdringenden Strahlung bei sieben

radiation with high
penetration power
impinges onto the
atmosphere from
above

no change during
solar eclipse
hence, Sun can be
excluded as origin

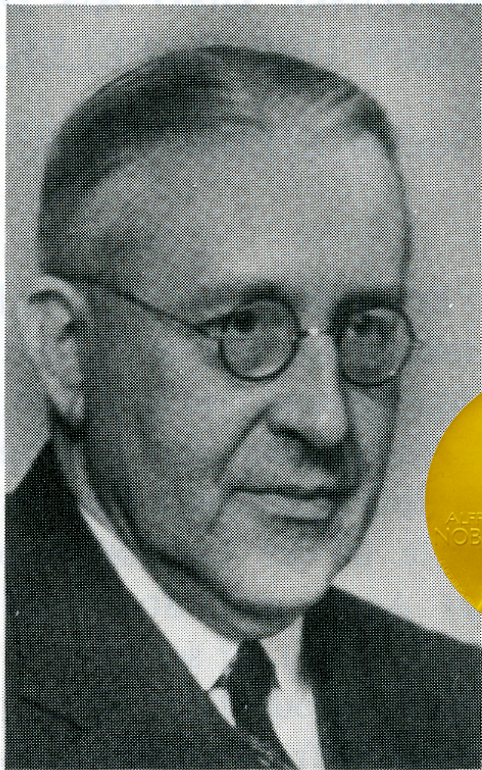
der Verringerung der radioaktiven Substanzen
der Atmosphäre zurückzuführen.

Die Ergebnisse der vorliegenden Beobachtungen scheinen am ehesten durch die Annahme erklärt werden zu können, daß eine Strahlung von sehr hoher Durchdringungskraft von oben her in unsere Atmosphäre eindringt, und auch noch in deren untersten Schichten einen Teil der in geschlossenen Gefäßen beobachteten Ionisation hervorruft. Die Intensität dieser Strahlung scheint zeitlichen Schwankungen unterworfen zu sein, welche bei einstündigen Ablesungsintervallen noch erkennbar sind. Da ich im Ballon weder bei Nacht noch bei einer Sonnenfinsternis eine Verringerung der Strahlung fand, so kann man wohl kaum die Sonne als Ursache dieser hypothetischen Strahlung ansehen, wenigstens solange man nur an eine direkte γ -Strahlung mit geradliniger Fortpflanzung denkt.

Daß die Zunahme der Strahlung erst jenseits 2000 m so stark merklich wird ist nicht

Aus der Abteilung für Geophysik, Meteorologie
und Erdmagnetismus:

Viktor F. Hess (Wien), Über Beobachtungen
der durchdringenden Strahlung bei sieben



V.F. Hess in 1936–37, on the occasion of Nobel prize.

Nobel Prize 1936

der Verringerung der radioaktiven Substanzen
der Atmosphäre zurückzuführen.

Die Ergebnisse der vorliegenden Beobachtungen scheinen am ehesten durch die Annahme erklärt werden zu können, daß eine Strahlung von sehr hoher Durchdringungskraft von oben her in unsere Atmosphäre eindringt, und auch noch in deren untersten Schichten einen Teil der in geschlossenen Gefäßen beobachteten Ionisation hervorruft. Die Intensität dieser Strahlung scheint zeitlichen Schwankungen unterworfen zu sein, welche bei einstündigen Ablesungsintervallen noch erkennbar sind. Da ich im Ballon weder bei Nacht noch bei einer Sonnenfinsternis eine Verringerung der Strahlung fand, so kann man wohl kaum die Sonne als Ursache dieser hypothetischen Strahlung ansehen, wenigstens solange man nur an eine direkte γ -Strahlung mit geradliniger Fortpflanzung denkt.

Daß die Zunahme der Strahlung erst jenseits 2000 m so stark merklich wird ist nicht

erweitertes Beobachtungsmaterial wurde.

**“Today in the Erice maze
Cosmic Rays are the craze
and this because a guy named Hess
ballooning up found more not less.”**

Poem from the 1st course of the ISCRA in 1978

Neue Untersuchungen über die durchdringende Hesssche Strahlung.

Von E. Steinke in Königsberg i. Pr.

Mit 25 Abbildungen. (Eingegangen am 23. März 1928.)

absorption measurements compatible with assumption of isotropic flux of cosmic rays

Es werde
Eigengang
konnte, I
ausgeführ
starken E
kleinsten
entspricht
Radiumst
was einer
kommt da

in Energie theoretisch freiwerdenden Mengen. In größerer Höhe die Strahlung wesentlich weiter (mittleres $\mu_{pb} = 0,018$ bis $0,025 \text{ cm}^{-1}$). Ein großer Teil der Ionisation wird durch noch erheblich weichere Streustrahlung hervorgerufen, die bei verschiedenen Absorptionsmedien variiert und dadurch die Messungen der Absorptionskoeffizienten leicht fälschen kann. Es wird der Nachweis erbracht, daß die Strahlung auch noch in Eisen merkliche Streustrahlung auslöst. — Die Richtungs- und Absorptionsmessungen gestatten eine einwandfreie Trennung der Hessschen von der Umgebungsstrahlung. Sie geben für die Richtungsverteilung der ersteren etwa solche Werte, wie sie sich aus der Annahme einer von allen Seiten aus dem Kosmos kommenden Strahlung unter Berücksichtigung der Absorption in den verschieden langen Luftwegen mit Hilfe eines mittleren Absorptionskoeffizienten errechnen lassen. Abweichungen hiervon deuten auf die Inhomogenität der Strahlung. — Mit allem Vorbehalt wird der Versuch einer Zerlegung der durchdringenden Strahlung in zwei Komponenten durchgeführt. — Informatorische Schwankungsmessungen in 2500 m Höhe unter Ausschluß der Erdstrahlung, jedoch mit fast der gesamten Intensität der durchdringenden Strahlung, ergeben von Tag zu Tag stark wechselnde Schwankungen (Extremdifferenzen 0,8 bis 4,2 %), welche aber im Mittel um ein Vielfaches kleiner sind als die von anderen Autoren in etwa gleicher Höhe gemessenen. — Auf Schwierigkeiten bei der Deutung des Barometer-effektes wird hingewiesen.

Absorption in the atmosphere

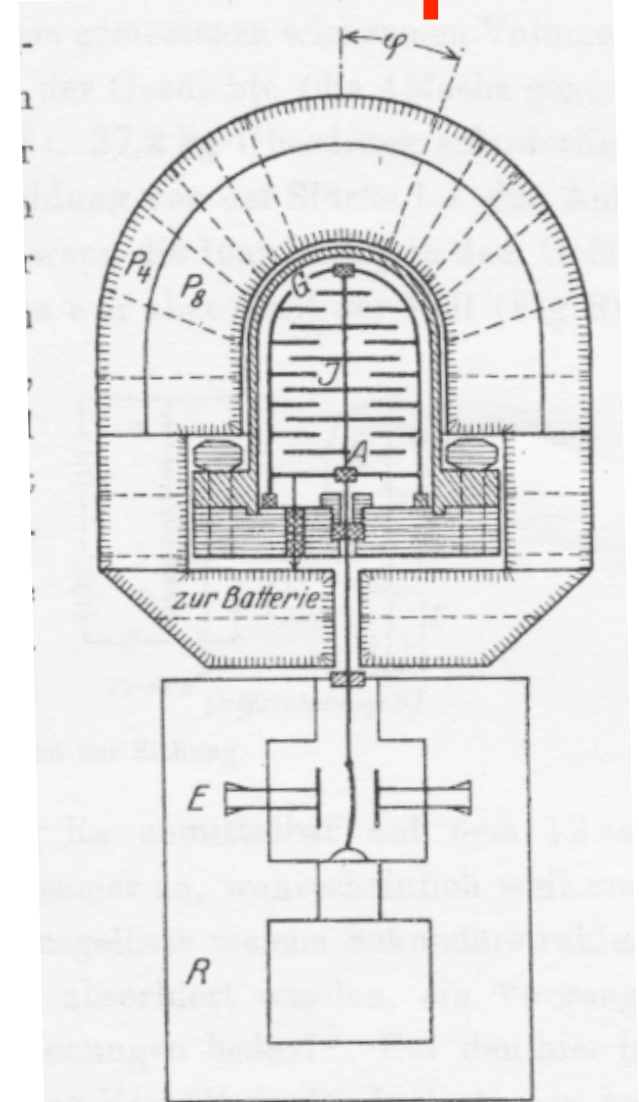


Fig. 1. Hochdruckionisationsgefäß *J* mit Drahtgitter *G* und Aufhänger *A*. Darum 12 cm Fe-Panzer (*P₄*, *P₈*), darunter Wulf'sches Einfeldenelektrometer *E* mit Registriervorrichtung *R*.

Neue Untersuchungen über die durchdringende Hesssche Strahlung.

Von E. Steinke in Königsberg i. Pr.

Absorption in the atmosphere

Intensity as function of zenith angle for different altitudes

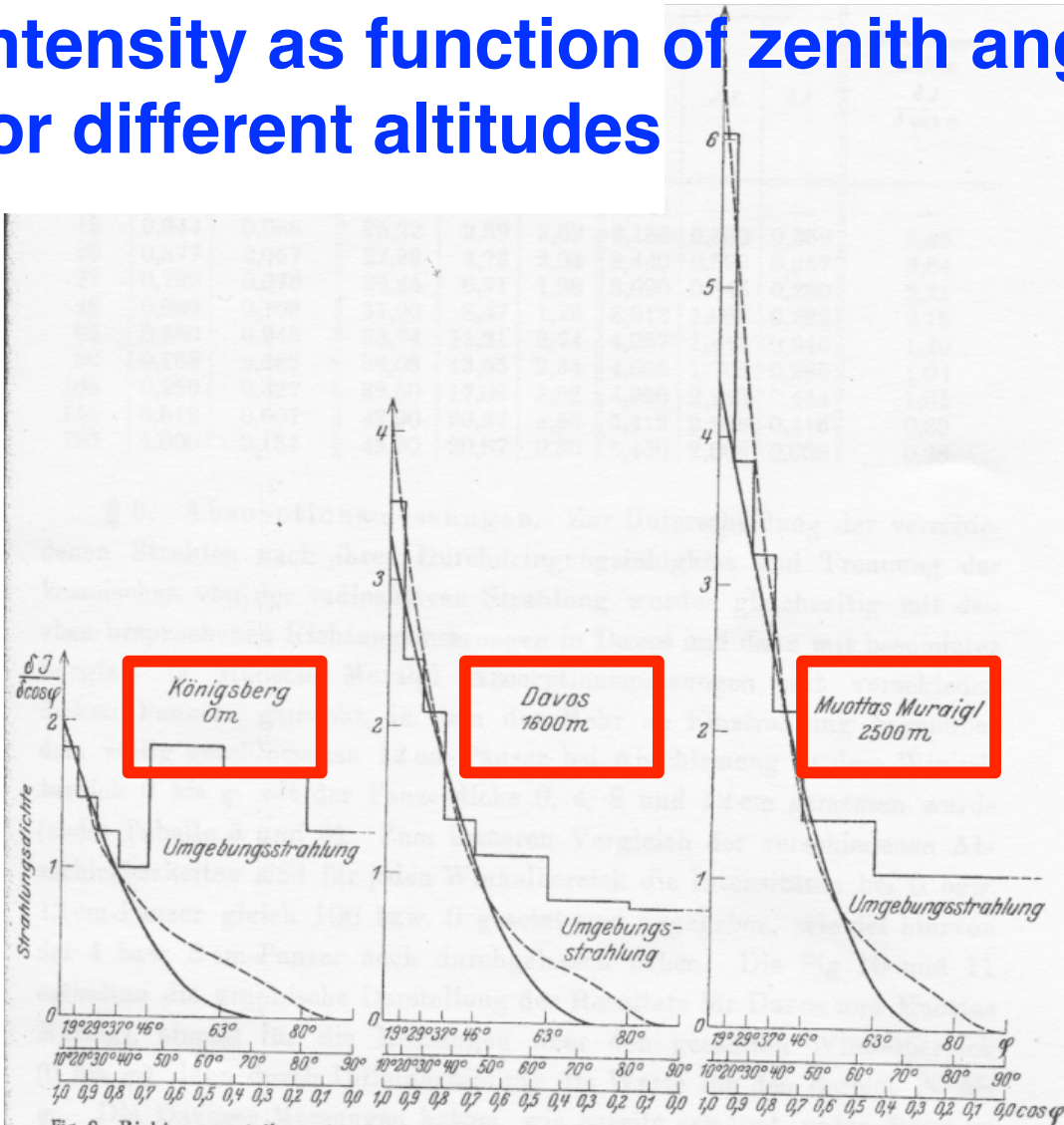


Fig. 9. Richtungsverteilung der durchdringenden Strahlung in verschiedenen Höhen. Die Treppenkurve gibt die beobachteten Werte an. Ferner bedeuten:
 — die berechnete Kurve für $\mu_{Pb} = 0,026 \text{ cm}^{-1}$,
 - - - " " " " $\mu_{Pb} = 0,0042$,
 — — — " " " " $\mu_{Pb} = 0,05 \text{ cm}^{-1}$ zusammen (§ 12).

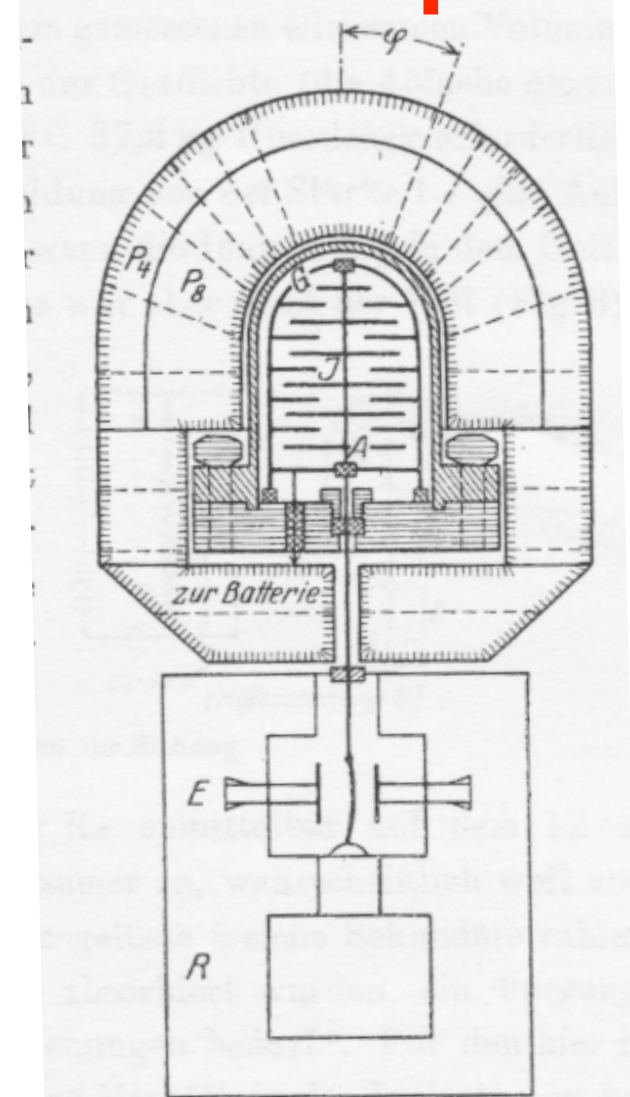


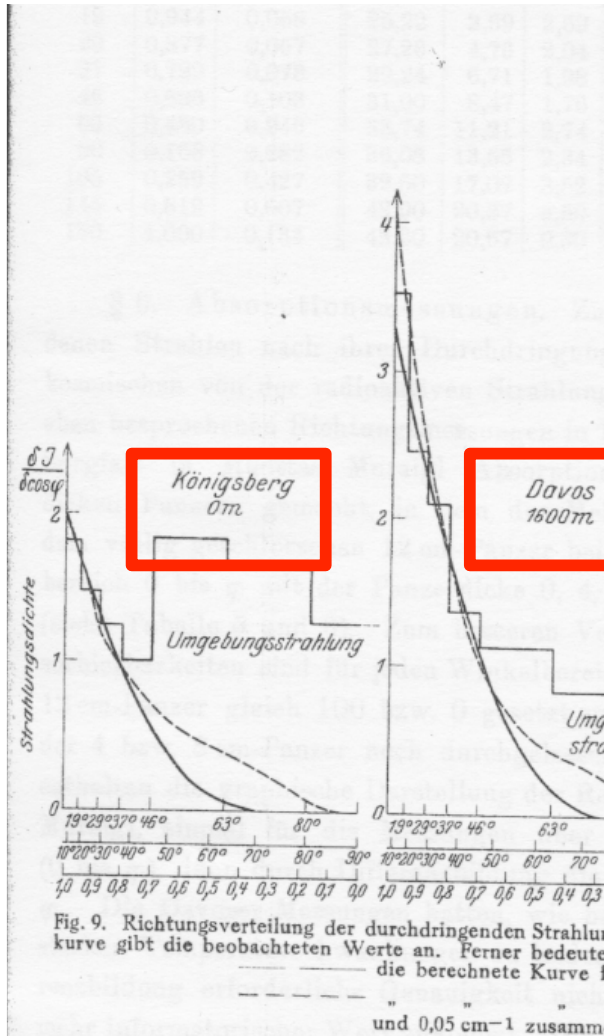
Fig. 1. Hochdruckionisationsgefäß J mit Drahtgitter G und Aufhänger A. Darum 12 cm Fe-Panzer (P_4, P_8), darunter Wulfsches Einfadenelektrometer E mit Registriervorrichtung R.

Neue Untersuchungen über die durchdringende Hesssche Strahlung.

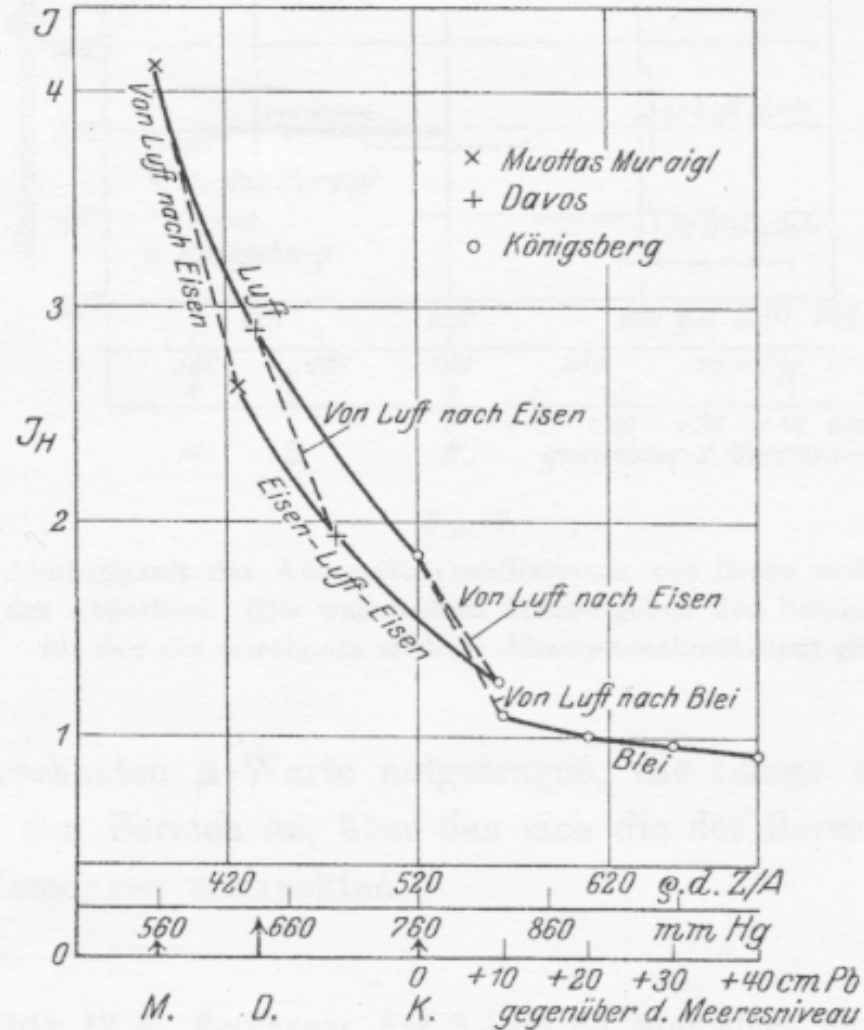
Von E. Steinke in Königsberg i. Pr.

Absorption in the atmosphere

Intensity as function for different altitudes



intensity



atmospheric overburden

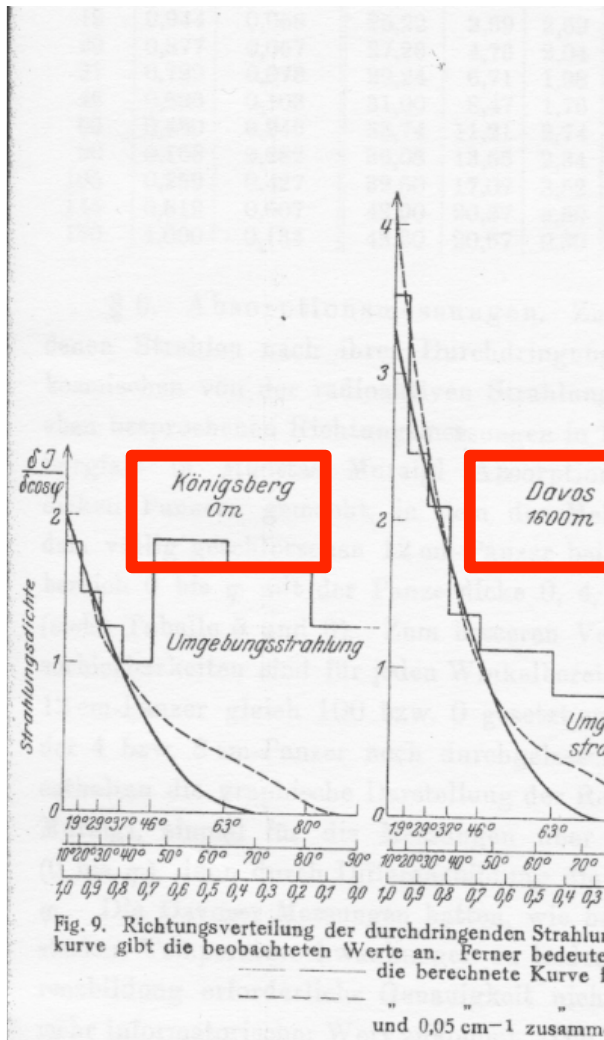
Intensität der durchdringenden Strahlung in Abhängigkeit von der Absorptionsschicht.

Neue Untersuchungen über die durchdringende Hesssche Strahlung.

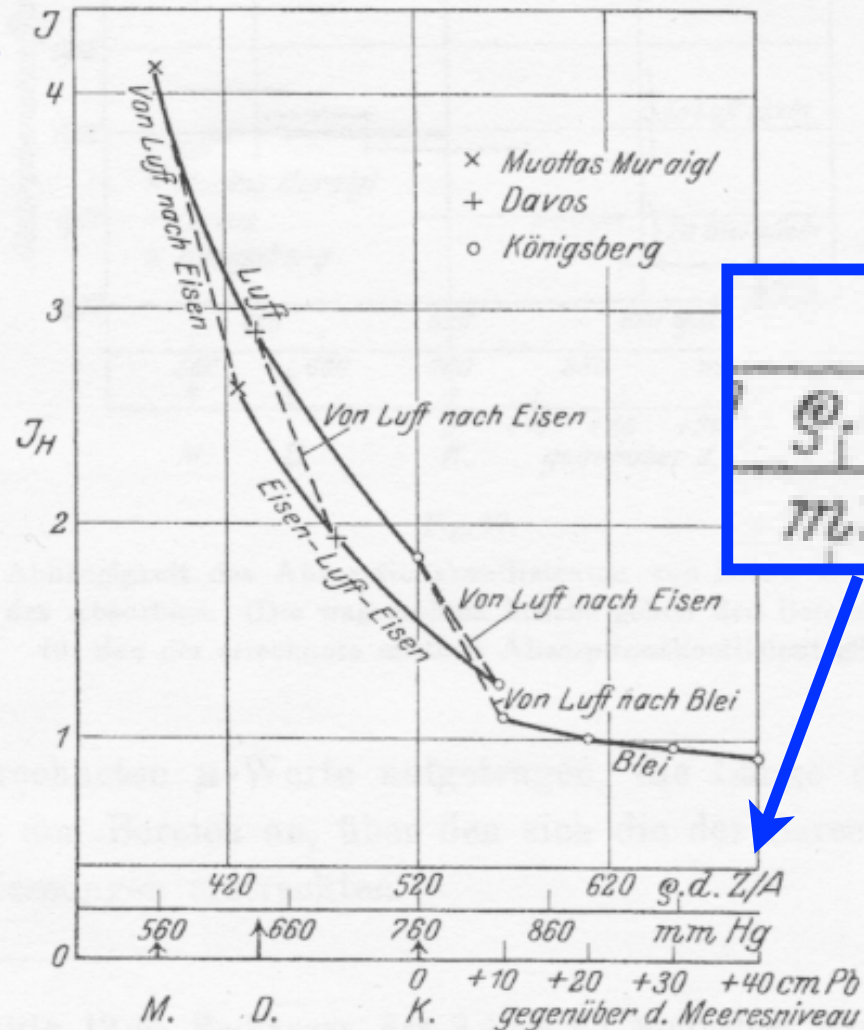
Von E. Steinke in Königsberg i. Pr.

Absorption in the atmosphere

Intensity as function for different altitudes



intensity



g.d. Z/A
mm Hg

atmospheric overburden

Intensität der durchdringenden Strahlung in Abhängigkeit von der Absorptionsschicht.

Über Schwankungen und Barometereffekt der kosmischen Ultrastrahlung im Meeresniveau.

Von E. Steinke in Königsberg i. Pr.

Mit 12 Abbildungen. (Eingegangen am 11. Juni 1930.)

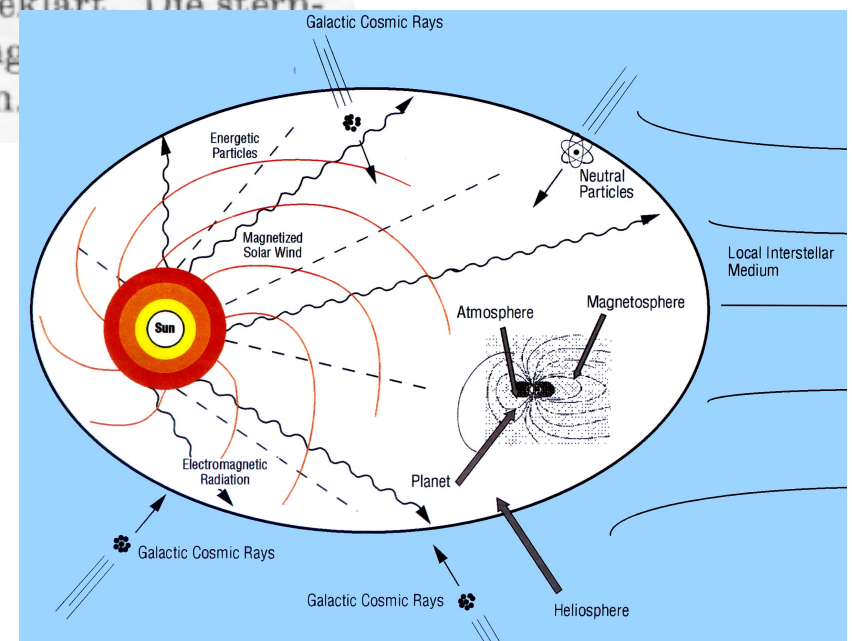
Während des Jahres 1929 wurden 20 Meßreihen mit insgesamt 100 Stunden aufgenommen teils mit allseitig geschlossenem Panzer, teils mit in der Mitte geöffnetem Panzer. Das Material ergab, daß neben zahlreichen Strahlungsschwankungen teils periodische teils unperiodische Intensitätsänderungen vorliegen, die teilweise spiegelbildlich zum Barometerstand („Barometereffekt“), periodische Schwankungen mit der Jahreszeit, ebensolche mit der Tageszeit, unregelmäßige Intensitätsänderungen über längere Zeiträume, endlich periodische Schwankungen nach Sternzeit. Der Verlauf während der einzelnen Meßreihen nicht konstant, sondern schwankend, im Mittel ist er aber ungefähr gleich dem Höhenbarometerstand. Für das Auftreten der anderen Schwankungen sind noch nicht geklärt. Die sternzeitlichen Schwankungen beweisen den kosmischen Ursprung der Ultrastrahlung und die Existenz besonderer Emissionszentren.

Barometric effect

**anti-correlation with pressure
-> barometric effect**

annual modulation of cosmic rays

siderial modulation of cosmic rays



Über Schwankungen und Barometereffekt der kosmischen Ultrastrahlung im Meeresniveau.

Von E. Steinke in Königsberg i. Pr.

Barometric effect

Während d
stunden au
Dicke, teil
material e
teils perio
Schwanku
mäßige In
endlich pe
während c
kungen, in
für das Au
zeitlichen

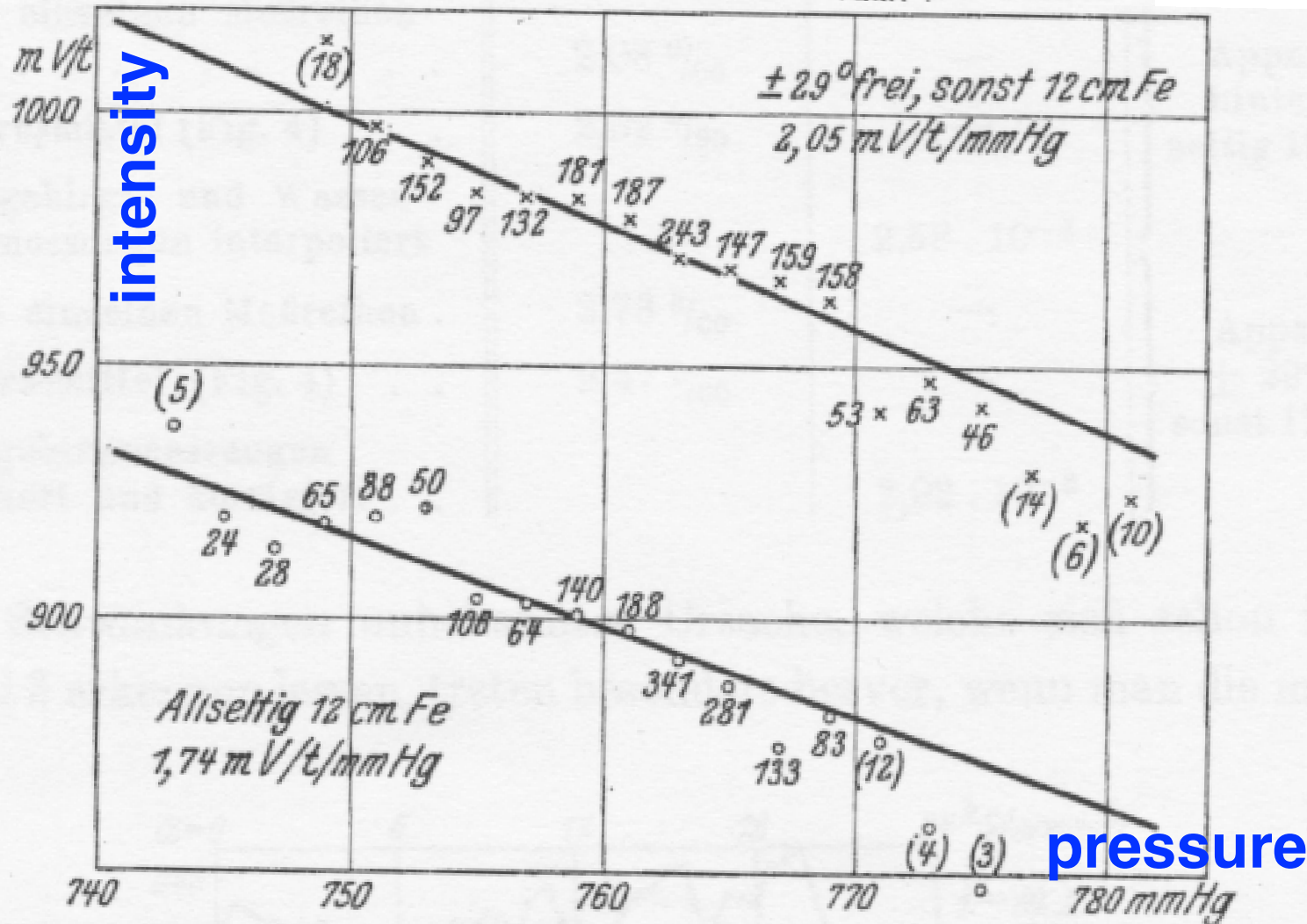
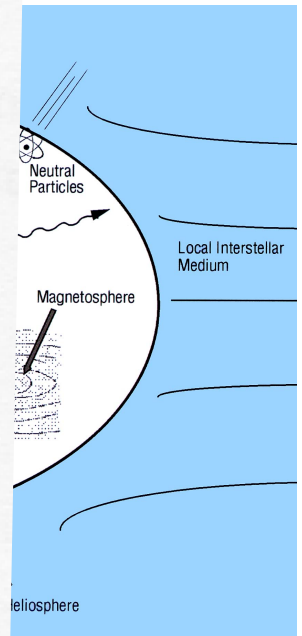


Fig. 4. Zusammenhang zwischen Barometerstand und Ionisation (Jahresmittel der Stundenwerte; die Zahlen geben die Anzahl der Stundenwerte an).

with
ect

on of



si

Absorption in Lake Constance 1928

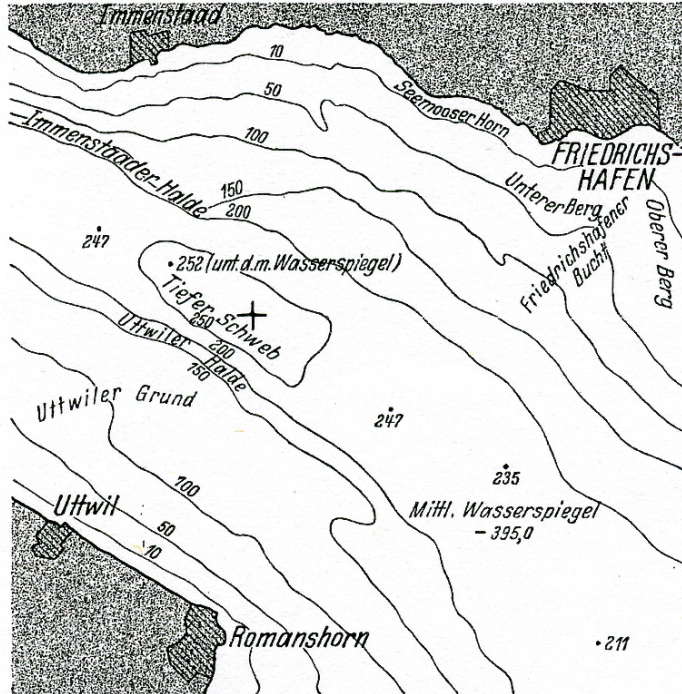


Fig. 5

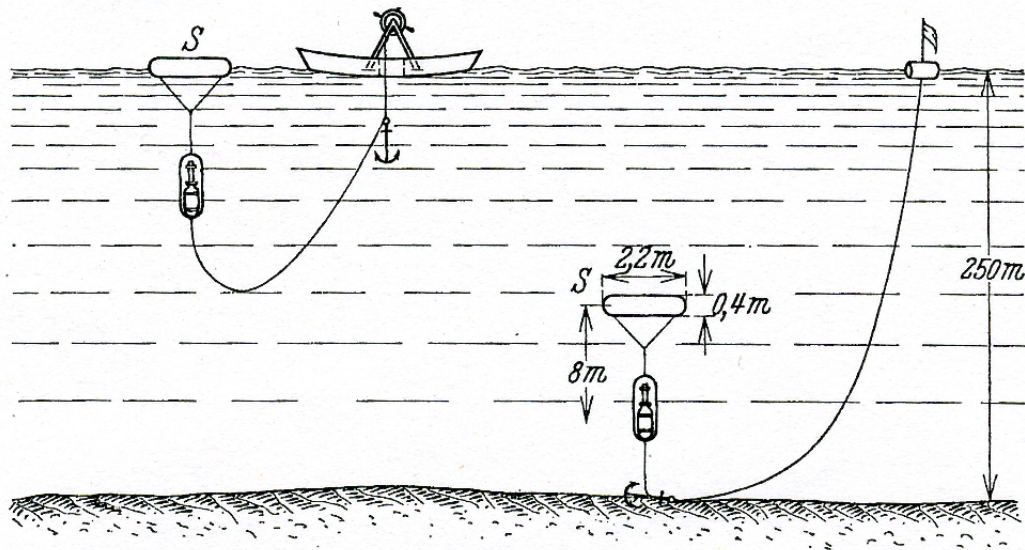
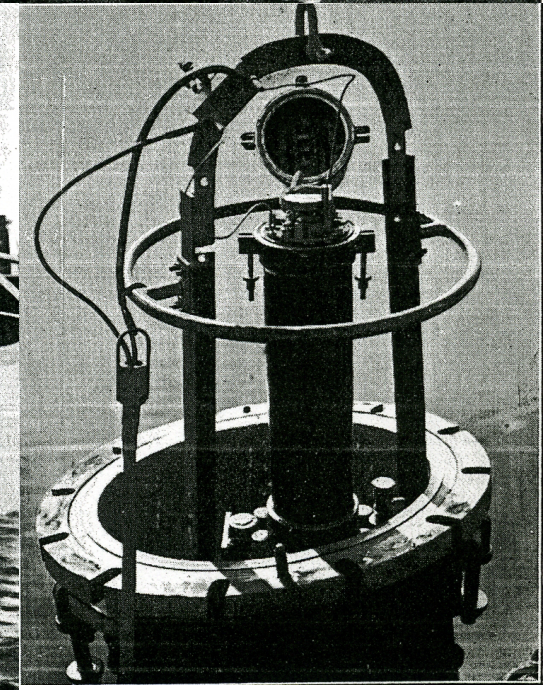
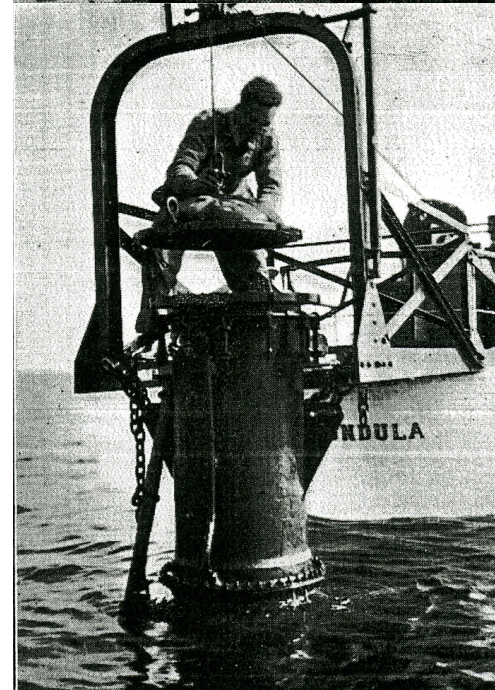


Fig. 6. Die „schwebende“ Verankerung des Apparates.



Absorption in Lake Constance 1928

Ionization chamber with electrometer read-out
automatic each hour, up to 8 days

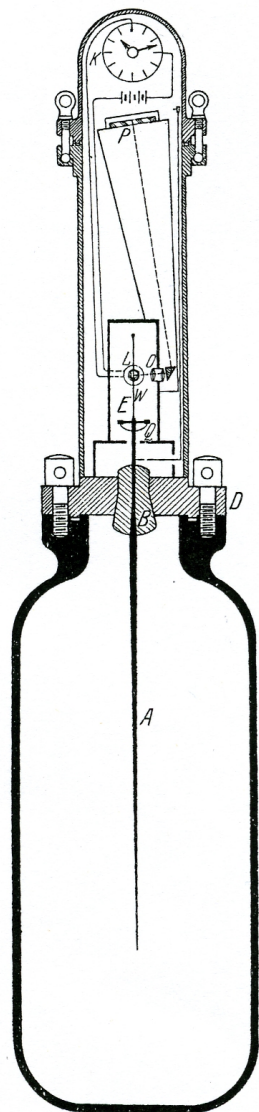


Fig. 1.
Aufbau des wasserdichten
Registrierapparates.

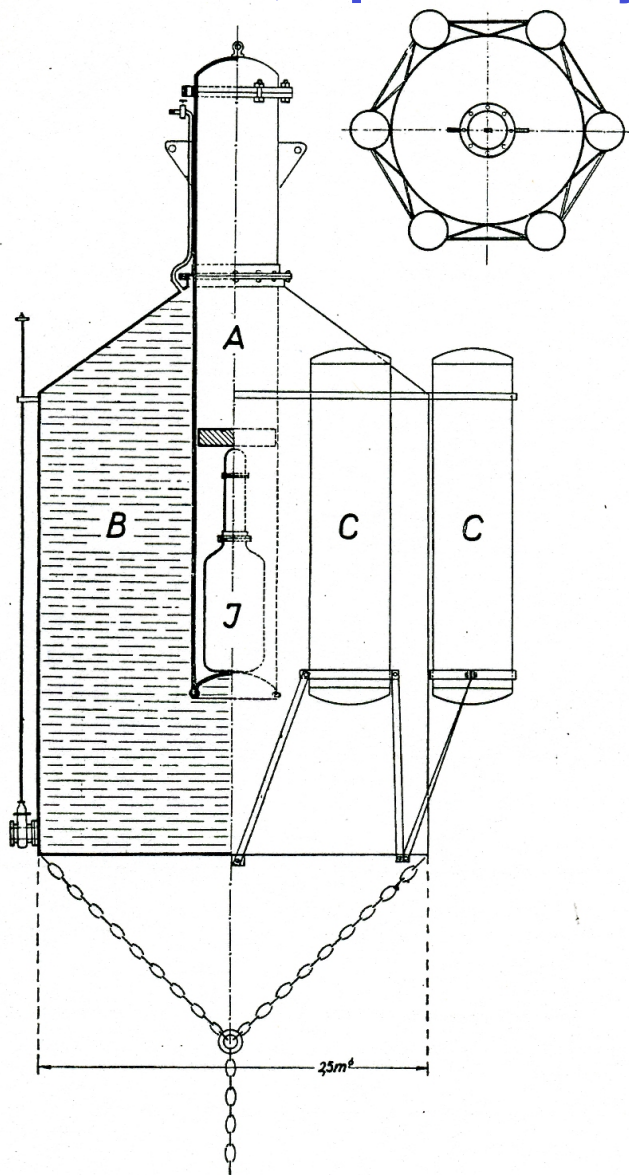


Fig. I.



Absorption in Lake Constance 1928

*Intensitätsverlauf der
Ultrastrahlung im Wasser*

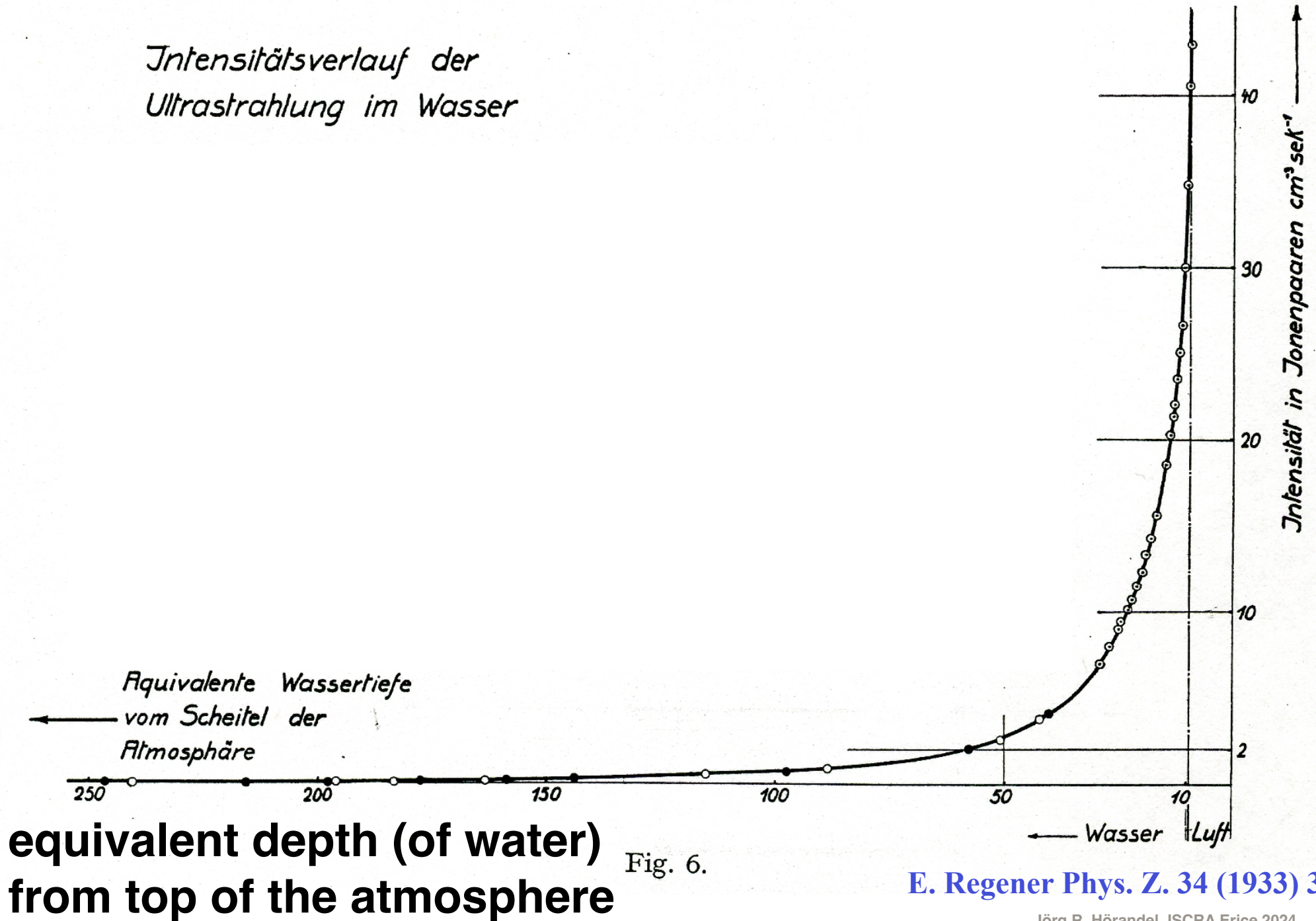


Fig. 6.

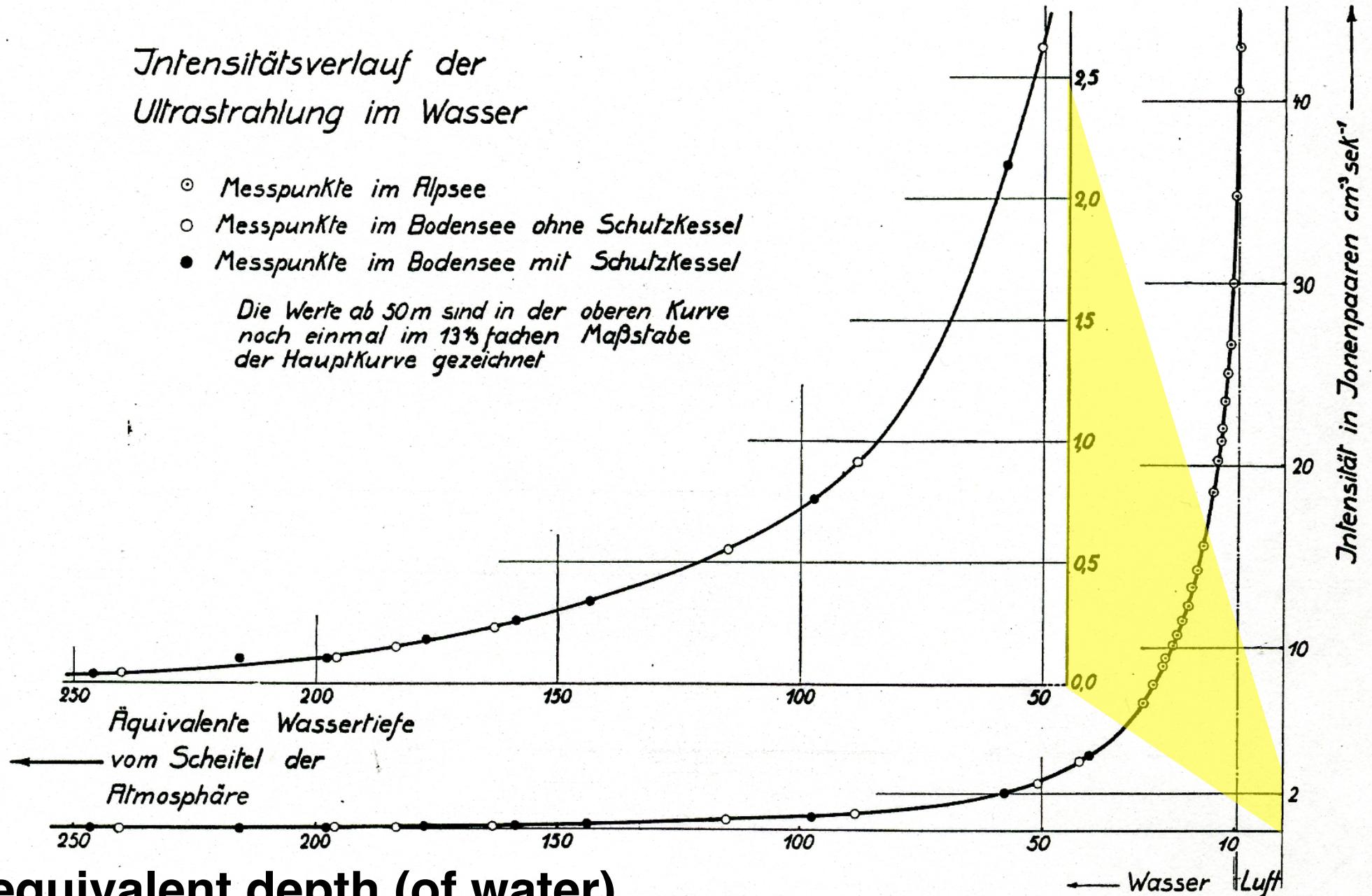
E. Regener Phys. Z. 34 (1933) 306

Absorption in Lake Constance 1928

*Intensitätsverlauf der
Ultrastrahlung im Wasser*

- ⊙ Messpunkte im Alpsee
- Messpunkte im Bodensee ohne Schutzkessel
- Messpunkte im Bodensee mit Schutzkessel

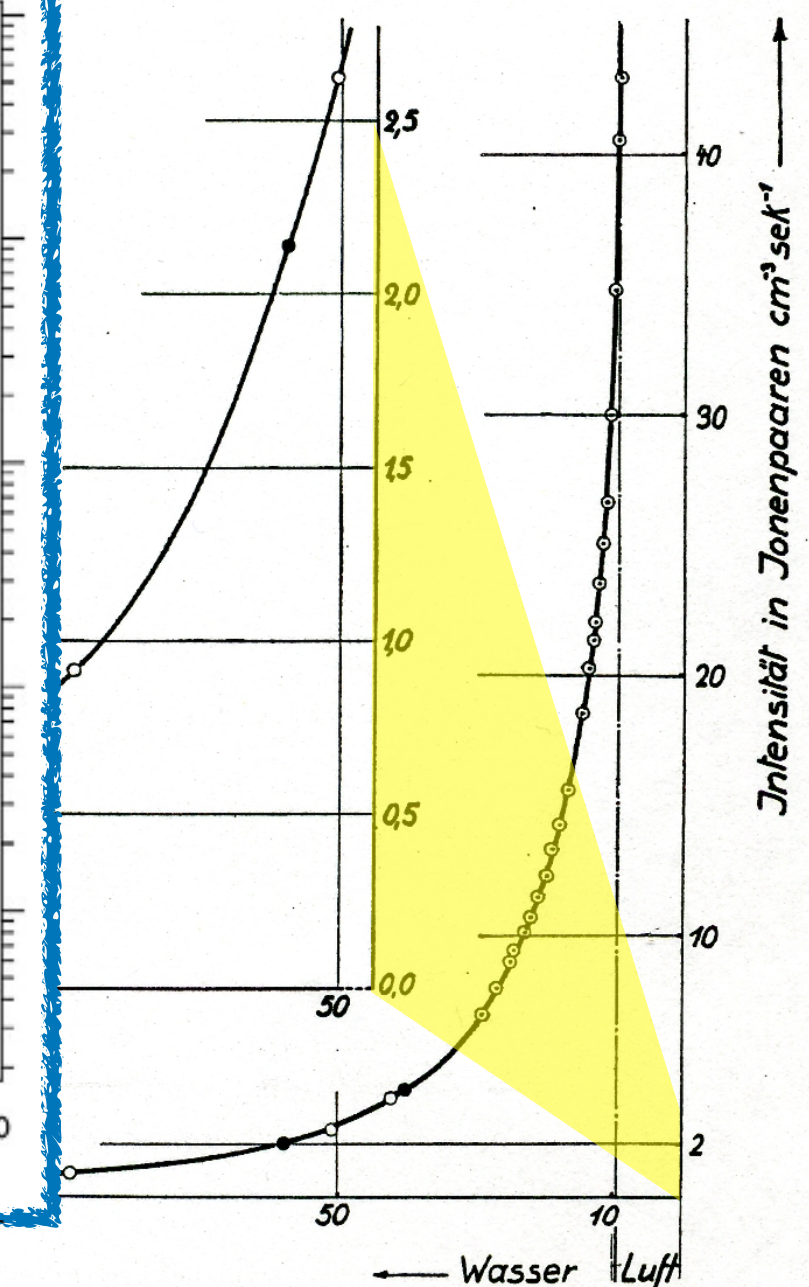
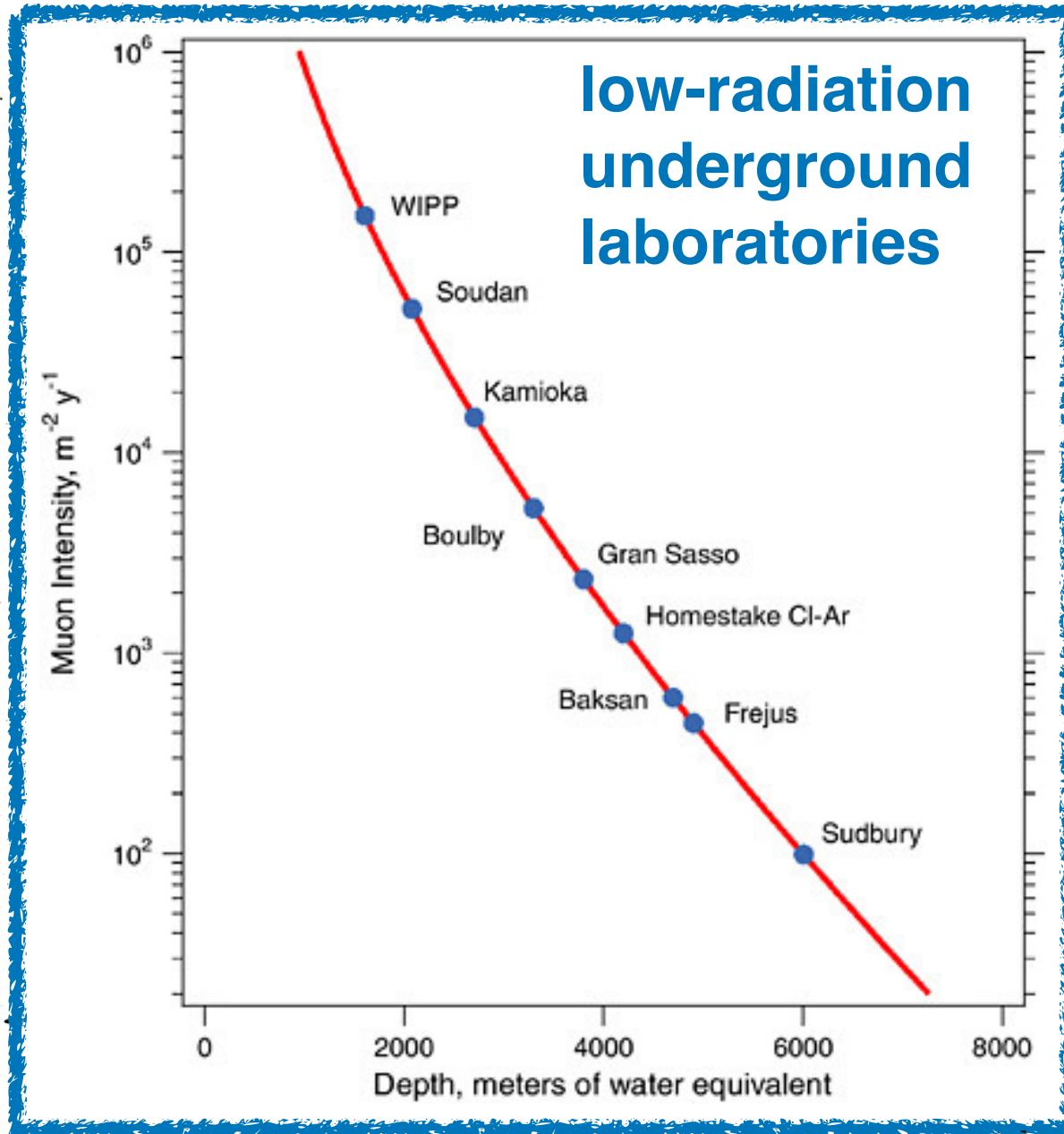
*Die Werte ab 50m sind in der oberen Kurve
noch einmal im 13% fachen Maßstabe
der Hauptkurve gezeichnet*



**equivalent depth (of water)
from top of the atmosphere**

Fig. 6.

Absorption in Lake Constance 1928



equivalent depth (of water)
from top of the atmosphere

Fig. 6.

E. Regener Phys. Z. 34 (1933) 306



Hess

Steinke

Regener

**Three pioneers of Cosmic Ray research
Regener demonstrates his balloon electrometer
(Immenstaad/Lake Constance, August 1932).**

derung zu danken, die sie in seinen Bestrebungen
von Anhe
groß

**Kolhörster
A new electrometer**

„Wer
gebracht hat“)

1) Oskar Taussig, aero-aeronautic research work
(The First World Conference, London 1924),
vgl. auch „Elektrotechnik und Maschinenbau“, Zeitschrift
des Elektrotechnischen Vereins in Wien, Heft 46, 1924.

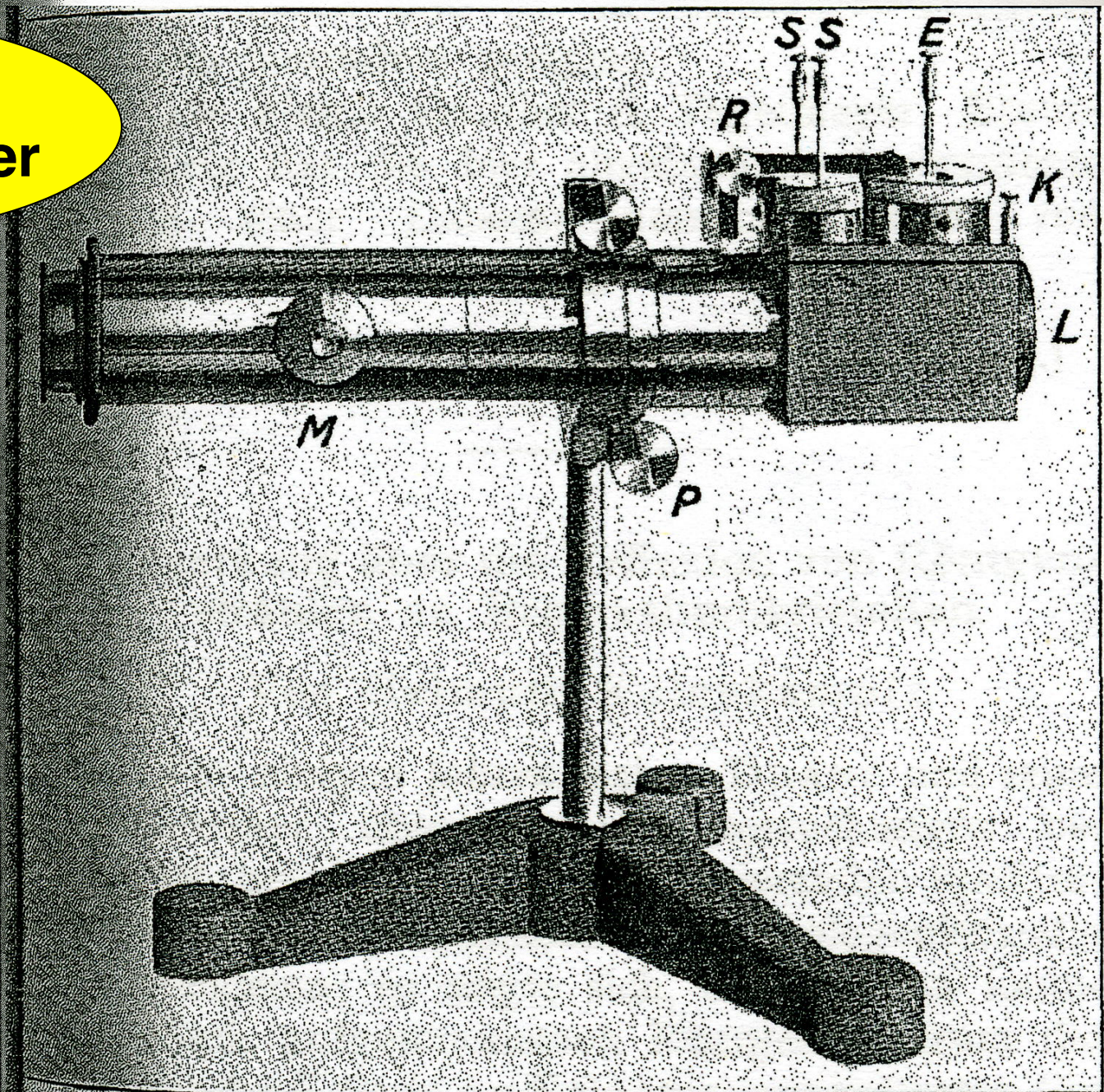
(Elektrische Anlagen 28. August 1925.)

Ein neues Fadenelektrometer.

Von Werner Kolhörster.

Zu Messungen der durchdringenden Strahlung
hatte ich für meine neuen Strahlungsapparate
ein Fadenelektrometer konstruiert¹⁾, das ohne
die bei derartigen Instrumenten notwendige
Temperaturkompensation arbeitet. Da es sich
auch für andere elektrostatische Messungen
seiner Vorzüge und allgemeinen Verwendbarkeit
halber als geeignet erwies, so seien hier einige
Angaben über die Instrumente²⁾ gemacht.

Prinzip: Als Gegenkraft gegen die elektro-
statischen Abstoßungskräfte dient allein die
Biegeelastizität der feinen Quarzfäden, die
die Form vertikal stehender, frei tragender
Schlingen haben und deren Enden in einigen
Millimetern Abstand voneinander an einem
Metallblech befestigt sind, das in den Isolator
eingesetzt wird. Entsprechend den Ein- und
Zweifadenelektrometern kann man Systeme mit
einer oder zwei kongruenten Schlingen ver-
wenden, die von einem Mikroskop mit Okular-
mikrometer am Scheitel der Schlingen abge-
lesen werden. Lädt man das System, so tritt
keine merkliche Formänderung der Schlingen
ein, diese bewegen sich vielmehr in der Horizon-



ladung zu verändern. Eine feinere Änderung
läßt sich durch Verschwenken der Schneiden
erzielen, die um die längere Rechteckseite dreh-
bar, mehr oder weniger den Fäden genähert



Kohlhörster - balloon flight 13. May 1934

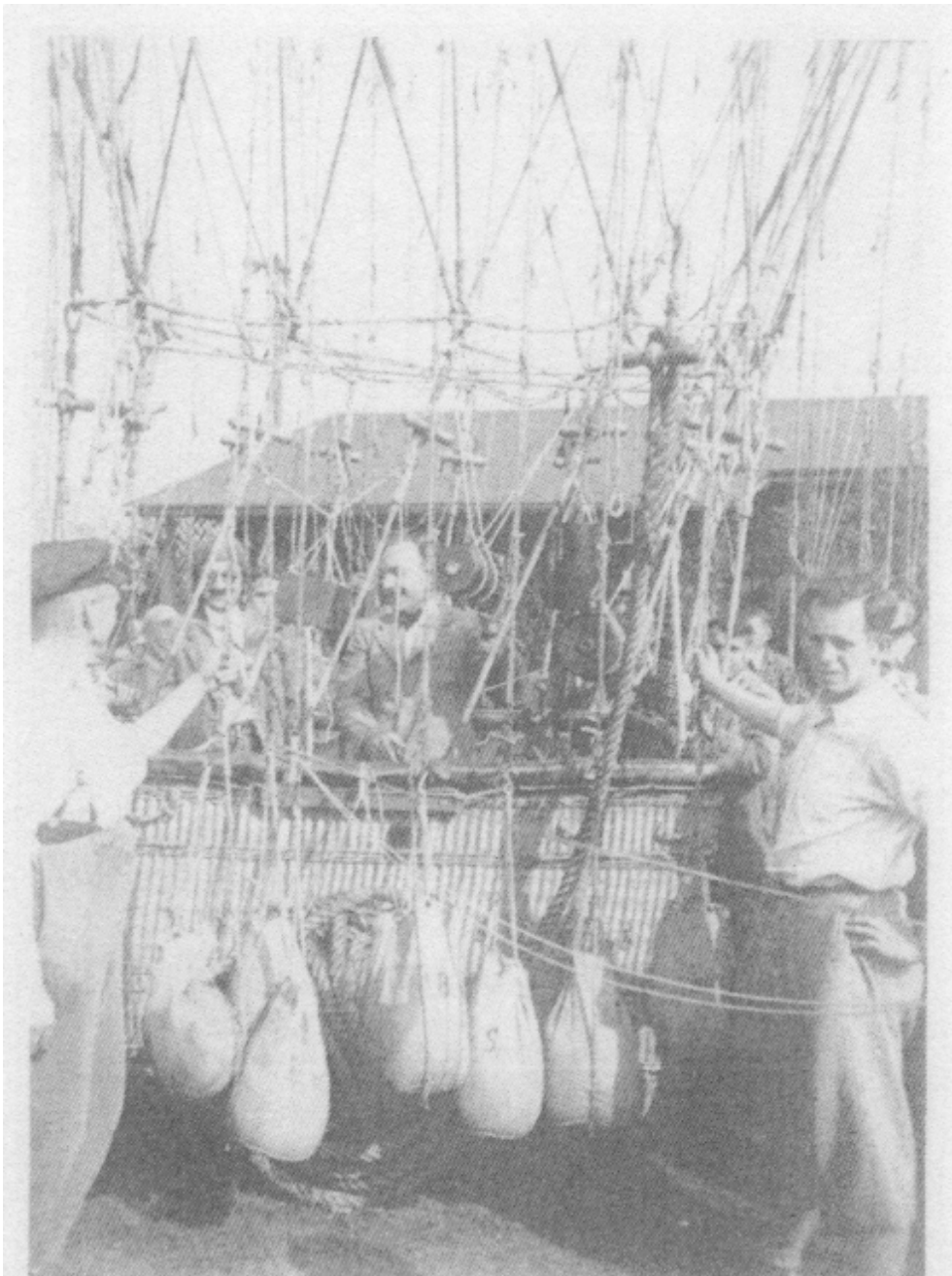


Abb.12 Vor dem Aufstieg

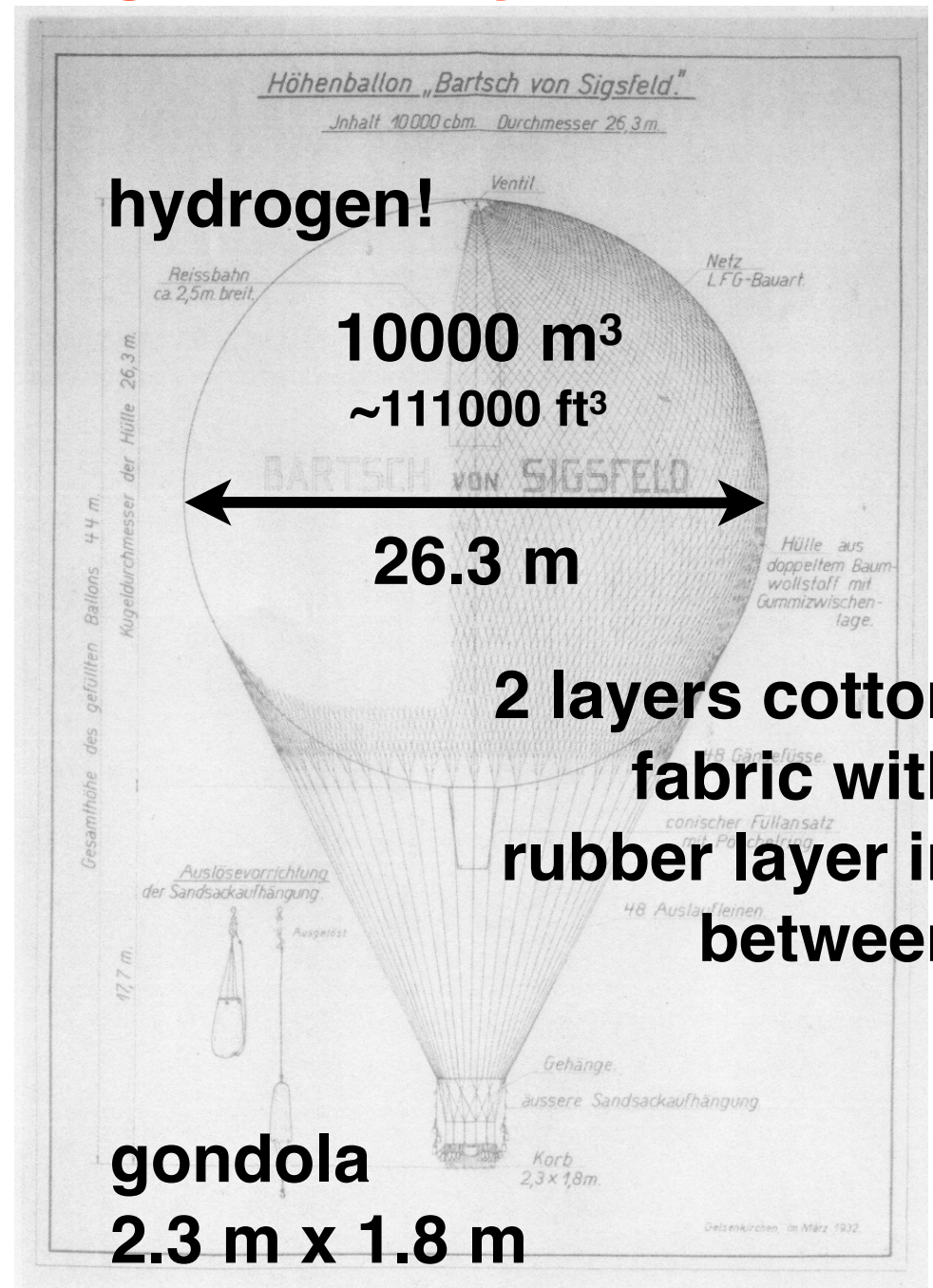


Abb.19 Höhenballon "Bartsch von Sigsfeld"
Füllung am 13.5.34 rd. 4400 m³.

Kohlhörster - balloon flight 13. May 1934



Abb. 17

Dr. Schrenk

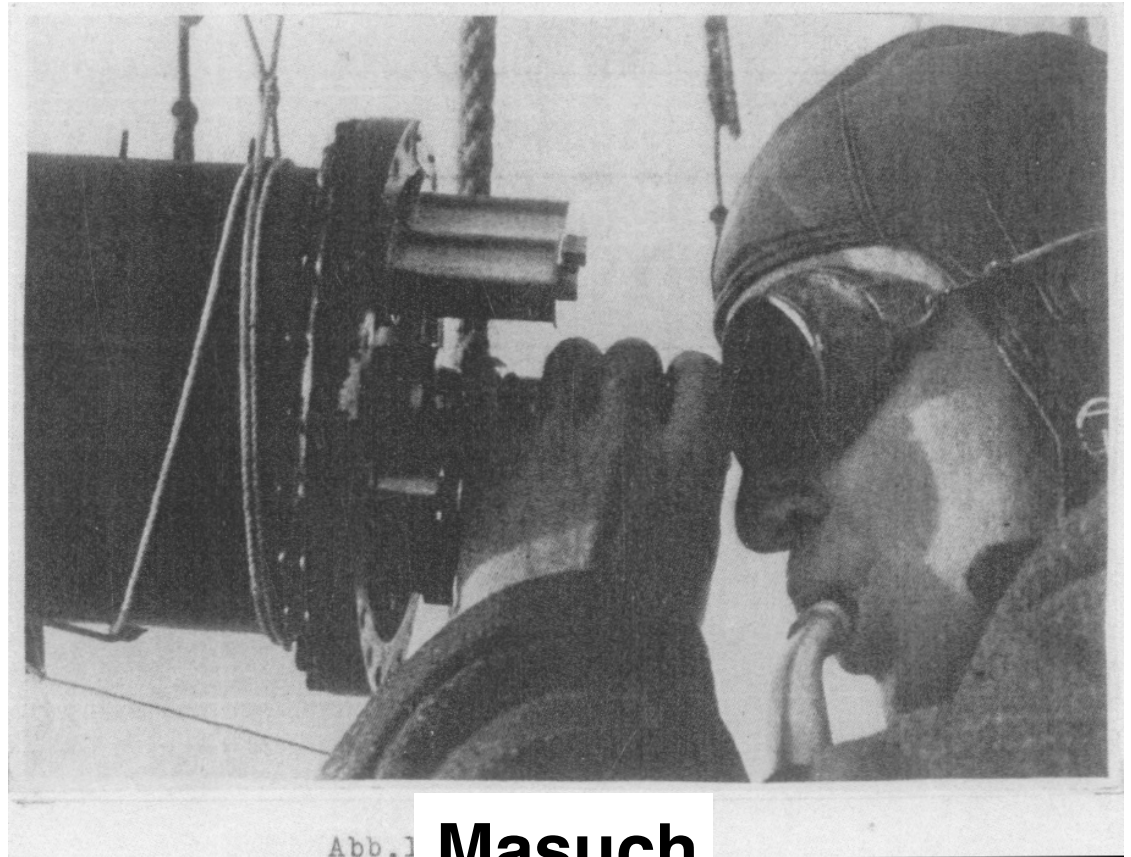


Abb. 1

Masuch

Measurements of the cosmic-ray intensity (Höhenstrahlung) up to 12000 m

Das Wesen der Höhenstrahlung.

Von **W. Bothe** und **W. Kolhörster**.

Mit 8 Abbildungen. (Eingegangen am 18. Juni 1929.)

the nature of the „high-altitude radiation“

ist die Höhenstrahlung

coincidence technique

absorber

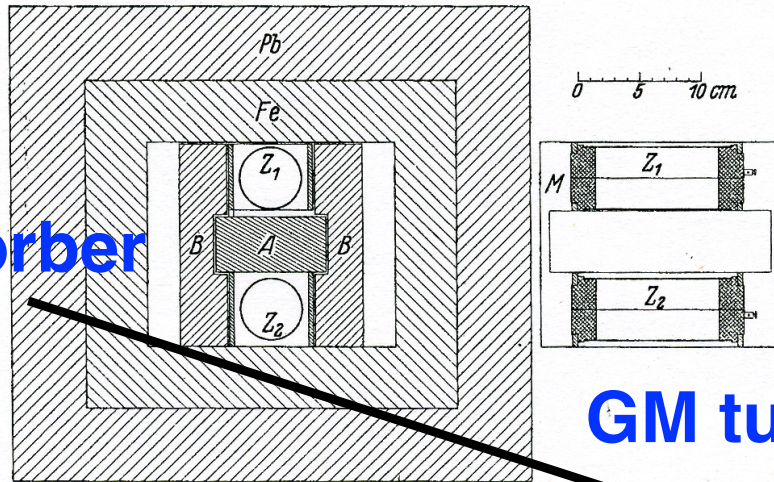


Fig. 1.

innen 5 cm Durchmesser und 10 cm Länge; sie waren anfangs aus 1 mm starkem Messing, später aus 1 mm starkem Zink hergestellt und an den Enden mit Ebonitstopfen verschlossen, welche die zentralen Drähte trugen. Diese waren nach dem Verfahren von Geiger und Müller präpariert. Die Zählrohre waren mit trockener, kohlenensäure- und emanationsfreier Luft von 4 bis 6 cm Hg-Druck gefüllt. Sie wurden durch ein Messinggestell *M* getragen, welches so eingerichtet war, daß Absorberschichten bis zu 45 mm Dicke zwischen die Zählrohre gebracht werden konnten. Seitlich waren die Rohre durch Bleiklötze *B* geschützt; diese hatten Nuten, in welche der Absorber eingriff. Die Dicke dieser Seitenblenden war stets so bemessen, daß ein Strahlenteilchen, welches etwa durch Streuung um den Absorber herum aus dem einen Zählrohr in das andere gelangen

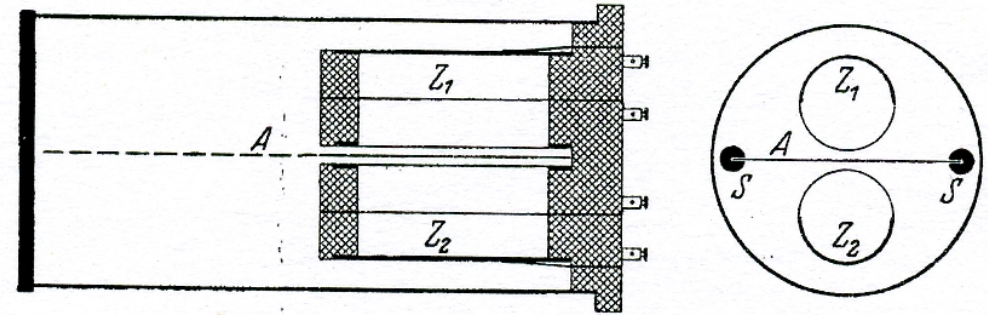


Fig. 2.

coinc./min

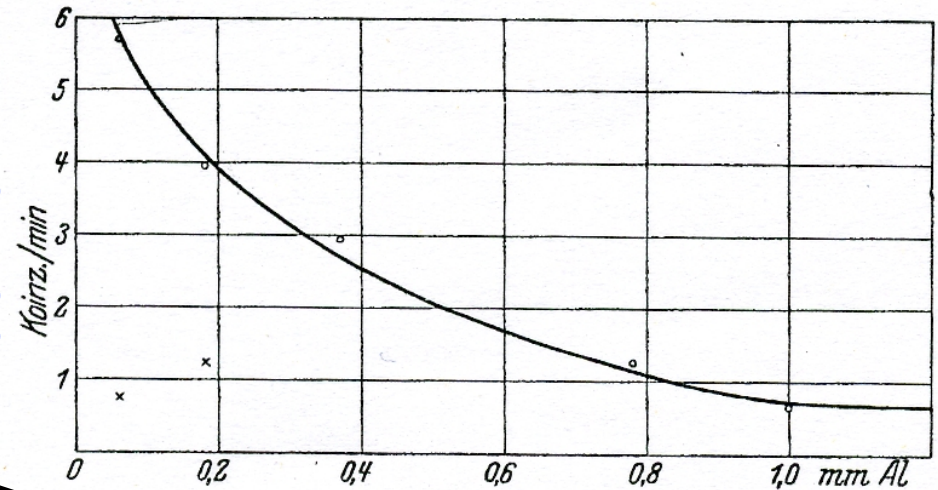


Fig. 3.

absorber thickness



W. Bothe
Nobel Prize 1954

W. Bothe & W. Kolhörster, *Z. f. Phys.* 56 (1929) 751

Dreifachkoinzidenzen der Ultrastrahlung aus vertikaler Richtung in der Stratosphäre *).

I. Meßmethode und Ergebnisse.

Von Georg Pfozter in Stuttgart.

Mit 11 Abbildungen. (Eingegangen am 9. Juni 1936.)

Mit einer selbstaufzeichnenden Apparatur werden bei drei Registrierballon-aufstiegen Dreifachkoinzidenzen der Ultrastrahlung aus vertikaler Richtung bis zu 10 mm Hg Luftdruck (29 km Höhe ü. M.) gemessen. Die Kurve der Zählrohrkoinzidenzen in Abhängigkeit vom Luftdruck zeigt ein Maximum bei 80 mm Hg und einen Buckel bei 300 mm Hg. Die Kurve kann gegen das Ende der Atmosphäre extrapoliert werden.

three-fold coincidences of the ultra rays from vertical direction in the stratosphere

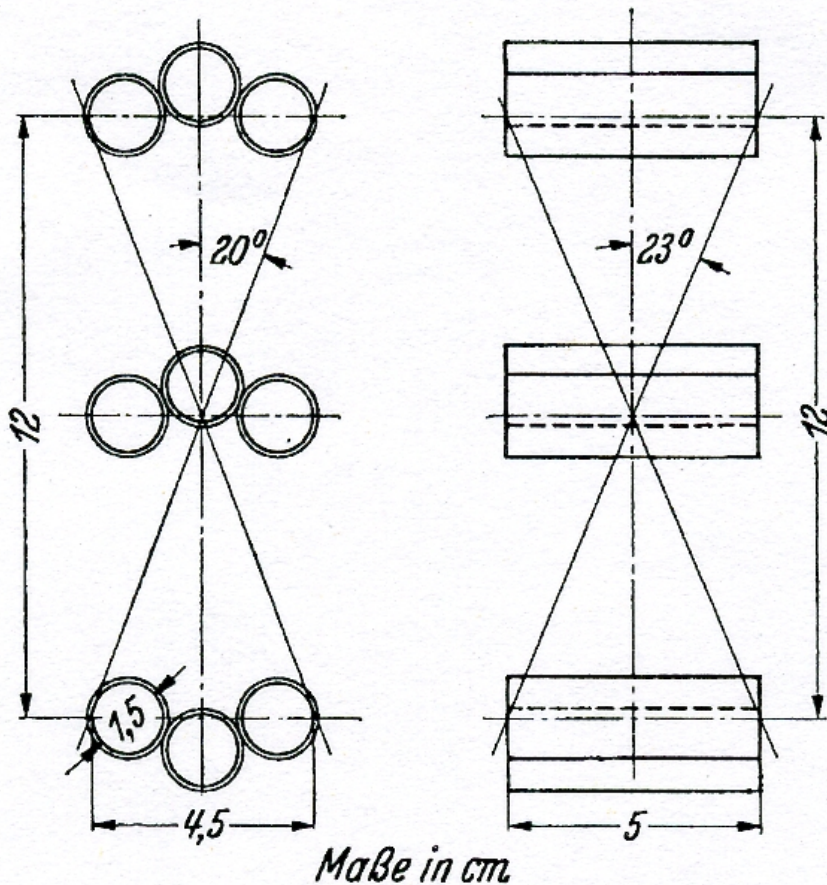


Fig. 1. Größenverhältnisse und Ausblendung der Koinzidenzanordnung.

Dreifachkoinzidenzen der Ultrastrahlung aus vertikaler Richtung in der Stratosphäre *).

I. Meßmethode und Ergebnisse.

Von Georg Pfozter in Stuttgart.

Mit 11 Abbildungen. (Eingegangen am 9. Juni 1936.)

Mit einer selbstaufzeichnenden Apparatur werden bei drei Registrierballon aufstiegen Dreifachkoinzidenzen der Ultrastrahlung aus vertikaler Richtung bis zu 10 mm Hg Luftdruck (29 km Höhe ü. M.) gemessen. Die Kurve der Zählrohrkoinzidenzen in Abhängigkeit vom Luftdruck zeigt ein Maximum bei 80 mm Hg und einen Buckel bei 300 mm Hg. Die Kurve kann gegen das Ende der Atmosphäre extrapoliert werden.

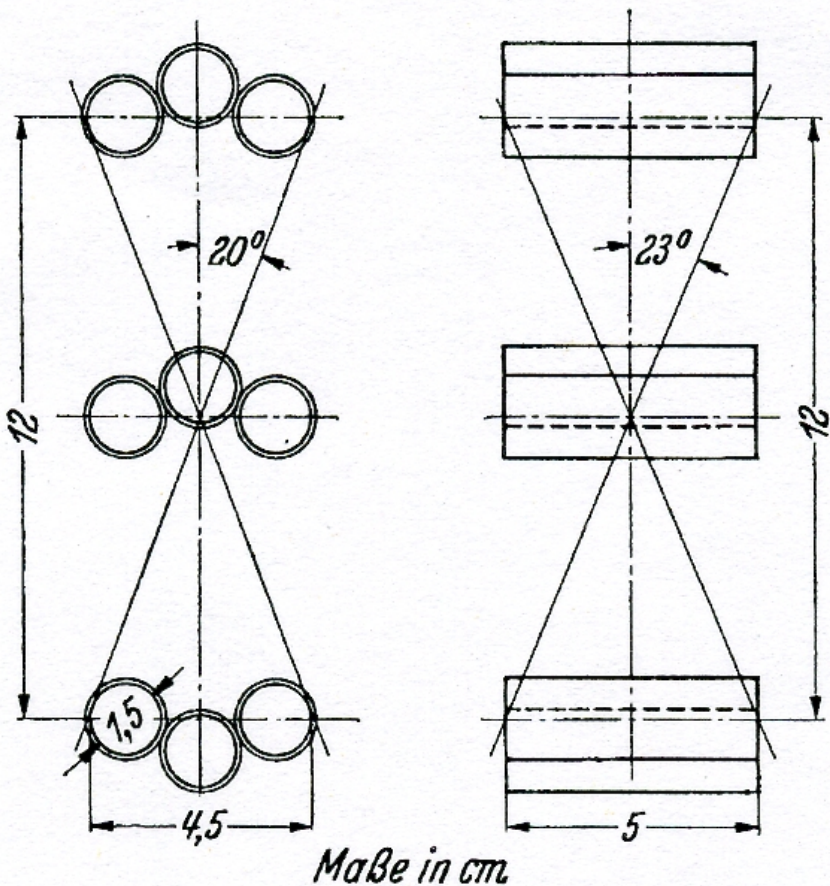


Fig. 1. Größenverhältnisse und Ausblendung der Koinzidenzanordnung.

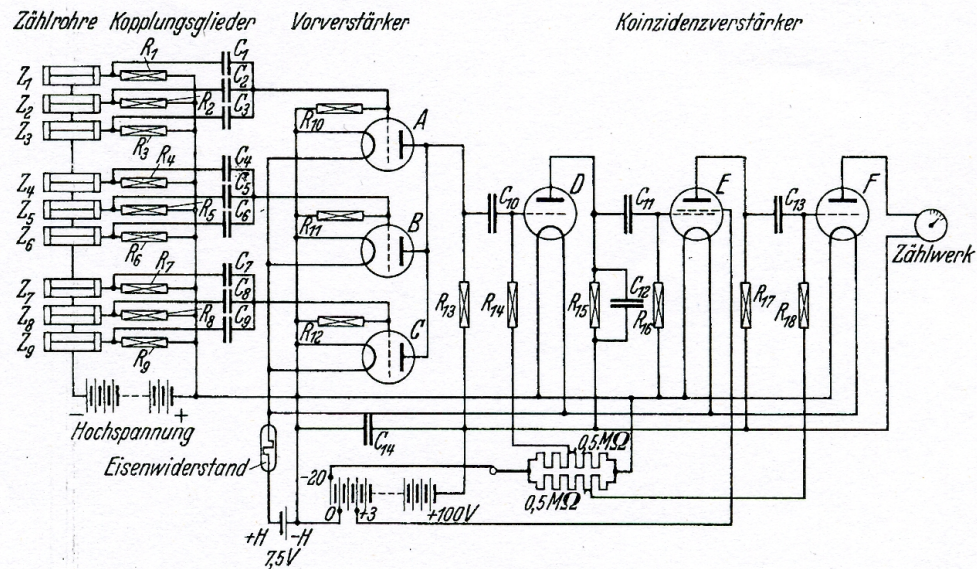


Fig. 2. Schaltschema des Koinzidenzenverstärkers.

- | | |
|--|---|
| $R_1 - R_9 = 10^9 \text{ Ohm,}$ | $C_1 - C_9 = 25 \mu\mu\text{F,}$ |
| $R_{10} - R_{12} = 10^8 \text{ Ohm,}$ | $C_{10} = 100 \mu\mu\text{F,}$ |
| $R_{13} = 0,5 \cdot 10^8 \text{ Ohm,}$ | $C_{11} = C_{12} = 250 \mu\mu\text{F,}$ |
| $R_{14} = R_{15} = R_{17} = 3 \cdot 10^8 \text{ Ohm,}$ | $C_{13} = 1000 \mu\mu\text{F,}$ |
| $R_{16} = R_{18} = 5 \cdot 10^8 \text{ Ohm.}$ | $C_{14} = 1 \mu\text{F.}$ |

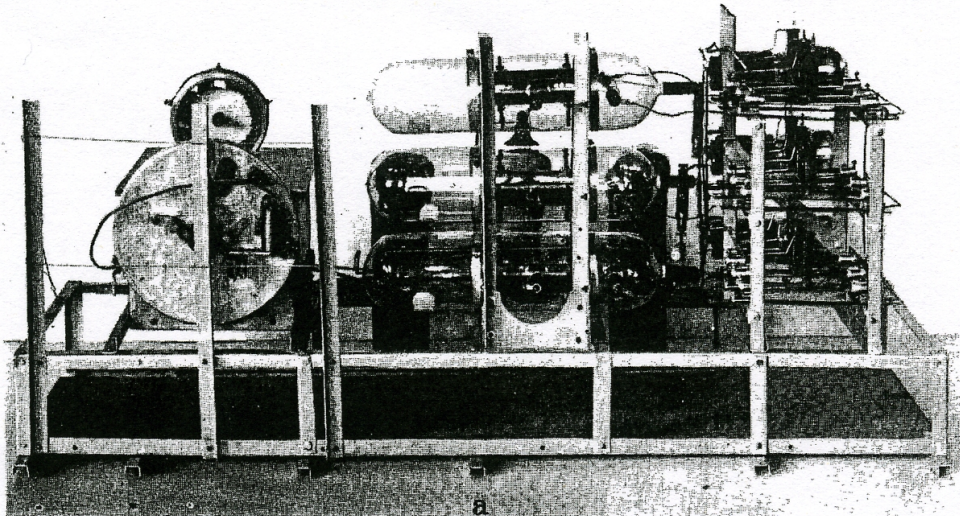
Dreifachkoinzidenzen der Ultrastrahlung aus vertikaler Richtung in der Stratosphäre *).

I. Meßmethode und Ergebnisse.

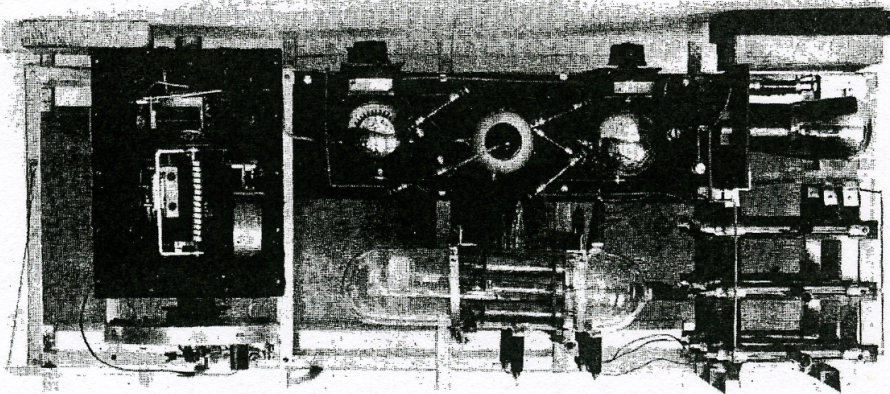
Von Georg Pfozter in Stuttgart.

Mit 11 Abbildungen. (Eingegangen am 9. Juni 1936.)

Mit einer selbstaufzeichnenden Apparatur werden bei drei Registrierballon aufstiegen Dreifachkoinzidenzen der Ultrastrahlung aus vertikaler Richtung bis zu 10 mm Hg Luftdruck (29 km Höhe ü. M.) gemessen. Die Kurve der Zählrohr koinzidenzen in Abhängigkeit vom Luftdruck zeigt ein Maximum bei 80 mm Hg und einen Buckel bei 300 mm Hg. Die Kurve kann gegen das Ende der Atmosphäre extrapoliert werden.



a



b

Fig. 5. Aufbau der Registrierapparatur. a) Von der Seite b) von oben gesehen.

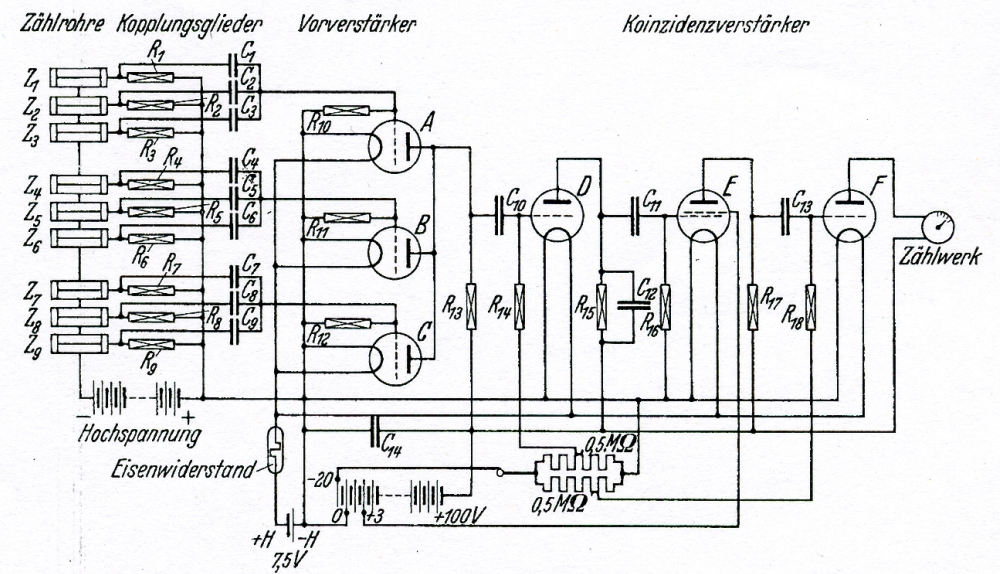


Fig. 2. Schaltschema des Koinzidenzenverstärkers.

- | | |
|--|---|
| $R_1 - R_9 = 10^9 \text{ Ohm,}$ | $C_1 - C_9 = 25 \mu\mu\text{F,}$ |
| $R_{10} - R_{12} = 10^8 \text{ Ohm,}$ | $C_{10} = 100 \mu\mu\text{F,}$ |
| $R_{13} = 0,5 \cdot 10^8 \text{ Ohm,}$ | $C_{11} = C_{12} = 250 \mu\mu\text{F,}$ |
| $R_{14} = R_{15} = R_{17} = 3 \cdot 10^8 \text{ Ohm,}$ | $C_{13} = 1000 \mu\mu\text{F,}$ |
| $R_{16} = R_{18} = 5 \cdot 10^8 \text{ Ohm.}$ | $C_{14} = 1 \mu\text{F.}$ |

Dreifachkoinzidenzen der Ultrastrahlung aus vertikaler Richtung in der Stratosphäre *).

I. Meßmethode und Ergebnisse.

Von **Georg Pfozter** in Stuttgart.

Mit 11 Abbildungen. (Eingegangen am 9. Juni 1936.)

Mit einer selbstaufzeichnenden Apparatur werden bei drei Registrierballon aufstiegen Dreifachkoinzidenzen der Ultrastrahlung aus vertikaler Richtung bis zu 10 mm Hg Luftdruck (29 km Höhe ü. M.) gemessen. Die Kurve der Zählrohrkoinzidenzen in Abhängigkeit vom Luftdruck zeigt ein Maximum bei 80 mm Hg und einen Buckel bei 300 mm Hg. Die Kurve kann gegen das Ende der Atmosphäre extrapoliert werden.

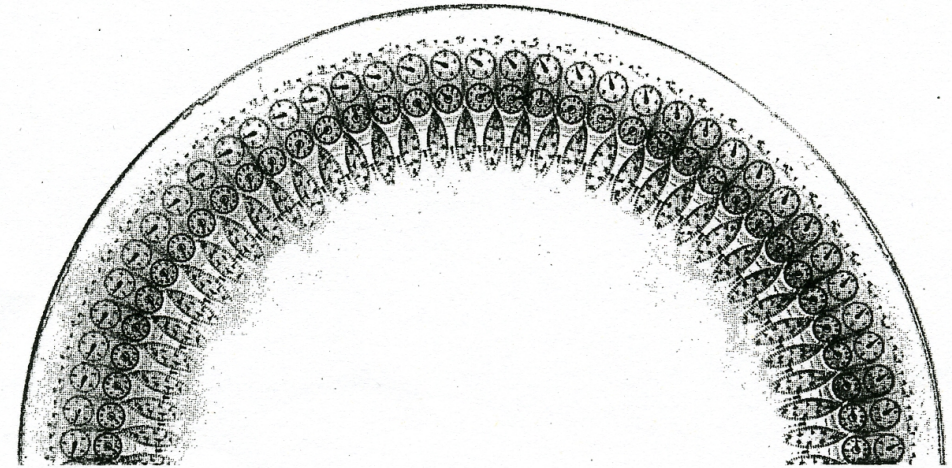
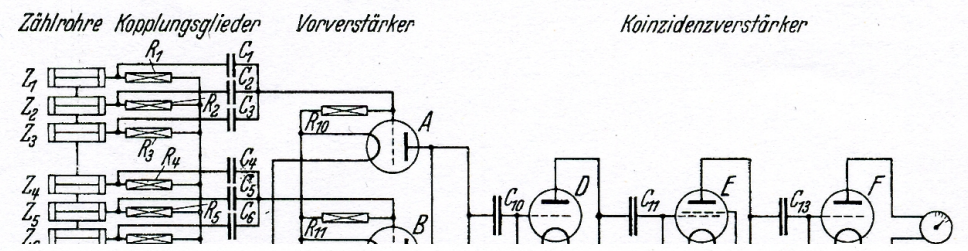
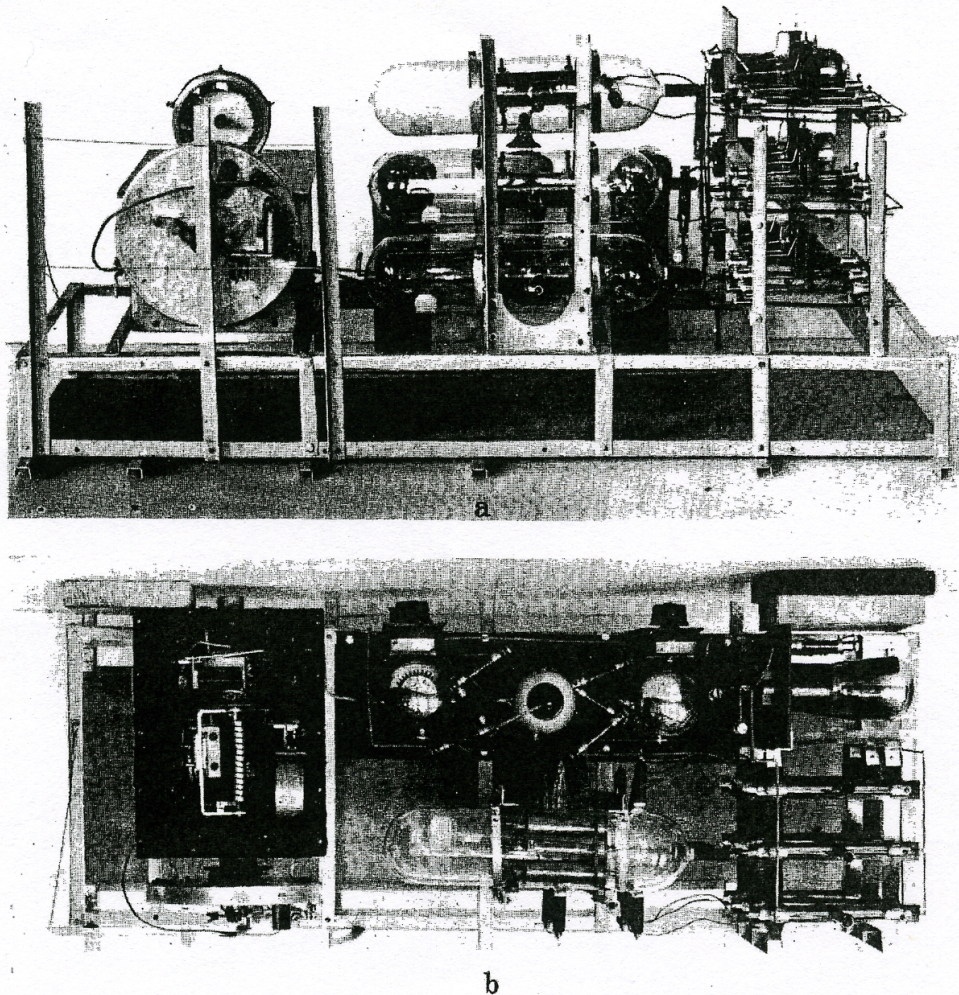


Fig. 4. a) Aufstiegplatte (nat. Größe, Hälfte); b) Vergrößerter Ausschnitt.

3) Die Firma Gebr. Junghans, Schramberg, hat uns freundlicherweise diese schönen Zählwerke hergestellt.

G. Pfozter, Z. f. Phys. 102 (1936) 23

Dreifachkoinzidenzen der Ultrastrahlung aus vertikaler Richtung in der Stratosphäre *).

I. Meßmethode und Ergebnisse.

Von Georg Pfozter in Stuttgart.

Mit 11 Abbildungen. (Eingegangen am 9. Juni 1936.)

Mit einer selbstaufzeichnenden Apparatur werden bei drei Registrierballon aufstiegen Dreifachkoinzidenzen der Ultrastrahlung aus vertikaler Richtung bis zu 10 mm Hg Luftdruck (29 km Höhe ü. M.) gemessen. Die Kurve der Zählrohr koinzidenzen in Abhängigkeit vom Luftdruck zeigt ein Maximum bei 80 mm Hg und einen Buckel bei 300 mm Hg. Die Kurve kann gegen das Ende der Atmosphäre extrapoliert werden.

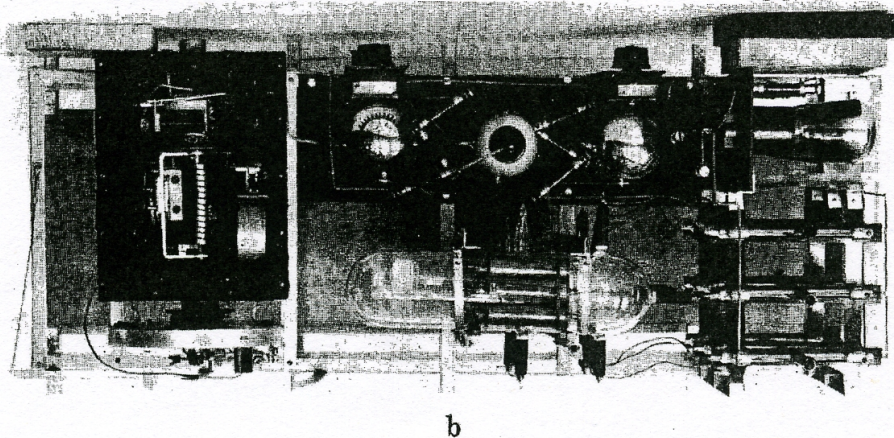
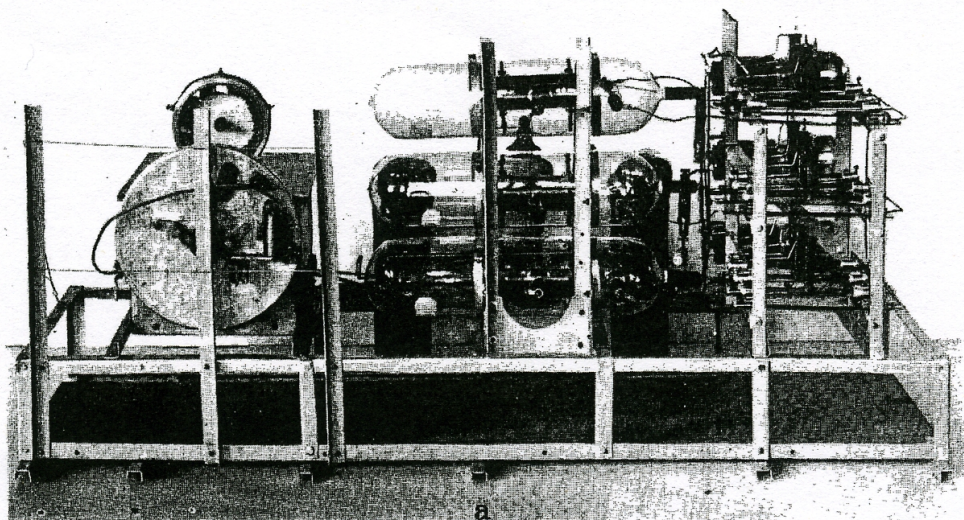


Fig. 5. Aufbau der Registrierapparatur. a) Von der Seite b) von oben gesehen.



Fig. 5. Launching of a balloon train from the courtyard of the institute.

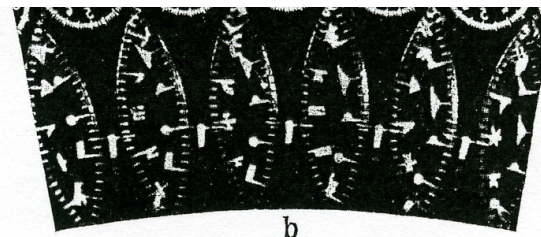


Fig. 4. a) Aufstiegplatte (nat. Größe, Hälfte); b) Vergrößerter Ausschnitt.

3) Die Firma Gebr. Junghans, Schramberg, hat uns freundlicherweise diese schönen Zählwerke hergestellt.

G. Pfozter, Z. f. Phys. 102 (1936) 23

Dreifachkoinzidenzen der Ultrastrahlung aus vertikaler Richtung in der Stratosphäre *).

I. Meßmethode und Ergebnisse.

Von Georg Pfozter in Stuttgart.

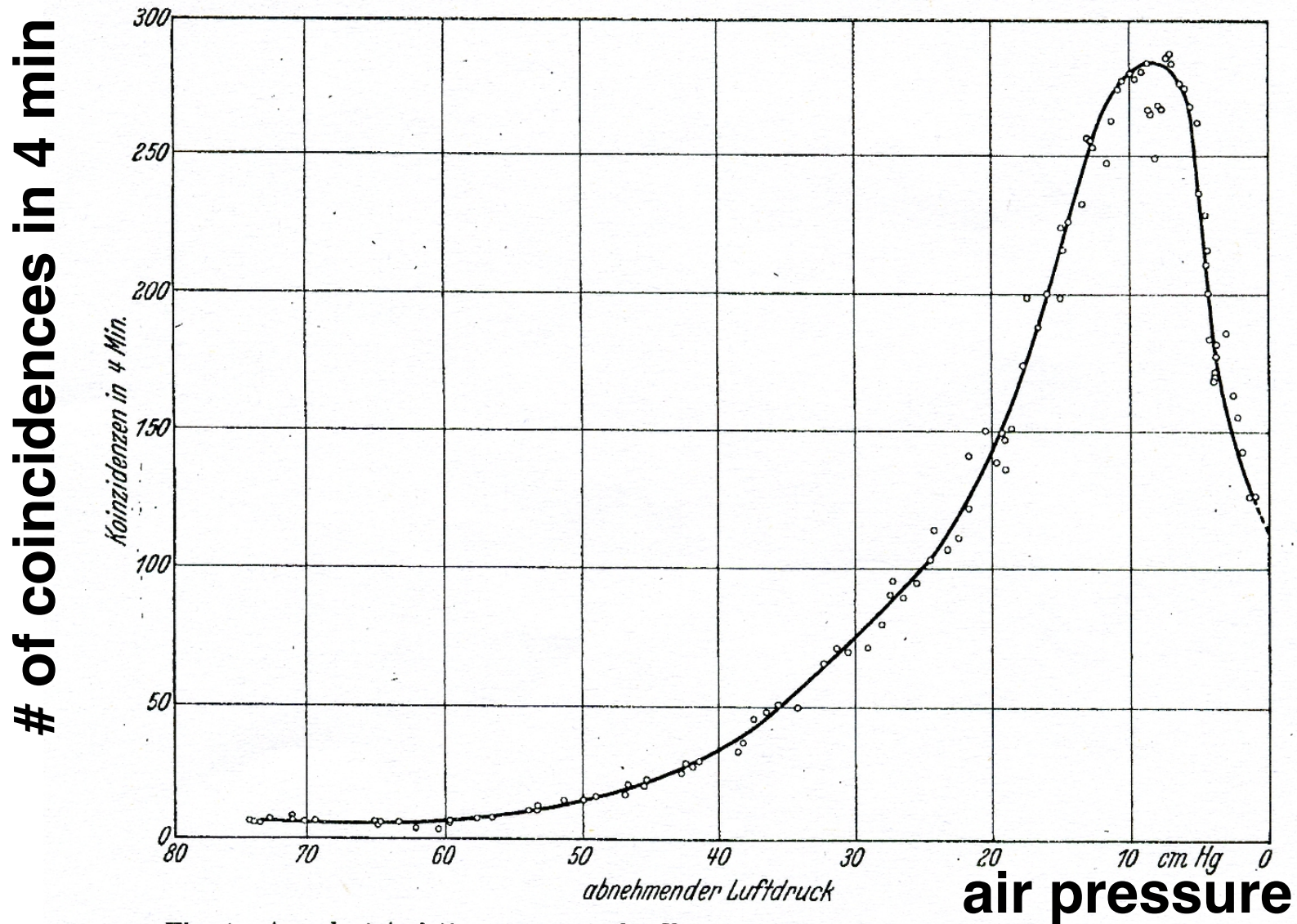


Fig. 1. Aus drei Aufstiegen gemittelte Kurve der Vertikalintensität der Ultrastrahlung in der Atmosphäre.

Dreifachkoinzidenzen der Ultrastrahlung aus vertikaler Richtung in der Stratosphäre *).

I. Meßmethode und Ergebnisse.

Von Georg Pfozter in Stuttgart.

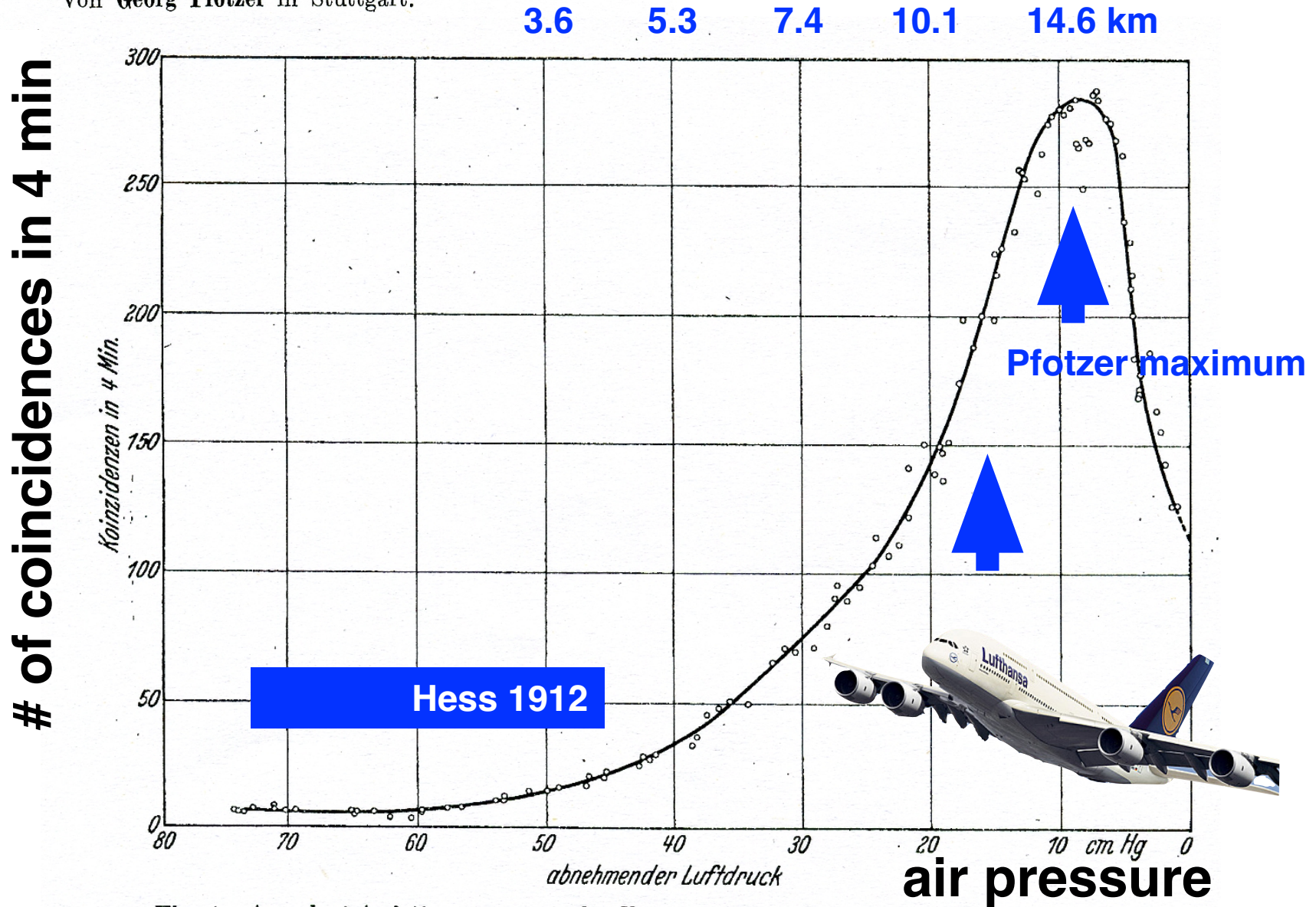


Fig. 1. Aus drei Aufstiegen gemittelte Kurve der Vertikalintensität der Ultrastrahlung in der Atmosphäre.

Letters to the Editor

[The Editor does not hold himself responsible for opinions expressed by his correspondents. Neither can he undertake to return, nor to correspond with the writers of, rejected manuscripts intended for this or any other part of NATURE. No notice is taken of anonymous communications.]

Latitude Effect of Cosmic Radiation

ON the expedition organised by the Deutscher und Oesterreichischer Alpenverein in 1932 to the Andes of Peru, observations of cosmic rays were made at several heights up to 6,100m. and during the sea-voyage. From Bremen to Peru one apparatus worked during March and April 1932 on board the M.S. Erfurt of the Norddeutscher Lloyd line. On the return voyage in January and February 1933, three apparatuses were in full action from Peru through the Strait of Magellan to Hamburg on board the M.S. Isis of the Hamburg-Amerika line. The self-recording electrometers were constructed by Prof. E. Regener on the same principle as those used for his researches in Lake Constance¹ and in the upper atmosphere². The electrometer wire is inside an ionisation chamber of 16 cm. diameter with 'deltametal' walls of 1 cm. thickness. The position of the wire is photographed every half-hour on a fixed photographic plate.

Instrument No. 1 was filled with carbon dioxide at 9.7 atmospheres pressure and 16° C. With a radium capsule, I found the temperature effect on ionisation to be + 0.13 per cent for every + 1° C. difference. The correction for barometric pressure was 0.29 per cent per millimetre of mercury. All data were reduced to 16° C. and 760 mm. pressure. The ionisation due to radioactivity in the chamber itself was allowed for as 0.8 volts per hour as found on the bottom of Lake Constance at a depth of 250 m. Eight hemispherical shells of iron were fitted round the chamber. The combined thickness of this iron wall was 10 cm.

In Fig. 1 are recorded the data of apparatus No. 1, the iron case of which was open on the upper side. The graph shows the intensity of cosmic radiation in volts per hour for different geomagnetic latitudes on the voyage from the Strait of Magellan to Hamburg. The geographical position of the geomagnetic north pole was taken to be 78° 32' N. and 69° 08' W. Each point of the curve corresponds to an average of a twenty hours' registration. The points give a smooth curve which shows the accuracy of the recording method employed. The intensity increases by about 12 per cent when going from the equatorial region to 55° N. geomagnetic latitude.

Apparatus No. 2 was wholly encased in the iron shell. Apparatus No. 3 worked without any iron shell. Every instrument shows substantially the same effect.

In general, the curves agree with the observations of Clay³ and with those of A. H. Compton⁴ made at about the same time. It is very interesting that the northern and southern parts of the curve are not

symmetrical with respect to either the geomagnetic or the geographical equator. Considering the accuracy of our uninterrupted registration, this result is quite trustworthy.

From the fact that a latitude effect of 12 per cent of the radiation exists, it must be concluded that this part of the radiation consists of corpuscles before entering the earth's atmosphere. For the magnitude of this part of the radiation, reference should be made to the analysis of the components of cosmic rays by Regener² and Lenz⁵.

A more detailed report of these observations and of the researches in the Andes will be published in the *Zeitschrift für Physik*.

H. HOERLIN.

Physikalisches Institut
der Technischen Hochschule,
Stuttgart. June 8.

- ¹ Regener, E., *Z. Phys.*, **74**, 433; 1932.
- ² Regener, E., *Phys. Z.*, **34**, 306; 1933.
- ³ Clay, J., *Naturwiss.*, **20**, 687; 1932.
- ⁴ Compton, A. H., *Phys. Rev.*, **43**, 387; 1933.
- ⁵ Lenz, E., *Z. Phys.*; in the press.

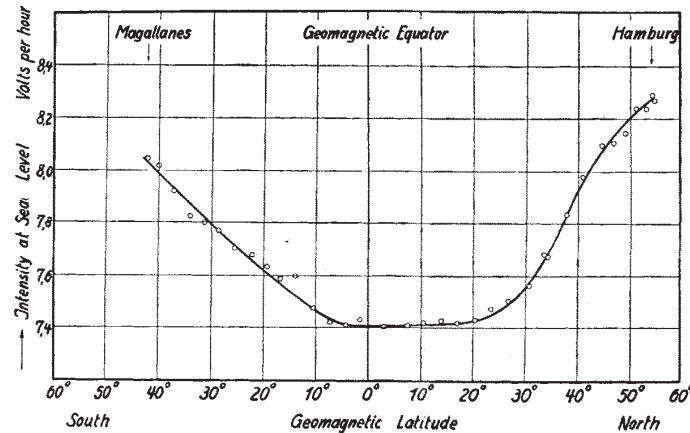


FIG. 1.

Latitude effect

Latitude effect

Letters to the Editor

[The Editor does not hold himself responsible for opinions expressed by his correspondents. Neither can he undertake to return, nor to correspond with the writers of, rejected manuscripts intended for this or any other part of NATURE. No notice is taken of anonymous communications.]

Latitude Effect of Cosmic Radiation

ON the expedition organised by the Deutscher und Oesterreichischer Alpenverein in 1932 to the Andes of Peru, observations of cosmic rays were made at several heights up to 6,100m. and during the sea-voyage. From Bremen to Peru one apparatus worked during March and April 1932 on board the M.S. *Erfurt* of the Norddeutscher Lloyd line. On the return voyage in January and February 1933, three apparatuses were in full action from Peru through the Strait of Magellan to Hamburg on board the M.S. *Isis* of the Hamburg-Amerika line. The self-recording electrometers were constructed by Prof. E. Regener on the same principle as those used for his researches in Lake Constance¹ and in the upper atmosphere². The electrometer wire is inside an ionisation chamber of 16 cm. diameter with 'deltametal' walls of 1 cm. thickness. The position of the wire is photographed every half-hour on a fixed photographic plate.

Instrument No. 1 was filled with carbon dioxide at 9.7 atmospheres pressure and 16° C. With a radium capsule, I found the temperature effect on ionisation to be + 0.13 per cent for every + 1° C. difference. The correction for barometric pressure was 0.29 per cent per millimetre of mercury. All data were reduced to 16° C. and 760 mm. pressure. The ionisation due

symmetrical with respect to either the geographical or the geographical equator. Considering the accuracy of our uninterrupted registration, this result is quite trustworthy.

From the fact that a latitude effect of 1 per cent of the radiation exists, it must be concluded that this part of the radiation consists of corpuscles before entering the earth's atmosphere. For the magnitude of this part of the radiation, reference should be made to the analysis of the components of cosmic rays by Regener² and Lenz⁵.

A more detailed report of these observations and of the researches in the Andes will be published in the *Zeitschrift für Physik*.

H. HOERLIN.

Physikalisches Institut
der Technischen Hochschule,
Stuttgart. June 8.

- ¹ Regener, E., *Z. Phys.*, **74**, 433; 1932.
- ² Regener, E., *Phys. Z.*, **34**, 306; 1933.
- ³ Clay, J., *Naturwiss.*, **20**, 687; 1932.
- ⁴ Compton, A. H., *Phys. Rev.*, **43**, 387; 1933.
- ⁵ Lenz, E., *Z. Phys.*; in the press.

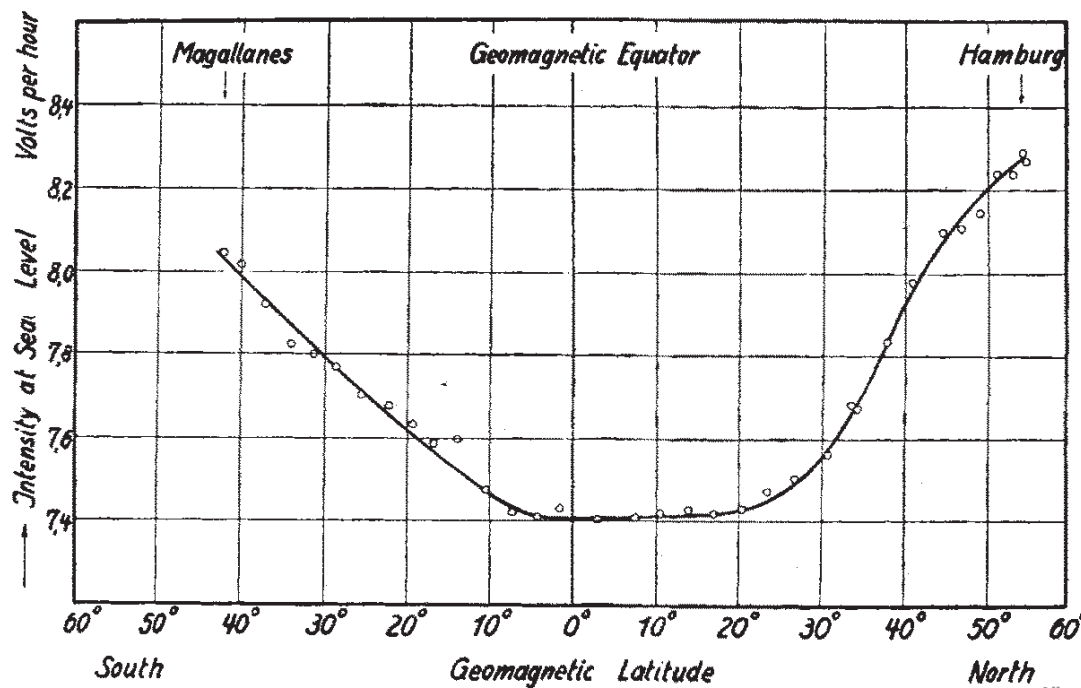


FIG. 1.

Clay: Latitude Effect

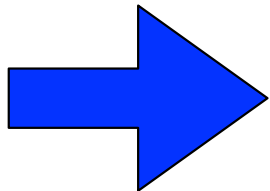
RESULTS OF THE DUTCH COSMIC RAY EXPEDITION 1933

II. THE MAGNETIC LATITUDE EFFECT OF COSMIC RAYS
A MAGNETIC LONGITUDE EFFECT

by J. CLAY, P. M. VAN ALPHEN and C. G. 'T HOOFT

Natuurkundig Laboratorium, Amsterdam

journey from Holland to Java
intensity varies with latitude



cosmic rays are
charged particles

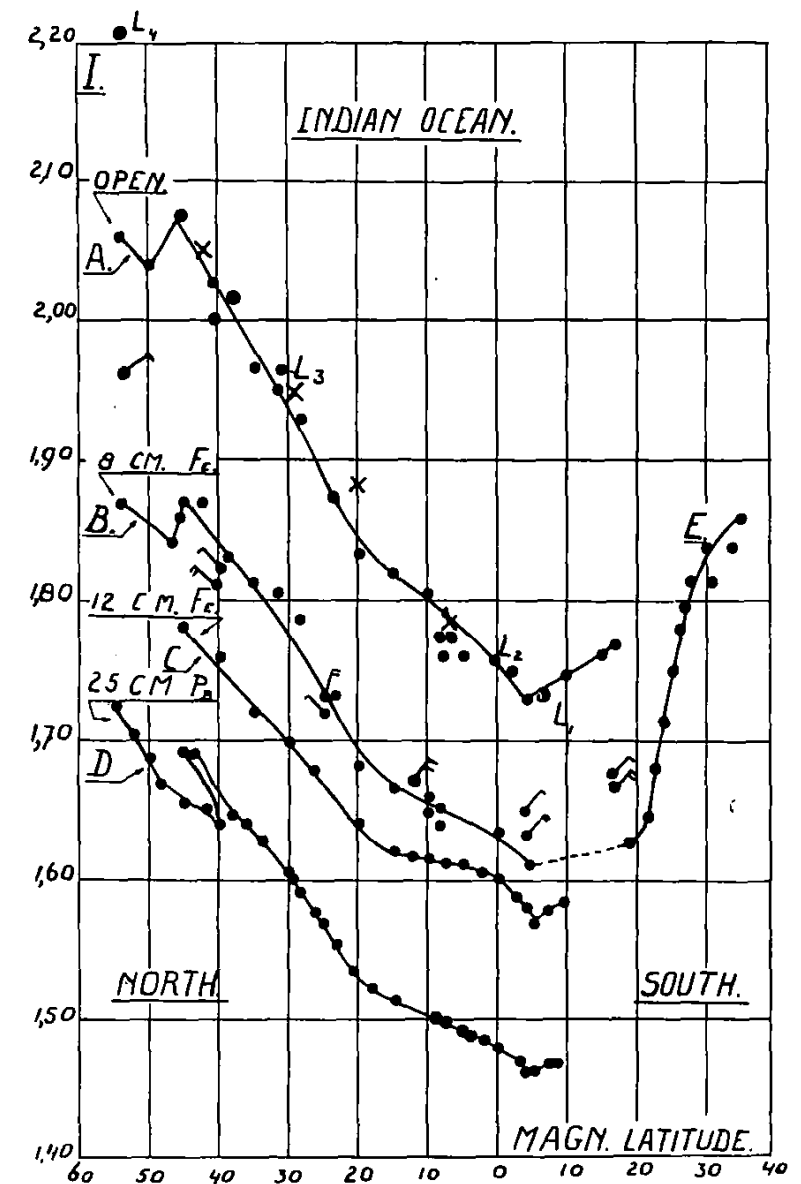


Fig. 1. Records of the variation of Cosmic Radiation with latitude on two different routes under different shielding with different instruments

x x results with instrument *D* open (Amsterdam—Batavia)
(*L*₁, *L*₂, *L*₃, *L*₄) results with instrument *D*₁ open (Batavia—Amsterdam)
! F Results 1928 and 1929.

Compton: World-wide survey of intensity of radiation

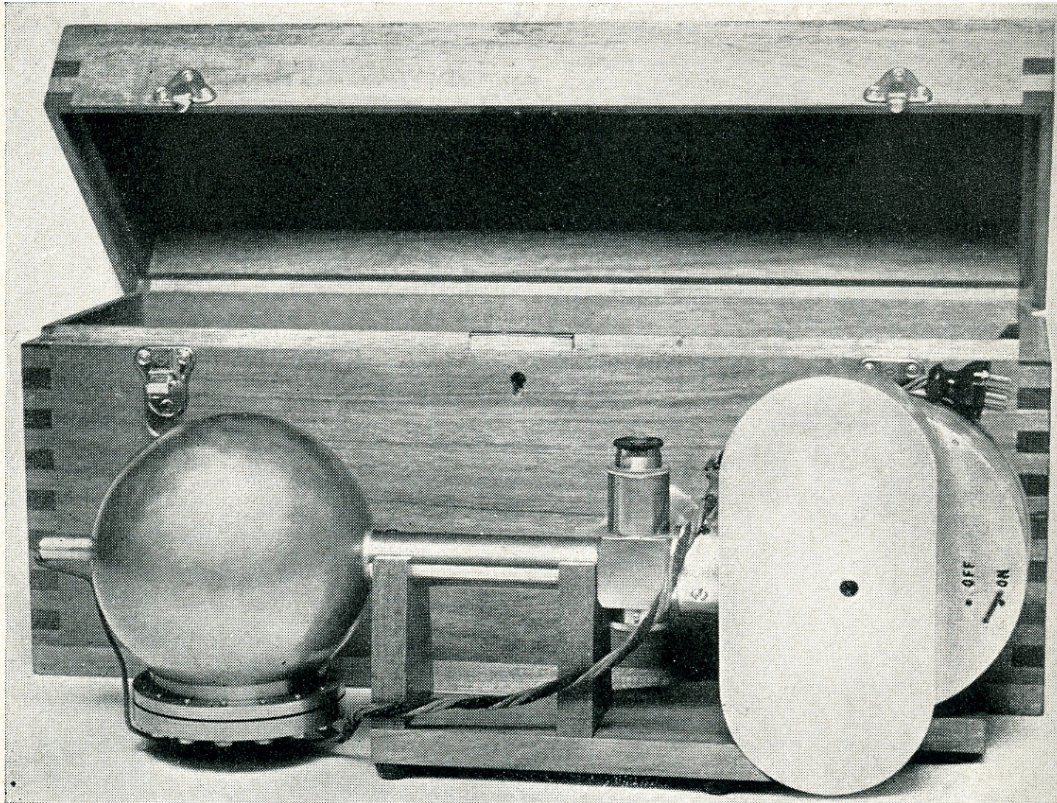


Fig. 24. The instrument used in this survey is usually shielded with lead and is placed in the box when used in most airplane flights.

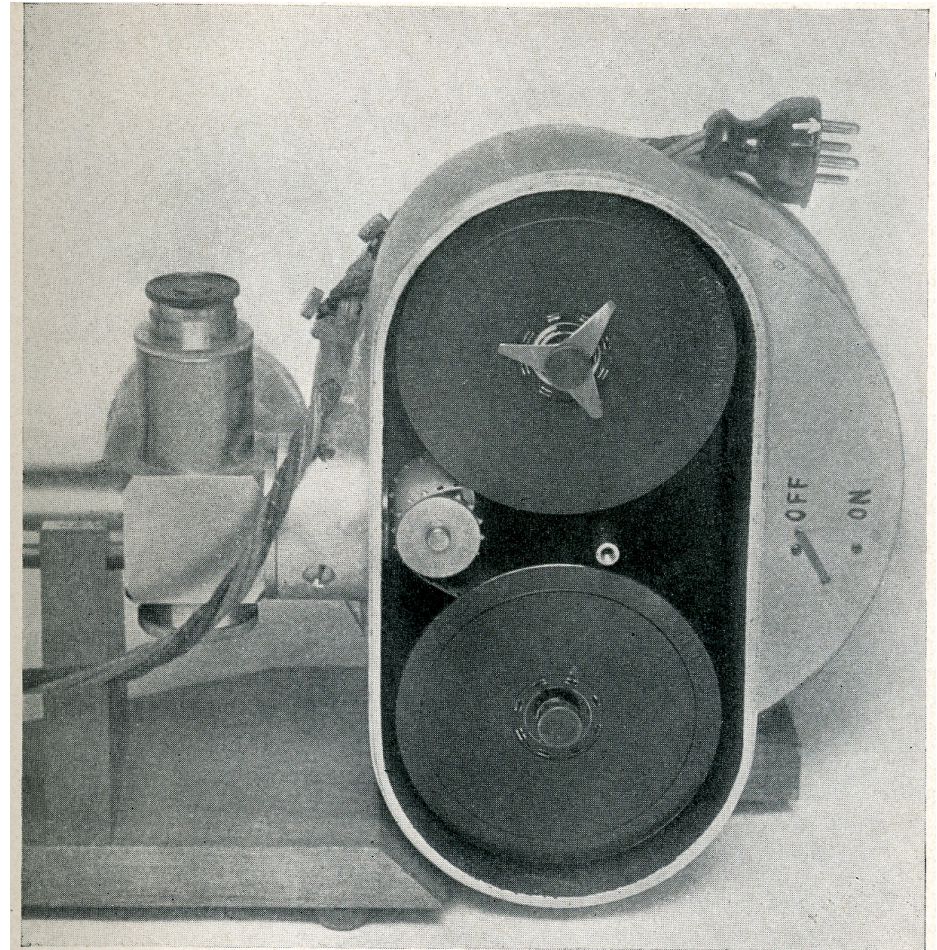


Fig. 25. The camera will take a one-hundred-foot reel of 35 mm motion picture film which is driven at a constant rate past the slit by a power clock. Changeable gears allow various rates of film speeds to be used, depending on the expected ionization.

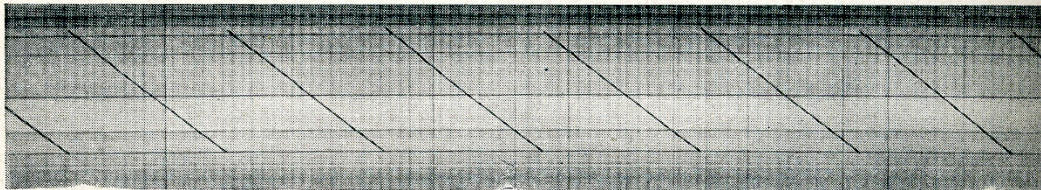


Fig. 27. Showing the type of record obtained at sea level in this world survey. Two of the horizontal lines give barometric and temperature records.

~1930

THE PHYSICAL REVIEW

A Journal of Experimental and Theoretical Physics

VOL. 43, No. 6

MARCH 15, 1933

SECOND SERIES

A Geographic Study of Cosmic Rays

ARTHUR H. COMPTON, *University of Chicago*

(Received January 30, 1933)

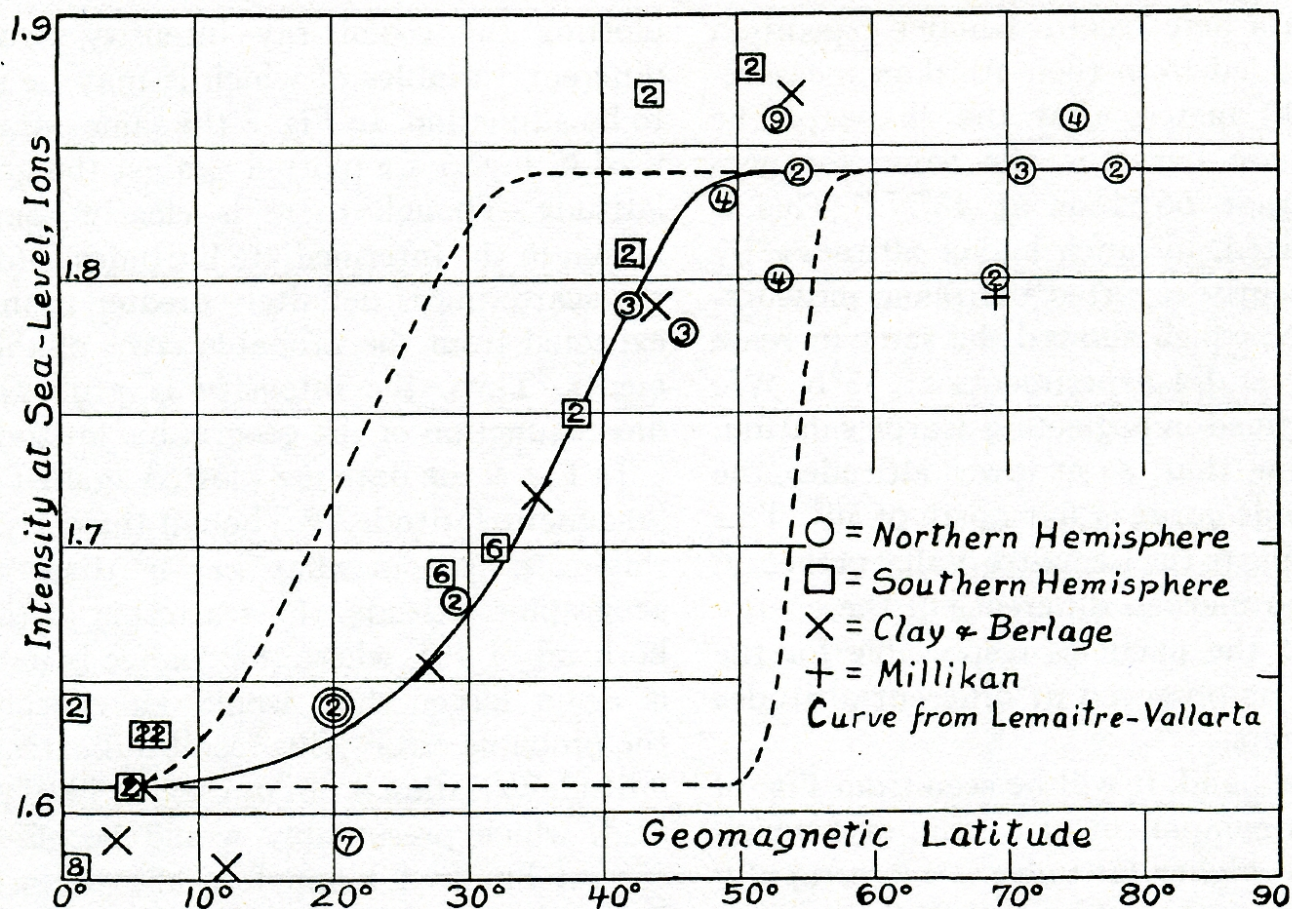


FIG. 7. Intensity vs. geomagnetic latitude at sea level, including data of Clay and Millikan.

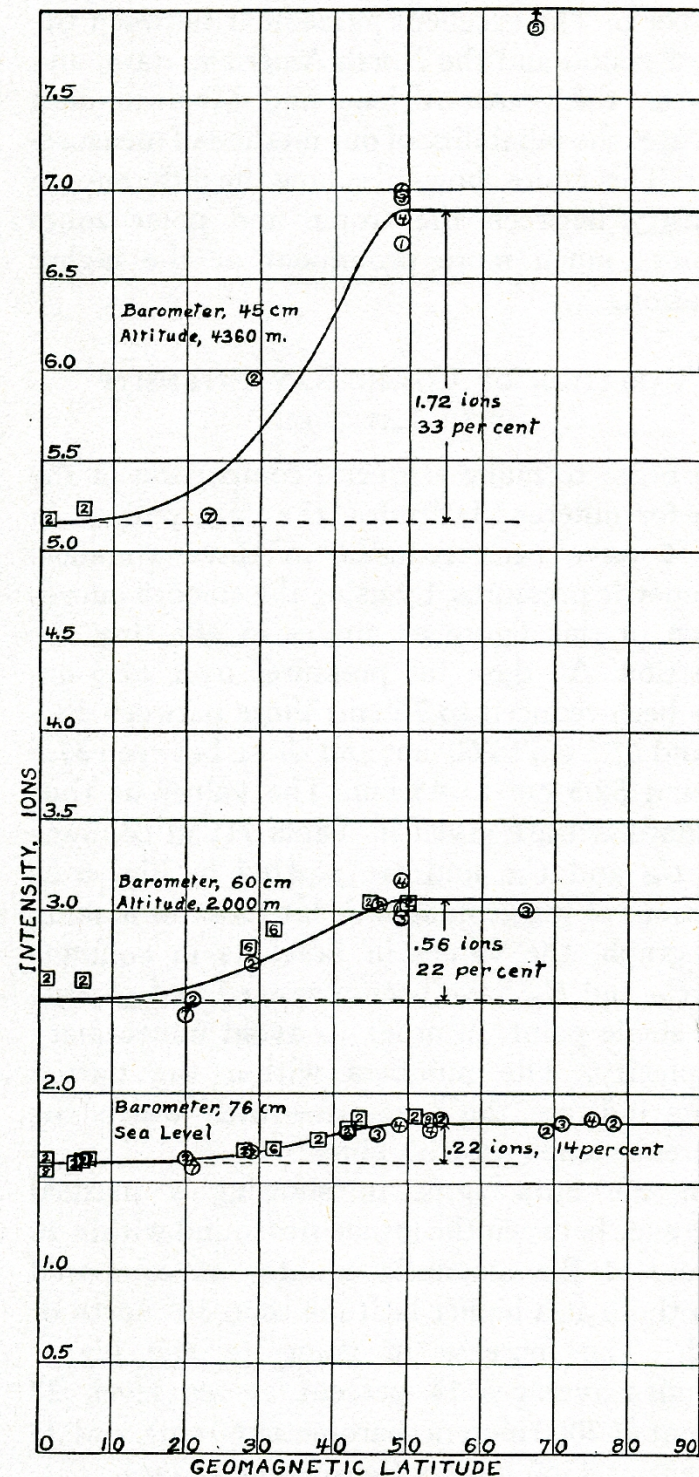


FIG. 6. Intensity vs. geomagnetic latitude for different elevations.

1931-34 A.H. Compton 12 expeditions → ~100 locations

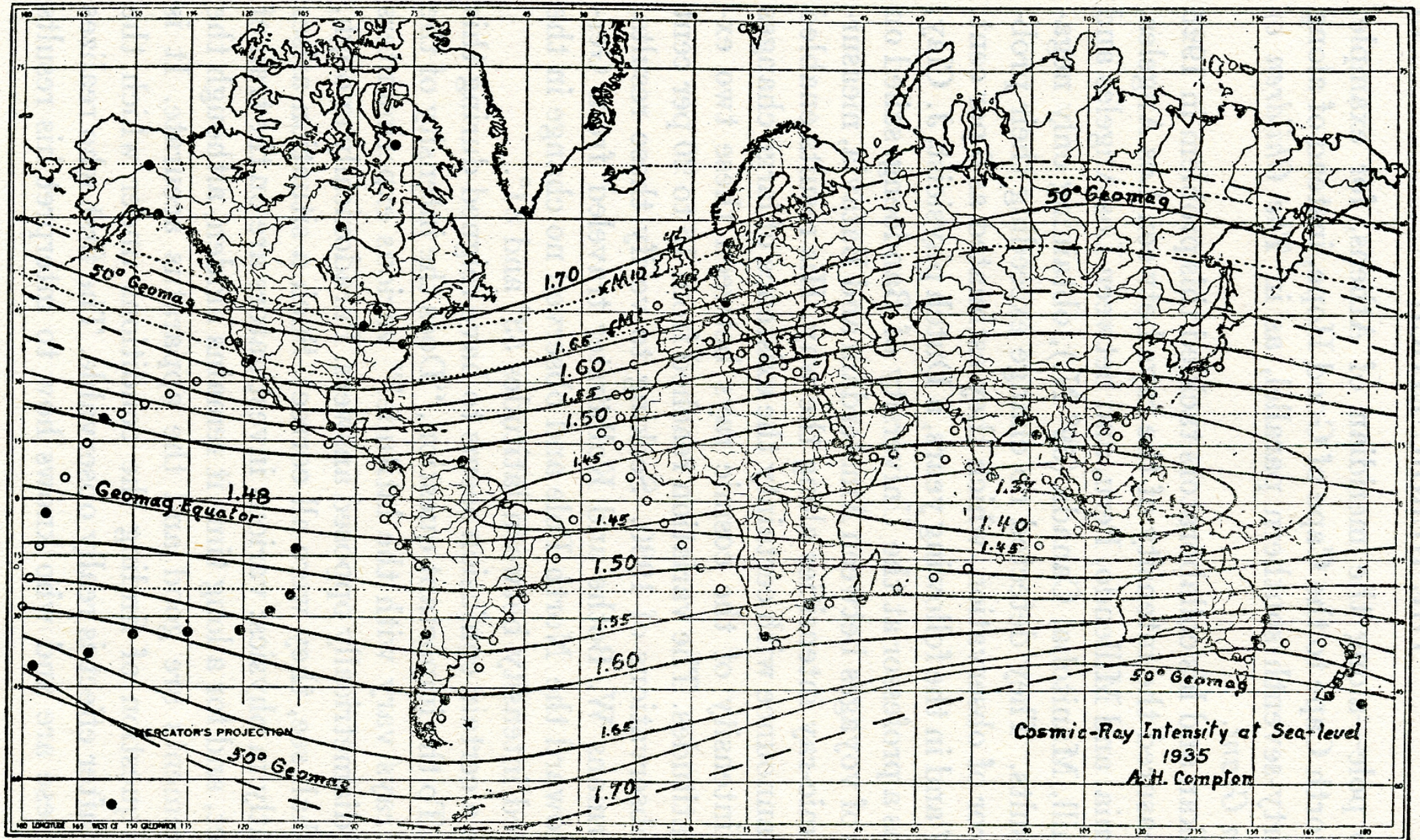


FIG. 6.—Compton's world map of isocosms. Note the parallelism of these lines of equal cosmic-ray intensity and the dotted curves of geomagnetic latitude (50° N. and S.).

➡ cosmic rays are charged particles

1931-34 A.H. Compton 12 expeditions → ~100 locations

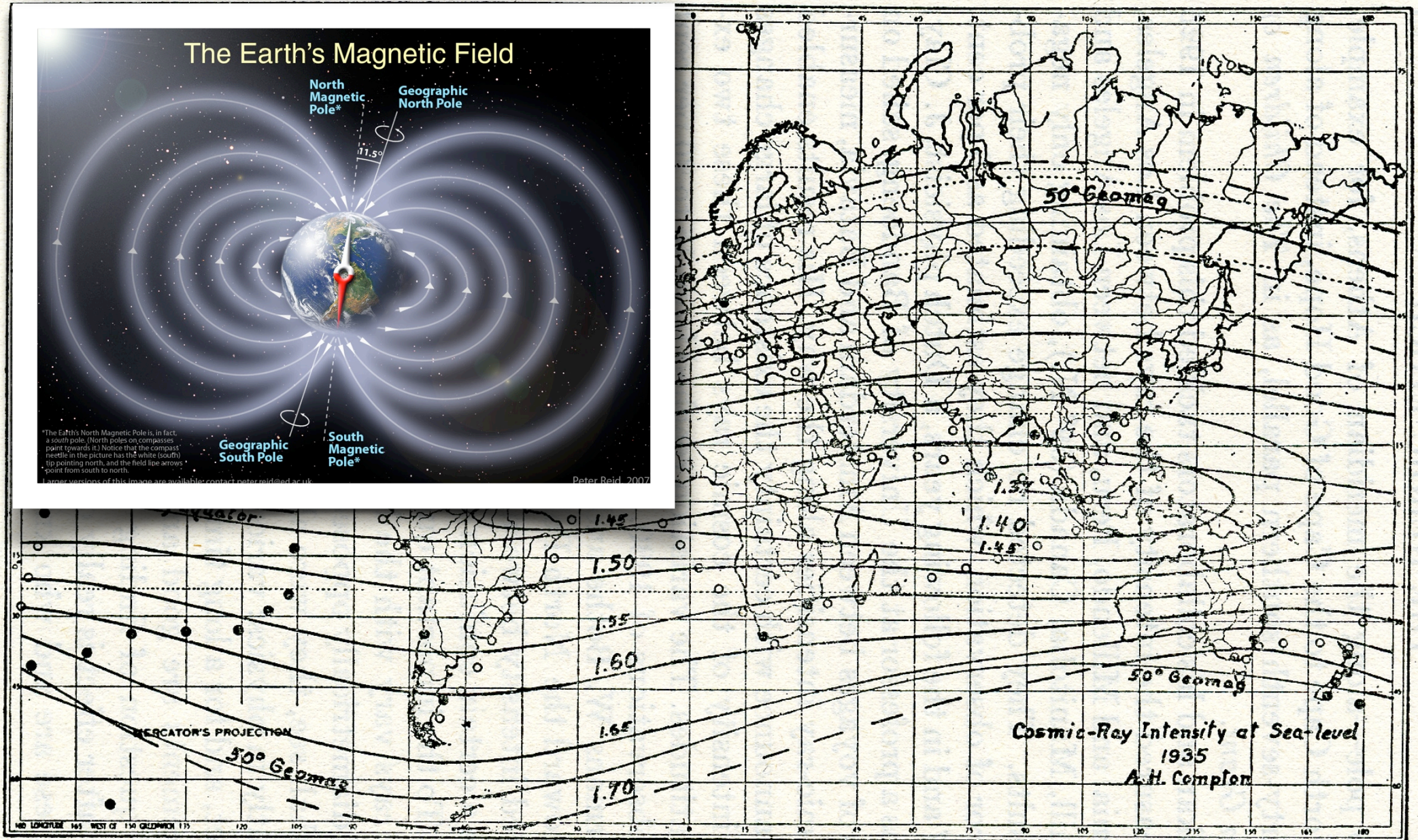


FIG. 6.—Compton's world map of isocosms. Note the parallelism of these lines of equal cosmic-ray intensity and the dotted curves of geomagnetic latitude (50° N. and S.).

➡ cosmic rays are charged particles

~1937 East-West Effect of Cosmic-Ray Intensity

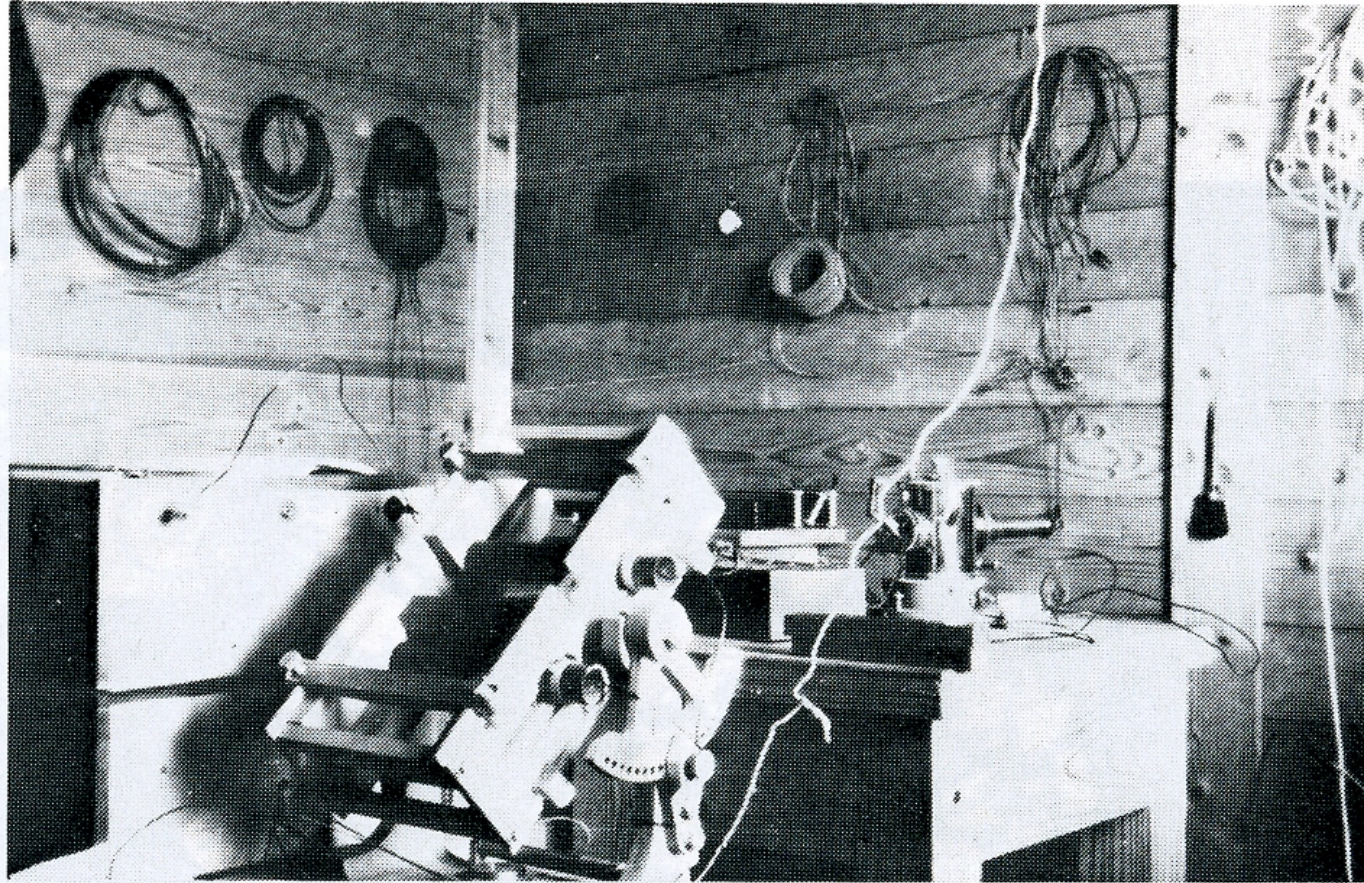


Fig. 14. The equipment for the E-W experiment.

Rossi and others

higher intensity from the west

➔ cosmic rays are mostly positively charged

~1930 „elementary particles“: charged neutral

Rutherford (1919)

p

n

(1932)

Chadwick



Thomson (1897)

e⁻

γ

(1905/26)

Einstein



~1930 „elementary particles“: charged neutral

Rutherford (1919) **p** **n** (1932) Chadwick

Thomson (1897) **e⁻** **γ** (1905/26) Einstein



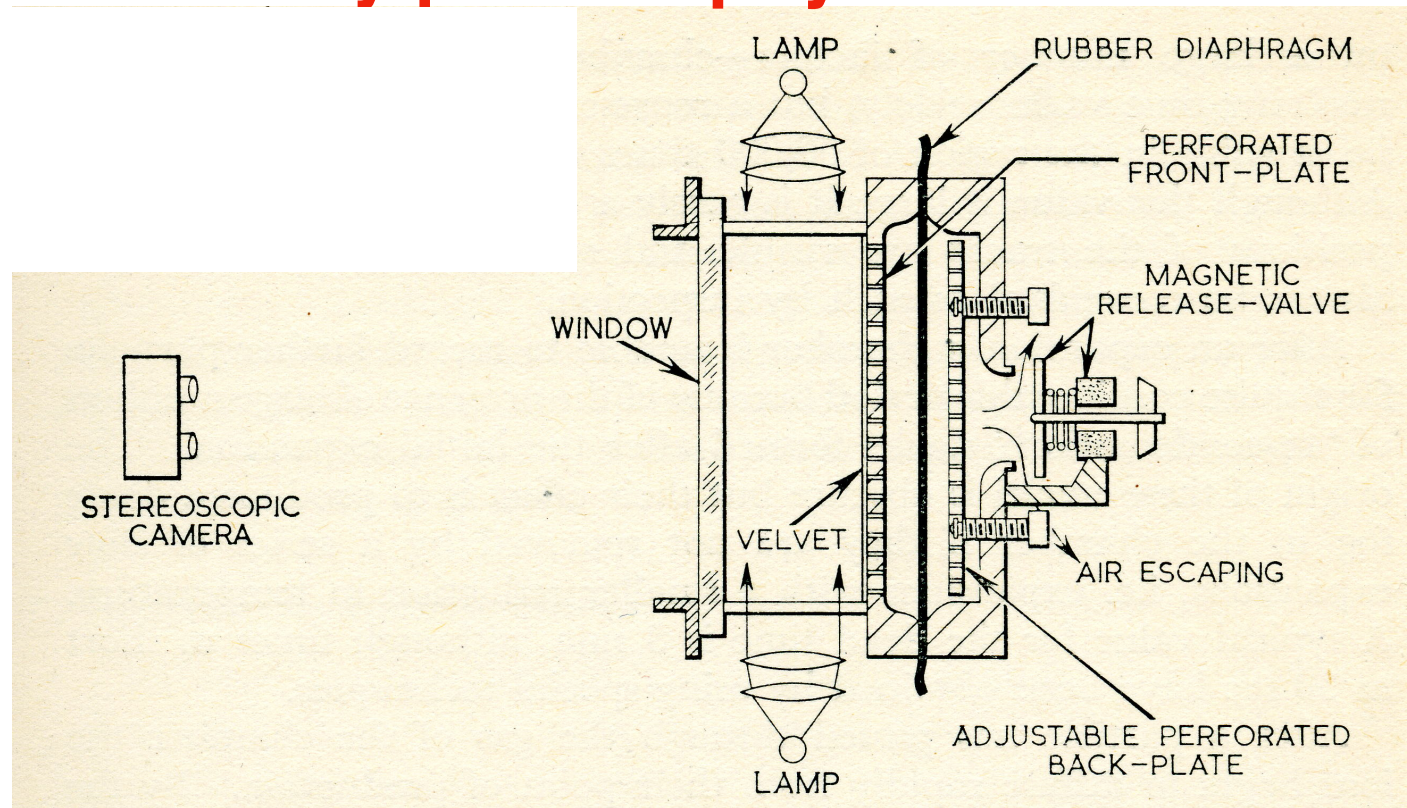
Discovery of new particles in cosmic rays

~1930 – 1950

birth of elementary particle physics



cloud chamber
C.T.R. Wilson
Nobel Prize 1927



The Positive Electron e^+

CARL D. ANDERSON, *California Institute of Technology, Pasadena, California*
(Received February 28, 1933)

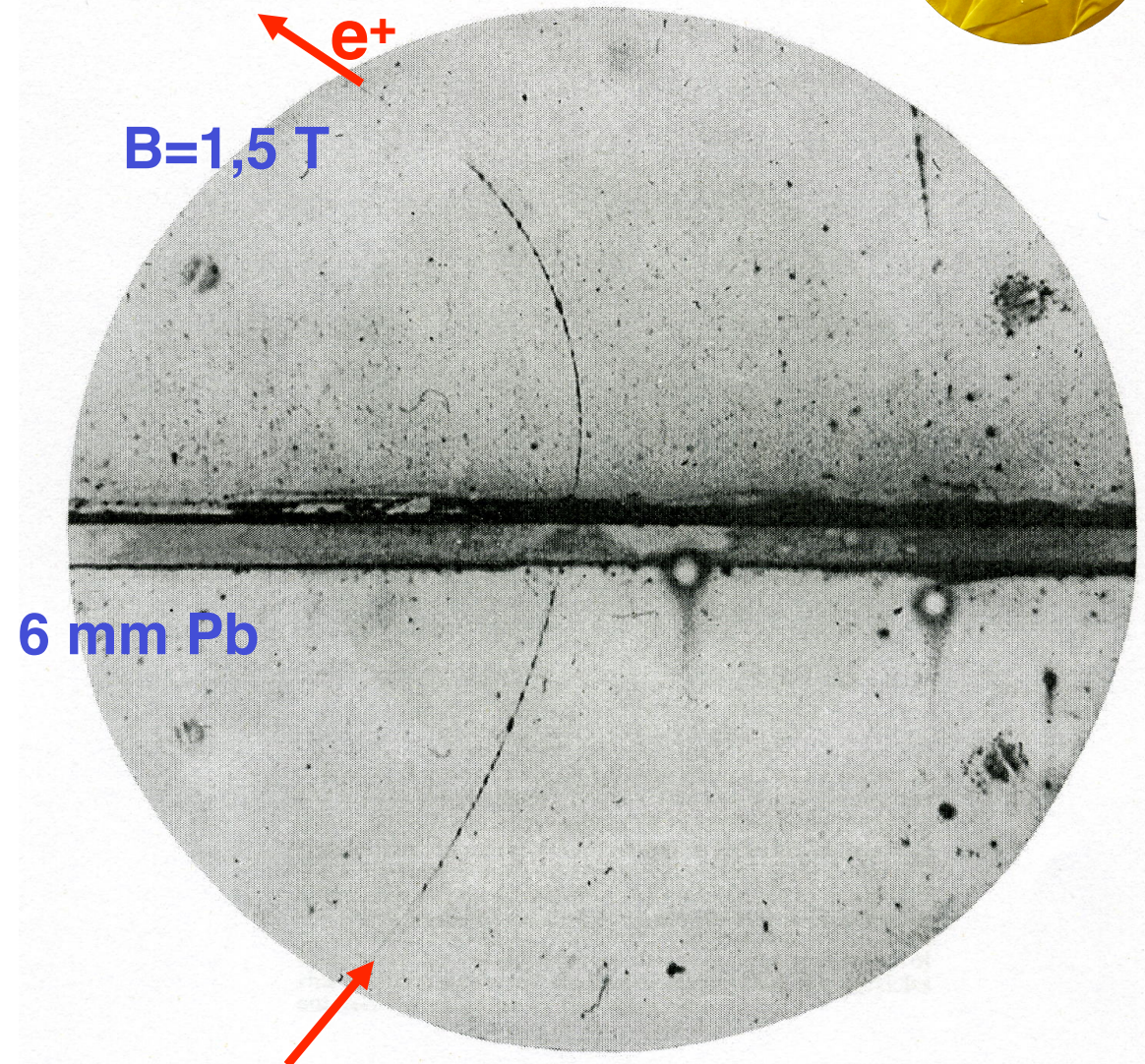
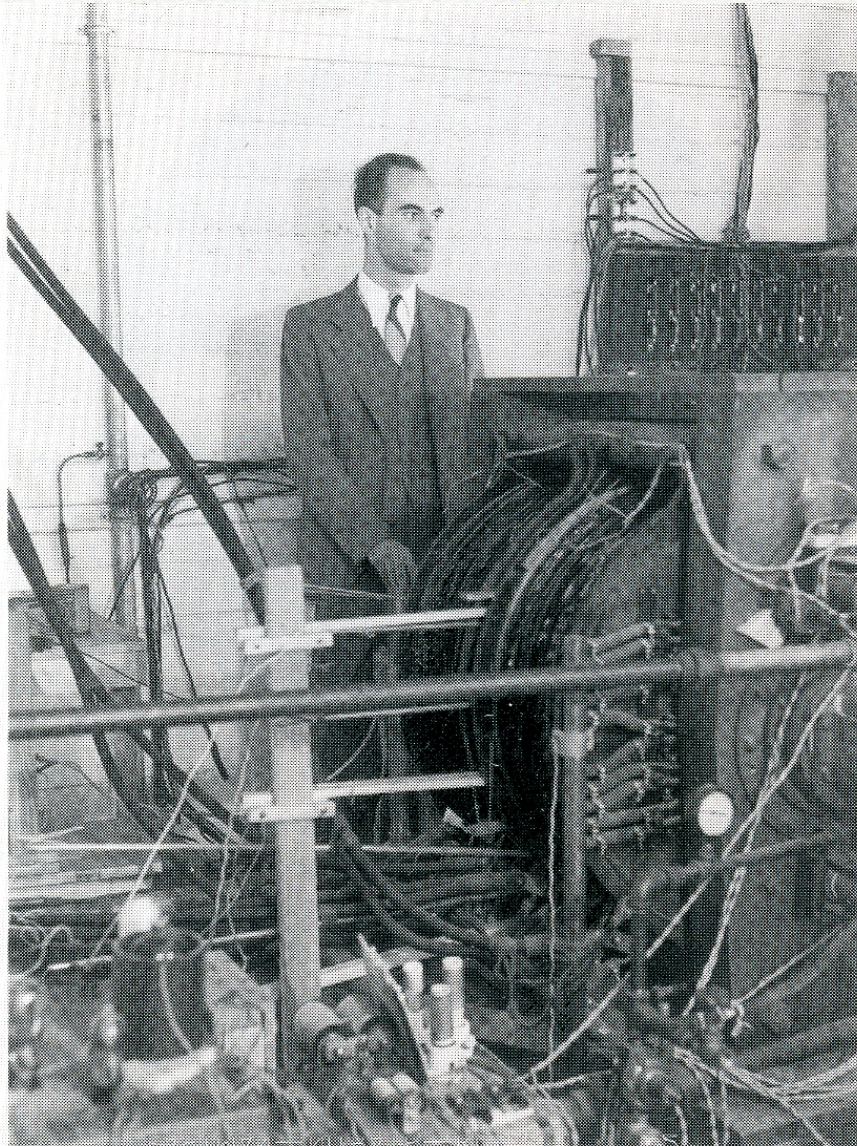
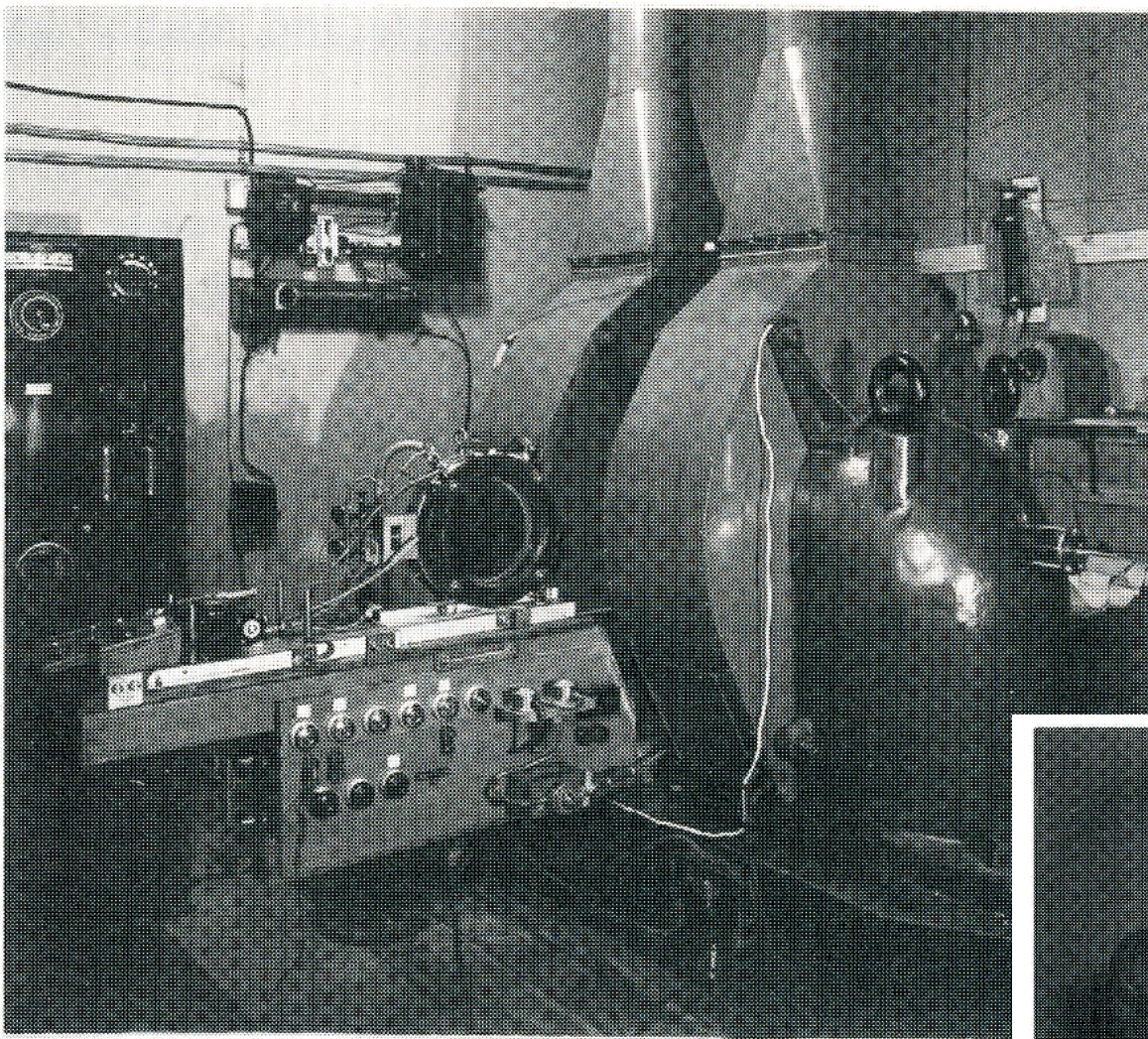


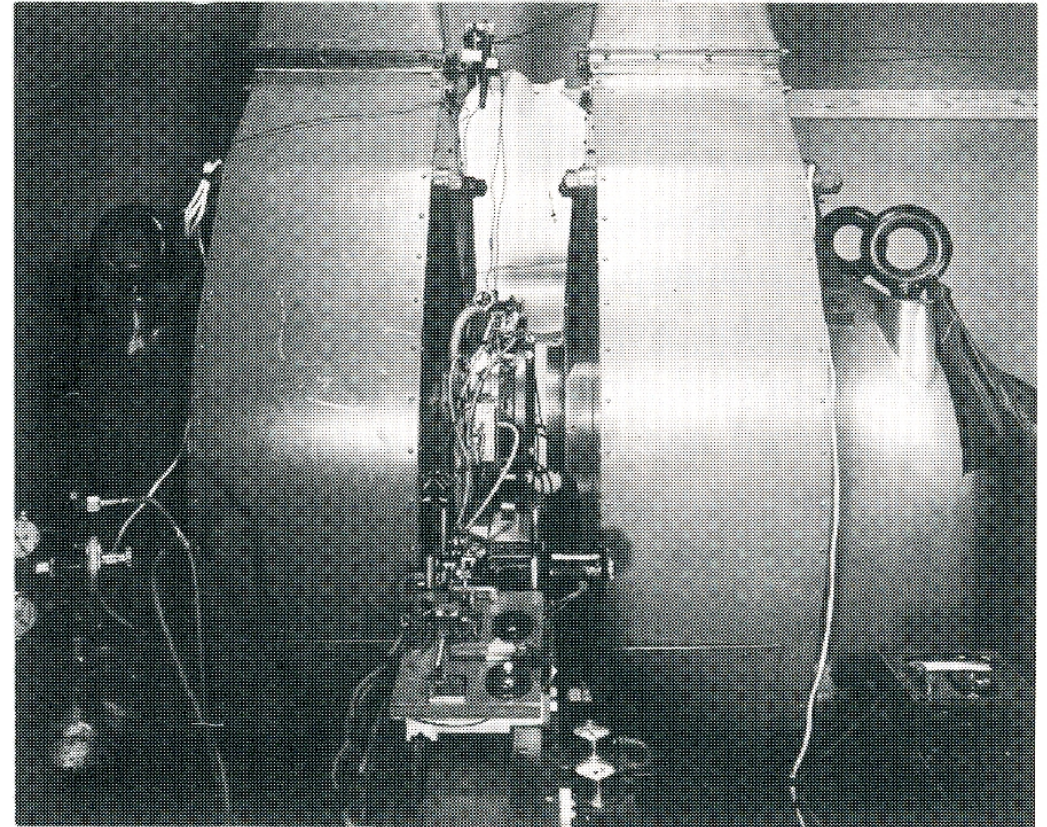
FIG. 1. A 63 million volt positron ($H_p = 2.1 \times 10^5$ gauss-cm) passing through a 6 mm lead plate and emerging as a 23 million volt positron ($H_p = 7.5 \times 10^4$ gauss-cm). The length of this latter path is at least ten times greater than the possible length of a proton path of this curvature.

P.M.S. Blackett
Nobel Prize 1948

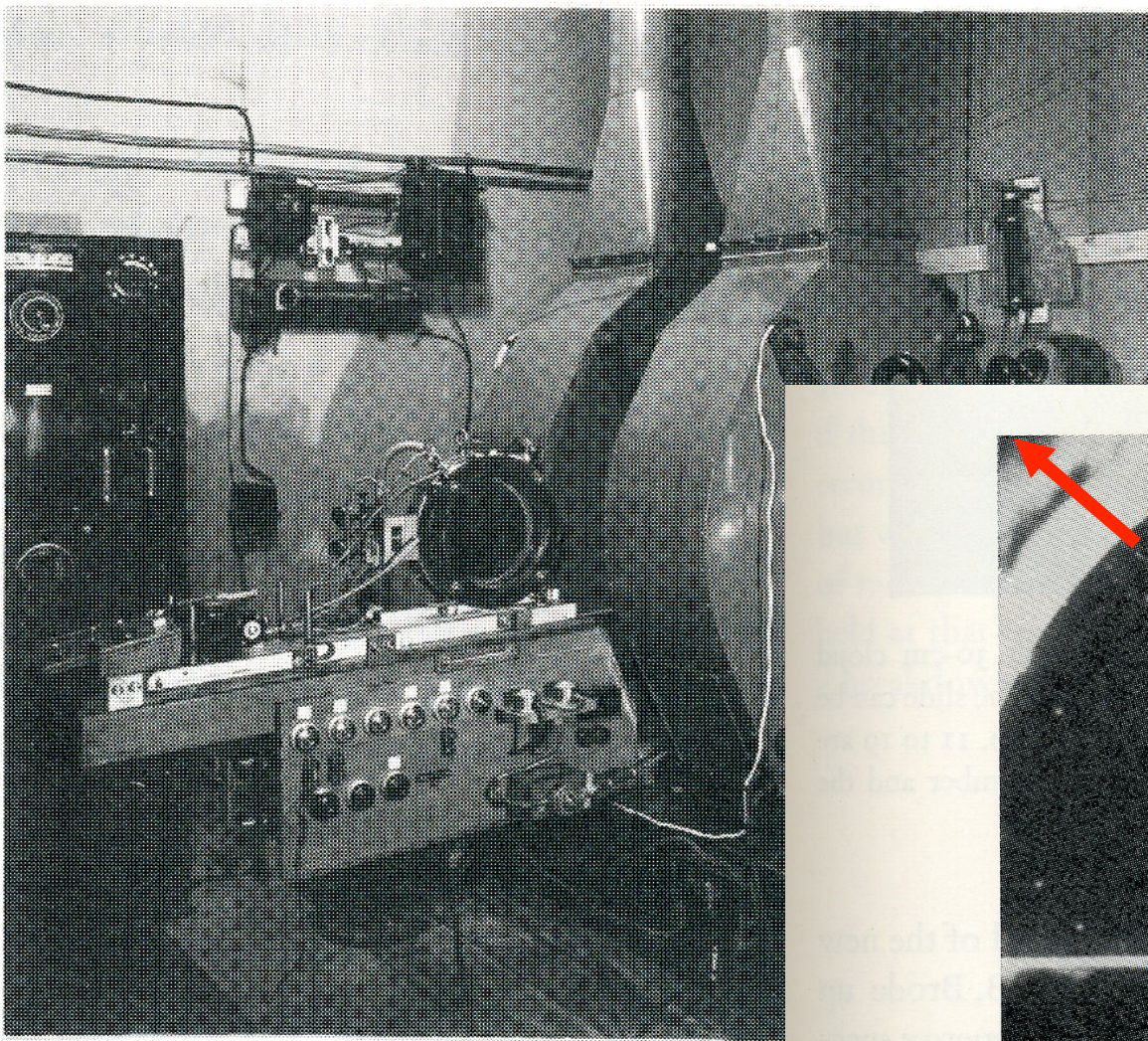


1933 Blackett & Occhialini

10 t electromagnet
30 cm cloud chamber



P.M.S. Blackett
Nobel Prize 1948



1933 Blackett & Occhialini

10 t electromagnet
30 cm cloud chamber

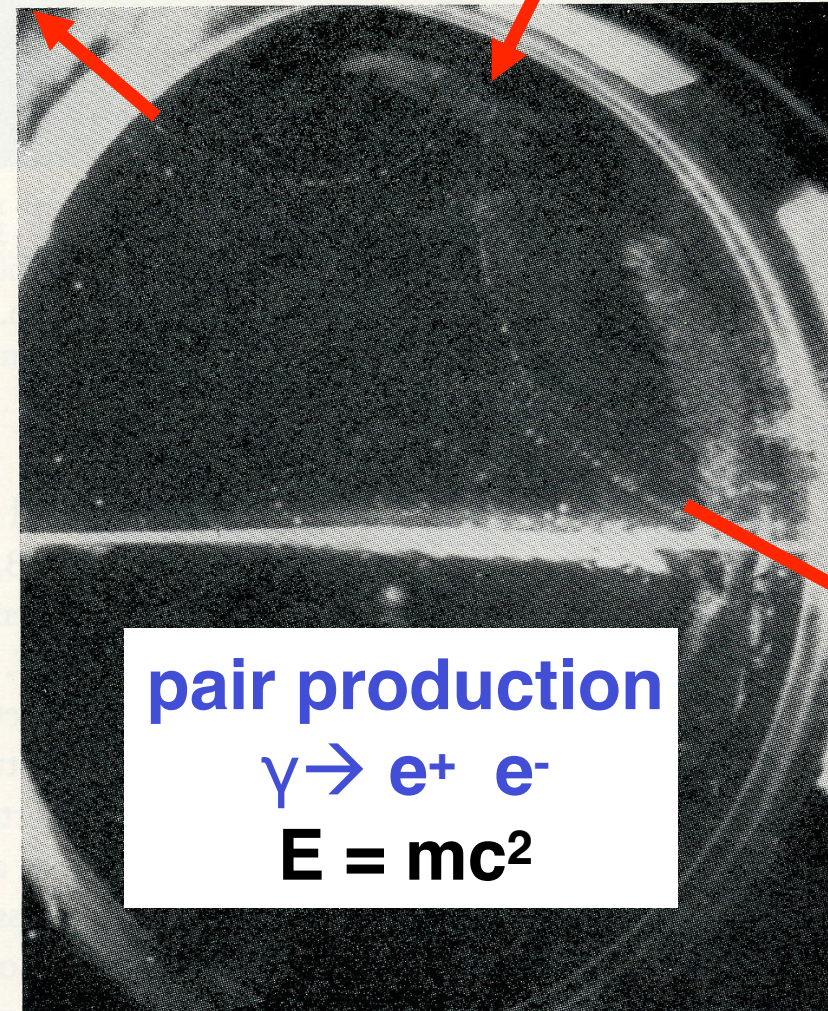


Fig. 9. Pair of positive and negative electrons produced by gamma rays. (Chadwick, Blackett, and Occhialini, 1934)

Electromagnetic Cascades B. Rossi 1933

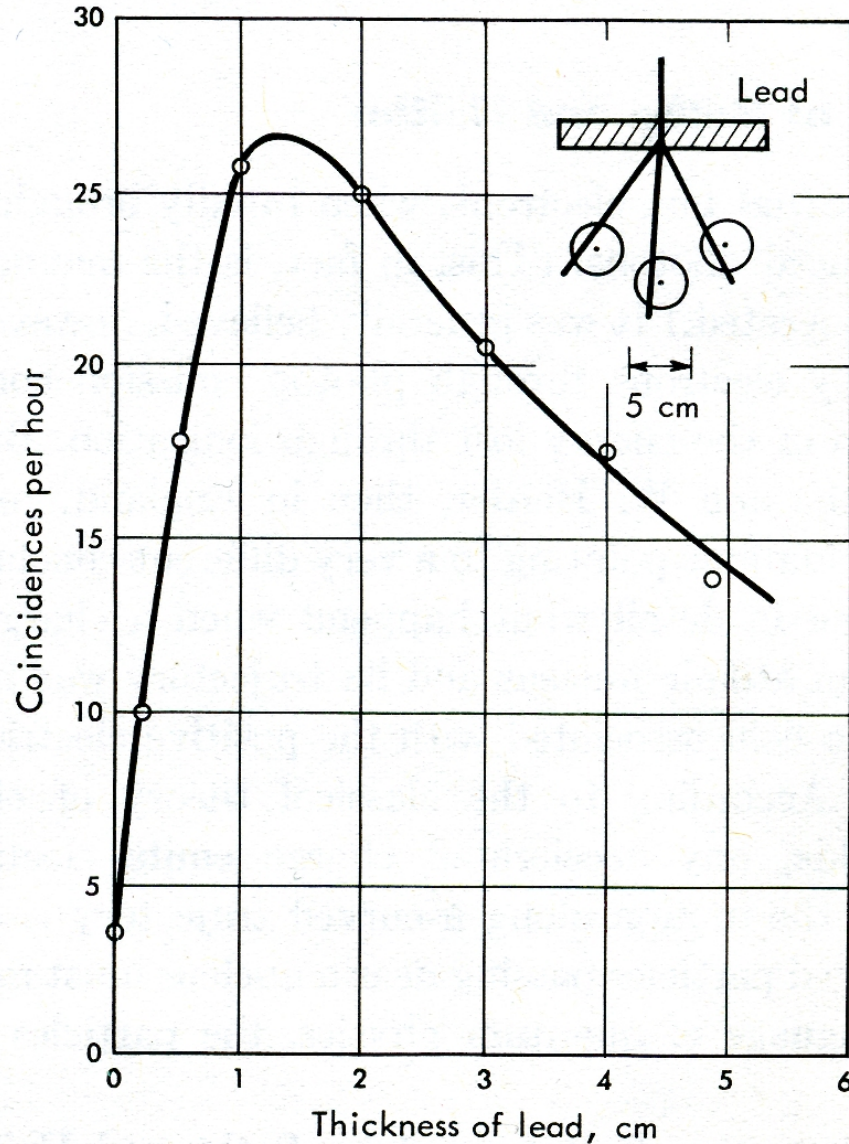


Fig. 7-1 Shower curve. The number of coincidences per hour is plotted as a function of the thickness of lead above the counters. The experimental arrangement is shown schematically in the inset. The circles are experimental points. (This figure is based on one appearing in a paper by the author in *Zeitschrift für Physik*, vol. 82, p. 151, 1933.)

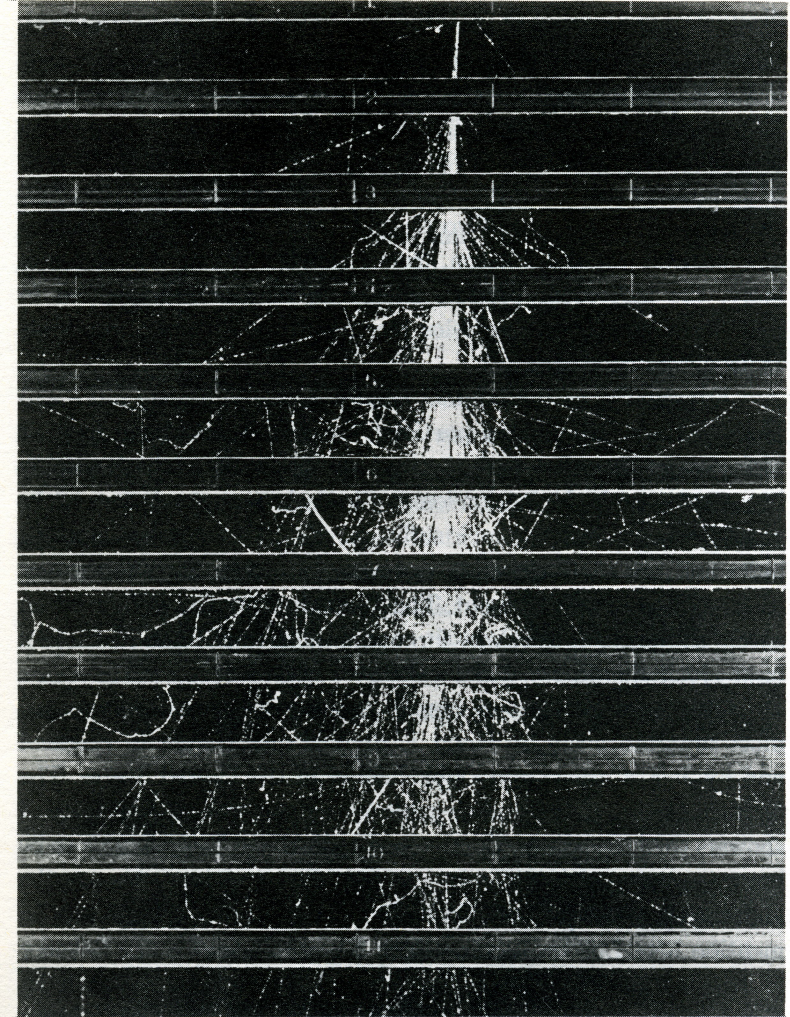
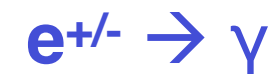


Fig. 7-5 A shower developing through a number of brass plates 1.25 cm thick placed across a cloud chamber. The shower was initiated in the top plate by an incident high-energy electron or photon. The photograph was taken by the MIT cosmic-ray group.



Discovery of the Muon

1937 Anderson & Neddermeyer: μ in cloud chamber

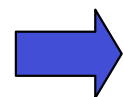
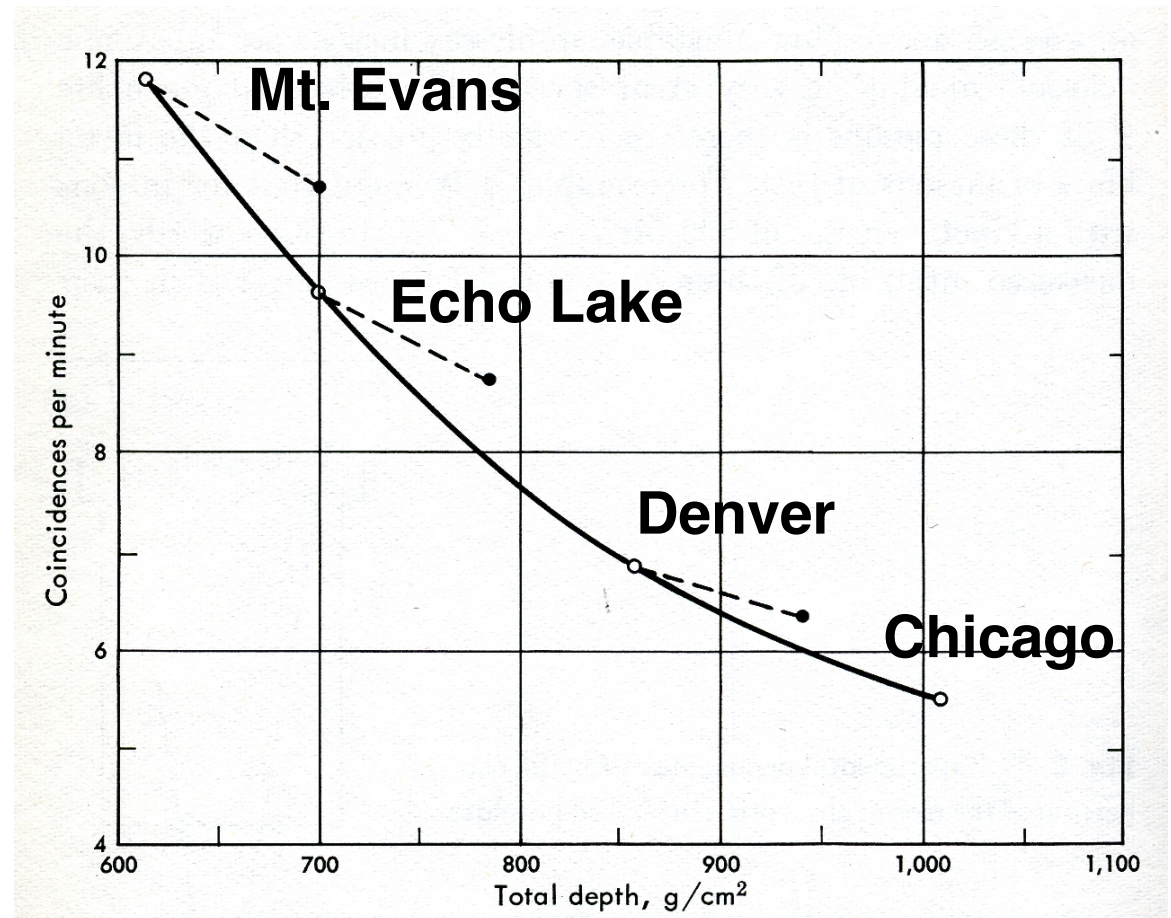
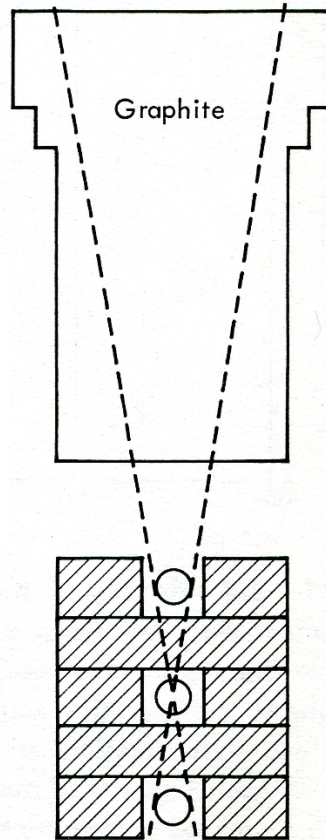
$$m_{\mu} \sim 200 m_e$$

Discovery of the Muon

1937 Anderson & Neddermeyer: μ in cloud chamber

$$m_{\mu} \sim 200 m_e$$

1939 B. Rossi: life time



life time $\tau \sim 2 \mu s$

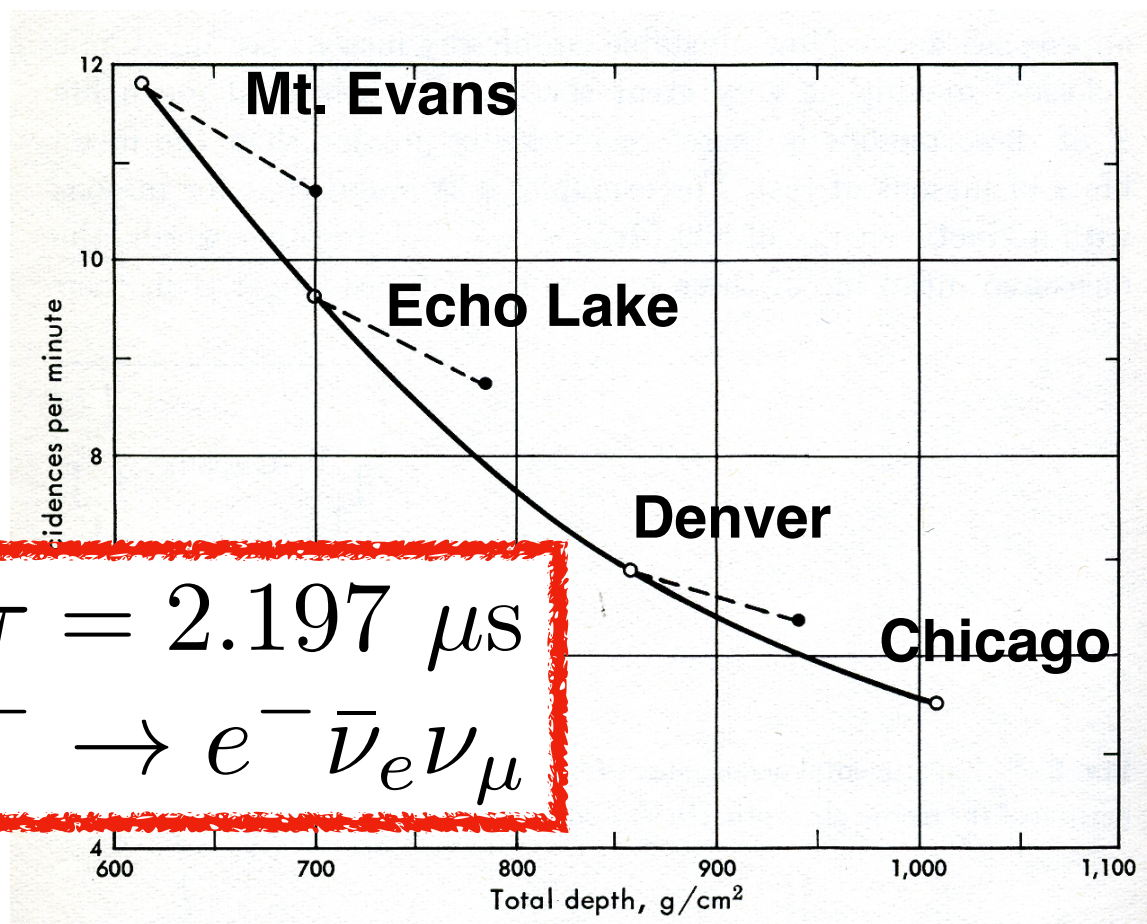
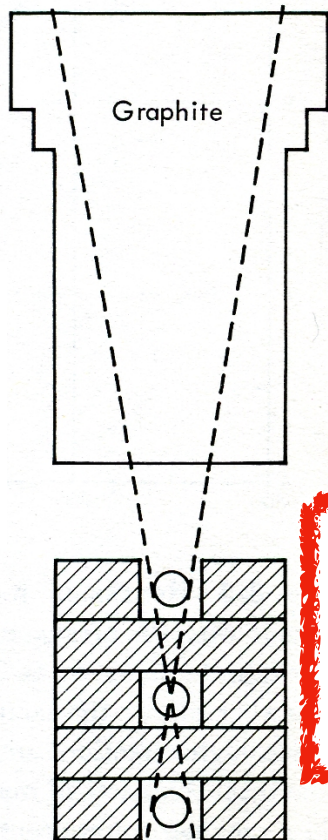
$\mu \rightarrow e + \dots$

Discovery of the Muon

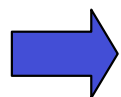
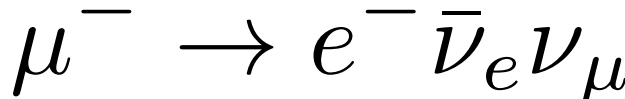
1937 Anderson & Neddermeyer: μ in cloud chamber

$$m_{\mu} \sim 200 m_e$$

1939 B. Rossi: life time



PDG: $\tau = 2.197 \mu\text{s}$



life time $\tau \sim 2 \mu\text{s}$

$\mu \rightarrow e + \dots$

pretation of the band at 4μ , and there seems to be no reason why it should not be interpreted as $\nu_3 + \nu_4$. An examination of the behavior of this band at low temperatures would perhaps be the best method of checking its assignment as a sum, rather than a difference, of two fundamentals.

The writer wishes to acknowledge the benefit of discussing this matter with Professor Barker.

G. B. B. M. SUTHERLAND*

Pembroke College,
Cambridge, England,
September 15, 1939

* At present Leverhulme Research Fellow.
† E. F. Barker, Phys. Rev. 55, 657 (1939).

Magnetic Anisotropy of Nickel at 20°K

The only experiments on ferromagnetic anisotropy that have been carried out below the temperature of liquid nitrogen are those of Honda, Masumoto and Shirakawa¹ on nickel in liquid hydrogen. Their results indicate a large increase in the anisotropy as the temperature is lowered from 77° to 20°K, the constant K_1 changing by a factor of about 5. On the other hand, Brukhatov and Kirensky² have found that in the temperature range from 77° to 350°K, the constant is given by the relation

$$K_1 = K_0 e^{-\alpha T^2} \quad (1)$$

which predicts an increase of but 20 percent in going from 77° to 20°K. Accordingly we have undertaken, with the kind cooperation of Drs. H. A. Boorse and S. L. Quimby of Columbia University, to measure again the anisotropy constants at 77° and 20°K, using this time the more accurate method of torques.

The ratio of the constants was found to be about 1.2 (accuracy about 10 percent), as compared with the ratio 5 derived from the data of Honda, Masumoto and Shirakawa and 1.21 from the equation of Brukhatov and Kirensky. Our absolute values at 77°K and above are very close to those of Brukhatov and Kirensky. Thus our work extends the validity of this equation to lower temperatures (see Fig. 1) and shows that there is no unusual behavior in the ferromagnetic anisotropy at these low temperatures. This clarifies the theoretical situation since Van Vleck³ in his discussion of the wave-mechanical theory of anisotropy, has not been able to find any basis for a difference in variation with temperature of the constants for iron and for nickel.

The crystal of nickel used was grown⁴ in pure hydrogen from high purity nickel kindly supplied by Mr. E. Wise of the International Nickel Company. It was cut in the form of a disk with planes parallel to (100) and edges rounded to a semi-circle. The thickness was 0.29 cm, the largest diameter 1.38 cm. The liquid hydrogen was introduced into the Dewar flask surrounding the crystal through a straight Dewar tube of stainless steel connected directly to the bottom of the liquefier. The flask and crystal and torsion-measuring apparatus⁵ were then removed to the electromagnet for measurement of the torque when the field was inclined at various angles to [011], the direction of easiest magnetization in the (100) plane. The highest field used was about 4000 oersteds. This was not sufficient

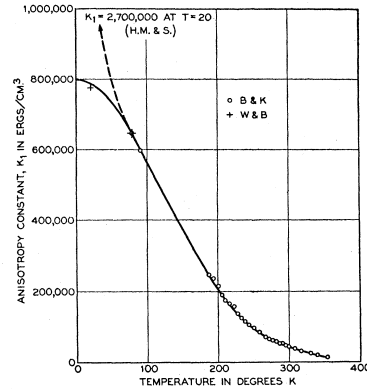


FIG. 1. Anisotropy constant of nickel as dependent on temperature. Data from 77° to 350°K, inclusive, by Brukhatov and Kirensky. Data at 77° and 20°K by the authors, adjusted slightly to fit the former data at 77°K. Curve calculated from Eq. (1) using $K_0 = 800,000$, $\alpha = 0.000034$.

to permit observation of the torque at saturation, but was so near this point that the ratio of the constants at 77° and 20°K could be determined with some accuracy. Saturation was later observed at 77°K in a field of about 5000 oersteds.

We wish to express our appreciation to Dr. Boorse and Dr. Quimby and others of the Cryogenic Laboratory of Columbia University, for supplying the liquid hydrogen.

H. J. WILLIAMS
R. M. BOZORTH

Bell Telephone Laboratories,
New York, New York,
September 25, 1939.

¹ K. Honda, H. Masumoto and Y. Shirakawa, Sci. Rep. Tohoku Imp. Univ. 24, 391 (1935). The anisotropy constant was derived by one of us (R. M. B., J. App. Phys. 8, 575 (1937)) from their magnetization curves for the [100] and [110] directions. A similar calculation by L. W. McKeehan (Phys. Rev. 52, 18 (1937)) yielded an even higher value of K_1 . Its value is somewhat uncertain on account of the extrapolation of the magnetization curves to saturation.

² N. L. Brukhatov and L. V. Kirensky, Soviet Phys. 12, 602 (1937).

³ J. H. Van Vleck, Phys. Rev. 52, 1178 (1937).

⁴ By Mr. O. L. Boothby, using the method described by P. P. Cioffi and O. L. Boothby, Phys. Rev. 55, 673 (1939).

⁵ Similar in design to that described by H. J. Williams, Rev. Sci. Inst. 8, 56 (1937).

The Disintegration of Mesotrons

In order to test the hypothesis of the spontaneous decay of mesotrons we have compared the absorption of the mesotron component of cosmic radiation in air and in carbon.

The mesotrons were detected by the coincidences of three Geiger-Müller tubes arranged in a vertical plane. The counters were shielded with 10 cm of lead on each side to prevent coincidences from the air showers. Also, 12.7 cm of lead was placed between the counters in order to cut off the soft component.

The absorption in air was measured by counting coincidences at different heights from Chicago up to the top of Mt. Evans, Colorado, (4300 m). The absorption in carbon

was measured by putting layers of graphite above the counters.

It was consistently found that the mass absorption in air was considerably larger than that in carbon. One set of measurements, for instance, gave the following results: Mt. Evans (4300 m, atmospheric pressure 618 g/cm²) without graphite: 11.9 ± 0.07 coinc./min. *Ibid.*, under 84 g/cm² graphite: 11.0 ± 0.057 coinc./min. Echo Lake (3240 m, atmospheric pressure 700 g/cm²) without graphite: 9.7 ± 0.046 coinc./min.

Thus the additional air layer of 82 g/cm² between Mt. Evans and Echo Lake reduced the intensity of the mesotrons by more than twice as much as did the graphite screen of 84 g/cm². It is obvious that this large difference cannot be ascribed to the difference in stopping power of air and carbon. We see, therefore, definite evidence for the disintegration of the mesotrons.

The above results show that 1.3 mesotrons out of 11 disintegrate while traveling a distance of $4.30 \times 10^8 - 3.24 \times 10^8 = 1.06 \times 10^8$ cm. Their mean-free-path for the disintegration is, therefore, $L = 1.06 \times 10^8 / \log(11/9.7) = 8.5 \times 10^6$ cm.

L is connected with the lifetime τ_0 by the formula: $c\tau_0 = \mu L / \beta$ where μ is the mass and β the momentum of the mesotrons. At sea level the average value of $\mu\beta$ was estimated to be about 0.07.¹ Assuming tentatively the same value in our case, one finds $\tau_0 = 2 \times 10^{-6}$ sec.

A fuller account of these experiments will be published later. The writers acknowledge with thanks the helpful discussions and support given to this work by Professor A. H. Compton. They also wish to express their appreciation for the facilities made available in Colorado by Dr. Joyce Stearns, as well as for the assistance of Mr. O. E. Polk and Mr. W. Bostick.

BRUNO ROSSI
H. VAN NORMAN HILBERRY
J. BARTON HOAG

Ryerson Physical Laboratory,
University of Chicago,
Chicago, Illinois,
September 30, 1939.

¹ B. Rossi, Cosmic Ray Symposium, Chicago, June, 1939; Rev. Mod. Phys. July-October (1939).

Gamma-Radiation

Richardson¹ has reported that the decay of N^{13} is accompanied by a gamma-ray of 280 ± 30 kev in addition to the well-known positron annihilation radiation. This gamma-ray is estimated to occur in 40 percent of all N^{13} disintegrations which take place.

The same radiation has been reported by Lyman² who estimates it to occur in 20 ± 15 percent of all disintegrations; and by Watase and Itoh³ who estimate it to occur in 20 percent of all disintegrations. The estimates of Richardson and of Watase and Itoh are uncertain by a factor of two. In view of the general interest in N^{13} , it has seemed desirable to make further observations on this radiation, using a method which is free from statistical errors.

The N^{13} gamma-ray spectrum was explored by measuring the energy and intensity of the secondary electrons ejected

from lead and aluminum foils of equivalent thickness. The magnetic spectrograph used was of the usual semi-circular focusing type, constructed largely of lead. The slit jaws and other parts nearest the radioactive source were faced with graphite in order to minimize the background. The radioactive sources were produced by bombarding 0.5 mm thick graphite plates with 4.3-Mev deuterons. Eastman "No Screen" x-ray film was used, and was developed for eight minutes in D19 developer at 66°F.

The Pb and Al foils were placed in contact with the radioactive sources. A particular gamma-ray will eject both photoelectrons and recoil electrons from lead, but only recoil electrons from aluminum because of the different Z dependence of the two effects. The photoelectrons from a particular gamma-ray, having an initially homogeneous velocity, will appear on the spectrogram as a group with a sharp upper energy limit followed by a gradual decrease in intensity toward lower energies because of a straggling in the emitting foil. The recoil electrons ejected by the same gamma-ray under these conditions will have a much less homogeneous energy distribution because their energies depend greatly upon their directions of emission with respect to those of the quanta.

In this experiment the photoelectron spectrum of lead irradiated by N^{13} gamma-radiation was isolated from the recoil spectrum as well as from the instrumental background. This was done by successive exposures with lead and aluminum secondary emitters, the latter distribution being subtracted from the former. The relative photoelectron intensities due to any gamma-rays present may thus be directly compared without making any estimates of the contribution of recoil electrons as was necessary in the experiments of Watase and Itoh.

The data are shown in Fig. 1, which gives the film opacity as a function of H_p . Because of differences in source intensities it was necessary to multiply the aluminum ordi-

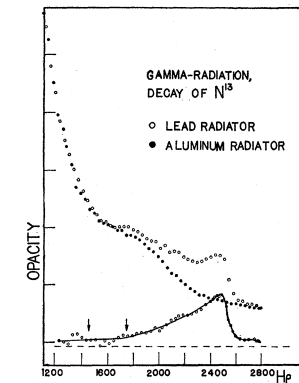


FIG. 1. Photometric measurements of films exposed to secondary electron spectra from lead and aluminum irradiated by N^{13} gamma-radiation. The lowermost curve is the difference between the Pb and Al data. The vertical arrows indicate the H_p region corresponding to lead K photoelectrons from a gamma-ray of 280 ± 30 kev.

The Disintegration of Mesotrons

In order to test the hypothesis of the spontaneous decay of mesotrons we have compared the absorption of the mesotron component of cosmic radiation in air and in carbon.

The mesotrons were detected by the coincidences of three Geiger-Müller tubes arranged in a vertical plane. The counters were shielded with 10 cm of lead on each side to prevent coincidences from the air showers. Also, 12.7 cm of lead was placed between the counters in order to cut off the soft component.

The absorption in air was measured by counting coincidences at different heights from Chicago up to the top of Mt. Evans, Colorado, (4300 m). The absorption in carbon

was measured by putting layers of graphite above the counters.

It was consistently found that the mass absorption in air was considerably larger than that in carbon. One set of measurements, for instance, gave the following results: Mt. Evans (4300 m, atmospheric pressure 618 g/cm²) without graphite: 11.9 ± 0.07 coinc./min. *Ibid.*, under 84 g/cm² graphite: 11.0 ± 0.057 coinc./min. Echo Lake (3240 m, atmospheric pressure 700 g/cm²) without graphite: 9.7 ± 0.046 coinc./min.

Thus the additional air layer of 82 g/cm² between Mt. Evans and Echo Lake reduced the intensity of the mesotrons by more than twice as much as did the graphite screen of 84 g/cm². It is obvious that this large difference cannot be ascribed to the difference in stopping power of air and carbon. We see, therefore, definite evidence for the disintegration of the mesotrons.

The above results show that 1.3 mesotrons out of 11 disintegrate while traveling a distance of $4.30 \times 10^5 - 3.24 \times 10^5 = 1.06 \times 10^5$ cm. Their mean-free-path for the disintegration is, therefore, $L = 1.06 \times 10^5 / \log(11/9.7) = 8.5 \times 10^5$ cm.

L is connected with the lifetime τ_0 by the formula: $c\tau_0 = \mu c L / p$ where μ is the mass and p the momentum of the mesotrons. At sea level the average value of $\mu c / p$ was estimated to be about 0.07.¹ Assuming tentatively the same value in our case, one finds $\tau_0 = 2 \times 10^{-6}$ sec.

A fuller account of these experiments will be published later. The writers acknowledge with thanks the helpful discussions and support given to this work by Professor A. H. Compton. They also wish to express their appreciation for the facilities made available in Colorado by Dr. Joyce Stearns, as well as for the assistance of Mr. O. E. Polk and Mr. W. Bostick.

BRUNO ROSSI
H. VAN NORMAN HILBERRY
J. BARTON HOAG

Ryerson Physical Laboratory,
University of Chicago,
Chicago, Illinois,
September 30, 1939.

¹ B. Rossi, Cosmic Ray Symposium, Chicago, June, 1939; Rev. Mod. Phys. July-October (1939).



REVIEWS OF
MODERN PHYSICS

VOLUME 11

JULY-OCTOBER, 1939

NUMBERS 3-4

SYMPOSIUM ON COSMIC RAYS

held at

THE UNIVERSITY OF CHICAGO

June, 1939

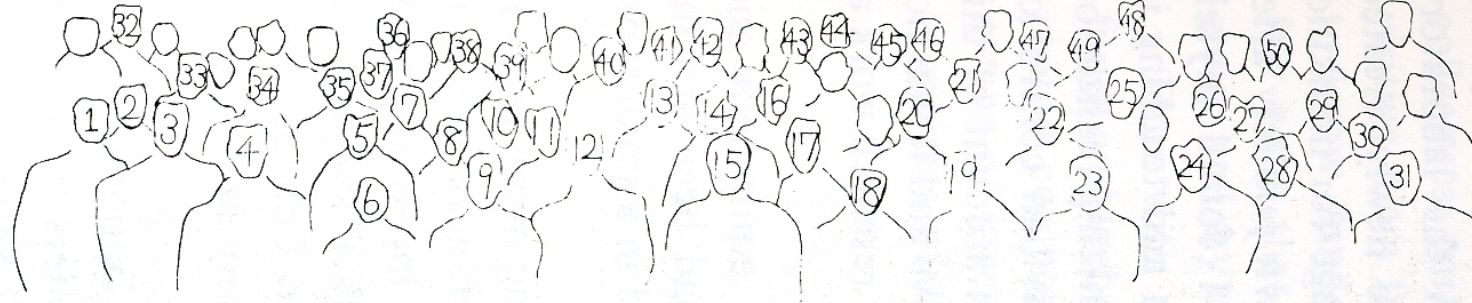


Fig. 1. Participants at the Cosmic Ray Conference (Symposium on Cosmic Rays, 1939) convened at the University of Chicago in the summer of 1939. The identification of participants is given by numbers in the over lay of this photograph as follows:

- | | | |
|--------------------------|-------------------------------|-------------------------------|
| 1. H. Bethe | 18. W. Bothe | 35. W. Bostick ⁺ |
| 2. D. Froman | 19. W. Heisenberg | 36. C. Eckart |
| 3. R. Brode | 20. P. Auger | 37. A. Code ⁺ |
| 4. A.H. Compton | 21. R. Serber | 38. J. Stearns (Denver?) |
| 5. E. Teller | 22. T. Johnson | 39. J. Hopfield |
| 6. A. Baños, Jr. | 23. J. Clay (Holland) | 40. E.O. Wollan* |
| 7. G. Groetzinger | 24. W.F.G. Swann | 41. D. Hughes ⁺ |
| 8. S. Goudsmit | 25. J.C. Street (Harvard) | 42. W. Jesse* |
| 9. M.S. Vallarta | 26. J. Wheeler | 43. B. Hoag |
| 10. L. Nordheim | 27. S. Neddermeyer | 44. N. Hillberry ⁺ |
| 11. J.R. Oppenheimer | 28. E. Herzog (?) | 45. F. Shonka ⁺ |
| 12. C.D. Anderson | 29. M. Pomerantz | 46. P.S. Gill ⁺ |
| 13. S. Forbush | 30. W. Harkins (U. of C.) | 47. A.H. Snell |
| 14. Nielsen (of Duke U.) | 31. H. Beutler | 48. J. Schremp |
| 15. V. Hess | 32. M.M. Shapiro ⁺ | 49. A. Haas? (Vienna) |
| 16. V.C. Wilson | 33. M. Schein* | 50. E. Dershem* |
| 17. B. Rossi | 34. C. Montgomery (Yale) | 51. H. Jones ⁺ |

*Then research associate of Compton.

+Then graduate student of Compton.



ICS

NUMBERS 3-4

RAY'S

AGO

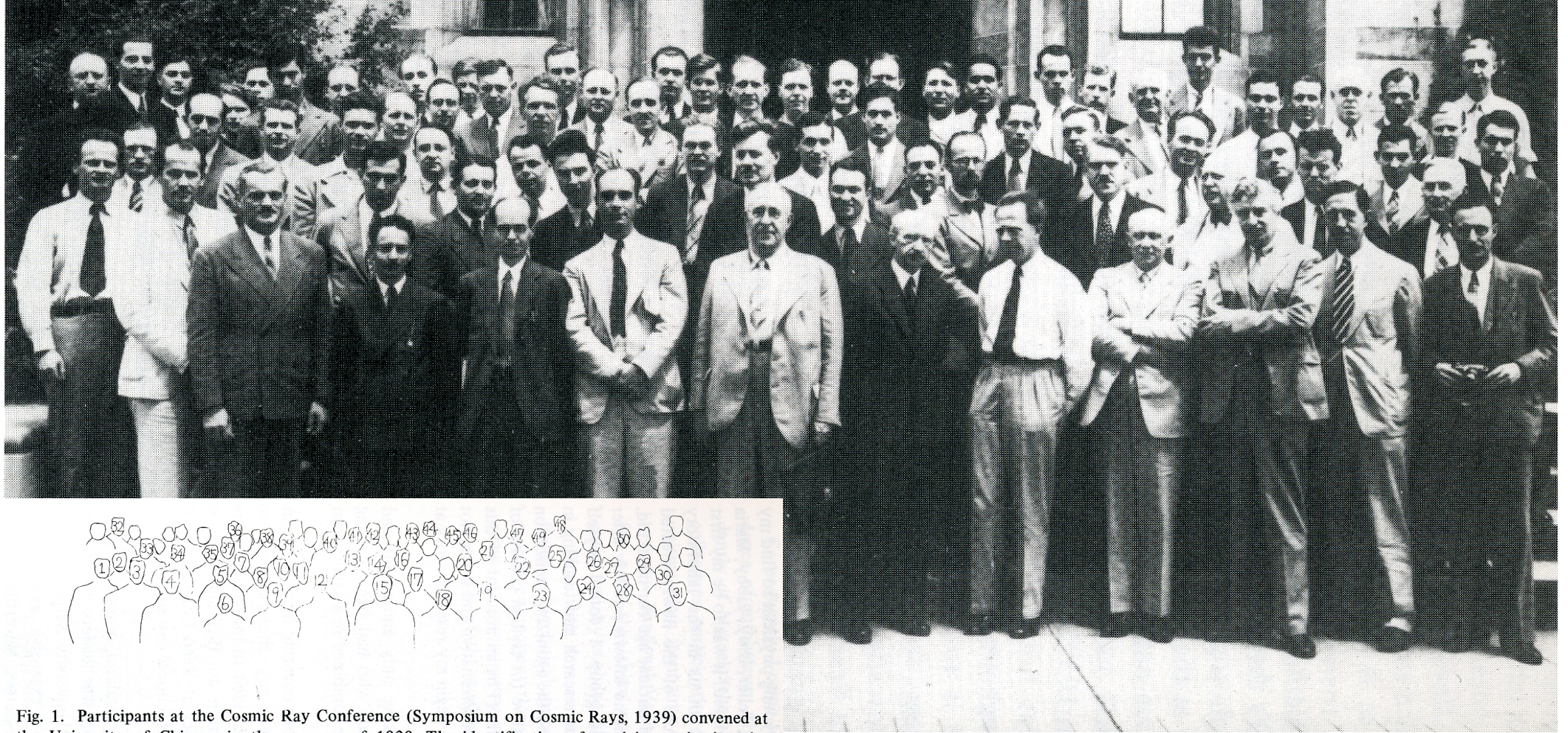


Fig. 1. Participants at the Cosmic Ray Conference (Symposium on Cosmic Rays, 1939) convened at the University of Chicago in the summer of 1939. The identification of participants is given by numbers in the over lay of this photograph as follows:

- | | | |
|--------------------------|-------------------------------|-------------------------------|
| 1. H. Bethe | 18. W. Bothe | 35. W. Bostick ⁺ |
| 2. D. Froman | 19. W. Heisenberg | 36. C. Eckart |
| 3. R. Brode | 20. P. Auger | 37. A. Code ⁺ |
| 4. A.H. Compton | 21. R. Serber | 38. J. Stearns (Denver?) |
| 5. E. Teller | 22. T. Johnson | 39. J. Hopfield |
| 6. A. Baños, Jr. | 23. J. Clay (Holland) | 40. E.O. Wollan ⁺ |
| 7. G. Groetzinger | 24. W.F.G. Swann | 41. D. Hughes ⁺ |
| 8. S. Goudsmit | 25. J.C. Street (Harvard) | 42. W. Jesse* |
| 9. M.S. Vallarta | 26. J. Wheeler | 43. B. Hoag |
| 10. L. Nordheim | 27. S. Neddermeyer | 44. N. Hillberry ⁺ |
| 11. J.R. Oppenheimer | 28. E. Herzog (?) | 45. F. Shonka ⁺ |
| 12. C.D. Anderson | 29. M. Pomerantz | 46. P.S. Gill ⁺ |
| 13. S. Forbush | 30. W. Harkins (U. of C.) | 47. A.H. Snell |
| 14. Nielsen (of Duke U.) | 31. H. Beutler | 48. J. Schrepf |
| 15. V. Hess | 32. M.M. Shapiro ⁺ | 49. A. Haas? (Vienna) |
| 16. V.C. Wilson | 33. M. Schein* | 50. E. Dershem* |
| 17. B. Rossi | 34. C. Montgomery (Yale) | 51. H. Jones ⁺ |

REVIEWS OF
MODERN PHYSICS

VOLUME 11

JULY-OCTOBER, 1939

NUMBERS 3-4

SYMPOSIUM ON COSMIC RAYS

held at

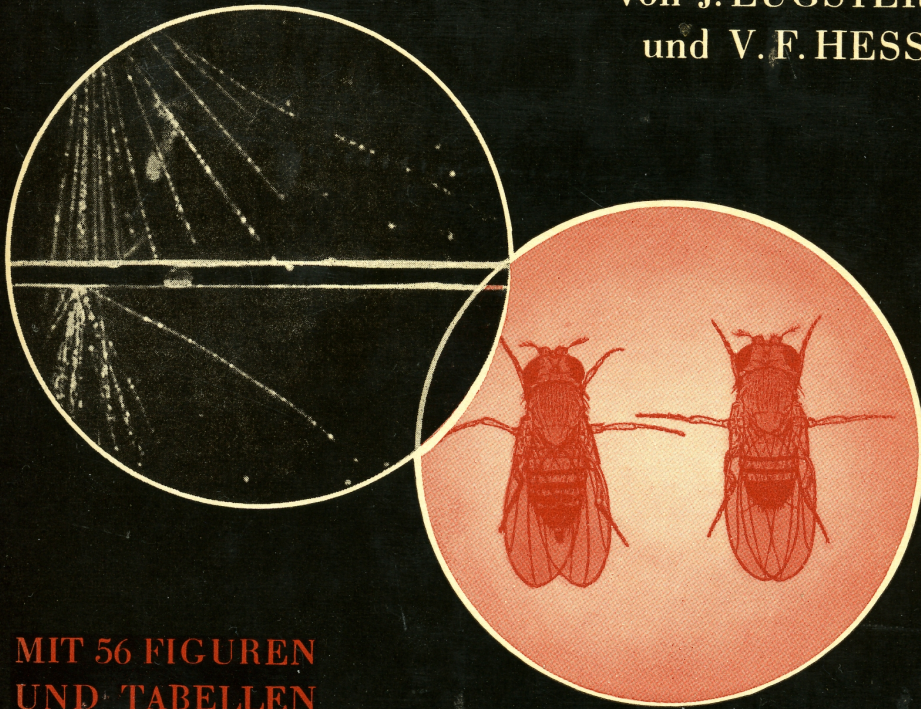
THE UNIVERSITY OF CHICAGO

June, 1939

*Then research associates of Compton

Die Weltraumstrahlung und ihre biologische Wirkung

von J. EUGSTER
und V. F. HESS



MIT 56 FIGUREN
UND TABELLEN

Die Kosmischen Strahlen, vor ca. 30 Jahren durch HESS entdeckt, und heute schon photographier- und meßbar, beeinflussen nachhaltig Wachstum, Fruchtbarkeit und Krebs, was EUGSTER in langjährigen Versuchen an Tieren und Pflanzen bewies. Das Buch gibt Physikern und Biologen, aber auch gebildeten Laien eine wertvolle Zusammenfassung der äußerst vielseitigen Forschungsergebnisse.

Space radiation and its biological effect

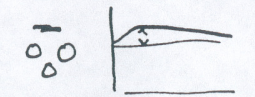
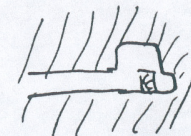
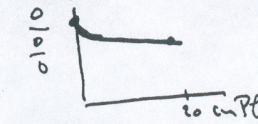
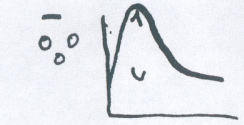
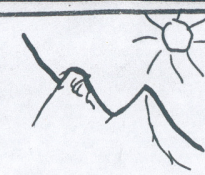
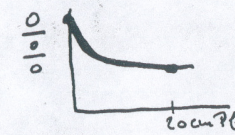
The cosmic rays, discovered ca. 30 years ago by Hess are today measurable. They influence the growth, fertility and cancer. Eugster has demonstrated this with experiments using plants and animals. The book gives physicists and biologists, as well as interested layman a valuable summary of a broad field of research.

1939

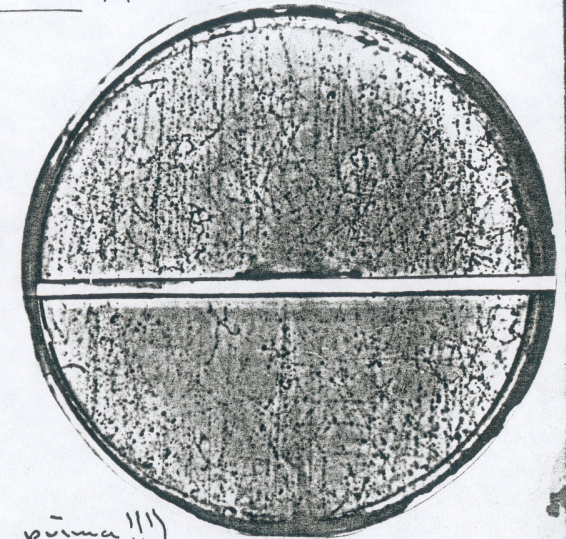
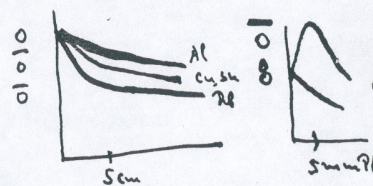
P. Auger Jungfraujoch

Pierre Auger }
Paul Elvenfest } 2-24 IX '34
Louis Leprieux Ringuet }

mit der Wilson kammer viele "shower"

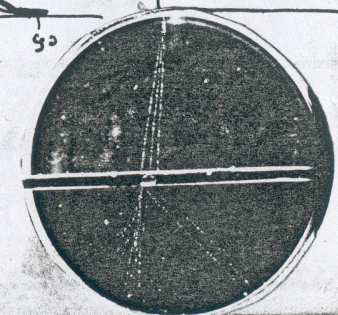
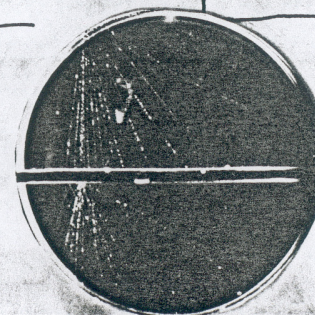
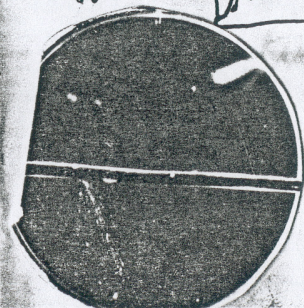
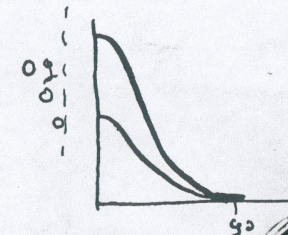


Pierre Auger }
Louis Leprieux Ringuet } 16 IV - 5 IV }
Francis Perrin } 28 IV - 5 IV } 1935
Paul Elvenfest } 16 IV - 12 IV }



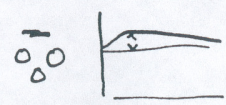
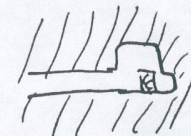
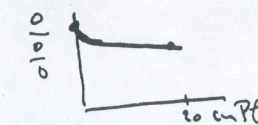
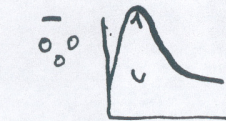
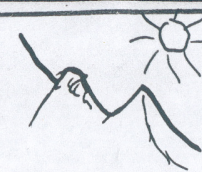
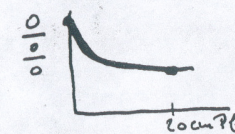
Doppelte Wilson Blasen kamme,
grosse shower

Pierre Auger } 2 IX - 23 IX, 3 IX - 6 IX }
Paul Elvenfest } 2 IX - 19 IX (15-19 IX) } keine Wiederholung, prima!!!
André Fréon } } 1936
Suzanne Auger } 5 IX - 15 IX }
Jean Mauvain } }
Solomon Rosenblum (Baby) } 20 IX - 3 IX }



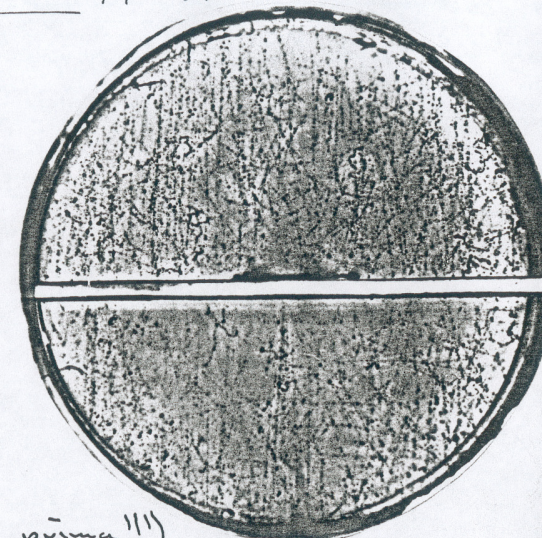
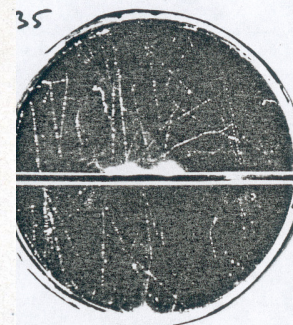
P. Auger Jungfraujoch

Pierre Auger
Paul Ehrenfest
Louis Leprince Ringuet } 2-24 IX '34
mit der Wilson kammer viele "shower"

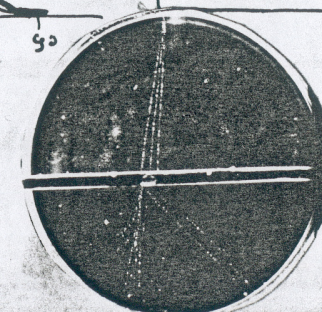
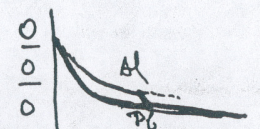
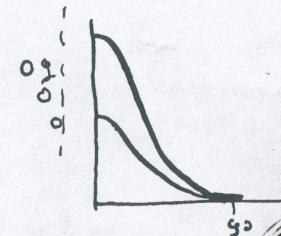


MEASURING COSMIC RAYS IN THE SWISS ALPS

The author (left) and his collaborator, P. Ehrenfest, set up their apparatus in the Jungfraujoch.



keine Wiederholung, prima!!!
1936



WHAT ARE COSMIC RAYS?

Revised and Enlarged American Edition

BY PIERRE AUGER

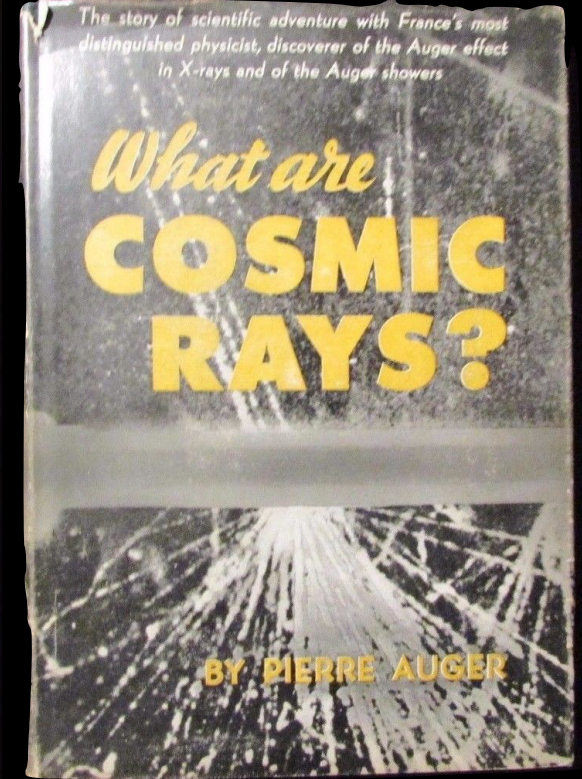
TRANSLATED FROM THE FRENCH

by

MAURICE M. SHAPIRO



UNIVERSITY OF CHICAGO PRESS · CHICAGO



Für die kurzen Originalmitteilungen ist ausschließlich der Verfasser verantwortlich.

Gekoppelte Höhenstrahlen.

Bei Bestimmungen der Zufallskoinzidenzen hoch auflösender Zählrohrverstärkeranordnungen (bis $5 \cdot 10^{-7}$ sec) ergab sich eine wesentlich größere Anzahl, als nach den elektrischen Konstanten der Anordnung zu erwarten war, ferner ihre Anzahl abhängig vom gegenseitigen Abstand der Zählrohre, wie z. B. für Zählrohre von 430 qcm wirksamer Oberfläche ($90 \cdot 4,8$) und $\tau = 5 \cdot 10^{-6}$ sec Tabelle 1 zeigt.

Tabelle 1. Anzahl der zusätzlichen Koinzidenzen je Stunde in Abhängigkeit vom gegenseitigen Abstand der ungepanzerten Zählrohre.

Rohrabstand in m:	1,25	3,75	5,00	7,50	10,00	20,00	75,00
Im Experimentierraum	$13,3 \pm 2,1$	$13,3 \pm 1,3$	$13,1 \pm 1,3$	$9,3 \pm 1,2$	$0,4 \pm 0,8$	—	—
Im Freien	$37,5 \pm 4,4$	—	$21,5 \pm 2,1$	—	$10,0 \pm 2,2$	$2,5 \pm 1,5$	$0,7 \pm 1,3$

Mit zunehmendem Abstand der Zählrohre voneinander nimmt die Anzahl der Zufallskoinzidenzen zunächst dauernd ab, bis sich bei über 10,0 m Abstand (Beobachtungen im Experimentierraum) konstante Werte einstellen und überschüssige Koinzidenzen nicht mehr nachweisbar sind. Wurde ein Bleipanzer ($10 \cdot 10 \cdot 40$ cm³) so zwischen die Zählrohre gebracht, daß er den Durchgang ein und desselben Strahles durch die beiden horizontal liegenden Rohre hinderte, so änderte sich wesentlich nichts, wie ja nach der Richtungsverteilung der Höhenstrahlen zu erwarten ist. Wohl aber machten sich die zusätzlichen Koinzidenzen nicht mehr bemerkbar, wenn die Rohre allseitig durch 10 cm Blei geschirmt wurden. Dann erhielt man auch bei nahe aneinander liegenden Rohren dieselben konstanten Werte für τ wie bei über 10 m Abstand ungepanzert. Die zusätzlichen Koinzidenzen mußten demnach von Strahlen herrühren, die durch 10 cm Blei weitgehend absorbiert werden. Bei starker Erhöhung der Stoßzahlen durch radioaktive Bestrahlung wird der Einfluß der Höhenstrahlen unwirksam. Dann ergab sich ebenfalls bei kleinerem Zählrohrabstande (5 m) der Wert des Auflösungsvermögens, der 1. nach den elektrischen Daten, 2. nach den Bestimmungen mit allseitigem Panzer und 3. nach den Messungen über 10 m Abstand ungepanzert das wahre Auflösungsvermögen der Anordnung darstellt.

Nur bei statistisch verteilten und voneinander unabhängigen Einzelstößen N_1 und N_2 der beiden Zählrohre gilt die Beziehung $K_z = 2N_1N_2\tau$ zur Bestimmung des Auflösungsvermögens τ . Es müssen also bei ungeschirmten und zu nahe

coupled
„high-altitude rays“

Strahlen im Schauer. Unter der Decke des Experimentierraumes sind diese Sekundärstrahlen über eine Fläche von mindestens 60 qm sicher nachweisbar.

Sollten sie bevorzugt in der Decke ausgelöst werden, so würden nach der Geometrie der Anordnung dabei Strahlen bis zu 80° aus ihrer ursprünglichen Richtung abgelenkt worden sein. Indessen ist bei einem Äquivalent der Decke von nur 1 cm Blei und dem Absorptionskoeffizienten der Strahlen von $\mu_{Pb} = 0,12 \text{ cm}^{-1}$ anzunehmen, daß die Strahlen überwiegend in der Atmosphäre bis zu großen Höhen über Boden erzeugt werden. Die Decke wird also mehr absorbierende als strahlenauslösende Wirkung haben, so daß im Freien eine größere Anzahl von Strahlen unter gleichen Bedingungen zu erwarten ist. Dies bestätigen die Messungen mit der 2-fach-Koinzidenzanordnung. Im Freien konnten die zusätzlichen Koinzidenzen bis auf Abstände von über 20 m sicher beobachtet werden, so daß gekoppelte Höhenstrahlen im Freien sogar bis über 400 qm Fläche auftreten (Tabelle 1). Selbst bei 75 m Abstand schien noch ein kleiner Überschuß vorhanden, der aber erst durch sehr lange Messreihen sichergestellt werden müßte.

Aus dem niedrigen Absorptionskoeffizienten ist zu folgern, daß selbst Schauerstrahlen, die bis etwa 2 km Höhe über dem Boden entstehen, diesen noch erreichen müssen. Diese würden dann über eine sehr große Fläche verteilt sein. Da für solche Schauer trotz des großen Strahlenreichtums die räumliche Dichte der Strahlen in Bodennähe nur außerordentlich gering sein kann, ist es durchaus verständlich, wenn sie als zusätzliche Koinzidenzen über größere Abstände nachweisbar sind. Bei den hier ausgeführten Messungen bis zu 20 m würde, die Entstehung dieser Höhenstrahlen vorausgesetzt, der Winkel kleiner als 1° sein.

Es besteht, den Ursprungsort der Höhenstrahlen genauer zu bestimmen dadurch, daß die Koinzidenzanordnungen angepeilt. Diese Angepeilung wird jedoch erst diskutiert werden, wenn es festgestellt sein wird.

Die ersten Untersuchungen hatte der Verfasser im Jahre 1938 in einem Vortrage im Institut für Höhenstrahlenforschung der Universität Berlin, den 25. August 1938.

W. KOLHÖRSTER. I. MATTHES. E. WEBER.

Neue Messungen der Fluoreszenz-Intensitäts-Änderungen grüner Pflanzen.

Ein günstiges Versuchsobjekt für quantitative Messungen ist die Meeresalge *Ulva lactuca*¹. Sie besteht aus blattartigen,

¹ Das Versuchsmaterial verdanken wir dem Entgegenkommen der Staatlichen Biologischen Anstalt auf Helgoland.

Kolhörster
discovery of air showers

wird sich also um Sekundärstrahlen handeln. Das zeigen auch folgende Versuche mit einer 3-fachen Koinzidenzapparatur, deren Auflösungsvermögen mit einer besonderen Anordnung zu $5 \cdot 10^{-6}$ sec bestimmt worden war. Bei Aufstellung der Zählrohre horizontal und radial auf einem Kreise ist dann überhaupt keine meßbare Anzahl von Zufallskoinzidenzen zu erwarten (höchstens 10^{-4} Koi/Std.). Es ergaben sich aber bei Zählrohren von 216 qcm wirksamer Fläche

Ungepanzert.	$2,7 \pm 0,4$ Koi/Std.
1 Rohr gepanzert.	$0,7 \pm 0,1$ Koi/Std.
3 Rohre gepanzert.	$0,68 \pm 0,02$ Koi/Std.

Für die kurzen Originalmitteilungen ist ausschließlich der Verfasser verantwortlich.

Gekoppelte Höhenstrahlen.

Bei Bestimmungen der Zufallskoinzidenzen hoch auflösender Zählrohrverstärkeranordnungen (bis $5 \cdot 10^{-7}$ sec) ergab sich eine wesentlich größere Anzahl, als nach den elektrischen Konstanten der Anordnung zu erwarten war, ferner ihre Anzahl abhängig vom gegenseitigen Abstand der Zählrohre, wie z. B. für Zählrohre von 430 qcm wirksamer Oberfläche ($90 \cdot 4,8$) und $\tau = 5 \cdot 10^{-6}$ sec Tabelle 1 zeigt.

Tabelle 1. Anzahl der zusätzlichen Koinzidenzen je Stunde in Abhängigkeit vom gegenseitigen Abstand der ungepanzerten Zählrohre.

Rohrabstand in m:	1,25	3,75	5,00	7,50	10,00	20,00	75,00
Im Experimentierraum	$13,3 \pm 2,1$	$13,3 \pm 1,3$	$13,1 \pm 1,3$	$9,3 \pm 1,2$	$0,4 \pm 0,8$	—	—
Im Freien	$37,5 \pm 4,4$	—	$21,5 \pm 2,1$	—	$10,0 \pm 2,2$	$2,5 \pm 1,5$	$0,7 \pm 1,3$

Mit zunehmendem Abstand der Zählrohre voneinander nimmt die Anzahl der Zufallskoinzidenzen zunächst dauernd ab, bis sich bei über 10,0 m Abstand (Beobachtungen im Experimentierraum) konstante Werte einstellen und überschüssige Koinzidenzen nicht mehr nachweisbar sind. Wurde ein Bleipanzer ($10 \cdot 10 \cdot 40$ cm³) so zwischen die Zählrohre gebracht, daß er den Durchgang ein und desselben Strahles durch die beiden horizontal liegenden Rohre hinderte, so änderte sich wesentlich nichts, wie ja nach der Richtungsverteilung der Höhenstrahlen zu erwarten ist. Wohl aber machten sich die zusätzlichen Koinzidenzen nicht mehr bemerkbar, wenn die Rohre allseitig durch 10 cm Blei geschirmt wurden. Dann erhielt man auch bei nahe aneinander liegenden Rohren dieselben konstanten Werte für τ wie bei über 10 m Abstand ungepanzert. Die zusätzlichen Koinzidenzen mußten demnach von Strahlen herrühren, die durch 10 cm Blei weitgehend absorbiert werden. Bei starker Erhöhung der Stoßzahlen durch radioaktive Bestrahlung wird der Einfluß der Höhenstrahlen unwirksam. Dann ergab sich ebenfalls bei kleinerem Zählrohrabstande (5 m) der Wert des Auflösungsvermögens, der 1. nach den elektrischen Daten, 2. nach den Bestimmungen mit allseitigem Panzer und 3. nach den Messungen über 10 m Abstand ungepanzert das wahre Auflösungsvermögen der Anordnung darstellt.

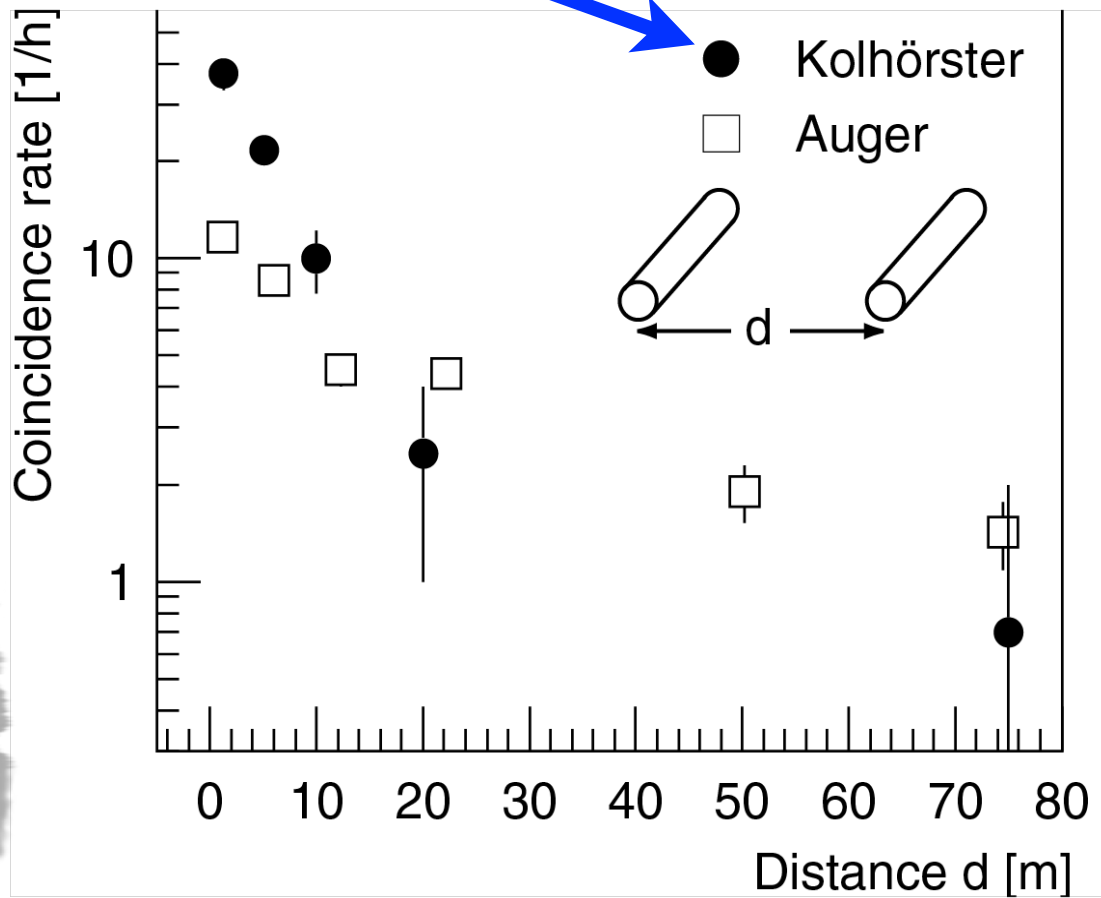
Nur bei statistisch verteilten und voneinander unabhängigen Einzelstößen N_1 und N_2 der beiden Zählrohre gilt die Beziehung $K_z = 2N_1N_2\tau$ zur Bestimmung des Auflösungsvermögens τ . Es müssen also bei ungeschirmten und zu nahe

coupled
„high-altitude rays“

Strahlen im Schauer. Unter der Decke des Experimentierraumes sind diese Sekundärstrahlen über eine Fläche von mindestens 60 qm sicher nachweisbar.

Sollten sie bevorzugt in die Richtung der Zählrohre bis zu 80° aus ihrer ursprünglichen Richtung abgelenkt worden sein. Indessen ist von nur 1 cm Blei und d. Strahlen von $\mu_{Pb} = 0,12$ cm⁻¹ überwiegend in der Atmosphäre erzeugt werden. Die durch den Bleipanzer hindere als strahlenauslösende Freien eine größere Anzahl Koinzidenzen zu erwarten ist, mit der 2-fach-Koinzidenzen die zusätzlichen Koinzidenzen 20 m sicher beobachtet werden können, wenn die Strahlen im Freien sogar bei 75 m (Tabelle 1). Selbst bei 75 m Überschuß vorhanden, der durch die Anordnung der Zählrohre sichergestellt werden kann.

Aus dem niedrigen Absolutwert der Koinzidenzen ist abzusehen, daß selbst Schauerstrahlen, die dem Boden entstehen, dies nicht bewirken würden, wenn sie über eine Fläche von 60 qm sich ausbreiten. Da für solche Schauer trotz der räumlichen Dichte der Koinzidenzen die räumliche Dichte der Koinzidenzen ordentlich gering sein kann, wenn sie als zusätzliche Koinzidenzen nachweisbar sind.



Kolhörster
discovery of air showers

wird sich also um Sekundärstrahlen handelt. Das zeigen auch folgende Versuche mit einer 3-fachen Koinzidenzapparatur, deren Auflösungsvermögen mit einer besonderen Anordnung zu $5 \cdot 10^{-6}$ sec bestimmt worden war. Bei Aufstellung der Zählrohre horizontal und radial auf einem Kreise ist dann überhaupt keine meßbare Anzahl von Zufallskoinzidenzen zu erwarten (höchstens 10^{-4} Koi/Std.). Es ergaben sich aber bei Zählrohren von 216 qcm wirksamer Fläche

Ungepanzert. $2,7 \pm 0,4$ Koi/Std.
1 Rohr gepanzert. $0,7 \pm 0,1$ Koi/Std.
3 Rohre gepanzert. $0,68 \pm 0,02$ Koi/Std.

Dresden kurz berichtet.
Berlin, Institut für Höhenphysik, den 25. August 1938.
W. KOLHÖRSTER

Neue Messungen der Fluoreszenz

Ein günstiges Versuchsobjekt für quantitative Messungen ist die Meeresalge *Ulva lactuca*¹. Sie besteht aus blattartigen

¹ Das Versuchsmaterial verdanken wir dem Entgegenkommen der Staatlichen Biologischen Anstalt auf Helgoland.

P. Auger et al., Comptes renduz 206 (1938) 1721

Für die kurzen Originalmitteilungen ist ausschließlich der Verfasser verantwortlich.

Gekoppelte Höhenstrahlen.

Bei Bestimmungen der Zufallskoinzidenzen hoch auflösender Zählrohrverstärkeranordnungen (bis $5 \cdot 10^{-7}$ sec) ergab sich eine wesentlich größere Anzahl, als nach den elektrischen Konstanten der Anordnung zu erwarten war, ferner ihre Anzahl abhängig vom gegenseitigen Abstand der Zählrohre, wie z. B. für Zählrohre von 430 qcm wirksamer Oberfläche ($90 \cdot 4,8$) und $\tau = 5 \cdot 10^{-6}$ sec Tabelle 1 zeigt.

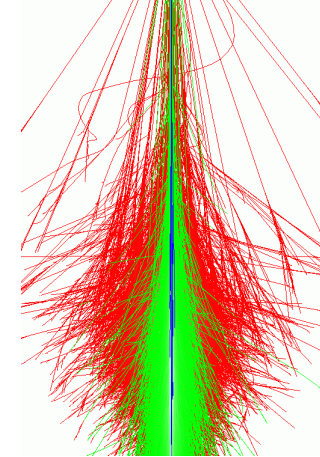
Tabelle 1. Anzahl der zusätzlichen Koinzidenzen je Stunde in Abhängigkeit vom gegenseitigen Abstand der ungepanzerten Zählrohre.

Rohrabstand in m:	1,25	3,75	5,00	7,50	10,00	20,00	75,00
Im Experimentierraum	$13,3 \pm 2,1$	$13,3 \pm 1,3$	$13,1 \pm 1,3$	$9,3 \pm 1,2$	$0,4 \pm 0,8$	—	—
Im Freien	$37,5 \pm 4,4$	—	$21,5 \pm 2,1$	—	$10,0 \pm 2,2$	$2,5 \pm 1,5$	$0,7 \pm 1,3$

Mit zunehmendem Abstand der Zählrohre voneinander nimmt die Anzahl der Zufallskoinzidenzen zunächst dauernd ab, bis sich bei über 10,0 m Abstand (Beobachtungen im Experimentierraum) konstante Werte einstellen und überschüssige Koinzidenzen nicht mehr nachweisbar sind. Wurde ein Bleipanzer ($10 \cdot 10 \cdot 40$ cm³) so zwischen die Zählrohre gebracht, daß er den Durchgang ein und desselben Strahles durch die beiden horizontal liegenden Rohre hinderte, so änderte sich wesentlich nichts, wie ja nach der Richtungsverteilung der Höhenstrahlen zu erwarten ist. Wohl aber machten sich die zusätzlichen Koinzidenzen nicht mehr bemerkbar, wenn die Rohre allseitig durch 10 cm Blei geschirmt wurden. Dann erhielt man auch bei nahe aneinander liegenden Rohren dieselben konstanten Werte für τ wie bei über 10 m Abstand ungepanzert. Die zusätzlichen Koinzidenzen mußten demnach von Strahlen herrühren, die durch 10 cm Blei weitgehend absorbiert werden. Bei starker Erhöhung der Stoßzahlen durch radioaktive Bestrahlung wird der Einfluß der Höhenstrahlen unwirksam. Dann ergab sich ebenfalls bei kleinerem Zählrohrabstand (5 m) der Wert des Auflösungsvermögens, der 1. nach den elektrischen Daten, 2. nach den Bestimmungen mit allseitigem Panzer und 3. nach den Messungen über 10 m Abstand ungepanzert das wahre Auflösungsvermögen der Anordnung darstellt.

Nur bei statistisch verteilten und voneinander unabhängigen Einzelstößen N_1 und N_2 der beiden Zählrohre gilt die Beziehung $K_z = 2N_1N_2\tau$ zur Bestimmung des Auflösungsvermögens τ . Es müssen also bei ungeschirmten und zu nahe

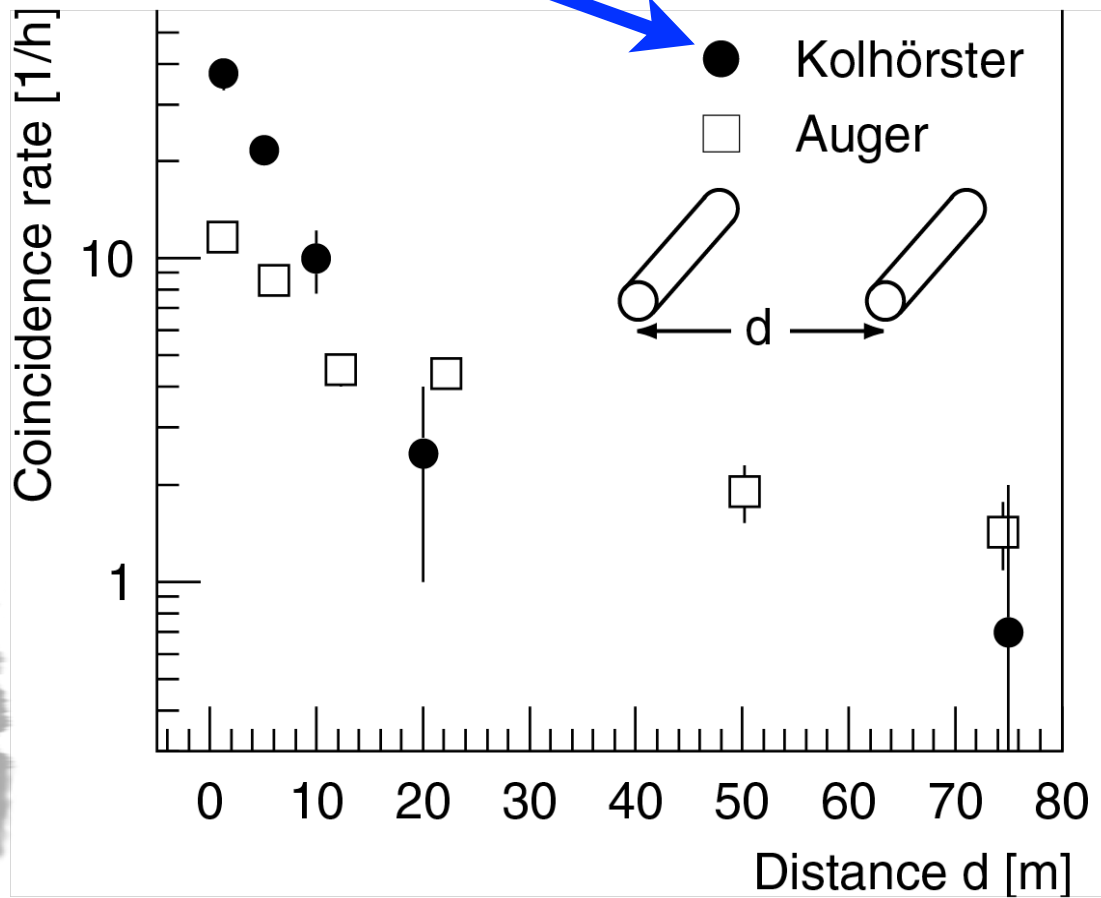
coupled „high-altitude rays“



Strahlen im Schauer. Unter der Decke des Experimentierraumes sind diese Sekundärstrahlen über eine Fläche von mindestens 60 qm sicher nachweisbar.

Sollten sie bevorzugt in die Richtung der Zählrohre bis zu 80° aus ihrer ursprünglichen Richtung abgelenkt worden sein. Indessen ist von nur 1 cm Blei und d. Strahlen von $\mu_{Pb} = 0,12$ cm⁻¹ überwiegend in der Atmosphäre erzeugt werden. Die durch den Bleipanzer hindurchgehende als strahlenauslösende Freien eine größere Anzahl Koinzidenzen zu erwarten ist, mit der 2-fach-Koinzidenzen die zusätzlichen Koinzidenzen 20 m sicher beobachtet werden können, wenn die Zählrohre im Freien sogar bei 75 m Abstand ein Überschuß vorhanden, der sich sicherstellen lassen würde.

Aus dem niedrigen Absolutwert der Koinzidenzen ist abzulesen, daß selbst Schauerstrahlen dem Boden entstehen, dies würde dann über eine Fläche von 60 qm. Da für solche Schauer trotz der räumlichen Dichte der Koinzidenzen die räumliche Dichte der Koinzidenzen ordentlich gering sein kann, wenn sie als zusätzliche Koinzidenzen nachweisbar sind.



Kolhörster discovery of air showers

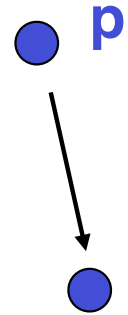
Dresden kurz berichtet. Berlin, Institut für Hörfähigkeit Berlin, den 25. August 1938. W. KOLHÖRSTER

Neue Messungen der Fluoreszenzgrüne

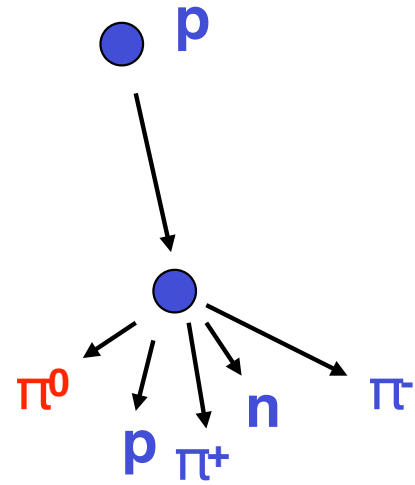
Ein günstiges Versuchsobjekt für quantitative Messungen ist die Meeresalge *Ulva lactuca*¹. Sie besteht aus blattartigen, grünen, zylindrischen, bis zu 1 cm langen, 0,5 mm breiten, 0,2 mm dicken Blättern. Das Versuchsmaterial verdanken wir dem Entgegenkommen der Staatlichen Biologischen Anstalt auf Helgoland.

P. Auger et al., Comptes renduz 206 (1938) 1721

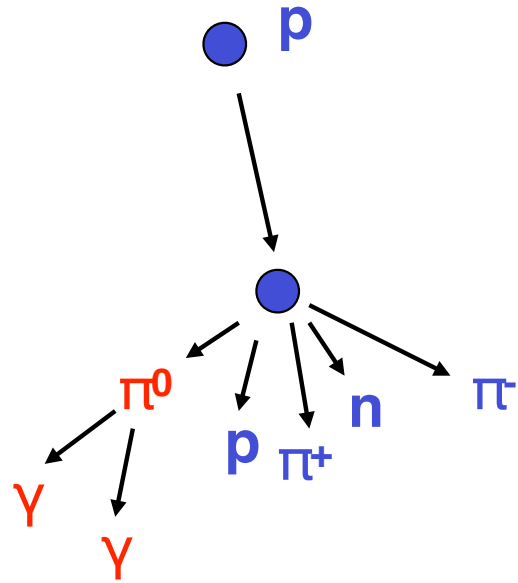
Extensive Air Shower



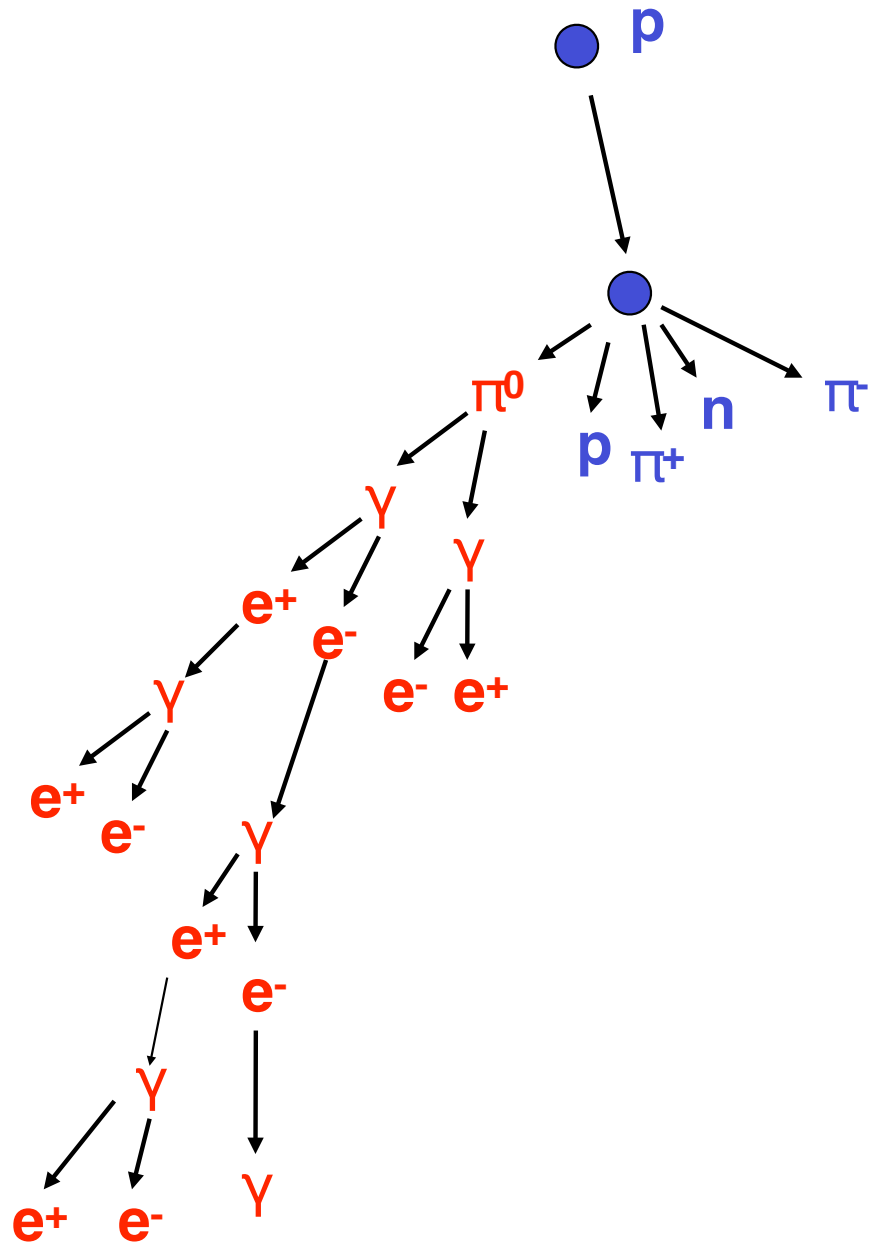
Extensive Air Shower



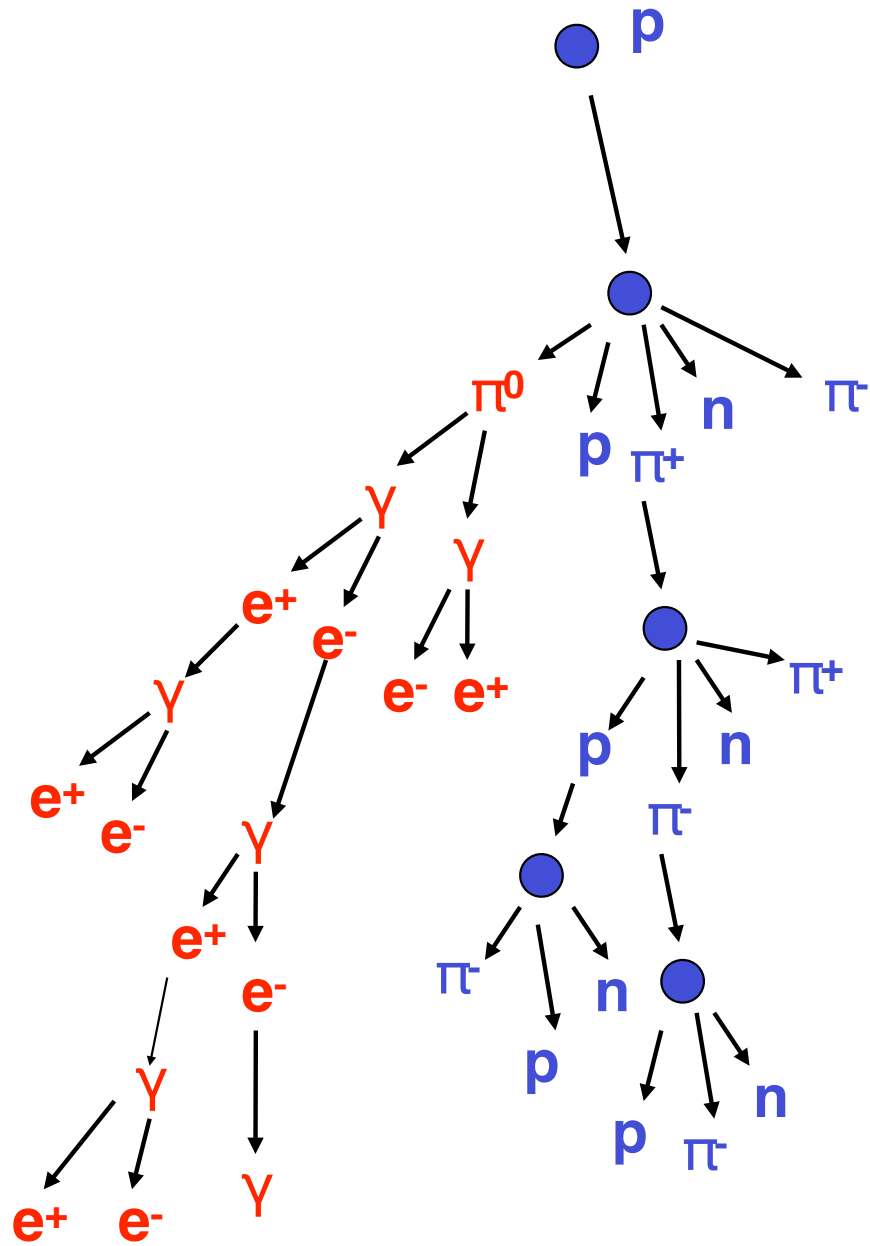
Extensive Air Shower



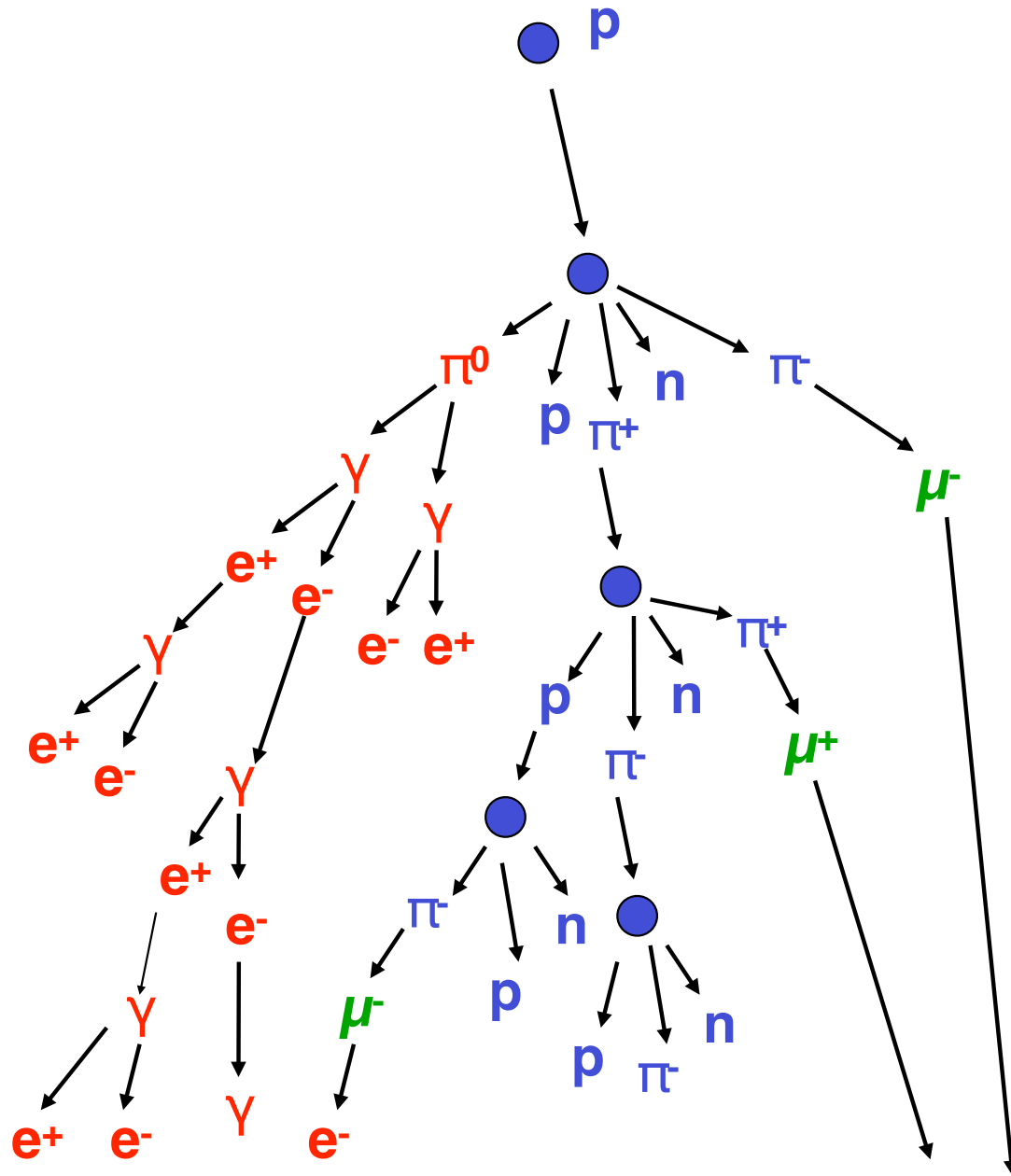
Extensive Air Shower



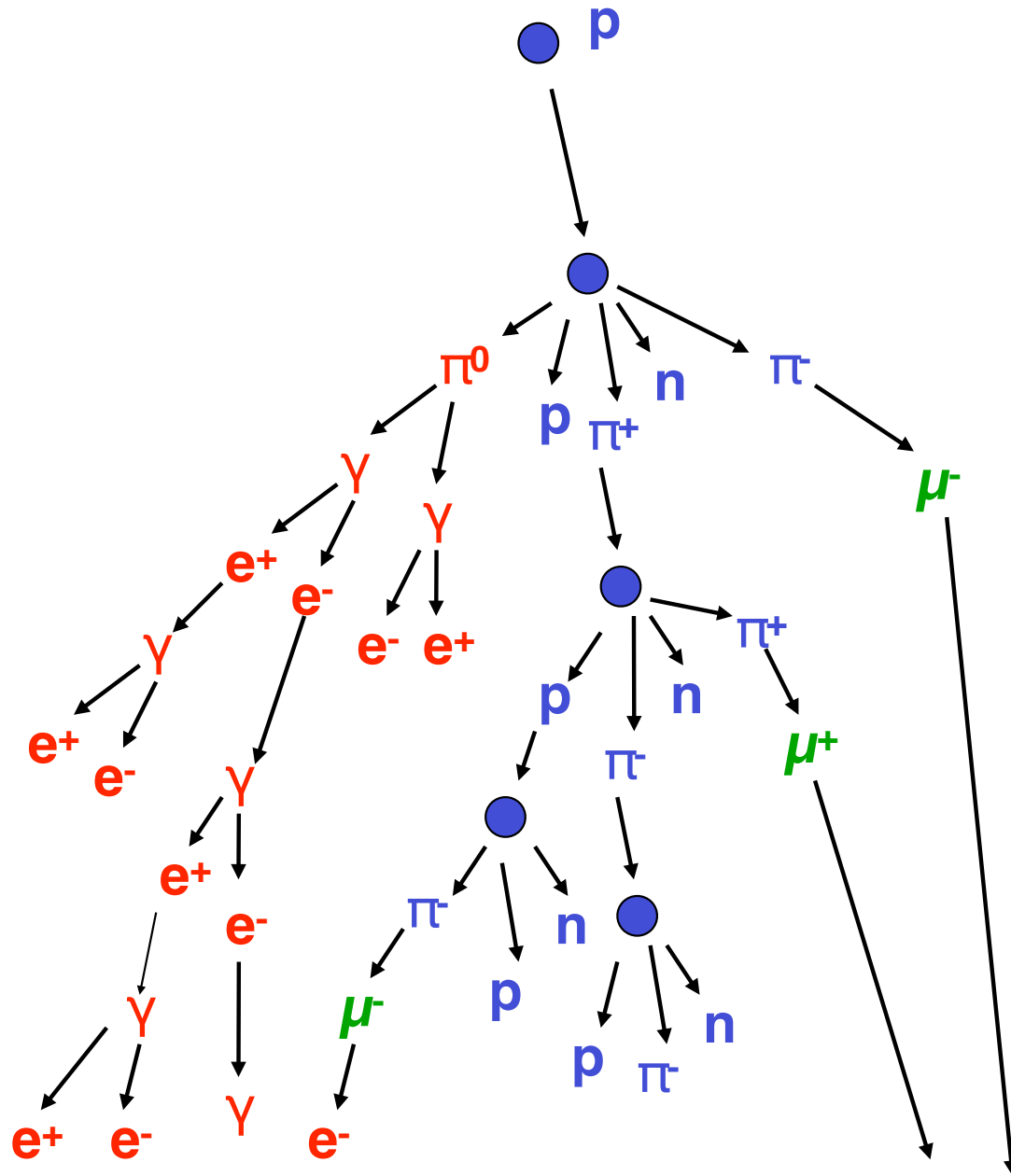
Extensive Air Shower



Extensive Air Shower



Extensive Air Shower



electromagnetic

hadronic

muonic

shower component

Extensive Air Shower

Proton 10^{15} eV:
on ground

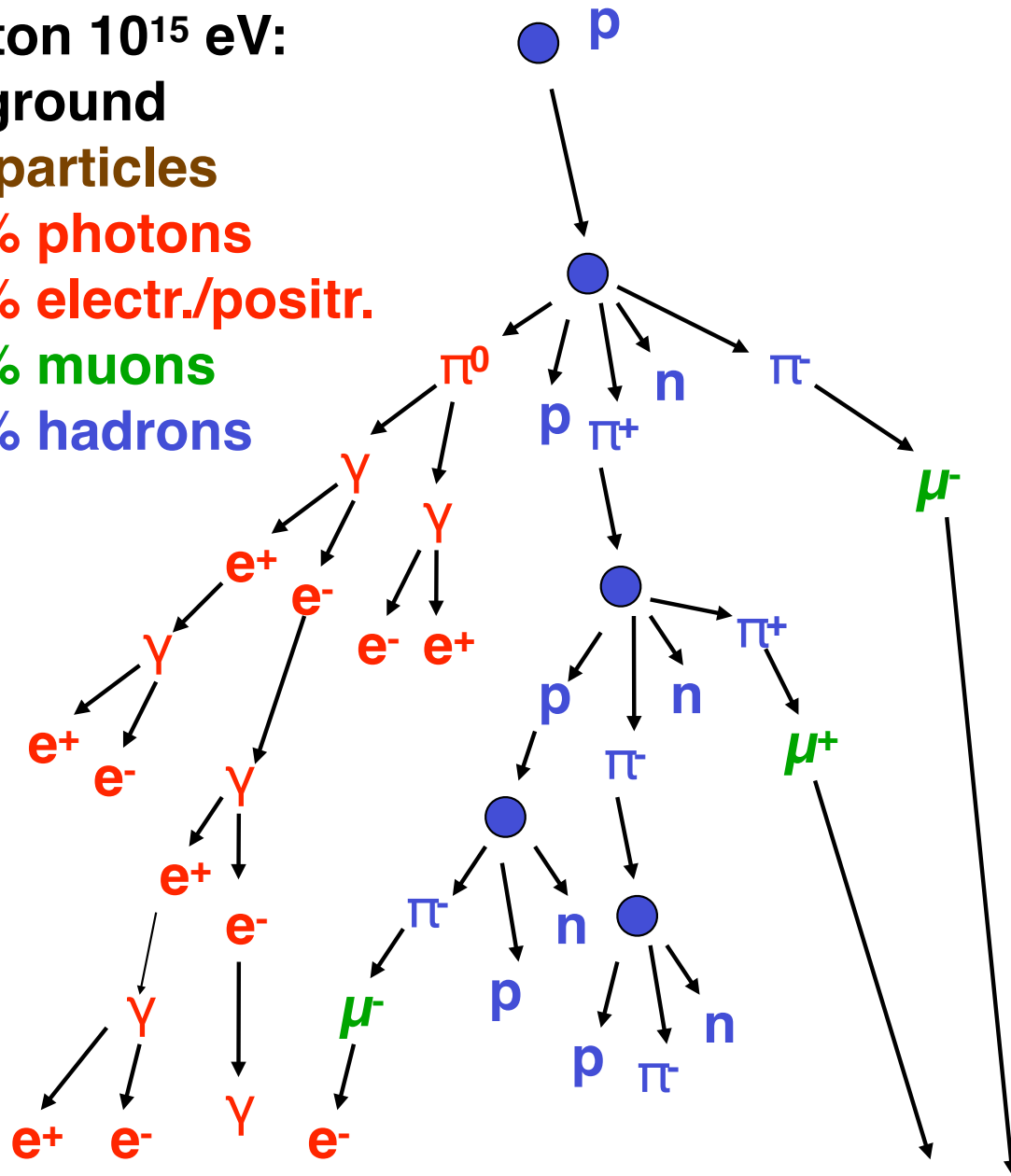
10^6 particles

80% photons

18% electr./positr.

1.7% muons

0.3% hadrons



electromagnetic

hadronic

muonic

shower component

Extensive Air Shower

Proton 10^{15} eV:
on ground

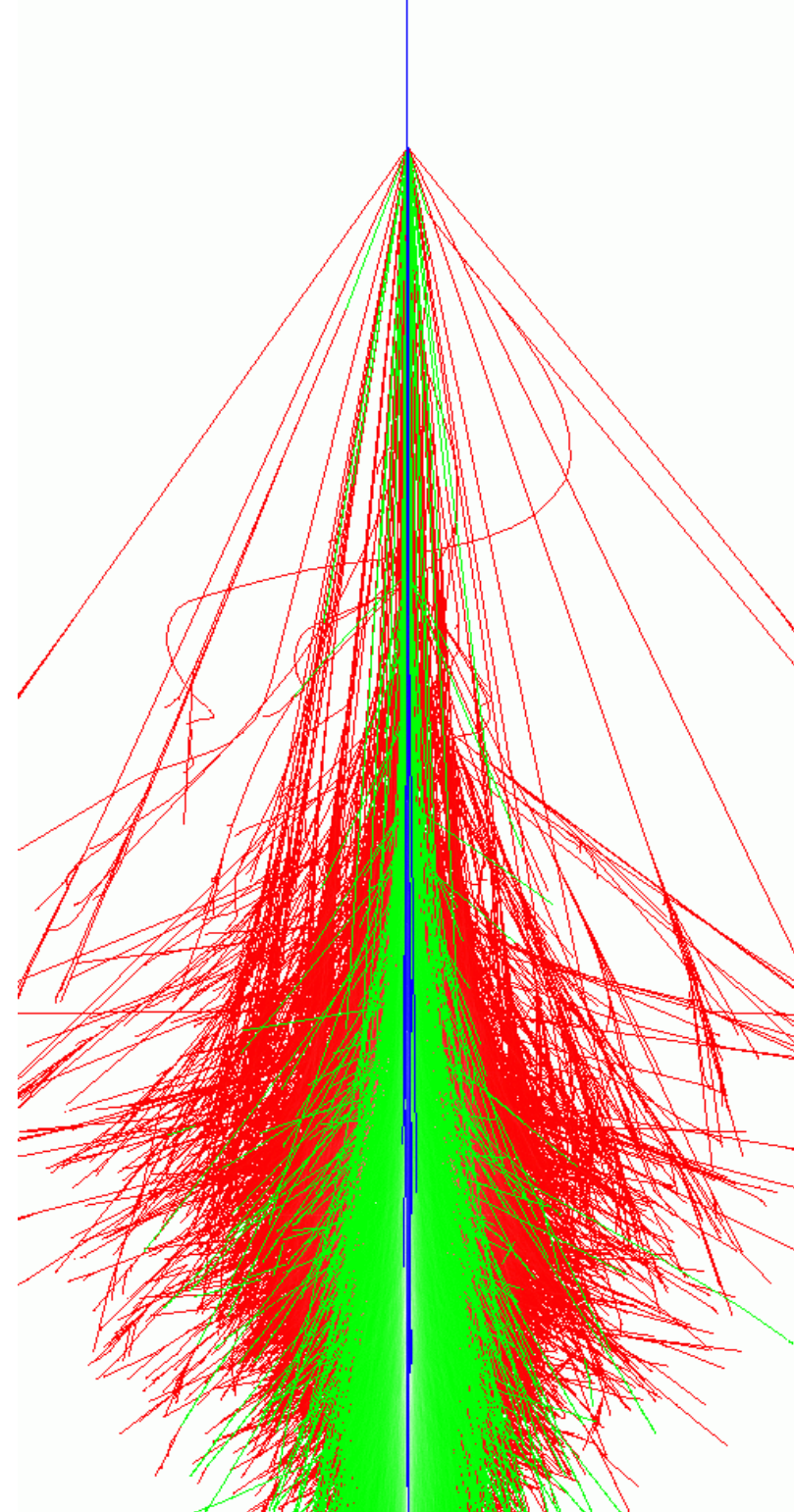
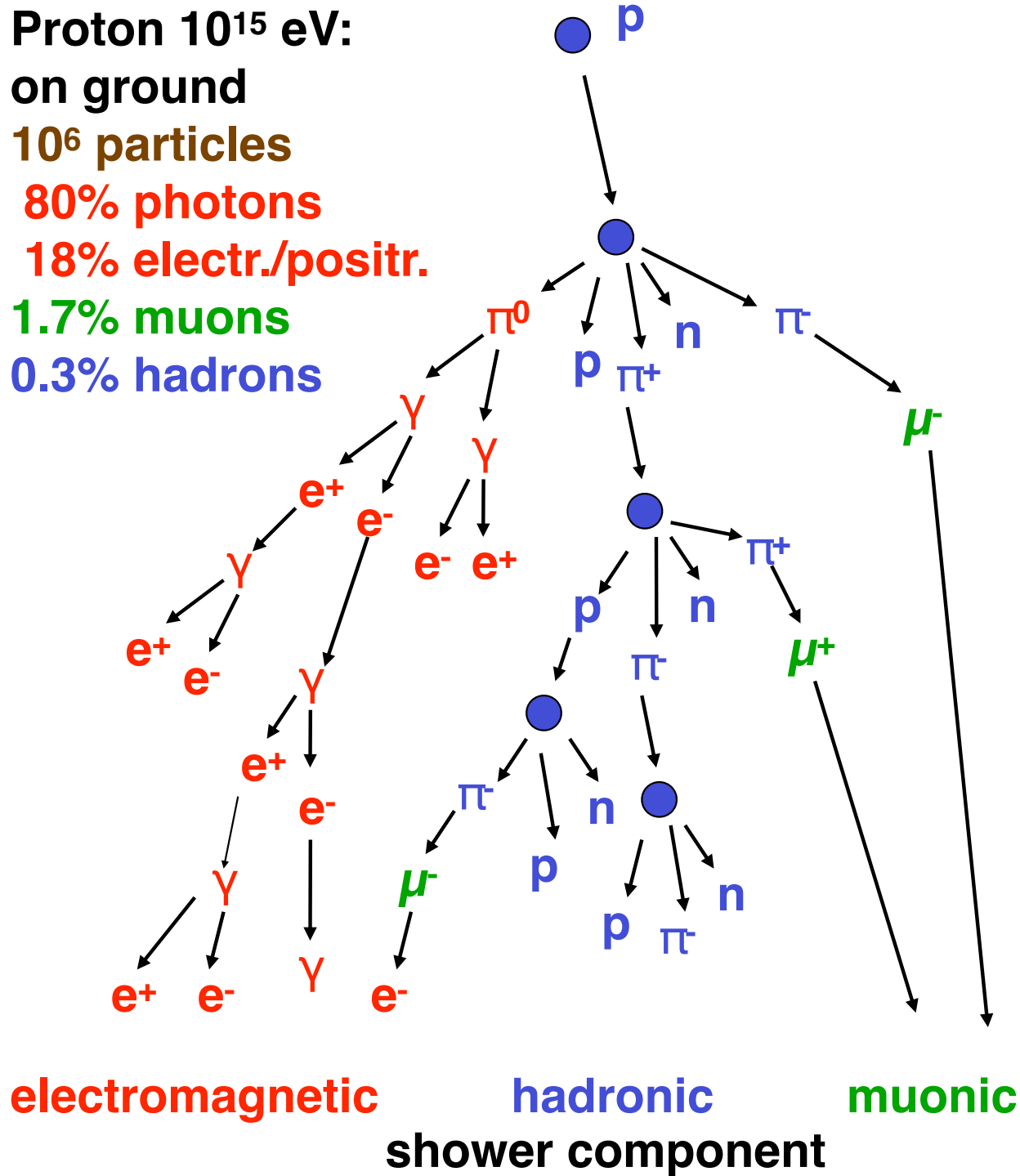
10^6 particles

80% photons

18% electr./positr.

1.7% muons

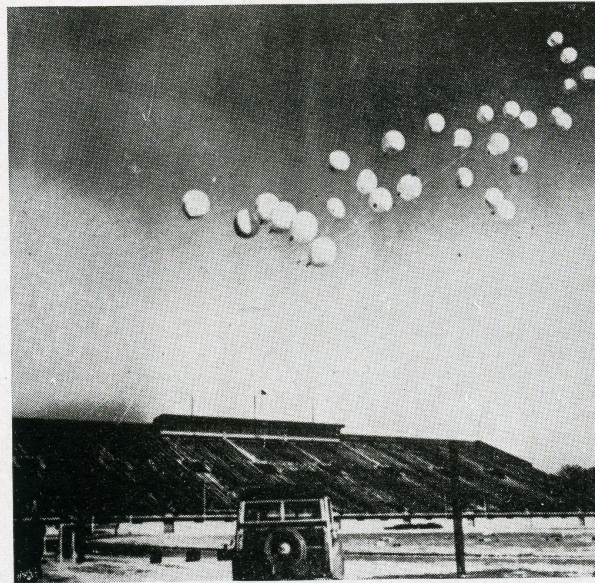
0.3% hadrons



1943

The University of Chicago

A



B



C
P. Auger

BALLOON FLIGHT OF JANUARY, 1943, CONDUCTED BY THE AUTHOR, SCHEIN, AND ROGOZINSKI FOR THE MEASUREMENT OF EXTENSIVE (OR AUGER-) SHOWERS IN THE STRATOSPHERE

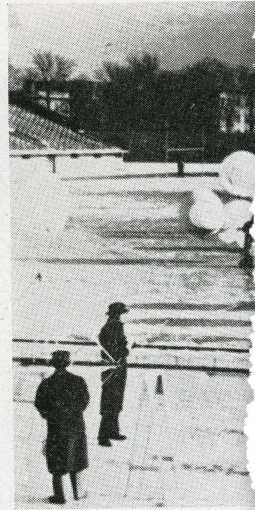
A. The balloons are assembled on Stagg Field at the University of Chicago, Chicago, Illinois. In the foreground can be seen the long frame which was required for the wide separation of the cosmic-ray counters.

B. The large cluster of balloons as it is about to be released.

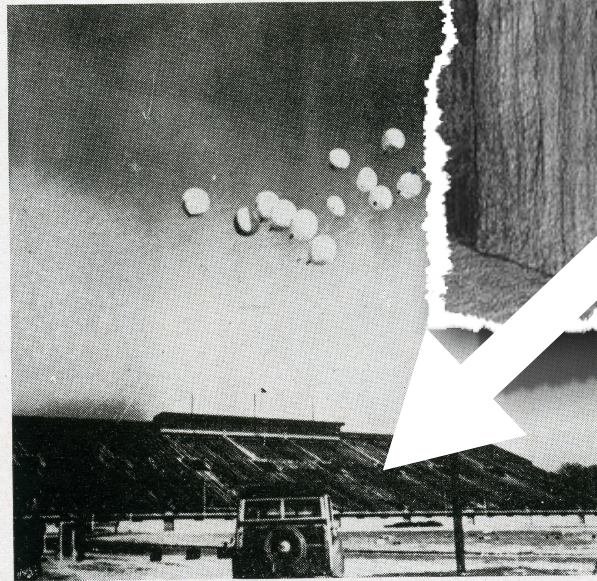
C. The balloon train sails into the sky after its release. Suspended below the balloons is the frame supporting the counters and recording apparatus.

1943

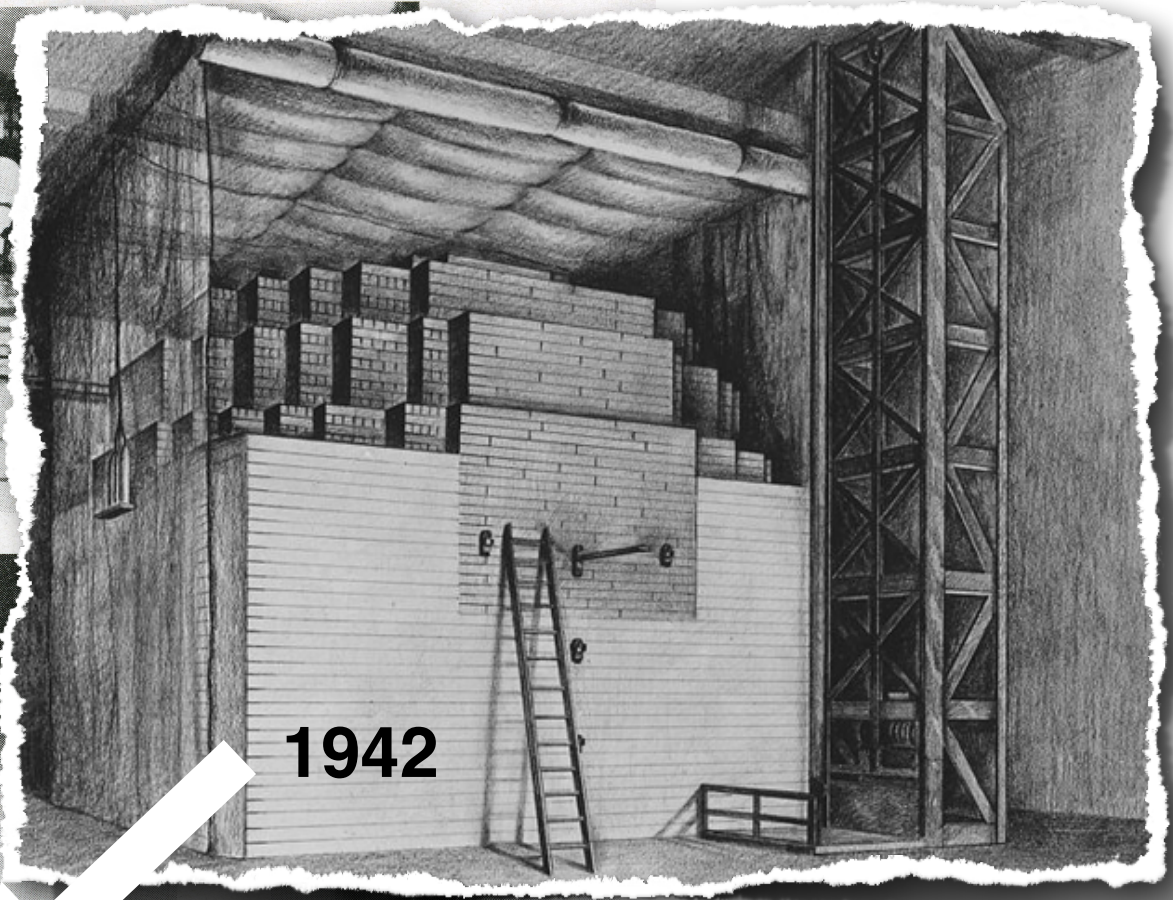
The University of Chicago



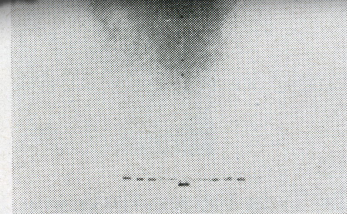
A



B



1942



C
P. Auger

BALLOON FLIGHT OF JANUARY, 1943, CONDUCTED BY THE AUTHOR, SCHEIN, AND ROGOZINSKI FOR THE MEASUREMENT OF EXTENSIVE (OR AUGER-) SHOWERS IN THE STRATOSPHERE

A. The balloons are assembled on Stagg Field at the University of Chicago, Chicago, Illinois. In the foreground can be seen the long frame which was required for the wide separation of the cosmic-ray counters.

B. The large cluster of balloons as it is about to be released.

C. The balloon train sails into the sky after its release. Suspended below the balloons is the frame supporting the counters and recording apparatus.

~ 1950 large detector arrays to measure extensive air showers

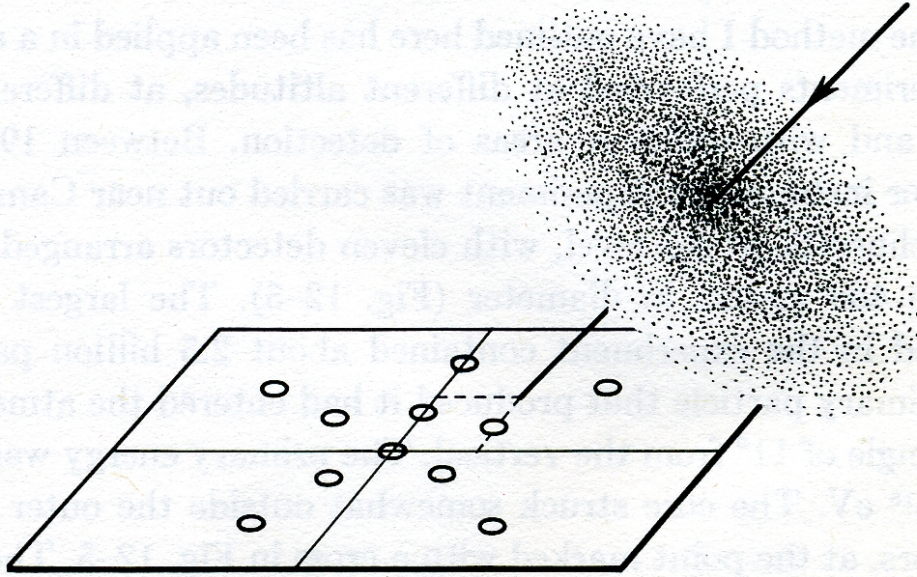


Fig. 12-4 Shower disk approaching detectors (represented by circles on a horizontal plane).

~ 1950 large detector arrays to measure extensive air showers

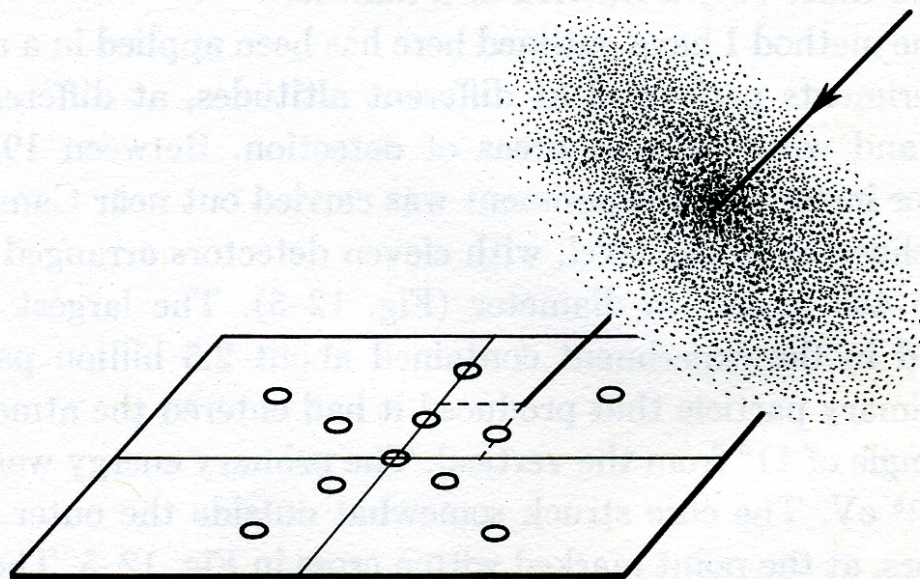


Fig. 12-4 Shower disk approaching detectors (represented by circles on a horizontal plane).

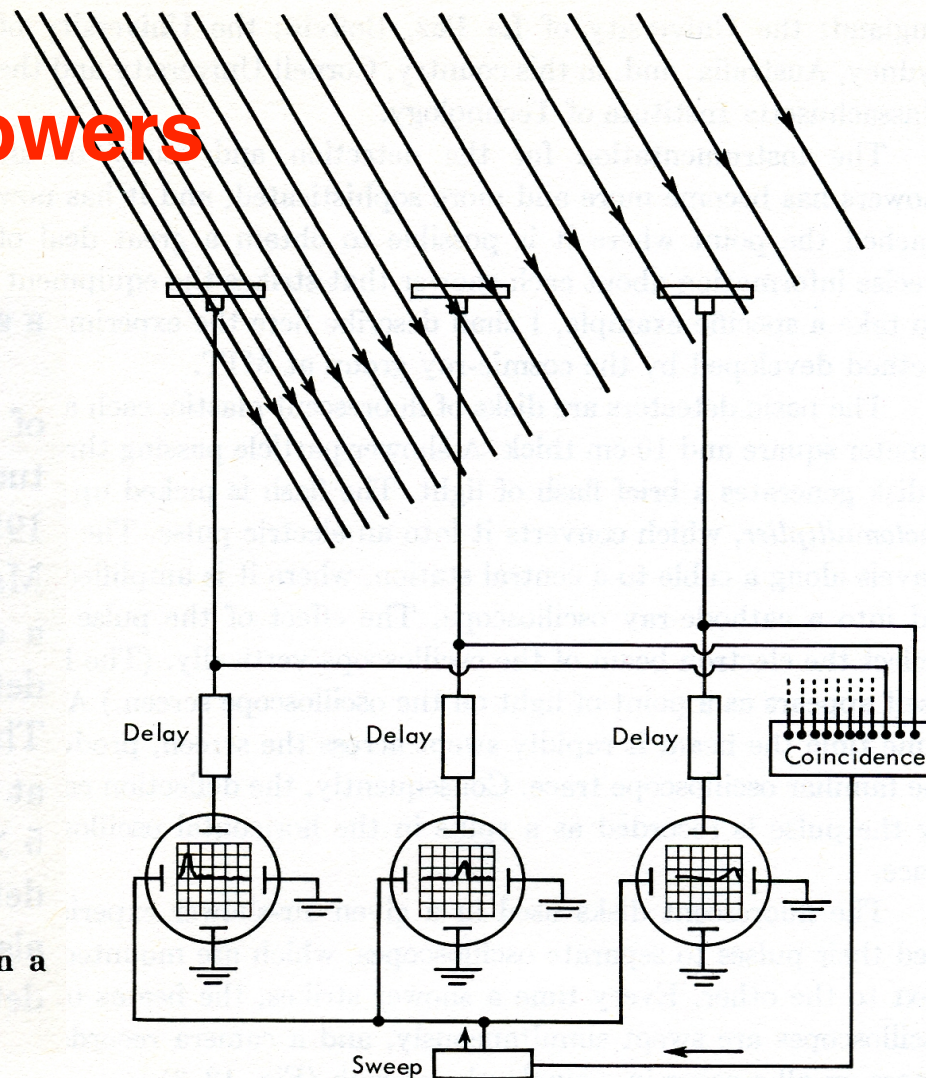


Fig. 12-3 Experimental arrangement used by the MIT cosmic-ray group to study air showers. Fluorescent plastic disks (thin rectangles at top) emit flashes of light when struck by charged particles. At the center of each disk is a photomultiplier tube that converts the light into an electrical pulse; the amplitude of the pulse is proportional to the brightness of the flash. Pulses travel to cathode-ray oscilloscopes (circles) through transmission lines containing delay circuits, which equalize the lengths of the electrical paths. Horizontal sweeps of all oscilloscope screens (grids) are triggered at the same time whenever three or more pulses pass through the coincidence circuit simultaneously. The amplitudes of the “spikes” (that is, the heights of the vertical deflections in the oscilloscope traces) indicate the numbers of particles striking the corresponding detectors. The positions of the spikes in the horizontal traces show the relative arrival times of the particles.

EVIDENCE FOR A PRIMARY COSMIC-RAY PARTICLE WITH ENERGY 10^{20} eV†

John Linsley

Laboratory for Nuclear Science, Massachusetts Institute of Technology, Cambridge, Massachusetts

(Received 10 January 1963)

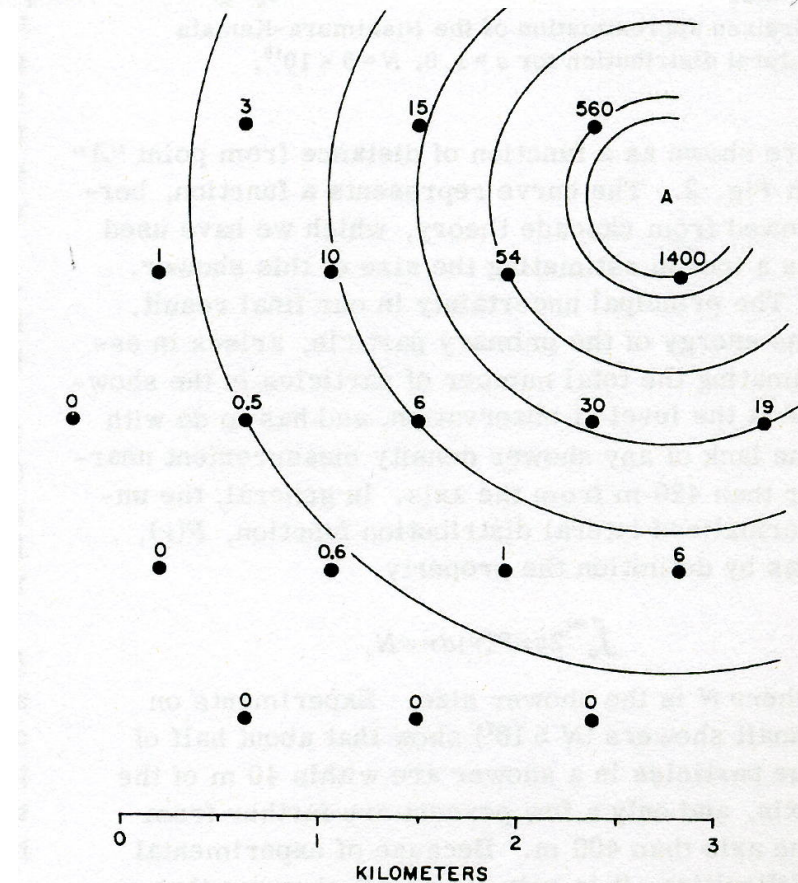


FIG. 1. Plan of the Volcano Ranch array in February 1962. The circles represent 3.3-m^2 scintillation detectors. The numbers near the circles are the shower densities (particles/ m^2) registered in this event, No. 2-4834. Point "A" is the estimated location of the shower core. The circular contours about that point aid in verifying the core location by inspection.

emulsion chambers at high-altitude lab above Innsbruck (Austria)

Disintegration Processes by Cosmic Rays with the Simultaneous Emission of Several Heavy Particles

ON photographic plates which had been exposed to cosmic radiation on the Hafelekar (2,300 m. above sea-level) near Innsbruck for five months, we found, apart from the very long tracks (up to 1,200 cm. in length) which have been reported recently in a note in the Wiener Akademie-Berichte, evidence of several processes described below.

From a single point within the emulsion several tracks, some of them having a considerable length, take their departure. We observed four cases with three particles, four with four and 'stars' with six, seven, eight and nine particles, one of each kind.

The longest track corresponded to a range in air (15°, 760 mm. Hg) of 176 cm. The ionization produced by the particles is different in the different cases. Most of the tracks show much larger mean grain-distances than α -particles and slow protons.

In Fig. 1 a 'star' with eight tracks is reproduced. On account of the rather steep angles at which some of the particles cross the emulsion-layer (approximately 70 μ thick) it is not possible to have all the tracks of a 'star' in focus simultaneously. Fig. 2 shows a sketch of the same 'star'. Measurement of the tracks gives the results in the accompanying table.

Track	Length in cm. of air (15°, 760 mm.)	Number of grains	Position of the end of the track
A	30.0 cm.	113	Within the emulsion
B	11.0 "	15	" " "
C	44.6 "	71	Glass
D	6.2 "	11	"
E	7.0 "	22	"
F	1.2 "	5	Within the emulsion
G	13.6 "	67	Surface of the emulsion
H	23.9 "	58	Glass

Centre of the 'star' 25 μ under the surface of the emulsion.

We believe that the process in question is a disintegration of an atom in the emulsion (probably Ag or Br) by a cosmic ray. The striking feature

about it is the simultaneous emission of so many heavy particles with such long ranges, which excludes any confusion with 'stars' due to radioactive contamination. A similar configuration of tracks by chance is equally out of question. Brode and others¹

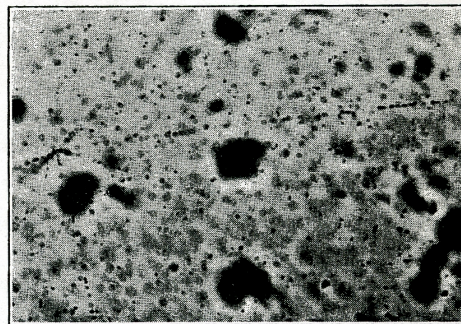


FIG. 1.

observed a single case of a disintegration with three heavy particles in a Wilson cloud chamber. The phenomenon which Wilkins believes was a shower of protons is perhaps a similar process, but he did not observe a centre².

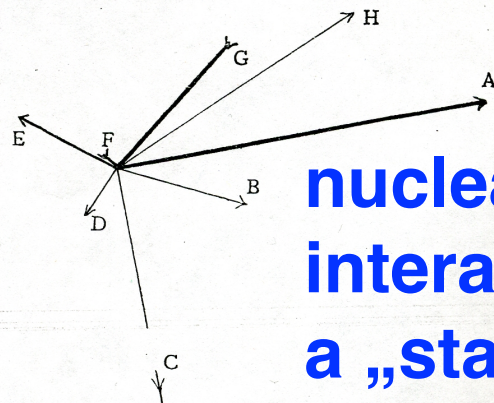


FIG. 2.

THICK LINES INDICATE A COMPARATIVELY LARGE NUMBER OF GRAINS PER UNIT OF LENGTH OF THE TRACK. AN INTERRUPTED LINE MEANS THAT THE TRACK IS TOO LONG TO BE REPRODUCED ON THE SAME SCALE. THE ARROWS INDICATE THE DIRECTION FROM THE SURFACE OF THE EMULSION TO THE GLASS.

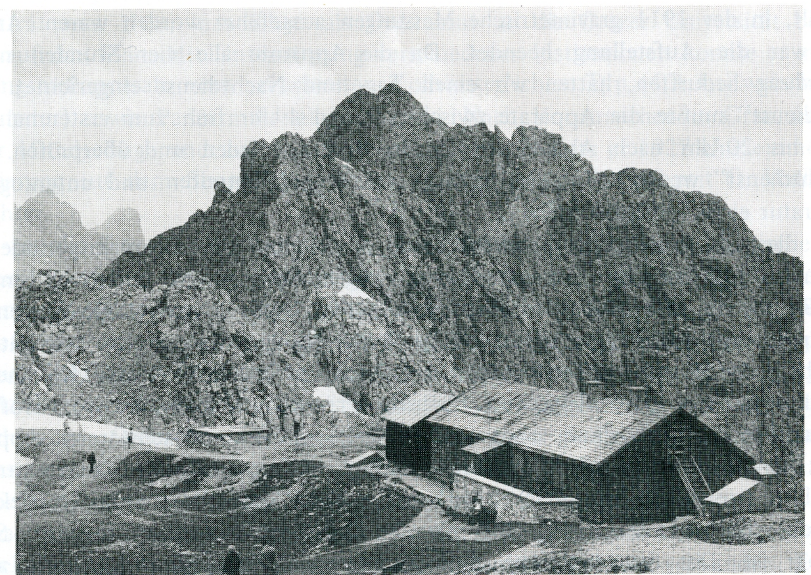
The total energy involved in the process cannot as yet be calculated as most of the particles do not end in the emulsion.

We hope to give further details before long in the Wiener Akademie-Berichte.

M. BLAU.
H. WAMBACHER.

Radium Institut
u. 2 Physik. Institut,
Wien.
Aug. 25.

¹ Brode, R. L., and others, *Phys. Rev.*, 50, 581 (October, 1936).
² Wilkins, *Nat. Geog. Soc.*, Stratosphere Series, No. 2, 37 (1936).



Die "Station für Ultrastrahlenforschung" auf dem Hafelekar bei Innsbruck (2300 m), 1960, vor dem späteren Ausbau.

nuclear interactions a „star“



emulsion chambers at high-altitude lab above Innsbruck (Austria)

about it is the simultaneous emission of so many heavy particles with such long ranges, which excludes any confusion with 'stars' due to radioactive contamination. A similar configuration of tracks by chance is equally out of question. Brode and others¹

Disintegration Processes by Cosmic Rays with the Simultaneous Emission of Several Heavy Particles

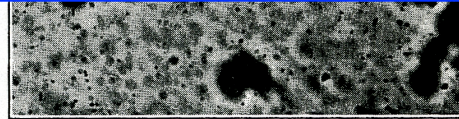


FIG. 1.

observed a single case of a disintegration with three heavy particles in a Wilson cloud chamber. The phenomenon which Wilkins believes was a shower of protons is perhaps a similar process, but he did not observe a centre².



Die "Station für Ultrastrahlenforschung" auf dem Hafelekar bei Innsbruck (2300 m), 1960, vor dem späteren Ausbau.

Disintegration Processes by Cosmic Rays with the Simultaneous Emission of Several Heavy Particles

ON photographic plates which had been exposed to cosmic radiation on the Hafelekar (2,300 m. above sea-level) near Innsbruck for five months, we found, apart from the very long tracks (up to 1,200 cm. in length) which have been reported recently in a note in the Wiener Akademie-Berichte, evidence of several processes described below.

From a single point within the emulsion several tracks, some of them having a considerable length take their departure. We observed four cases with three particles, four with four and 'stars' with seven, eight and nine particles, one of each kind.

The longest track corresponded to a range in (15°, 760 mm. Hg) of 176 cm. The ionization produced by the particles is different in the different cases. Most of the tracks show much larger mean grain-distances than α -particles and slow protons.

In Fig. 1 a 'star' with eight tracks is reproduced. On account of the rather steep angles at which some of the particles cross the emulsion-layer (approximately 70 μ thick) it is not possible to have all the tracks of a 'star' in focus simultaneously. Fig. 2 shows a sketch of the same 'star'. Measurement of the tracks gives the results in the accompanying table.

Track	Length in cm. of air (15°, 760 mm.)	Number of grains	Position of the end of the track
A	30.0 "	113	Within the emulsion
B	11.0 "	15	" " "
C	44.6 "	71	Glass
D	6.2 "	11	"
E	7.0 "	22	"
F	1.2 "	5	Within the emulsion
G	13.6 "	67	Surface of the emulsion
H	23.9 "	58	Glass

Centre of the 'star' 25 μ under the surface of the emulsion.

We believe that the process in question is a disintegration of an atom in the emulsion (probably Ag or Br) by a cosmic ray. The striking feature



M. BLAU.
H. WAMBACHER.

Radium Institut
u. 2 Physik. Institut,
Wien.
Aug. 25.

TRACK. AN INTERRUPTED LINE MEANS THAT THE TRACK IS TOO LONG TO BE REPRODUCED ON THE SAME SCALE. THE ARROWS INDICATE THE DIRECTION FROM THE SURFACE OF THE EMULSION TO THE GLASS.

The total energy involved in the process cannot as yet be calculated as most of the particles do not end in the emulsion.

We hope to give further details before long in the Wiener Akademie-Berichte.

M. BLAU.
H. WAMBACHER.

Radium Institut
u. 2 Physik. Institut,
Wien.
Aug. 25.

¹ Brode, R. L., and others, *Phys. Rev.*, 50, 581 (October, 1936).

² Wilkins, *Nat. Geog. Soc.*, Stratosphere Series, No. 2, 37 (1936).



emulsion chambers at high-altitude lab above Innsbruck (Austria)

Disintegration Processes by Cosmic Rays with the Simultaneous Emission of Several Heavy Particles

ON photographic plates which had been exposed to cosmic radiation on the Hafelekar (2,300 m. above sea-level) near Innsbruck for five months, we found, apart from the very long tracks (up to 1,200 cm. in length) which have been reported recently in a note in the Wiener Akademie-Berichte, evidence of several processes described below.

From a single point within the emulsion several tracks, some of them having a considerable length, take their departure. We observed four cases with three particles, four with four and 'stars' with six, seven, eight and nine particles, one of each kind.

The longest track corresponded to a range in air (15°, 760 mm. Hg) of 176 cm. The ionization produced by the particles is different in the different cases. Most of the tracks show much larger mean grain-distances than α -particles and slow protons.

In Fig. 1 a 'star' with eight tracks is reproduced. On account of the rather steep angles at which some of the particles cross the emulsion-layer (approximately 70 μ thick) it is not possible to have all the tracks of a 'star' in focus simultaneously. Fig. 2 shows a sketch of the same 'star'. Measurement of the tracks gives the results in the accompanying table.

Track	Length in cm. of air (15°, 760 mm.)	Number of grains	Position of the end of the track
A	30.0 cm.	113	Within the emulsion
B	11.0 "	15	" " "
C	44.6 "	71	Glass
D	6.2 "	11	"
E	7.0 "	22	"
F	1.2 "	5	Within the emulsion
G	13.6 "	67	Surface of the emulsion
H	23.9 "	58	Glass

Centre of the 'star' 25 μ under the surface of the emulsion.

We believe that the process in question is a disintegration of an atom in the emulsion (probably Ag or Br) by a cosmic ray. The striking feature

about it is the simultaneous emission of so many heavy particles with such long ranges, which excludes any confusion with 'stars' due to radioactive contamination. A similar configuration of tracks by chance is equally out of question. Brode and others¹

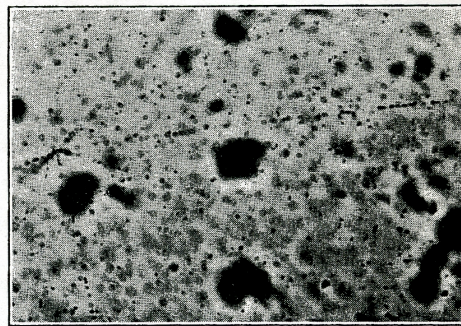


FIG. 1.

observed a single case of a disintegration with three heavy particles in a Wilson cloud chamber. The phenomenon which Wilkins believes was a shower of protons is perhaps a similar process, but he did not observe a centre².

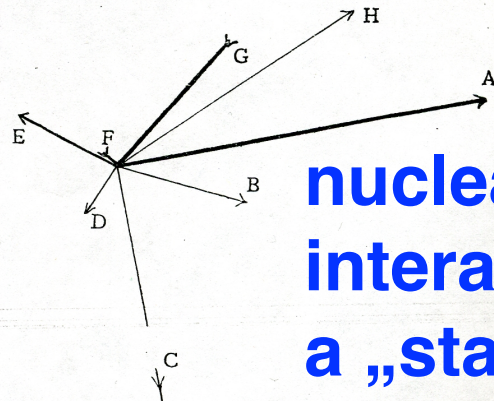


FIG. 2.

THICK LINES INDICATE A COMPARATIVELY LARGE NUMBER OF GRAINS PER UNIT OF LENGTH OF THE TRACK. AN INTERRUPTED LINE MEANS THAT THE TRACK IS TOO LONG TO BE REPRODUCED ON THE SAME SCALE. THE ARROWS INDICATE THE DIRECTION FROM THE SURFACE OF THE EMULSION TO THE GLASS.

The total energy involved in the process cannot as yet be calculated as most of the particles do not end in the emulsion.

We hope to give further details before long in the Wiener Akademie-Berichte.

M. BLAU.
H. WAMBACHER.

Radium Institut
u. 2 Physik. Institut,
Wien.
Aug. 25.

¹ Brode, R. L., and others, *Phys. Rev.*, 50, 581 (October, 1936).
² Wilkins, *Nat. Geog. Soc.*, Stratosphere Series, No. 2, 37 (1936).



Die "Station für Ultrastrahlenforschung" auf dem Hafelekar bei Innsbruck (2300 m), 1960, vor dem späteren Ausbau.

nuclear interactions a „star“



Tracks of Nuclear Particles in Photographic Emulsions

MAURICE M. SHAPIRO

Ryerson Laboratory, University of Chicago, Chicago, Illinois

CONTENTS

I. Early history of the direct photographic method	58
II. Nature of the photographic technique—its advantages and limitations	61
III. Contributions of the photographic method in the field of cosmic rays	63
IV. Contributions of the photographic method to other problems in nuclear physics	68

1947 Discovery of the Pion

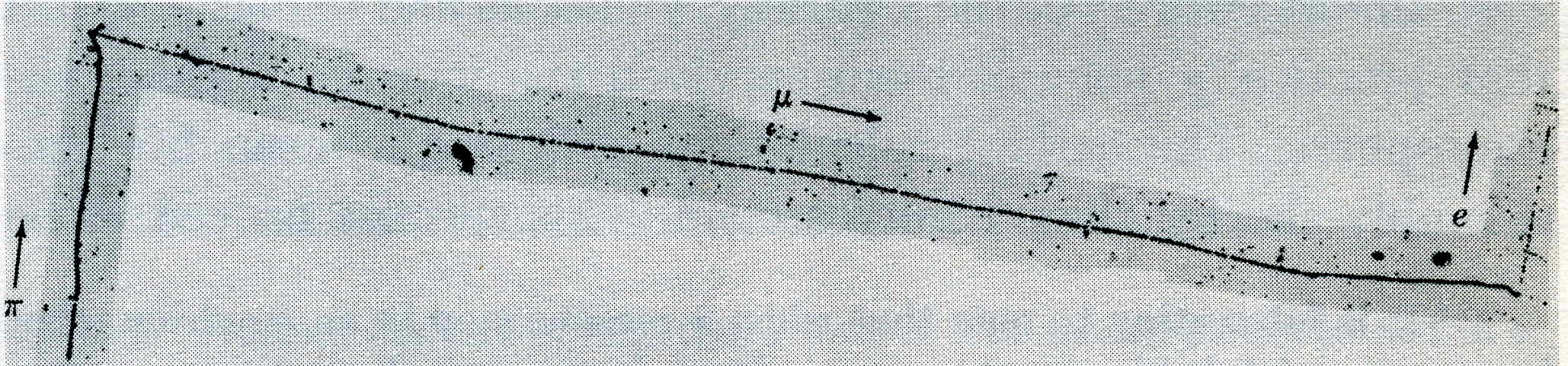


Fig. 9-4 Photomicrograph of tracks in a nuclear emulsion, showing a π meson (π) that comes to rest and decays into a μ meson (μ). The μ meson in turn comes to rest and decays into an electron (e). (From R. H. Brown, U. Camerini, P. Fowler, H. Muirhead, C. F. Powell, and D. M. Ritson, *Nature*, vol. 163, p. 47, 1949.)

C.F. Powell
Nobel Prize 1950

$$m_{\pi} \sim 280 m_e$$

Pion: nuclear interaction

$$\text{decay } \pi^{+/-} \rightarrow \mu^{+/-} \rightarrow e^{+/-}$$

$$\pi^0 \rightarrow \gamma\gamma$$



End 1940s plastic balloons



Fig. 1. Inflation of balloon of polyethylene just after dawn. The balloon has a total length of about 120 ft. and most of the fabric is on the ground. Such a balloon can in favourable conditions give level flight at about 90,000 ft. for many hours with a load of 40 kg.

1941 protons (M. Schein)

End 1940s plastic balloons



Fig. 1. Inflation of balloon of polyethylene just after dawn. The balloon has a total length of about 120 ft. and most of the fabric is on the ground. Such a balloon can in favourable conditions give level flight at about 90,000 ft. for many hours with a load of 40 kg.

1941 protons (M. Schein)

1948 heavy nuclei (Brandt & Peters)

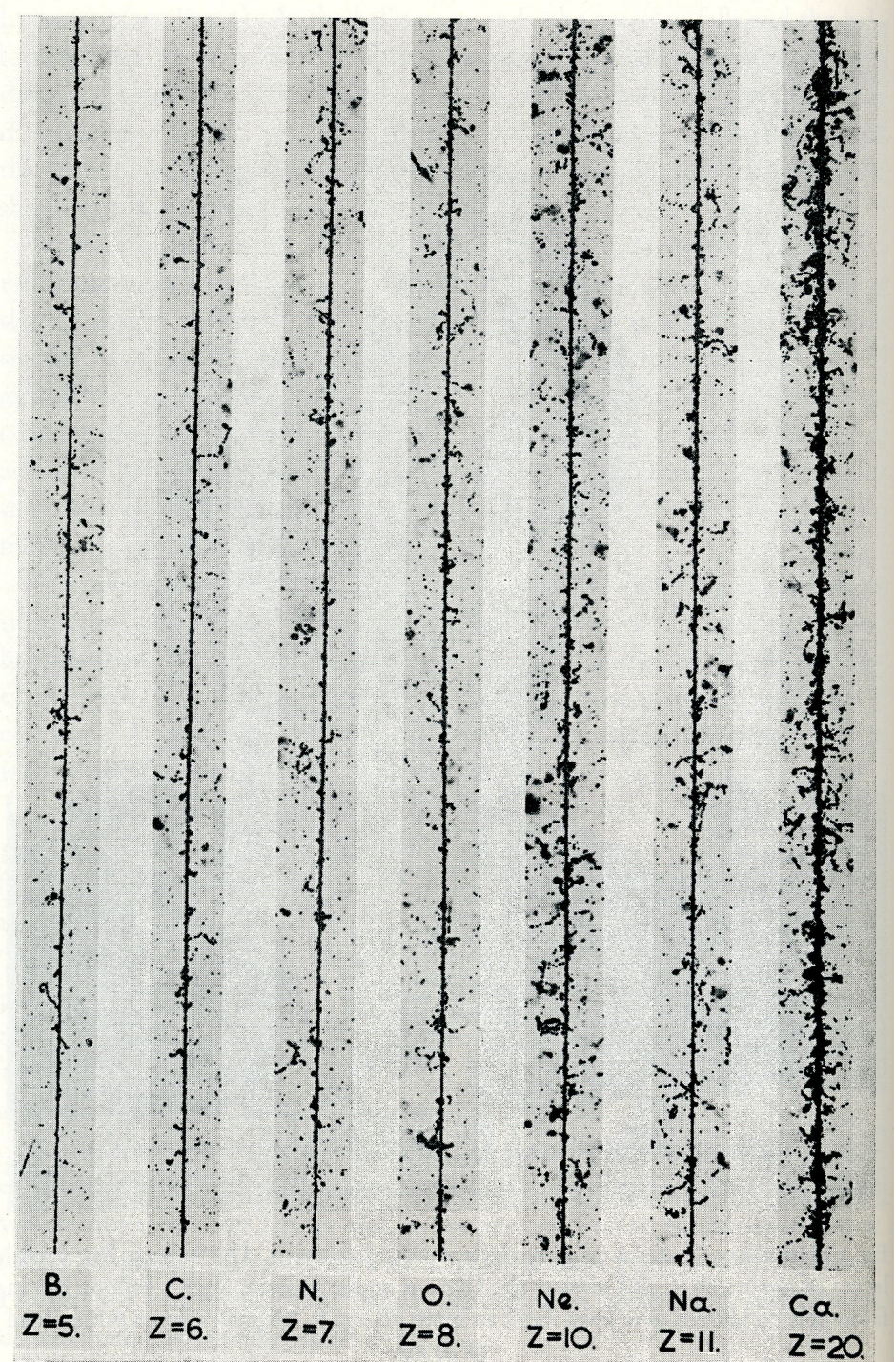


Fig. 2. Examples of the tracks in photographic emulsions of primary nuclei of the cosmic radiation moving at relativistic velocities.

The Cosmic-Ray Counting Rate of a Single Geiger Counter from Ground Level to 161 Kilometers Altitude

J. A. VAN ALLEN AND H. E. TATEL*

Applied Physics Laboratory, Johns Hopkins University, Silver Spring, Maryland

(Received October 16, 1947)

The cosmic-ray counting rate of a single Geiger counter has been measured from ground level to an altitude of 161 kilometers. The equipment was carried in a V-2 rocket at geomagnetic latitude $\lambda=41^\circ\text{N}$. Especial care was taken to avoid multiplicative effects from surrounding material. A value of the charged primary cosmic-ray flux of $j=0.12/\text{sec.}/\text{cm}^2/\text{steradian}$, averaged over the upper hemisphere, is implied by the data above 55 km. This interpretation of the counting rate must be qualified by the as yet unknown contribution from secondaries which emerge from the atmosphere and execute orbits in the earth's magnetic field.

1. INTRODUCTION

THIS is one of a series of reports on cosmic ray experiments conducted by this laboratory in flights of V-2 rockets during the past year and a half. The data presented herein were obtained during the flight of July 29, 1947.

* Now at Department of Terrestrial Magnetism, Carnegie Institution of Washington, Washington, D. C.

They are believed to be reliable but are provisional in the sense that they have been obtained in only one flight. A description of this work is thought to be worthwhile at this time for several reasons:

- (a) No previous comparable data are known to us.
- (b) The technique of using high altitude rockets as vehicles for scientific measurements is not as yet very generally known.

The Cosmic-Ray Counting Rate of a Single Geiger Counter from Ground Level to 161 Kilometers Altitude

J. A. VAN ALLEN AND H. E. TATEL*

Applied Physics Laboratory, Johns Hopkins University, Silver Spring, Maryland

(Received October 16, 1947)

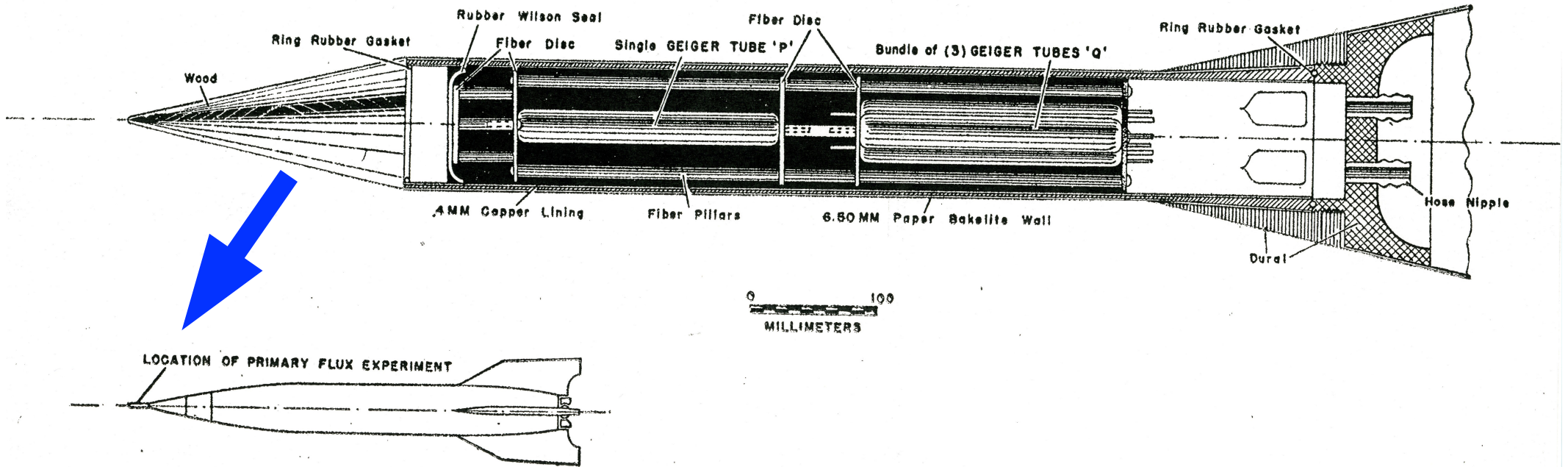


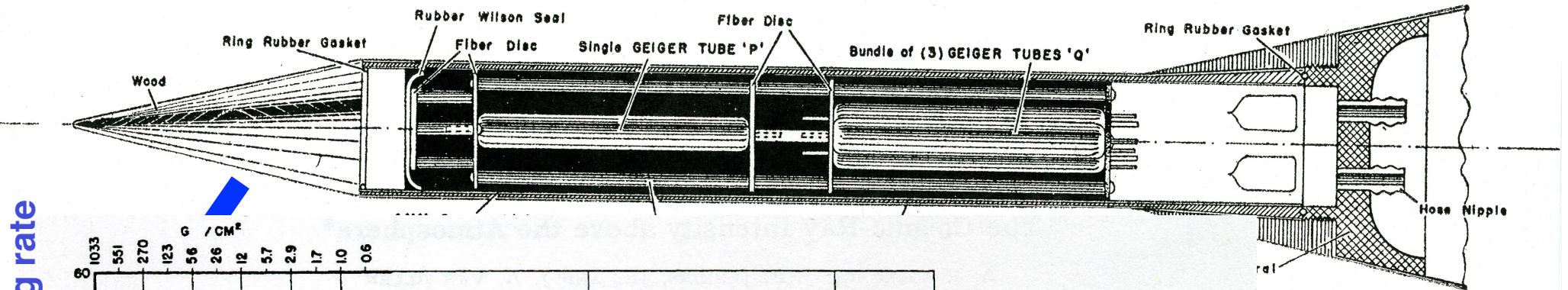
FIG. 1. Diagram of equipment, drawn to scale and showing location in V-2 rocket.

The Cosmic-Ray Counting Rate of a Single Geiger Counter from Ground Level to 161 Kilometers Altitude

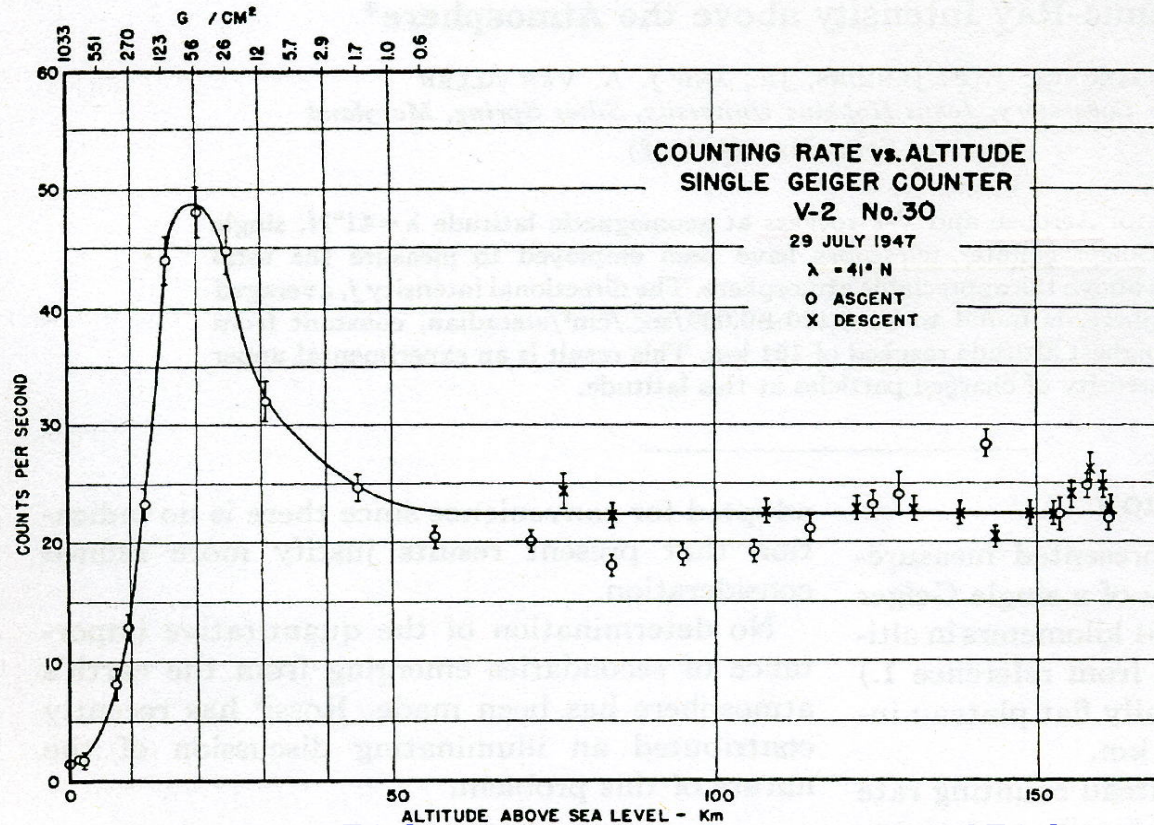
J. A. VAN ALLEN AND H. E. TATEL*

Applied Physics Laboratory, Johns Hopkins University, Silver Spring, Maryland

(Received October 16, 1947)



counting rate

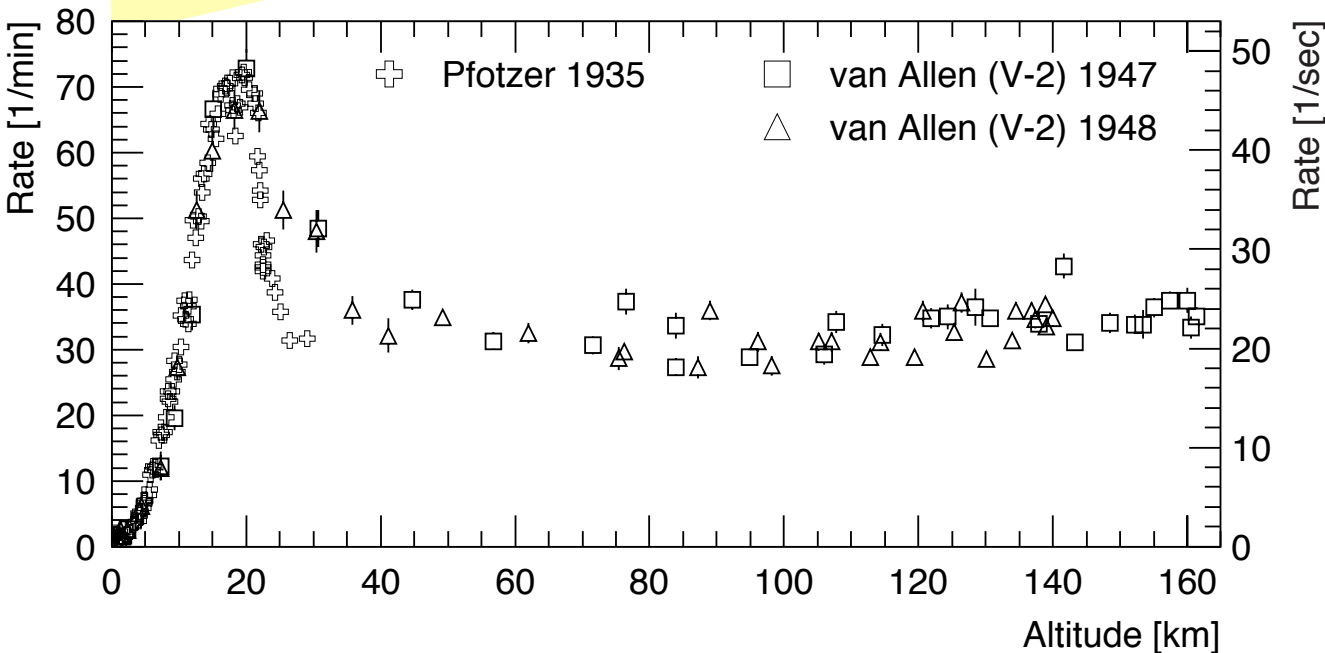
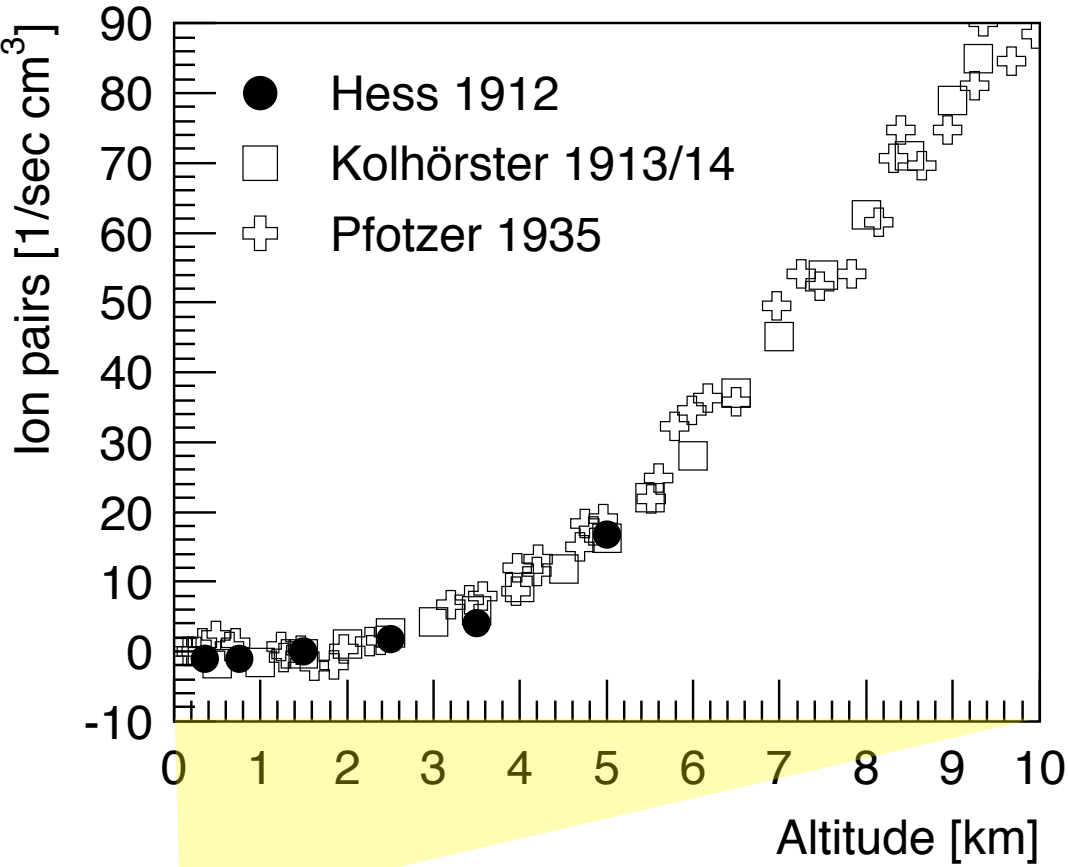


50 km

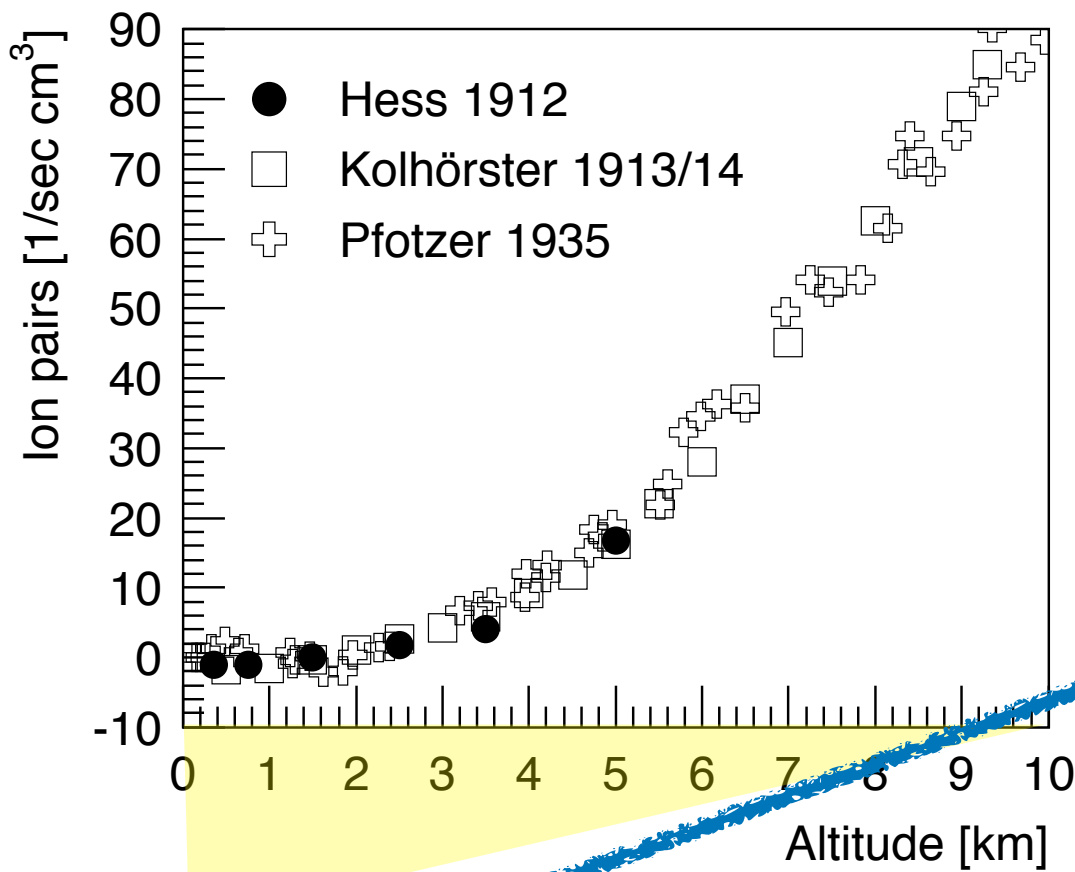
150 km

FIG. 1. Counting rate of a single Geiger counter as a function of altitude.

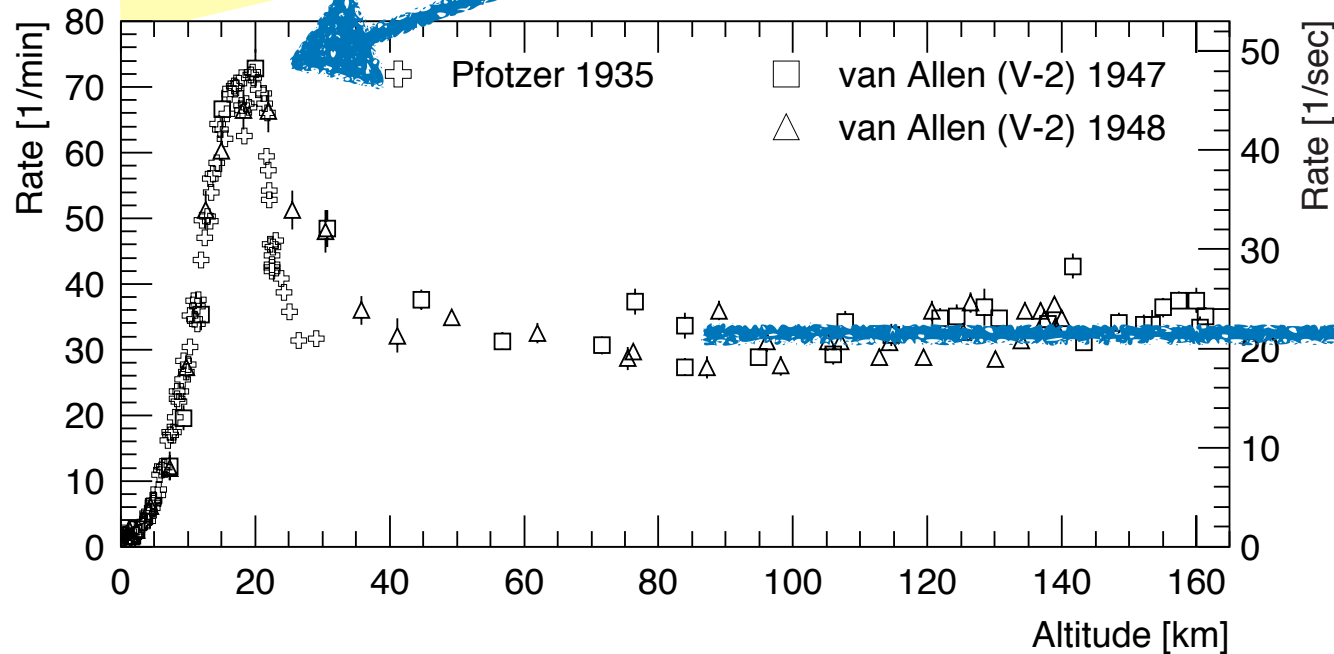
Intensity vs. height



Intensity vs. height



cosmic rays with
~GeV energies
initiate cascades in
the atmosphere



(galactic)
cosmic rays

Stars and Heavy Primaries Recorded during a V-2 Rocket Flight

HERMAN YAGODA, HERVASIO G. DE CARVALHO,* AND NATHAN KAPLAN
*Laboratory of Physical Biology, Experimental Biology and Medicine Institute,
 National Institutes of Health, Bethesda, Maryland*

(Received February 23, 1950)

Plates flown to an altitude of 150.7 km in a V-2 rocket exhibit a differential star population of 5000 ± 800 per cc per day and a flux of heavy primaries of about 0.03 per cm^2 per min. above the stratosphere. The star intensity is about 3.6 times greater than that recorded by plates exposed in the stratosphere, the increment being attributable to secondary star forming radiations created by interaction of cosmic-ray primaries with the massive projectile. The flux of heavy primaries is essentially of the same order of magnitude as reported for elevations of 28 km.

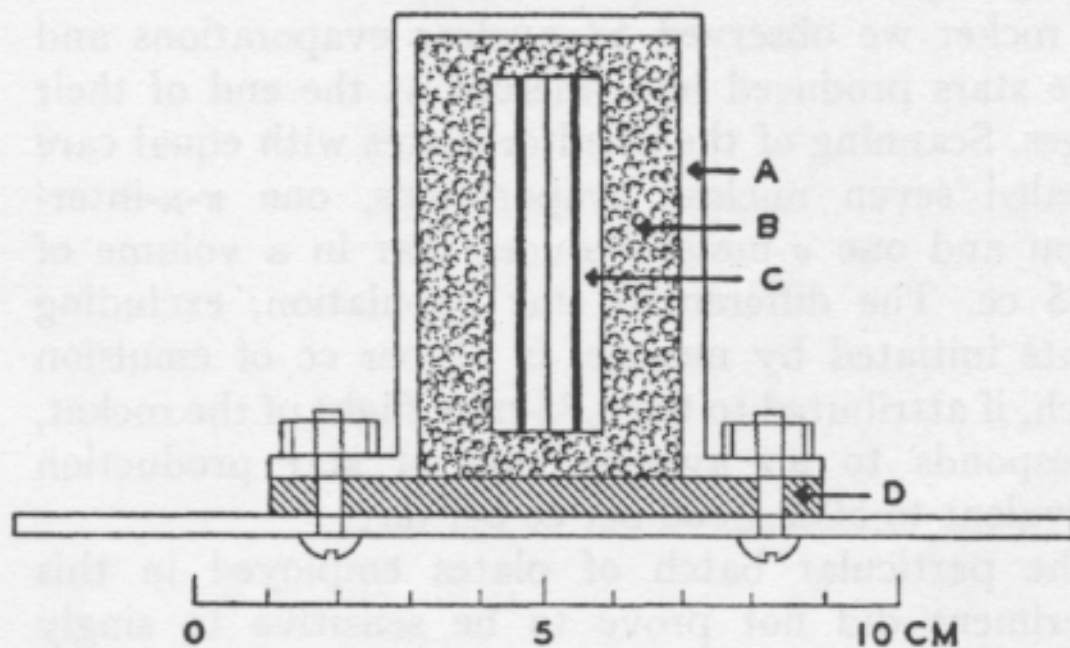


FIG. 1. Cross section of plate holder. A. Aluminum jacket 3 mm thick. B. Sponge rubber packing. C. Plates assembled with emulsion layers adjacent to each other. D. Rubber gasket.

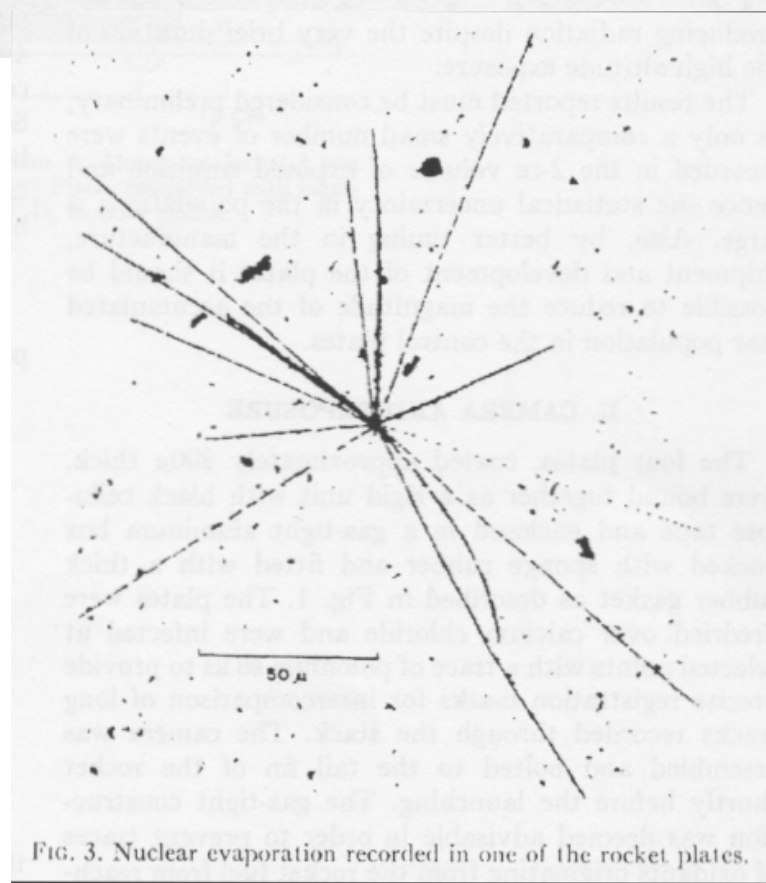


FIG. 3. Nuclear evaporation recorded in one of the rocket plates.

1953 Cosmic-Ray Conference

birth of particle physics

particles discovered in cosmic rays:

- 1932 e^+ Anderson
- 1937 μ Anderson/
Neddermeyer
- 1947 π Lattes,
Occhialini, Powell
- 1947 K Rochester,
Butcher, Powell
- 1951-53 hyperons

Λ Ξ Σ

CONGRÈS INTERNATIONAL
SUR LE
RAYONNEMENT COSMIQUE

ORGANISÉ PAR
L'UNIVERSITÉ DE TOULOUSE
SOUS LE PATRONAGE DE L'UIPPA
AVEC L'APPUI DE L'U.N.E.S.C.O.

BAGNÈRES DE BIGORRE JUILLET 1953

Rocket Determination of the Ionization Spectrum of Charged Cosmic Rays at $\lambda = 41^\circ\text{N}$

G. J. PERLOW,* L. R. DAVIS, C. W. KISSINGER, AND J. D. SHIPMAN, JR.
U. S. Naval Research Laboratory, Washington, D. C.

(Received June 30, 1952)

In a V-2 rocket measurement at $\lambda = 41^\circ\text{N}$ an analysis has been made of the various components of the charged particle radiation on the basis of ionization and absorption in lead. The ionization was determined by two proportional counters, the particle paths through which were defined by Geiger counters. With increasing zenith angle toward the north, the intensity is found to be substantially constant until the earth ceases to cover the under side of the telescope. The intensity of all particles with range $\geq 7 \text{ g/cm}^2$ is $0.079 \pm 0.005 \text{ (cm}^2 \text{ sec steradian)}^{-1}$. Of this an intensity 0.012 ± 0.002 is absorbed in the next 14 g/cm^2 . The ionization measurement is consistent with $\frac{2}{3}$ of these soft particles being electrons of $< \sim 60 \text{ Mev}$, the remainder being slow protons and alpha-particles. For the particles with greater range an ionization histogram is plotted, the smaller of the two ionization measurements for a single event being used to improve the resolution. The particles divide into protons, alpha-particles, and one carbon nucleus, with $N_p/N_\alpha = 5.3 \pm 1.0$. Their absorption is exponential with mean free path $440 \pm 70 \text{ g/cm}^2 \text{ Pb}$. Extrapolating to zero thickness, the total primary intensity is $0.070 \pm 0.005 \text{ (cm}^2 \text{ sec steradian)}^{-1}$ with 0.058 ± 0.005 as protons, 0.011 ± 0.002 as alpha-particles, and 0.001 ± 0.001 as $Z > 2$.

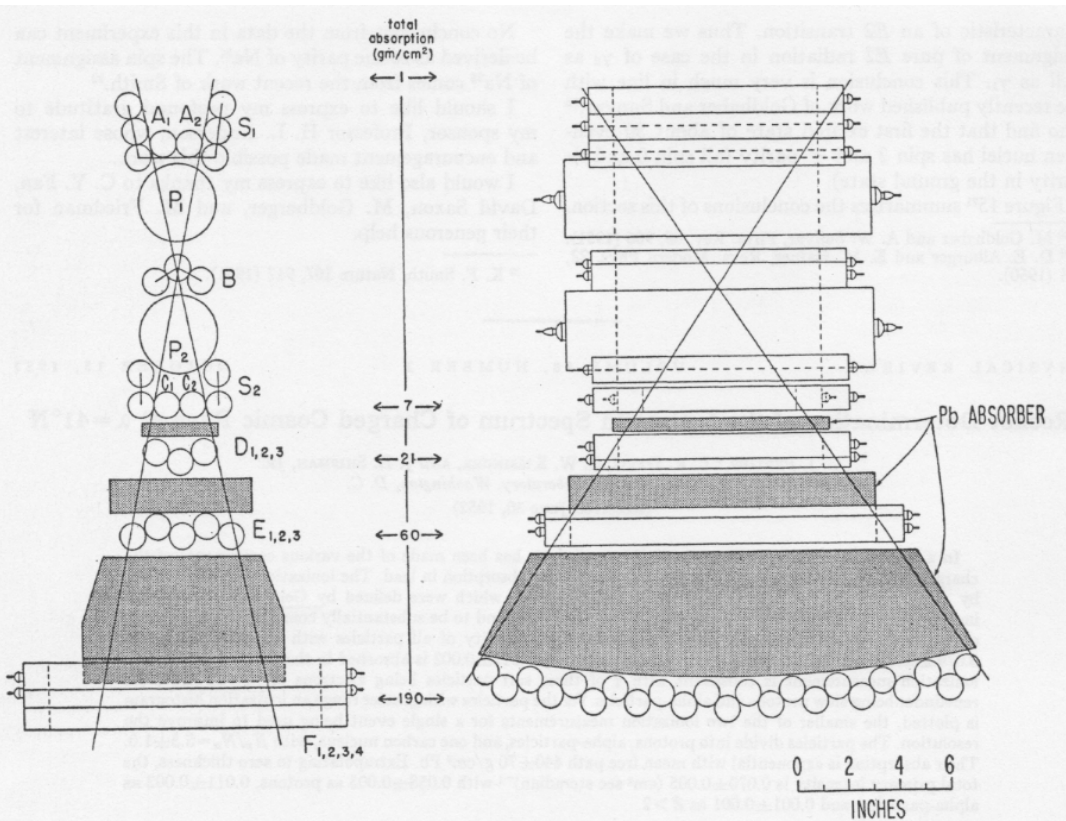


FIG. 1. Diagram of telescope.

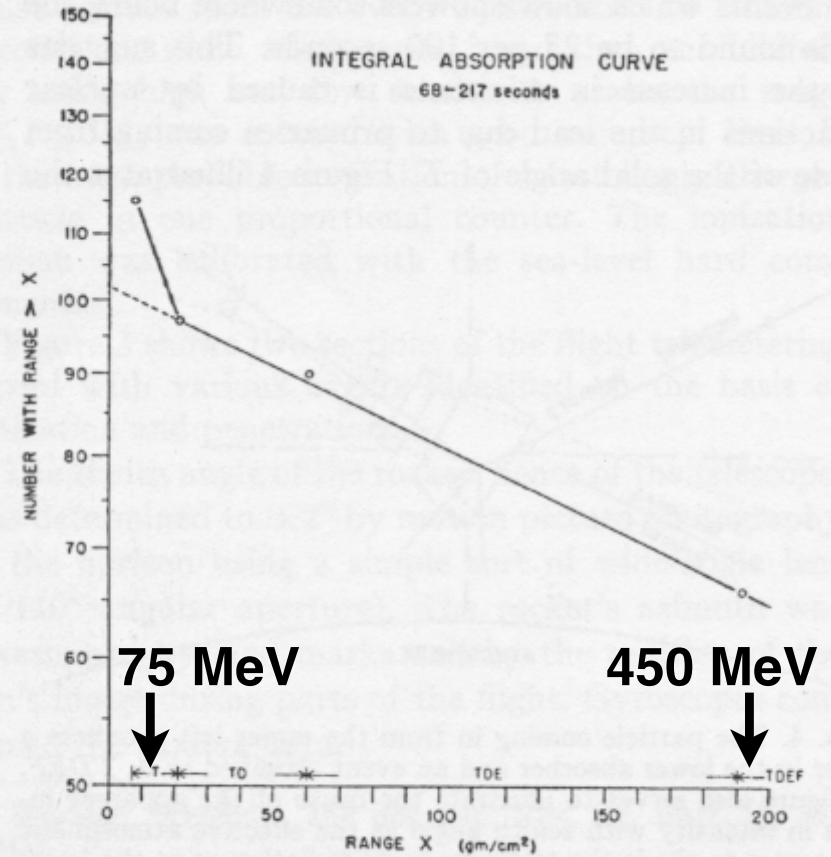


FIG. 6. Absorption in lead of the total radiation.

Van Allen Belts

Radiation Around the Earth to a Radial Distance of 107,400 km.

JAMES A. VAN ALLEN & LOUIS A. FRANK

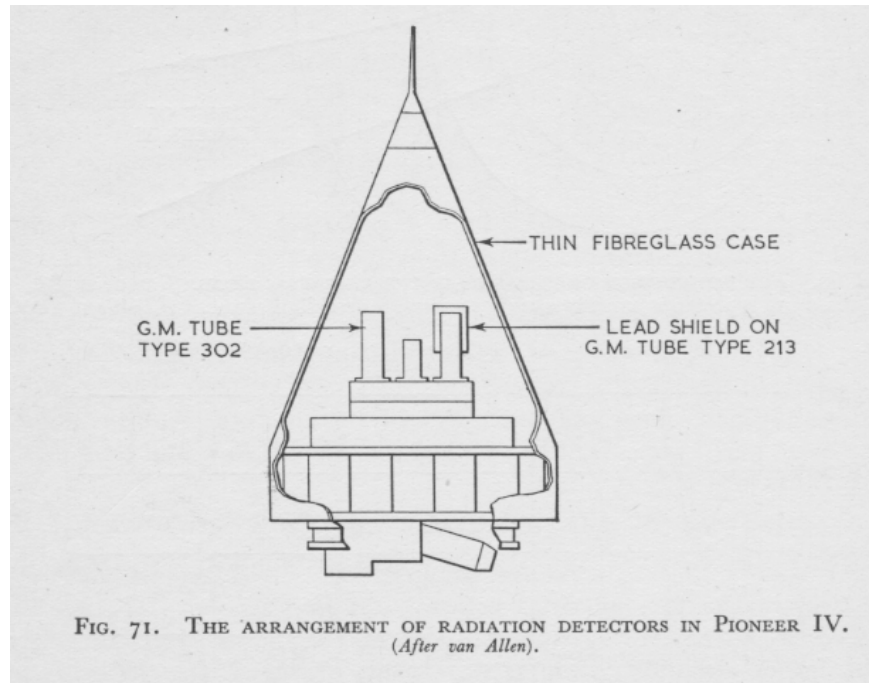


FIG. 71. THE ARRANGEMENT OF RADIATION DETECTORS IN PIONEER IV. (After van Allen).

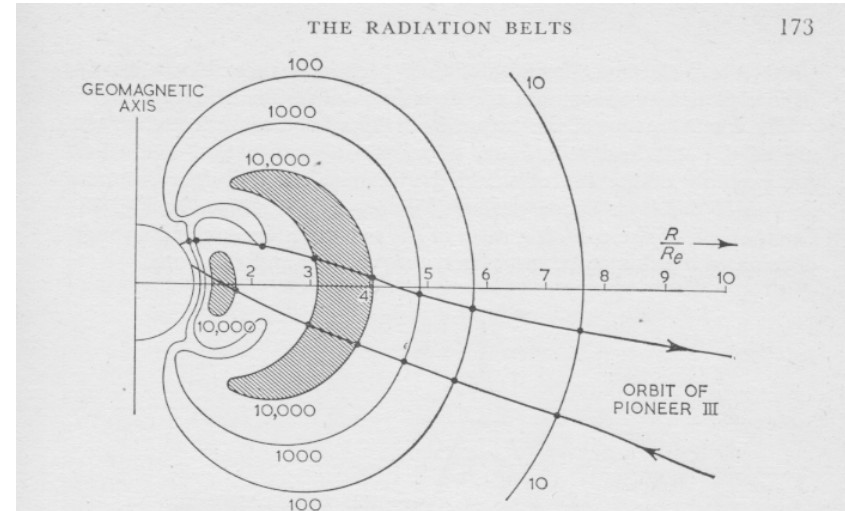


Fig. 69. THE DISTRIBUTION OF INTENSITY IN THE RADIATION BELTS. (6 DEC. 1958). The diagram represents a cross section through a meridian plane. R_e (~ 6400 km) is the radius of the earth. (After van Allen and Frank, *Nature*, 183, 430 (1959)).

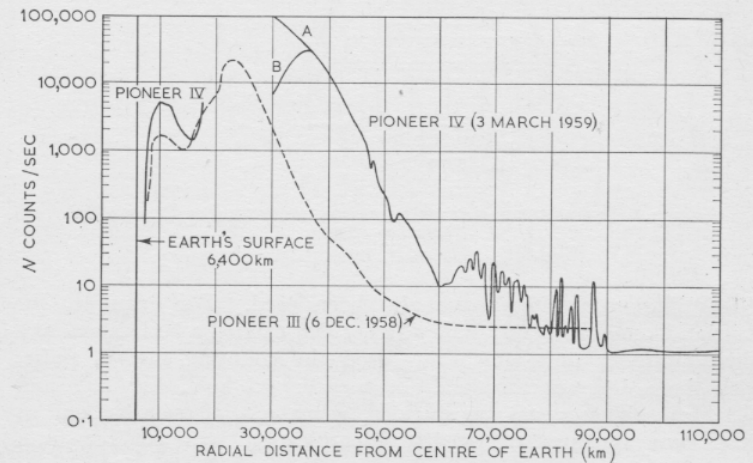


FIG. 70. A COMPARISON OF THE INTENSITIES OF RADIATION FOUND WITH NEARLY IDENTICAL COUNTERS IN PIONEER III AND PIONEER IV. The trajectories of the two probes were almost, but not quite, the same. At the peak of the second belt the readings of the intensity from Pioneer IV were ambiguous and followed either curve A or curve B. Curve A is more probable. (After van Allen and Frank, *Nature* 184, 219 (1960)).



John Simpson (Chicago)

**Precision
measurements of CR
abundances**

**dE/dx vs. E technique with
solid state detectors in
space**

1958 PIONEER 2

1959 EXPLORER 6

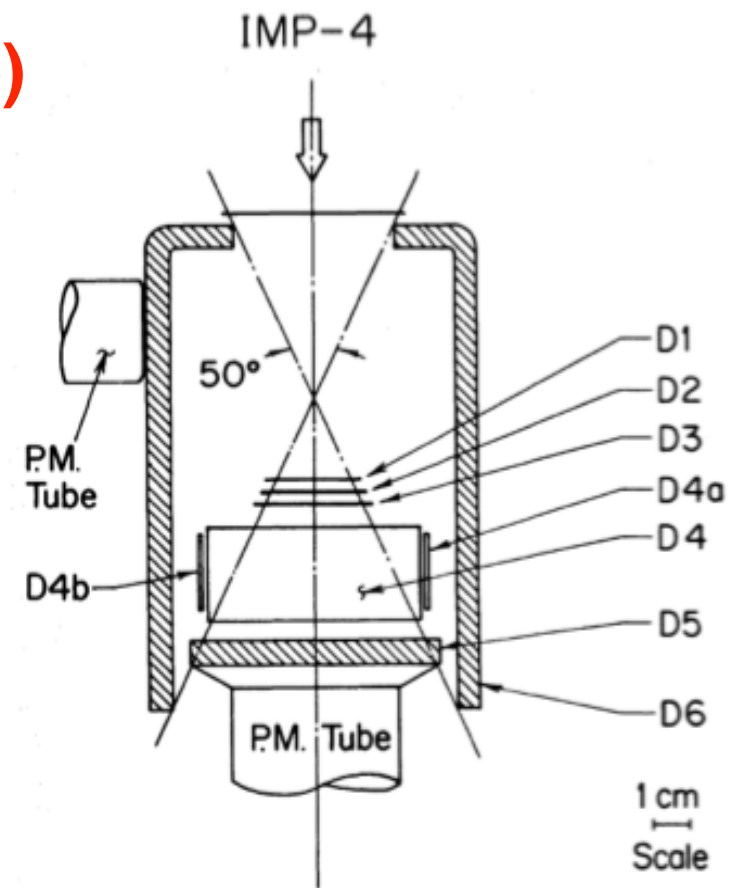
**subsequently, more than 20 other space missions
including: IMP1-8; OGO 1,3,5 - Earth orbit**

PIONEER 5,6,7 - Solar orbit;

PIONEER 10,11 - out of Solar System

ULYSSES - out of ecliptic plane (Jupiter flyby)

- **Elemental composition of cosmic rays**
- **Isotopic composition**
- **Measurement of anomalous cosmic rays**
- **Particles and fields in the Heliosphere**
- **Planetary magnetospheres**
- **Solar modulation to outer Heliosphere**





John Simpson

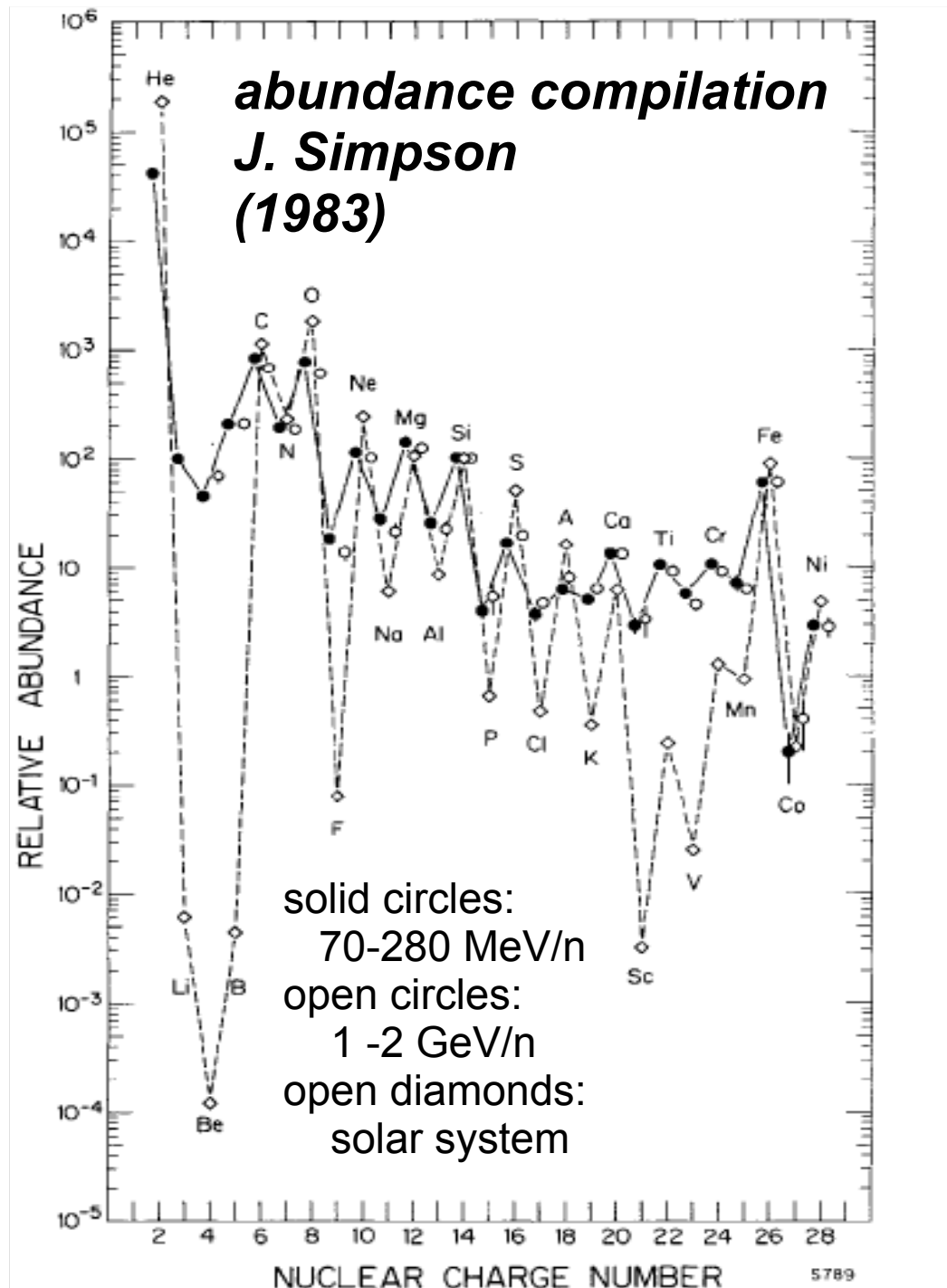
**Precise
measurements of
abundance
 dE/dx vs
solid state
space**

1958 PIONEER 2

1959 EXPLORER 6

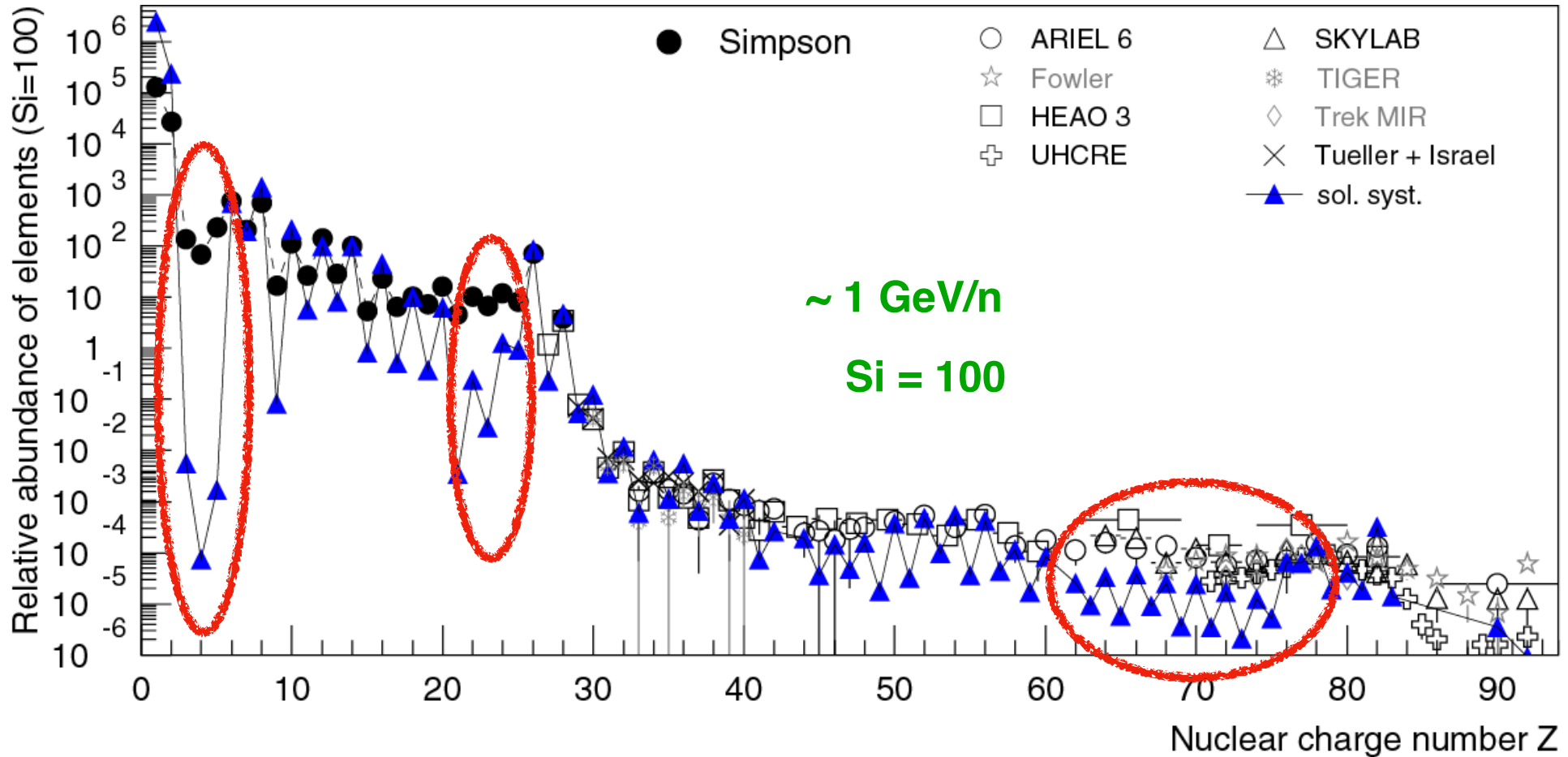
subsequently, more than 20 other missions
including: *IMP1-8; OGO 1,3,5*
PIONEER 5,6,7 - Solar Wind 3
PIONEER 10,11 - Voyager 1,2
ULYSSES - out of the solar system

- Elemental composition of cosmic rays
- Isotopic composition
- Measurement of anomalous light nuclei
- Particles and fields in the Heliosphere
- Planetary magnetospheres
- Solar modulation to outer Heliosphere



Formation of the chemical composition

Relative abundance of elements at Earth



JRH, Adv. Space Res. 41 (2008) 442

abundance of elements in CRs and solar system mostly similar

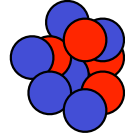
but few differences, e.g. Li, Be, B → important to understand propagation of cosmic rays in Galaxy → column density of traversed matter

primary cosmic rays generated at source e.g. p, He, Fe
spallation products → **secondary cosmic rays**, e.g. Li, Be, B

THE AGE OF THE GALACTIC COSMIC RAYS DERIVED
FROM THE ABUNDANCE OF ^{10}Be *

M. GARCIA-MUNOZ, G. M. MASON, AND J. A. SIMPSON†
Enrico Fermi Institute, University of Chicago
Received 1977 March 14; accepted 1977 April 21

^{10}Be



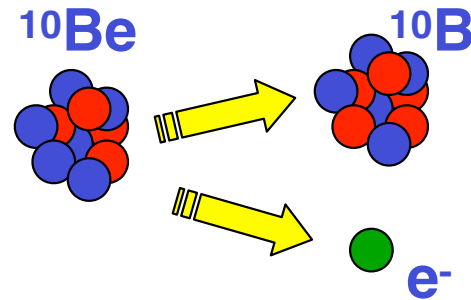
Age of cosmic rays



THE AGE OF THE GALACTIC COSMIC RAYS DERIVED
FROM THE ABUNDANCE OF $^{10}\text{Be}^*$

M. GARCIA-MUNOZ, G. M. MASON, AND J. A. SIMPSON†
Enrico Fermi Institute, University of Chicago
Received 1977 March 14; accepted 1977 April 21

Age of cosmic rays



THE AGE OF THE GALACTIC COSMIC RAYS DERIVED FROM THE ABUNDANCE OF $^{10}\text{Be}^*$

M. GARCIA-MUNOZ, G. M. MASON, AND J. A. SIMPSON†

Enrico Fermi Institute, University of Chicago

Received 1977 March 14; accepted 1977 April 21

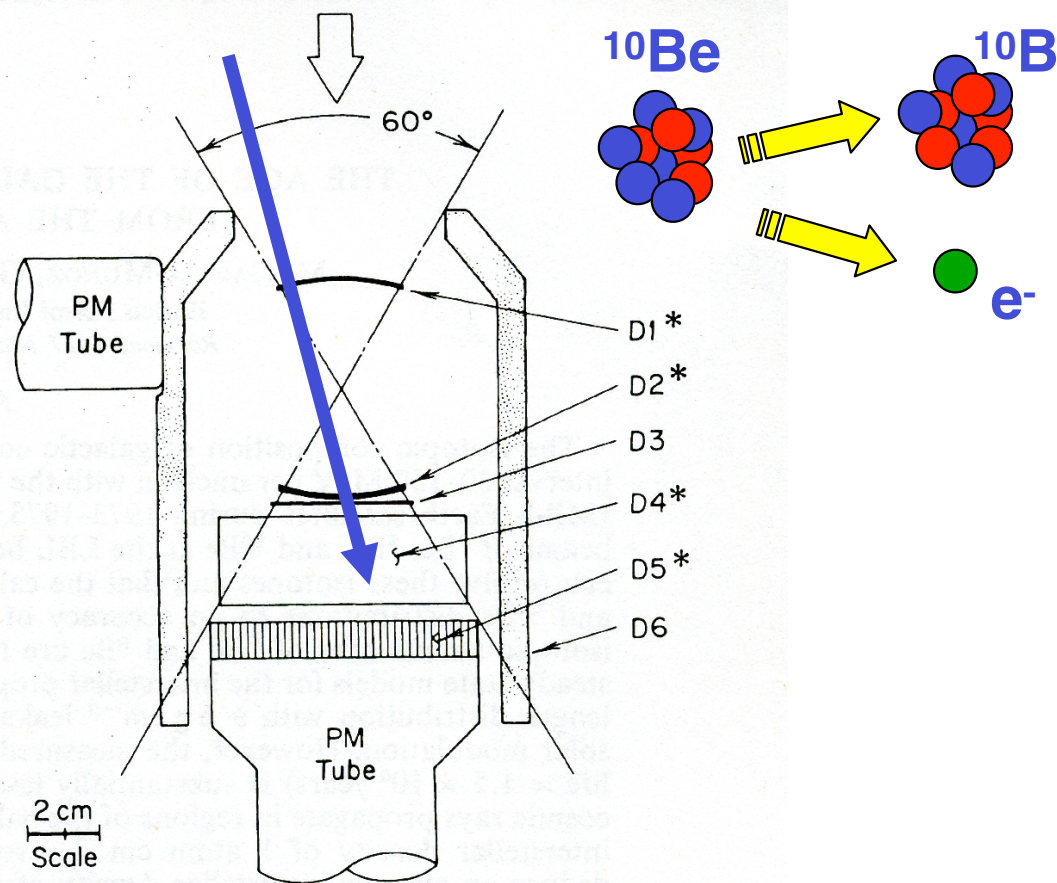


FIG. 1.—Cross section of the IMP-7 and IMP-8 telescopes. D1, D2, and D3 are lithium-drifted silicon detectors of thickness 750, 1450, and 800 μm , respectively. D4 is an 11.5 g cm^{-2} thick CsI (T1) scintillator viewed by four photodiodes. D5 is a sapphire scintillator/Cerenkov radiator of thickness 3.98 g cm^{-2} , and D6 is a plastic scintillation guard counter viewed by a photomultiplier tube. Asterisks denote detectors whose output is pulse-height analyzed.

Age of cosmic rays

THE AGE OF THE GALACTIC COSMIC RAYS DERIVED FROM THE ABUNDANCE OF $^{10}\text{Be}^*$

M. GARCIA-MUNOZ, G. M. MASON, AND J. A. SIMPSON†
 Enrico Fermi Institute, University of Chicago
 Received 1977 March 14; accepted 1977 April 21

$$\tau = 17 \cdot 10^6 \text{ a}$$

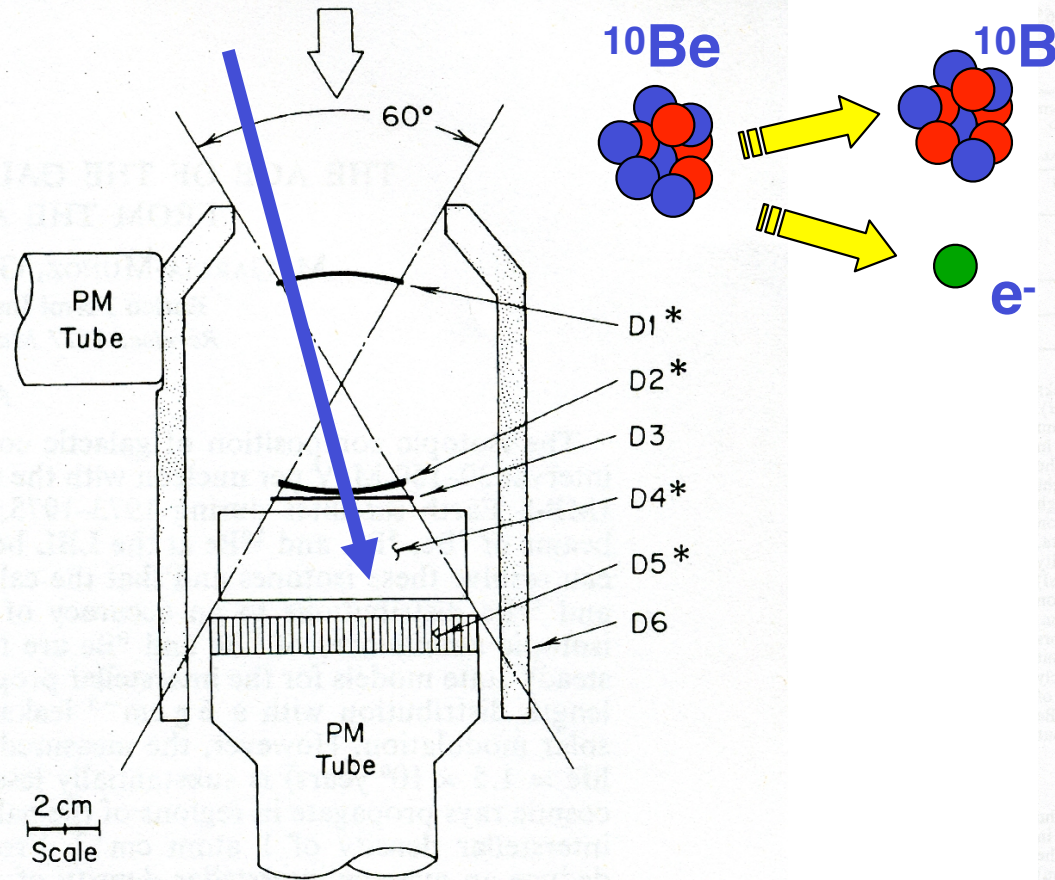


FIG. 1.—Cross section of the IMP-7 and IMP-8 telescopes. D1, D2, and D3 are lithium-drifted silicon detectors of thickness 750, 1450, and 800 μm , respectively. D4 is an 11.5 g cm^{-2} thick CsI (T1) scintillator viewed by four photodiodes. D5 is a sapphire scintillator/Cerenkov radiator of thickness 3.98 g cm^{-2} , and D6 is a plastic scintillation guard counter viewed by a photomultiplier tube. Asterisks denote detectors whose output is pulse-height analyzed.

calibration

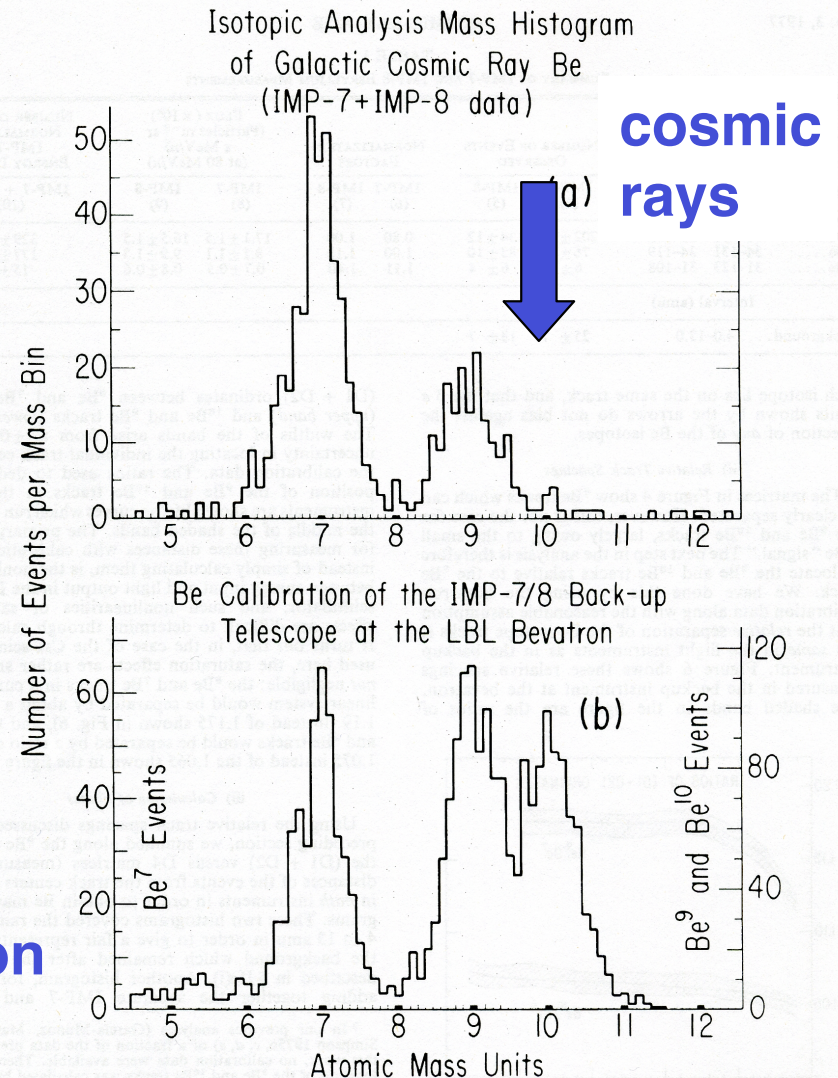


FIG. 7.—(a) Mass histogram of beryllium data from IMP-7 and IMP-8 summed together. (b) Corresponding mass histogram obtained with the backup instrument at the bevatron calibration.

Path length of cosmic rays

Composition of Cosmic-Ray Nuclei at High Energies*

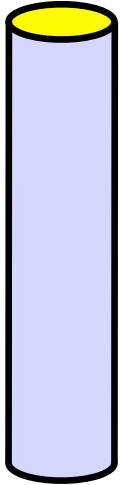
Einar Juliusson, Peter Meyer, and Dietrich Müller

Enrico Fermi Institute and Department of Physics, University of Chicago, Chicago, Illinois 60637

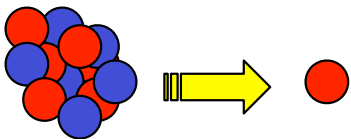
(Received 26 May 1972)

We have measured the charge composition of cosmic-ray nuclei from Li to Fe with energies up to about 100 GeV/nucleon. A balloon-borne counter telescope with gas Cherenkov counters for energy determination was used for this experiment. Our first results show that, in contrast to low-energy observations, the relative abundances change as a function of energy. We find that the ratio of the galactic secondary nuclei to primary-source nuclei decreases at energies above about 30 GeV/nucleon.

g/cm²



spallation



Path length of cosmic rays

Composition of Cosmic-Ray Nuclei at High Energies*

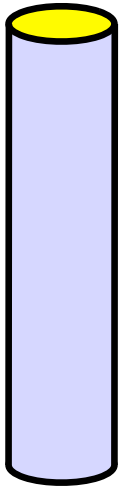
Einar Juliusson, Peter Meyer, and Dietrich Müller

Enrico Fermi Institute and Department of Physics, University of Chicago, Chicago, Illinois 60637

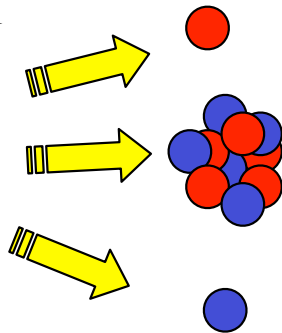
(Received 26 May 1972)

We have measured the charge composition of cosmic-ray nuclei from Li to Fe with energies up to about 100 GeV/nucleon. A balloon-borne counter telescope with gas Cherenkov counters for energy determination was used for this experiment. Our first results show that, in contrast to low-energy observations, the relative abundances change as a function of energy. We find that the ratio of the galactic secondary nuclei to primary-source nuclei decreases at energies above about 30 GeV/nucleon.

g/cm²



spallation



Path length of cosmic rays

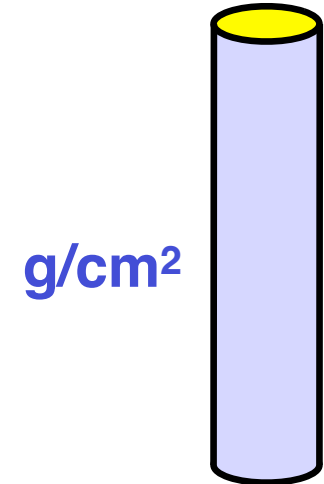
Composition of Cosmic-Ray Nuclei at High Energies*

Einar Juliusson, Peter Meyer, and Dietrich Müller

Enrico Fermi Institute and Department of Physics, University of Chicago, Chicago, Illinois 60637

(Received 26 May 1972)

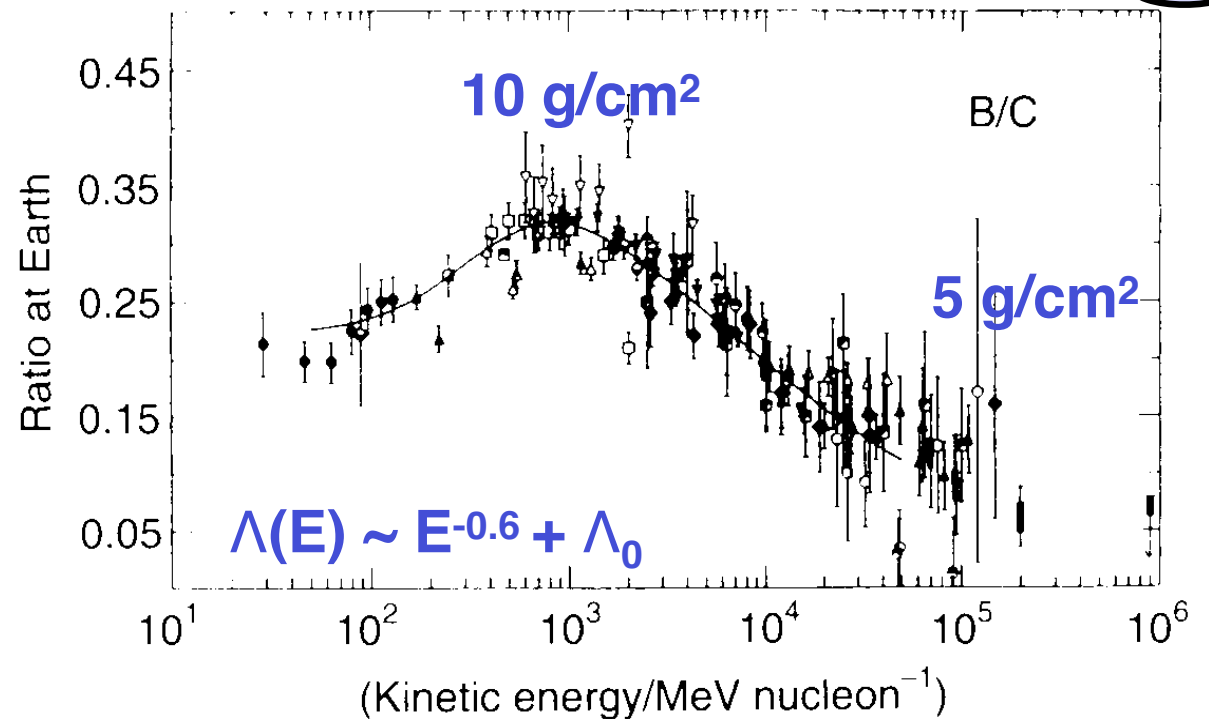
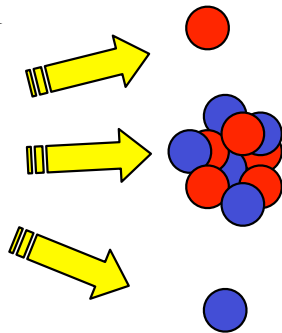
We have measured the charge composition of cosmic-ray nuclei from Li to Fe with energies up to about 100 GeV/nucleon. A balloon-borne counter telescope with gas Cherenkov counters for energy determination was used for this experiment. Our first results show that, in contrast to low-energy observations, the relative abundances change as a function of energy. We find that the ratio of the galactic secondary nuclei to primary-source nuclei decreases at energies above about 30 GeV/nucleon.



g/cm²

B/C-ratio

spallation

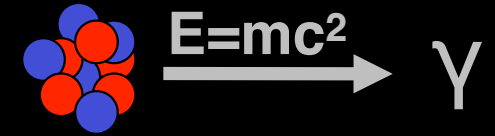


(a)

Origin of Cosmic Rays?

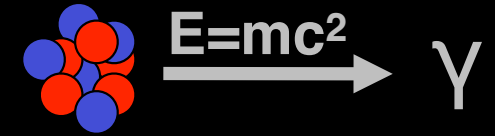
Origin of Cosmic Rays?

1927 R.A. Millikan: „death cries of atoms“



Origin of Cosmic Rays?

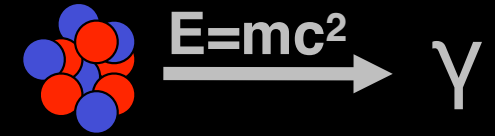
1927 R.A. Millikan: „death cries of atoms“



1933 Regener: E density in CRs \sim E density of B field in Galaxy

Origin of Cosmic Rays?

1927 R.A. Millikan: „death cries of atoms“



1933 Regener: E density in CRs \sim E density of B field in Galaxy

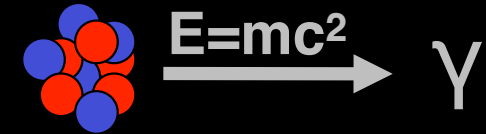
1934 Supernovae



Walter Baade Fritz Zwicky

Origin of Cosmic Rays?

1927 R.A. Millikan: „death cries of atoms“



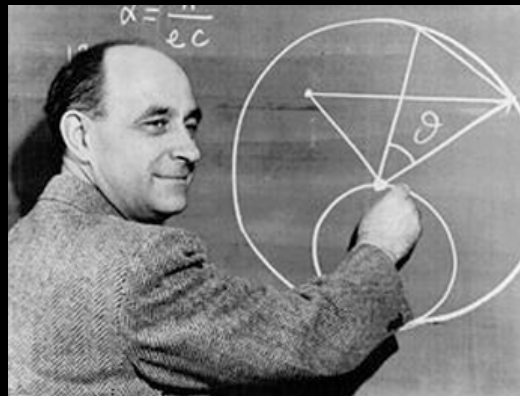
1933 Regener: E density in CRs \sim E density of B field in Galaxy

1934 Supernovae



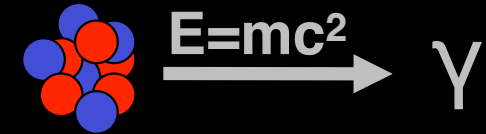
Walter Baade Fritz Zwicky

1949 E. Fermi: acceleration at magnetic clouds



Origin of Cosmic Rays?

1927 R.A. Millikan: „death cries of atoms“



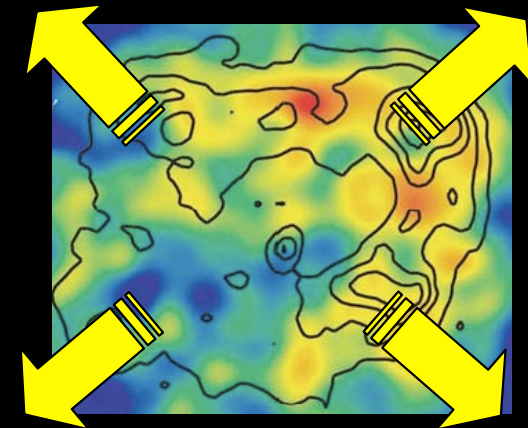
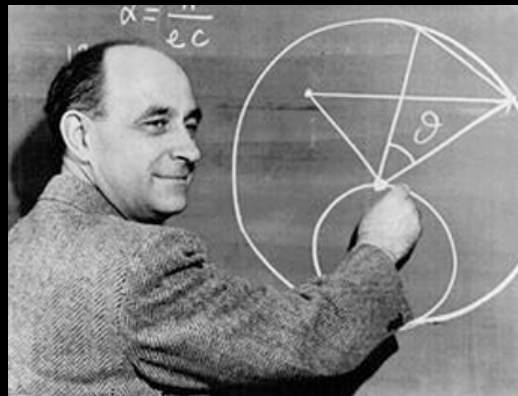
1933 Regener: E density in CRs \sim E density of B field in Galaxy

1934 Supernovae



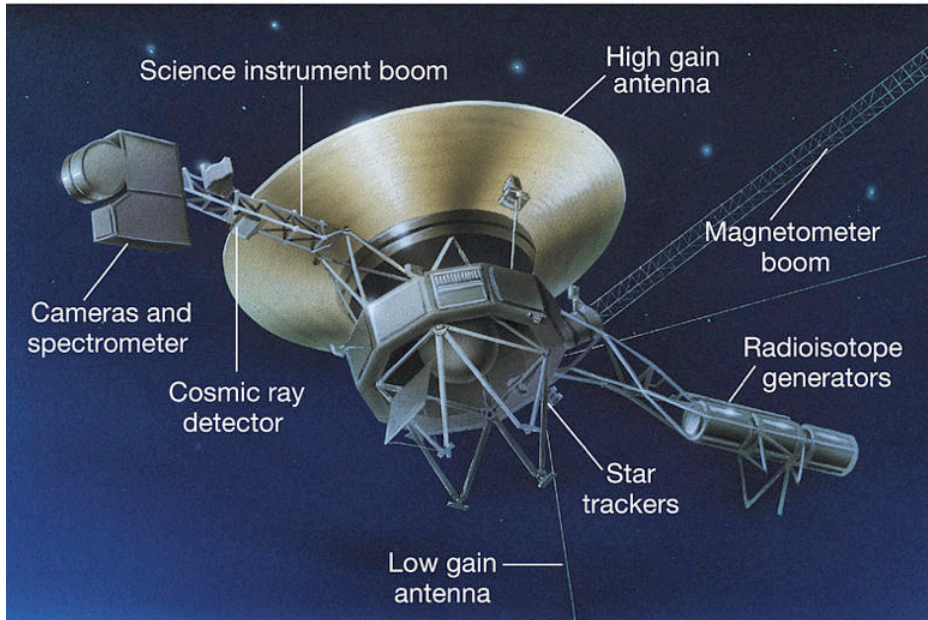
Walter Baade Fritz Zwicky

1949 E. Fermi: acceleration at magnetic clouds



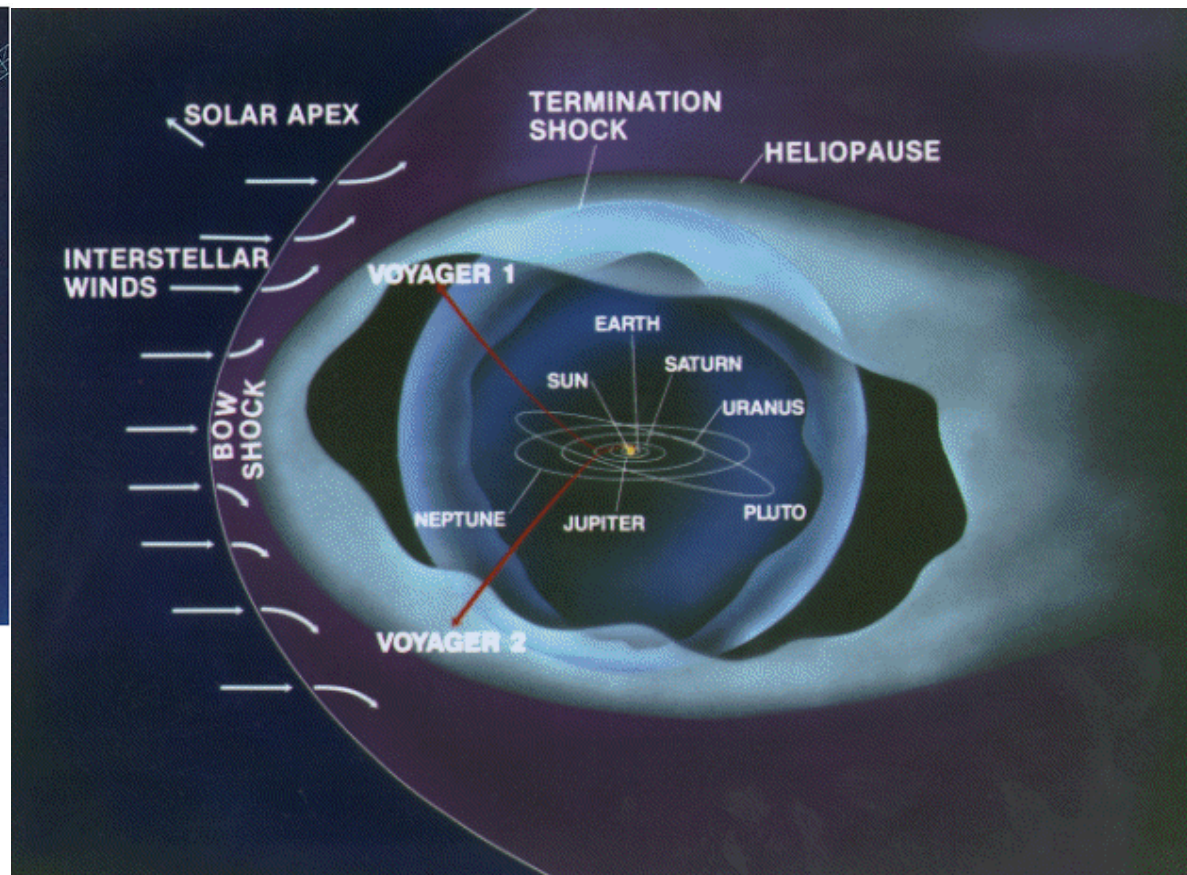
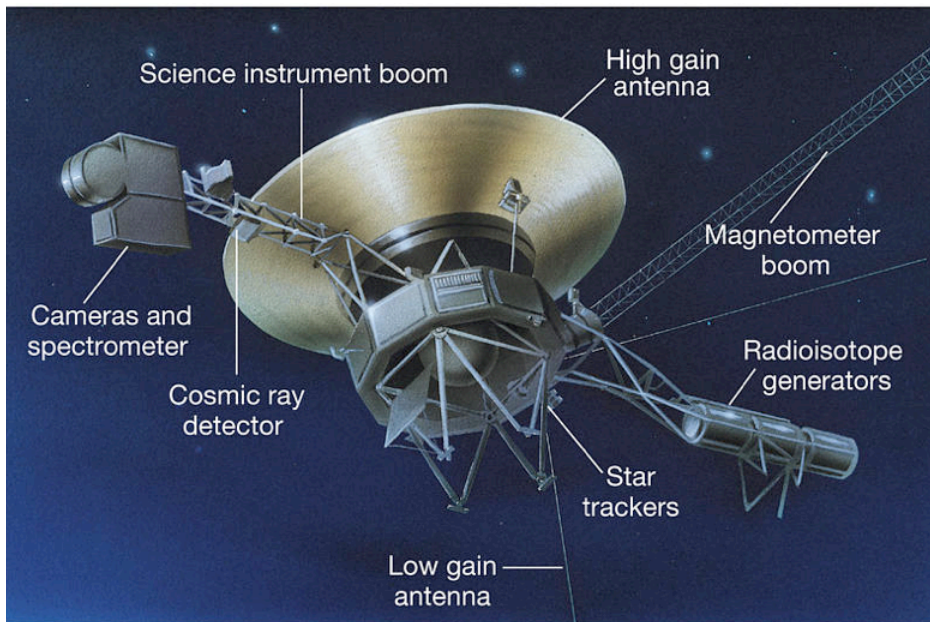
1978 R.D. Blanford, J.P. Ostriker: acceleration at strong shock front
(1st order Fermi acceleration)

Beyond the boundaries of our Solar System



Voyager 2: 20 August 1977
Voyager 1: 5 September 1977
Kenedy Space Center

Beyond the boundaries of our Solar System



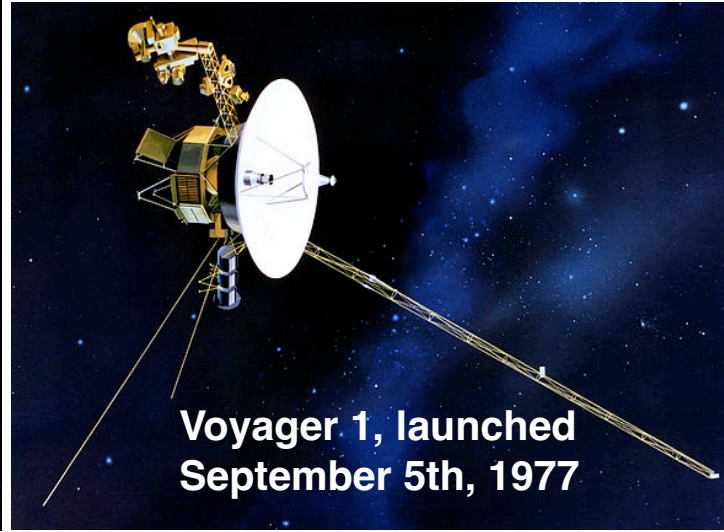
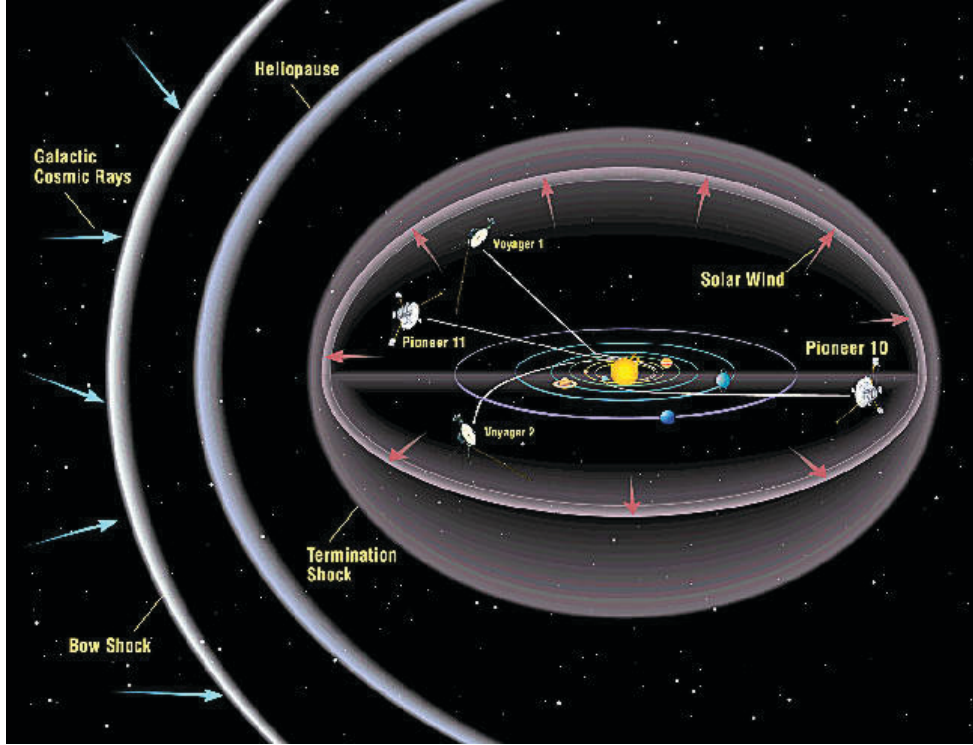
passage through termination shock ended
Voyager 1: 94 AU, December 2004
Voyager 2: 84 AU, August 2007

July 2024: Voyager 1: 164 AU from Sun
Voyager 2: 137 AU from Sun

$$\Delta T = c d \approx 22 \text{ h}$$

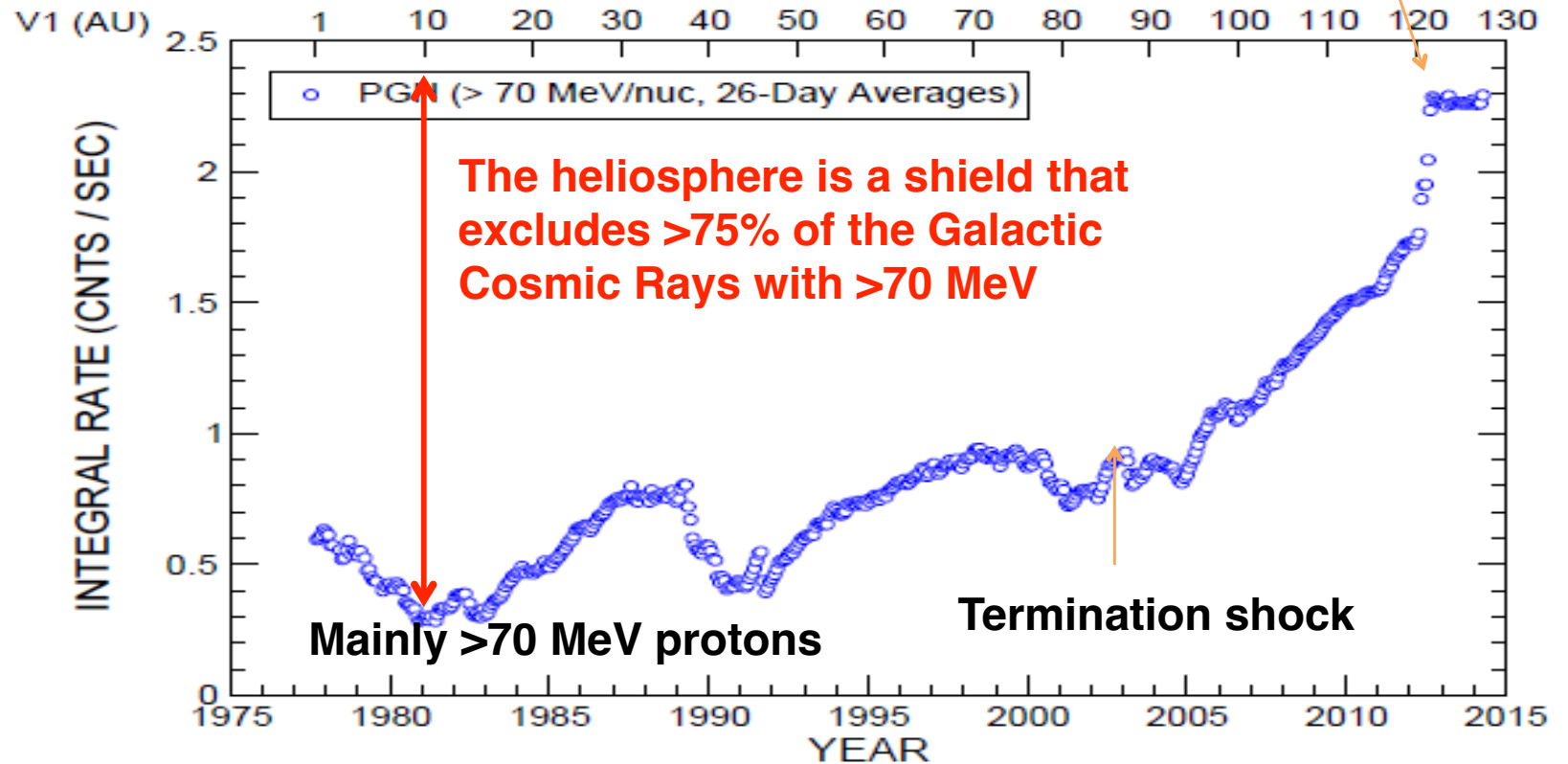
Voyager 2: 20 August 1977
Voyager 1: 5 September 1977
Kennedy Space Center

Galactic Cosmic Rays and the Heliosphere



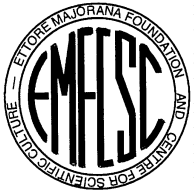
Voyager 1, launched September 5th, 1977

August 25th, 2012
Interstellar Space



Historical introduction

Basic properties of Cosmic Rays



«ETTORE MAJORANA» FOUNDATION AND CENTRE FOR SCIENTIFIC CULTURE

INTERNATIONAL SCHOOL OF COSMIC-RAY ASTROPHYSICS
«MAURICE M. SHAPIRO»

23rd Course: “Multi-Messenger Astroparticle Physics” 20 – 28 July 2024

PRESIDENT AND DIRECTOR OF THE CENTRE: PROFESSOR A. ZICHICHI

DIRECTORS OF THE COURSE: PROFESSORS J.R. HÖRANDEL, T. STANEV, R. SPARVOLI - J.P. WEFEL (director emeritus)



Nikhef



Jörg R. Hörandel

RU Nijmegen, Nikhef, VU Brussel

<http://particle.astro.ru.nl>